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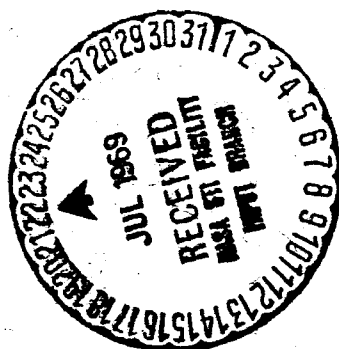
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# MULTI-EXPERIMENT DETECTION OF THE PLASMAPAUSE FROM EOGO SATELLITES AND ANTARCTIC GROUND STATIONS

D. L. CARPENTER  
C. G. PARK  
H. A. TAYLOR, Jr.  
H. C. BRINTON



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FROM EOGO SATELLITES AND ANTARCTIC GROUND STATIONS

D. L. Carpenter and C. G. Park

Radioscience Laboratory, Stanford University

Stanford, California, 94305

and

H. A. Taylor, Jr. and H. C. Brinton

Laboratory for Atmospheric and Biological Sciences

Goddard Space Flight Center

Greenbelt, Maryland, 20771

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ABSTRACT

Three independent methods of detecting the plasmopause have been compared:

- 1) the GSFC ion mass spectrometers on OGO 1 and OGO 3;
- 2) Stanford University/Stanford Research Institute broadband vlf receivers (0.3 - 12.5 kHz) on OGO 1 and OGO 3;
- 3) Stanford University broadband vlf receivers at ground stations Eights and Byrd, Antarctica near the 90°W meridian.

In the satellite vlf data the plasmopause crossings are identified by abrupt changes in observed whistler and vlf noise activity, and by noise bands of limited duration.

In two cases of simultaneous vlf and ion data from the same EOGO satellite, plasmopause crossings were detected by both experiments within less than  $0.1R_E$  in L value. In eight cases of OGO 1 ion data and simultaneous ground whistler data spaced from 1 to 12 hours from OGO 1 in local time, good agreement was found between the measured plasmopause positions. The comparisons provide new verification of the essentially world-wide extent of the plasmopause, and also verify previous indications that the radius of the plasmopause is frequently about constant over large local-time sectors in the range 0-18 LT.

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# MULTI-EXPERIMENT DETECTION OF THE PLASMAPAUSE FROM EOGO SATELLITES AND ANTARCTIC GROUND STATIONS

## INTRODUCTION

Early emphasis in plasmopause studies has been upon plasmopause crossings, since detection of the boundary is often simpler than a description of the regions on either side. Many boundary crossings have been identified in satellite data, including results from ion mass spectrometers on OGO 1 and 3 [Taylor et al., 1965; Taylor et al., 1968; Brinton et al., 1968], Faraday cups on IMP 2 [Binsack, 1967] and on OGO 1 and 3 [Vasyliunas, 1968], and traps on Electron 2 and 4 [Bezrukikh, 1968]. The reported plasmopause position measurements agree statistically both among themselves and with ground-based whistler measurements of the plasmopause radius [Carpenter, 1966; 1967].

In spite of the extent of this work, the plasmopause remains a phenomenon of great mystery both in matters of detail and of large scale structure. The vast size of the plasmasphere and the apparent abruptness of plasma density variations at the boundary suggest that further work include correlation studies, so as not only to provide new descriptive information, but also to encourage development of new detection schemes and provide means of verifying the interpretation previously placed on data from individual experiments.

A recent correlation study showed close agreement between vlf data on plasmopause crossings from the Alouette 1 and 2 satellites and ground-based whistler measurements of the plasmopause radius [Carpenter et al., 1968]. On the satellites the plasmopause was detected as an abrupt termination or sudden decrease in the rate of whistlers propagating from the conjugate hemisphere. Accompanying this whistler cutoff there was frequently a sudden breakup in the vlf noise called the lower hybrid resonance (LHR) noise band.

The purpose of the present paper is to describe another correlation study of plasmopause position, this time involving various combinations of three independent experiments, 1) the GSFC ion mass spectrometers on OGO 1 and OGO 3; 2) the Stanford University/Stanford Research Institute broadband vlf receivers (0.3 - 12.5 kHz) on OGO 1 and OGO 3; 3) Stanford University broadband vlf recordings at ground stations Eights and Byrd, Antarctica near the 90°W meridian. Two basic types of comparison have been made, involving either: 1) ion and vlf data recorded simultaneously at high altitudes on the same OGO satellite, or 2) OGO 1 ion data and simultaneous whistler ground data spaced 1 to 12 hours from OGO in local time.

In the ion mass spectrometer data, the plasmopause is identified by characteristic large decreases in the ion currents of  $H^+$  and  $He^+$ . Examples of plasma density profiles deduced from the ion currents have been shown by Taylor et al. [1965], Taylor et al. [1968], and by Brinton et al. [1968]. In the ground vlf data the plasmopause and its approximate position in space are identified from the dispersion characteristics of multicomponent whistlers in the manner described by Carpenter [1963, 1966] and by Angerami and Carpenter [1966]. In the broadband vlf data recorded on OGO 1 and OGO 3, plasmopause crossings are identified by abrupt changes in observed whistler and vlf noise activity, and by noise bands of limited duration. These effects, although different in matters of detail, appear closely related to some of the low-altitude crossing phenomena detected on Alouette 1 and 2 [Carpenter et al., 1968].

Three purposes of the present note are: 1) to illustrate the vlf method of detecting plasmopause crossings at high altitudes; 2) to show two cases of close agreement between the vlf and ion mass spectrometer methods of detecting the

plasma-pause; and 3) to demonstrate through local-time separation of simultaneous measurements the essentially worldwide longitudinal extent of the plasma-pause.

The data described below were selected on the basis of availability and suitability for comparison. There are seven cases from May - June 1965 involving OGO 1, and one case for July 1966 involving OGO 3. The supporting ground whistler data were obtained from Eights Station (64°S geomagnetic latitude) and Byrd Station (71°S geomagnetic latitude) in Antarctica.

#### SIMULTANEOUS DETECTION OF THE PLASMAPAUSE BY TWO EXPERIMENTS ON THE SAME EOGO SATELLITE

Figures 1 and 2 illustrate simultaneous detection of the plasma-pause from the EOGO satellites by an ion mass spectrometer and broadband vlf receiver. Figure 1 represents the OGO 3 outbound pass of July 5, 1966 and Figure 2 the inbound OGO 1 pass of June 9, 1965. In both figures the satellite orbit is projected on a geomagnetic meridian plane. The solid and dashed portions of the curves represent satellite positions inside and outside the plasmasphere respectively, as determined from the two experiments. Ion concentration as determined from the mass spectrometer (by the GSFC group) is plotted versus time and is to be compared with the vlf spectra recorded simultaneously. In Figure 3 the sections of vlf data marked A through F in Figures 1 and 2 are shown on an expanded time scale.

##### Details of the July 5, 1966 Event (Figure 1)

On the OGO 3 outbound pass of July 5, 1966, a plasma-pause crossing was identified in the ion data at approximately 1250 UT, at which time the concentration



of  $H^+$  was observed to decrease rapidly by about an order of magnitude. At this time the satellite was in the afternoon sector at  $\sim 1640$  LT, a geomagnetic latitude of  $30^\circ N$  and dipole  $L = 4.6$  (see also Figure 5, lower left). The period was one of low to moderate magnetic agitation, with maximum  $K_p$  values in the preceding 12 and 24 hours of 2+ and 4+, respectively. Arrows on the satellite trajectory identify the boundary crossing and indicate a spatial range of uncertainty of about 0.1  $L$  in determining a point location of the plasmopause from the associated ion concentration profile. Corresponding time intervals are indicated by arrows on the vlf and ion records.

The vlf data begin in the top panel at  $\sim 1220$  UT, and continue outbound toward the plasmopause in the bottom panel. Within the plasmasphere there is high whistler activity, and also a well defined noise band in the range  $\sim 400 - 900$  Hz. (The intense signal appearing near 1238:30 and decreasing in frequency with time is interference from the magnetometer aboard the spacecraft.) After  $\sim 1235$  the whistlers exhibit well defined echoes, a phenomenon that may be dependent on the proximity of the boundary (interval A is shown expanded in Figure 3-A). After  $\sim 1245$ , the whistler activity diminishes in intensity, and at 1247:50 UT, within roughly two minutes of the time of plasmopause crossing indicated in the ion data, there appear two band limited noises with center frequencies 3.7 and 4.7 kHz (only the lower one is shown in Figure 3-B so as to avoid AGC effects on the analysis by the magnetometer signal). The onset of the noise bands is quite sudden and can be resolved within 10 seconds (cf. Figure 3-B). At  $\sim 1249$  the ion data indicate the beginning of a plasmopause profile. At about this time the noise bands at 3.7 and 4.7 kHz decay, and at  $\sim 1250$ , faint whistler activity and the 400 - 900 Hz noise cease to be detected. The behavior of the 400 - 900 Hz band

is consistent with reports on magnetic field measurements in the 0 - 1000 Hz range on OGO 1 by Russell et al. [1968].

At the approximate time of the OGO 3 observations summarized in Figure 1, ground based whistler data on the plasmopause radius were obtained at Byrd Station, Antarctica. When OGO 3 crossed the plasmopause at about 1640 LT, Byrd Station was 'viewing' a region of the magnetosphere in the local time sector 0600 - 0700 (see Figure 5, lower left). The radius measured from Byrd data was in the range  $4.6 - 5.6 R_E$ , in good agreement with the crossing value of  $L = 4.6$  found by OGO 3. Certain longitudinal and local-time differences in plasmopause radius are to be expected, and further details of this comparison will be presented in a later paragraph.

#### Details of the June 9, 1965 Event (Figure 2)

The ion and vlf data observed on the OGO 1 inbound pass of June 9, 1965 are shown in Figure 2. The vlf data begin in the top panel at 0938 UT ( $L = 4.5$ ) and continue inbound toward the plasmopause in the bottom panel. The ion data begin at 10:04 UT, with the first detectable concentration of  $H^+$  observed above the experimental limiting sensitivity of 1 ion/cm<sup>3</sup>. The ion spectrometer detected the plasmopause crossing at approximately 1013 UT, at which time the satellite was in the late morning sector at ~1100 LT, geomagnetic latitude 38°S, and dipole  $L = 4.02$  (see also Figure 5, upper middle). The magnetic condition was one of moderate, steady agitation, with  $K_p$  in the range 2-4.

While the data of Figure 1 were obtained from a 3-axis attitude stabilized satellite, OGO 1 was spin stabilized and consequently the raw  $H^+$  ion currents used for Figure 2 were modulated by the variation in angle-of-attack between the spectrometer orifice and the velocity vector. Although this modulation does

not obscure the fundamental evidence of the plasmopause crossing, the modulation has been eliminated in the final data by normalizing the measurements to zero angle-of-attack using inflight calibration of the dependence of sensor efficiency upon spacecraft attitude. Subsequent conversion of ion currents to ion concentrations followed the plasma probe technique described in previous papers by Taylor et al. [1965] and Brinton et al. [1968].

In Figure 2, connected arrows again indicate the spatial range of the plasmopause, spanning the period 1009 UT ( $L = 4.1$ ) to 1014 UT ( $L = 4.0$ ). During this interval and earlier in the pass, the  $H^+$  currents occasionally fell below the detectable level, so that concentrations cannot be determined. In comparing the data of Figures 1 and 2, note that differences in orbit result in differences in the time rate of change of  $L$  near the plasmopause.

In contrast to the case of Figure 1, the vlf records of Figure 2 exhibit complex noise activity outside the plasmopause. The principal crossing effect is a transition from an outer regime of intense noise activity in the range 5 - 8 kHz and rising noise tones near 2 kHz to an inner regime of whistler activity and diffuse noise bands in the range  $\sim 0.6 - 3$  kHz. (The narrow vertical lines on the spectrogram represent telemetry fading at the spacecraft spin period of 12 seconds.)

Reviewing details of the vlf activity, the rising chorus tones near 2 kHz appear to increase in intensity with time, being particularly well defined after  $\sim 0951$  (upper vlf record). As the plasmopause is approached, the rising tones begin to disappear during intervals when the log compressor receiver is apparently captured by intense noise bursts near 5 - 6 kHz. Figure 3-C shows

details of the 1 - 2 kHz rising tones and also complex activity near 5 kHz at the beginning of an intense noise event.

On the upper record of Figure 2 noise activity near 5 - 6 kHz exhibits a succession of short bursts with irregular structure but relatively smoothly increasing overall frequency limits. As the plasmopause is approached, the bursts become more frequent, until at ~1004, when the first detectable ionization is encountered, there begins a prolonged band-like noise containing many closely spaced elements. At ~1009:30 another long noise event begins, this time exhibiting relatively little structure (cf. Figure 3-C). The noise band increases in intensity, and then at 1012:09 UT begins a series of irregular variations in bandwidth and center frequency, finally fading away near 1015 UT. Figure 3-D shows the band at a high intensity level near 1014 UT, only a minute prior to its disappearance (shown in Figure 3-E).

A noise band of center frequency ~2 kHz appears near 1014 UT and continues after the disappearance of the higher frequency noise (cf. Figures 3-D, 3-E). Within a few minutes the upper part of the new band fades and the noise becomes a steady narrow band centered at ~1 kHz (Figure 3-F).

The last detectable rising-tone emission is at 1014:55, and the first tentative evidence of a whistler is at 1015:03 UT (cf. Figure 3-E), at which point the ion concentration has significantly increased. After this time whistlers become more numerous and better defined; an example is shown in Figure 3-F at 1022:03.

Some of the details of the vlf noise and the apparent patchiness in the ion concentration observed outside the plasmopause on this OGO 1 pass may eventually be explained by the grazing nature of the orbit and the associated

encounters with boundary fine structure. Figure 5, upper middle, shows the orbit near plasmapause crossing in coordinates of L versus local time, and indicates the rapid longitudinal variation experienced by OGO 1 as the magnetic shell parameter changed only slowly.

Ground whistler data for the June 9 OGO 1 pass were obtained at Eights Station, Antarctica. The plasmapause radius as determined from Eights data for 1230-1300 UT (about 0730-0830 LT) was in the range  $L = 3.8 - 4.1$ , and again there is good agreement with the OGO 1 value ( $L \sim 4$ , ~1100 LT). The comparison is discussed further in a later section.

#### COMPARISONS OF GROUND AND SATELLITE MEASUREMENTS OF THE PLASMAPAUSE POSITION

A comparison of OGO 1 ion data on plasmapause crossings and ground data on the plasmapause radius is illustrated in Figure 4. The satellite orbit is presented in coordinates of L value versus local time and plasmapause crossings are again shown by the transition from solid to dashed curves.

The combination of a fixed ground station and a satellite in nonsynchronous orbit affords the opportunity for: 1) simultaneous measurements at various local time spacings, and 2) nearly simultaneous measurements in which the ground data are plotted as a function of local time before and after the plasmapause crossing by the satellite. In the case illustrated in Figure 4, the diurnal variation of the plasmapause position deduced from whistlers recorded at Eights and Byrd Stations on May 8, 1965, is plotted with two OGO 1 crossings of the plasmapause on the same day.

Ground whistler data on the equatorial radius of the plasmapause are represented by triangles. The value of equatorial radius was calculated assuming

propagation in a centered dipole geometry. A triangle with apex pointing toward the earth represents a whistler component that propagated in the tenuous region outside and near the plasmapause, while a triangle with apex pointing away from the earth is interpreted as having propagated inside and near the boundary. The filled triangles represent cases in which the plasmapause is known to be within about 0.2 L in the direction of the apex. The open triangles represent cases of less certainty in which the plasmapause is interpreted as lying within about 0.5 L in the direction of the apex. (The use of whistler data in which propagation is detected on only one side of the plasmapause has been discussed in a previous paper [Carpenter, 1966]. When suitable caution is observed, such as by using periods preceded and followed by more complete data, and by restriction to periods of high whistler activity, the method has been found to be very useful in plasmapause studies.)

In considering the details of Figure 4 the ground measurements may be thought of as made by two closely spaced 'synchronous' satellites, corotating with the earth near the dawn meridian. Meanwhile, OGO 1 arrived at the plasmapause boundary in the forenoon sector and entered the plasmasphere at 0912 UT. At 1146 UT, about 2 ½ hours later, it emerged from the plasmasphere near the midnight meridian. During this interval the ground stations moved between the local times indicated by the arcs of circles, so that the whistler measurements near these arcs represent the plasmapause measurements most nearly time-coincident with those of OGO 1. The series of triangles extending back across the night side shows in local time the measurements made at Eights and Byrd during 9 hours prior to the arrival of OGO 1 at the plasmapause. The series of data points across the day side represents similar observations following the emersion of OGO 1 from the plasmasphere on the night side.

The general appearance of Figure 4 is one of substantial agreement between the ground and satellite data. The expected worldwide extent of the plasmapause is verified by the simultaneous detection of the boundary, first with an  $\sim 04 - 09$  LT separation (ground and satellite, respectively) and then with an  $\sim 06 - 23$  LT comparison. The observed values of plasmapause radius in the range  $4.5 - 5.5 R_E$  are consistent with expectations for the moderate level of geomagnetic agitation on May 8, 1965 ( $K_p = 1-3$ ) [Carpenter, 1966; Taylor et al., 1965; Binsack, 1967]. The asymmetry shown in the ground data (triangles), with a dawn minimum and dusk maximum in the plasmapause radius, is also in agreement with expectations from previous work [Carpenter, 1966].

The difference between the OGO 1 measurement of the plasmapause position at 0912 UT and the value obtained from whistlers when the ground station reached the same local-time region 5 hours later is probably explicable in terms of magnetospheric substorm activity. Substorm activity began to increase at about 0440 UT on May 8, 1965, in time to effect a decrease in the plasmapause radius near Byrd and Eights, but apparently not in time to affect the position of the boundary that was eventually crossed at 0912 UT by OGO 1. The effect on Eights' data is indicated by the inward shift of data points near 2330 LT.

#### Additional Case Studies

A group of six ground-satellite comparisons of the plasmapause radius is presented in Figure 5. The symbolism and coordinates are essentially the same as in Figure 4, but only single satellite crossings of the boundary are shown. An arrow along the periphery of each polar plot shows the approximate location of the ground station at the time of the plasmapause crossing by the satellite. The lower left and upper middle plots provide additional documentation

on the cases illustrated in Figure 1 and 2, respectively. In the case of June 1, 1965, upper left, the boundary identification is tentative, since the ion data coverage began just beyond  $L = 5$ , in time to observe only part of a plasmopause-like profile.

The six cases of Figure 5 represent relatively large variations in magnetic activity, local time of plasmopause crossing by the satellite, and local-time separation of the ground and satellite measurements. Moving from left to right at the top, there is a progression from quiet magnetic conditions (June 1, 1965) to moderate disturbance (June 9, 1965) to severe disturbance (June 17, 1965), and a corresponding shrinkage is evident in the plasmopause radius. In all three cases there is good agreement between the ground and satellite data.

Moving from left to right along the top and then the bottom row of Figure 5, there is a progression in local time of satellite crossing from morning to pre-midnight, and also variations in satellite-ground separation from 1 to 10 hours. Good agreement is obtained for all the separations, indicating the essential persistence of the plasmopause over large intervals of time and space.

The number of case studies presented here was limited by several factors, including seasonal variations in ground whistler data, low-power cycles on OGO 1, and the need to compare special purpose OGO vlf data with digital data. Because of these restrictions, the results should be considered as indicative only of a class of cases, with the size of the class still unknown.

## CONCLUSIONS

Within the limits of the cases illustrated thus far, it is concluded that:

- 1) On two occasions of observations on an EOGO satellite, plasmopause crossings detected by vlf techniques and by an rf mass spectrometer were coincident within 0.1  $L$ .



2) Broadband vlf data from satellites in eccentric, equatorial orbits may on some occasions be useful in identifying the plasmapause. Significant changes in vlf activity have been observed to take place within intervals of ten seconds or less.

3) New verification of the worldwide extent and persistence of the plasmapause has been obtained from eight cases of simultaneous satellite-ground observations. The measurements involved local-time separations ranging from 1 to 10 hours, varying levels of magnetic agitation, and varying local times of boundary-crossing by the satellite.

4) The ground-satellite measurements strongly support previous suggestions, based on measurements on a single meridian [Carpenter, 1966], that the radius of the plasmapause is frequently about constant over large local time sectors in the range 0 - 18 LT.

The results presented above may be deceptive in their simplicity. Data not presented here indicate that the plasmapause is extremely complex, with regions of irregular behavior, periods of rapid expansion or compression, and variations in details of the plasma profile at the boundary. Further generations of correlation studies are needed in order to obtain a proper description of these effects.

#### ACKNOWLEDGMENTS

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## ILLUSTRATIONS

Figure 1. Comparison of simultaneous vlf and ion mass spectrometer data on a plasmopause crossing during the OGO 3 outbound pass of July 5, 1966. The satellite orbit is shown in a geomagnetic meridian projection. A pair of arrows along the orbit indicates the position of the plasmopause as detected in the ion data. The corresponding time interval is shown by arrows on the ion and vlf records. The vlf data begin in the top panel and continue outbound toward the plasmopause in the bottom panel. Horizontal lines on the records are of instrumental origin. The vlf signal descending in frequency after 1238:30 is interference from the onboard magnetometer.

Figure 2. Comparison of simultaneous vlf and ion mass spectrometer data on a plasmopause crossing during the OGO 1 inbound pass on June 9, 1965. A pair of arrows along the orbit indicate the position of the plasmopause as detected in the ion data. The corresponding time interval is shown by arrows on the ion and vlf records. The vlf data begin in the upper record and continue inbound toward the plasmopause on the bottom panel.

Figure 3. Expanded vlf spectra for intervals marked A and B in Figure 1 and C through F in Figure 2. The frequency scale of panels A and B is 0 - 4 kHz, instead of 0 - 7 kHz as in Figure 1.

Figure 4. Example of comparison between OGO 1 ion data and ground whistler data on the plasmopause position. The OGO 1 orbit of May 8, 1965, is projected in coordinates of L value versus local time. Triangles represent ground whistler information on plasmopause radius, while the transition from dashed to solid curve indicates a plasmopause crossing in the ion data. During the 2 ½ hours of the OGO 1 passage through the plasmasphere, the ground stations Eights and Byrd moved through the arcs of circles indicated near the 05-06 LT meridians. Further details of the comparison and meaning of symbols are given in the text.

Figure 5. Comparisons of OGO 1 and OGO 3 ion data on plasmopause crossings with simultaneous ground whistler measurements of the plasmopause radius and also ground measurements in the preceding and succeeding hours. The transition from a solid to a dashed curve indicates a plasmopause crossing as identified in the ion data. An arrow along the periphery of the plot indicates the approximate region of whistler observations at the time of plasmopause crossing by the satellite.

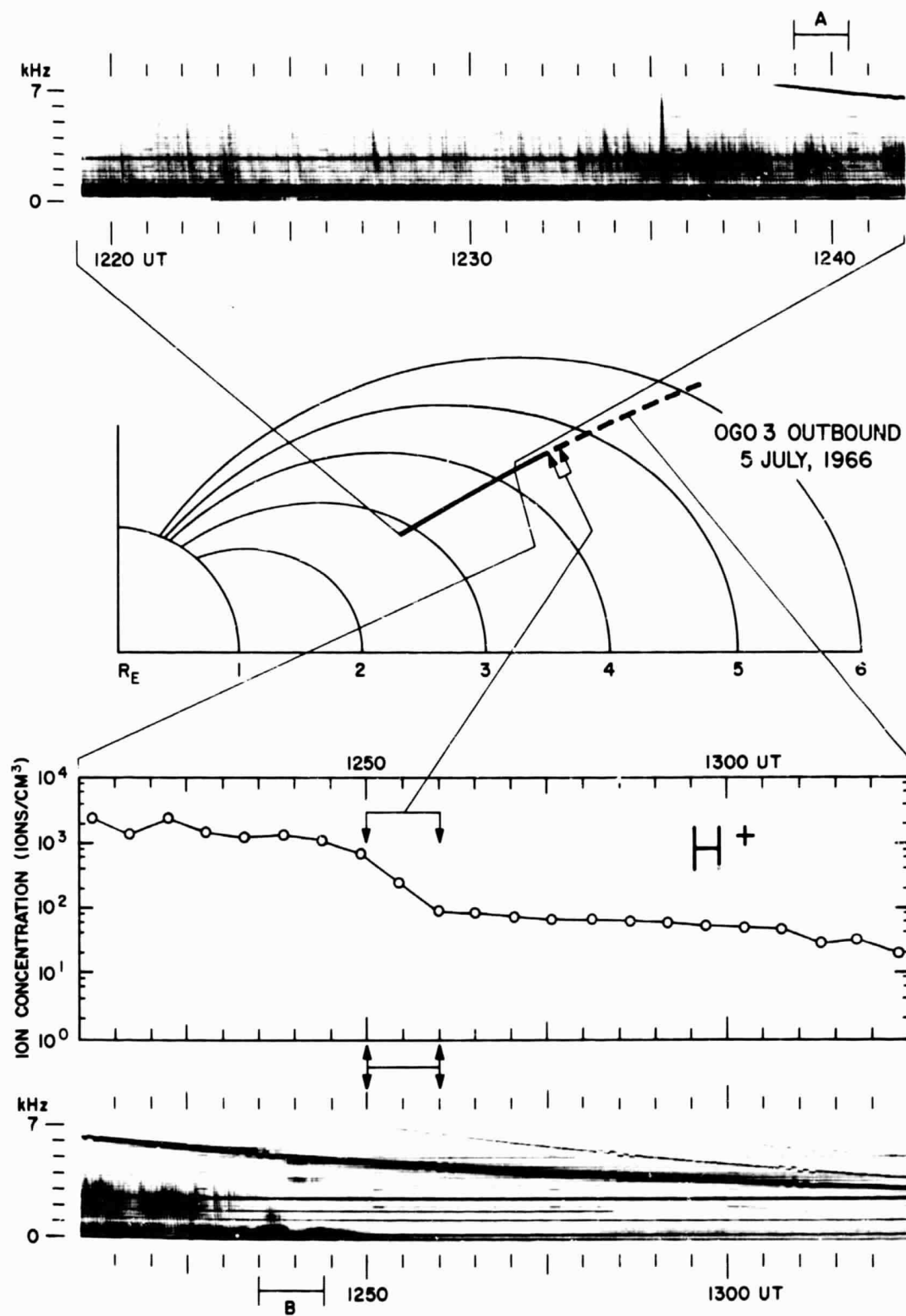


Figure 1

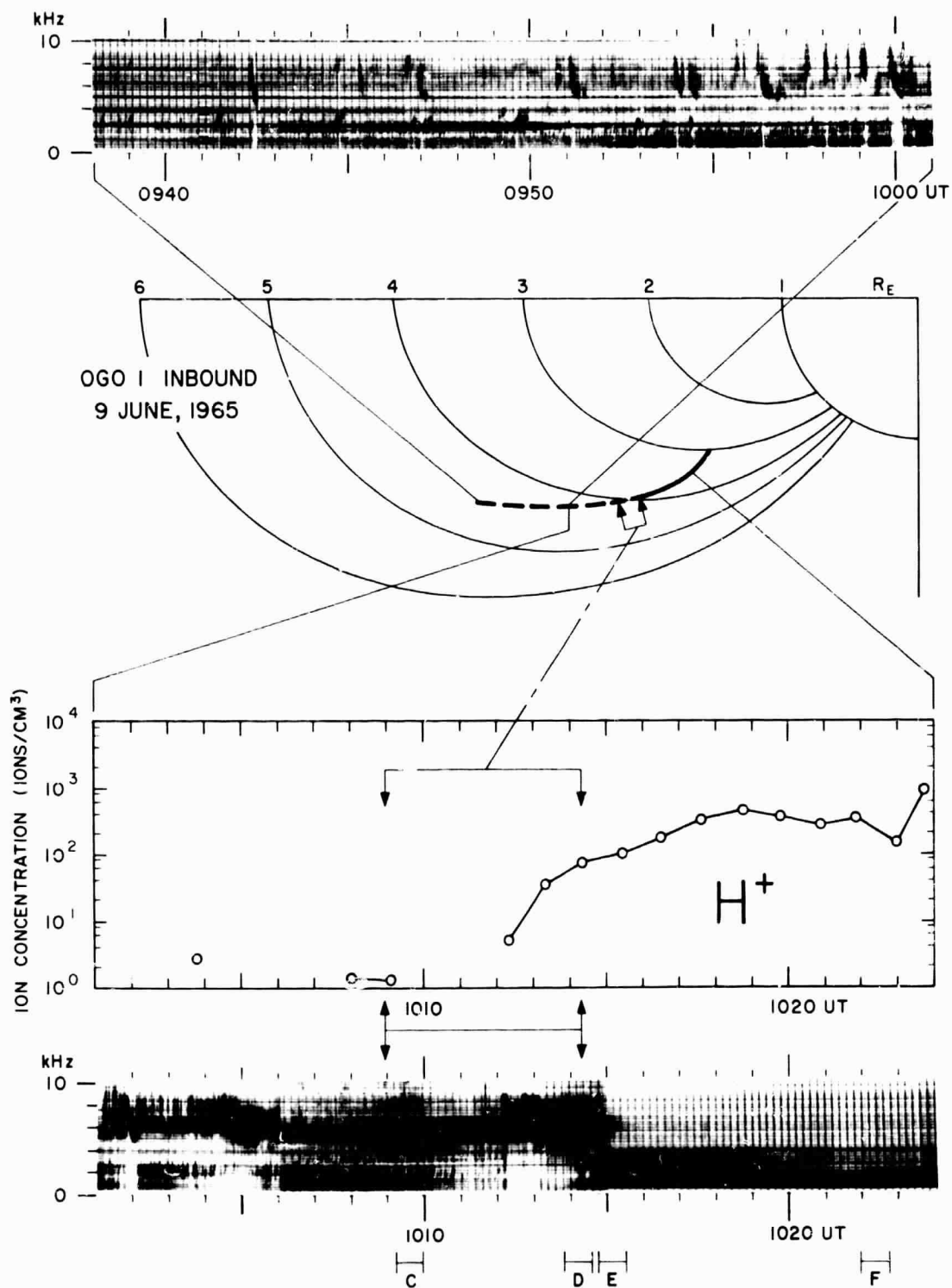


Figure 2

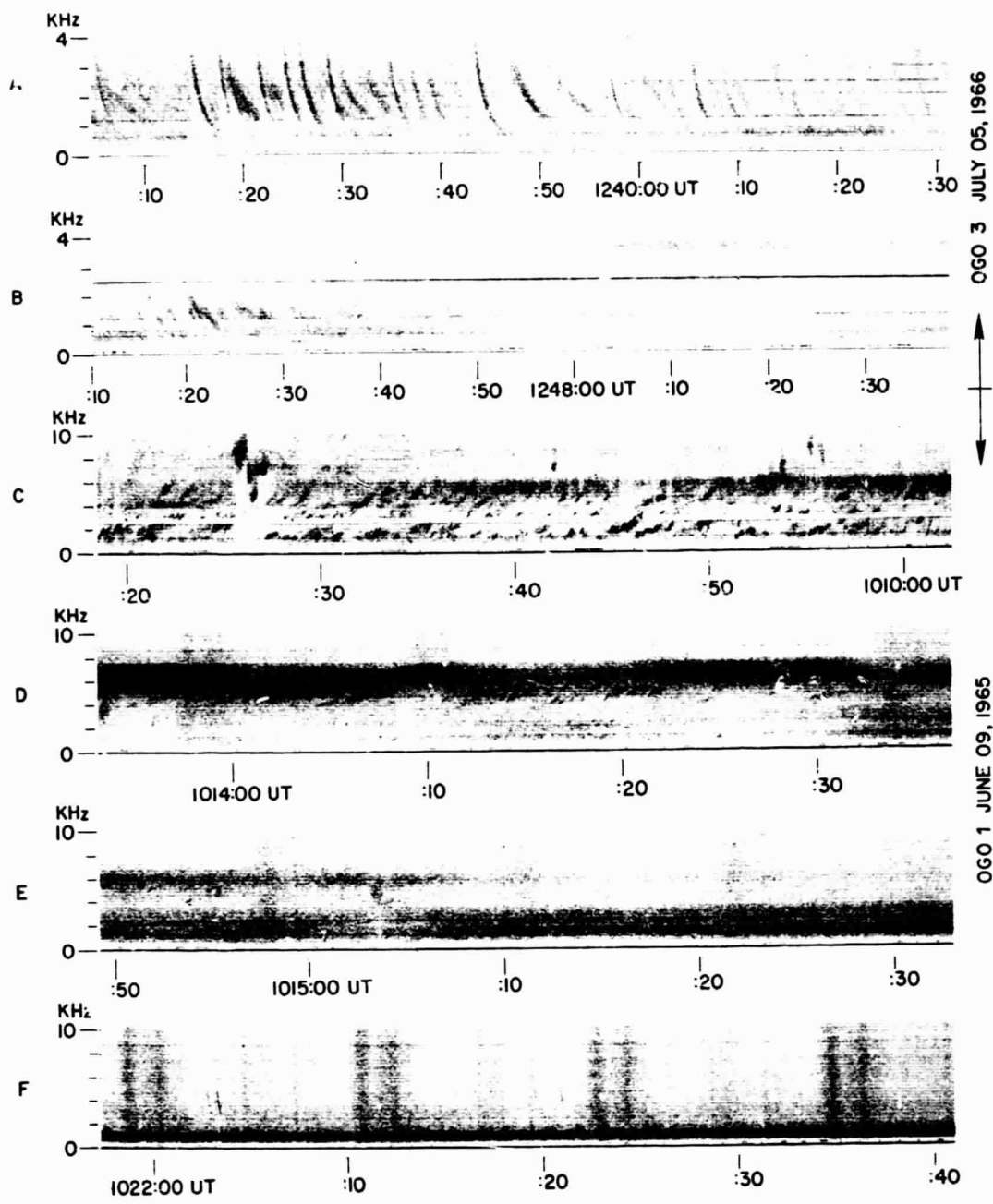


Figure 3



MAY 08, 1965  
OGO 1

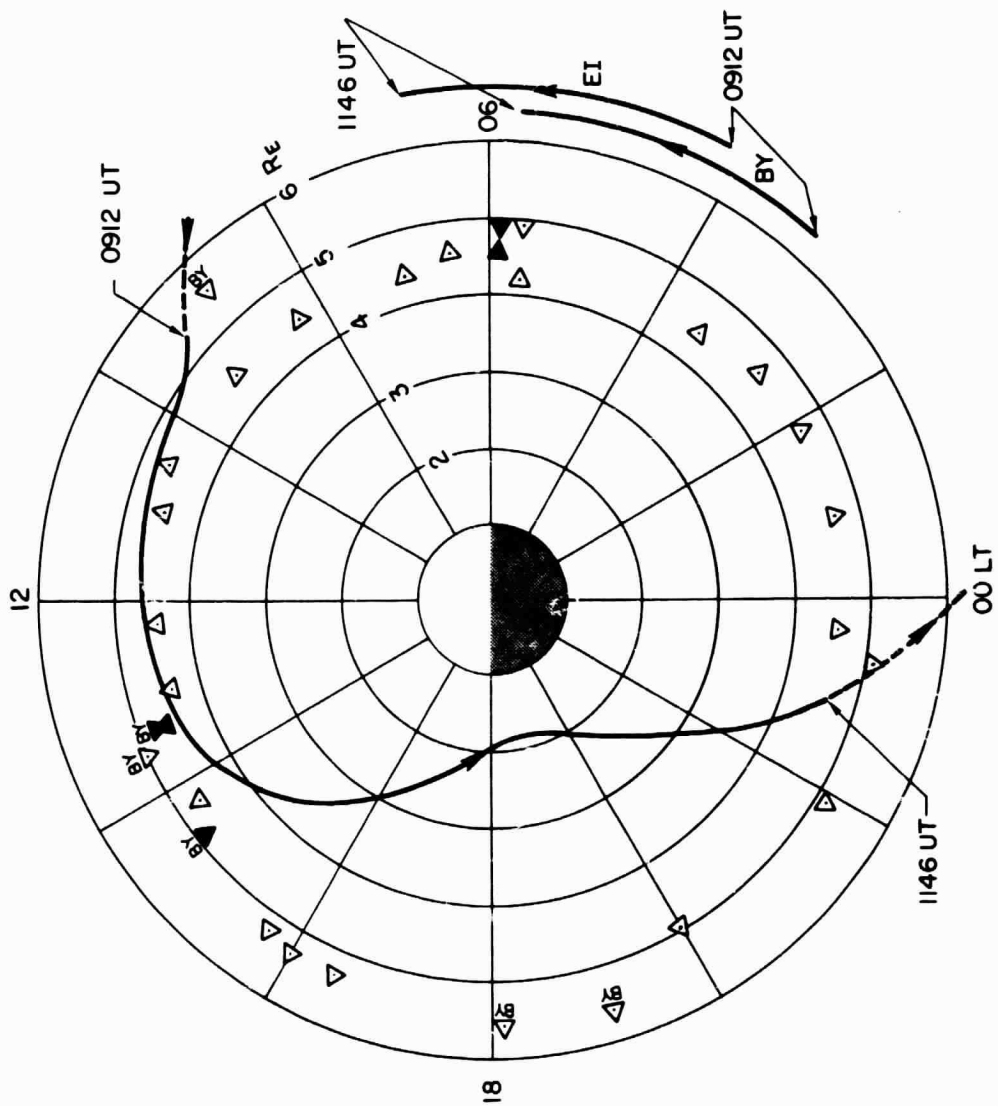


Figure 4

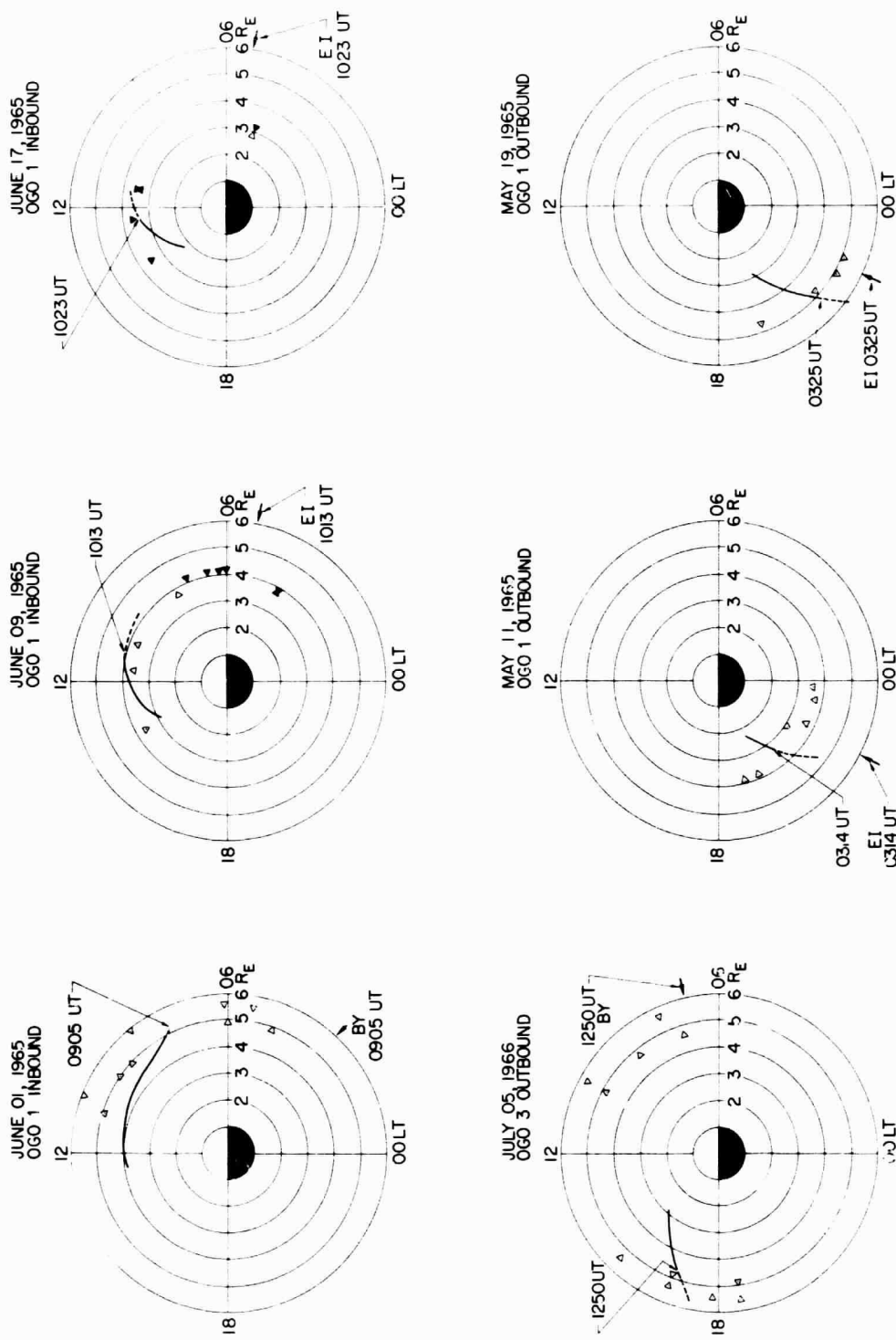


Figure 5