NASA
SPACE VEHICLE
DESIGN CRITERIA
ENVIRONMENT)

THE PLANET JUPITER (1970)

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DECEMBER 1971

NATIONAL AERONAUTICS AND SPACE ADMINISTRATION
NASA experience has indicated a need for uniform criteria for the design of space vehicles. Accordingly, criteria are being developed in the following areas of technology:

- Environment
- Structures
- Guidance and Control
- Chemical Propulsion

Individual components of this work are issued as separate monographs as soon as they are completed. A list of the monographs published in this series can be found on the last page.

These monographs are to be regarded as guides to design and not as NASA requirements, except as may be specified in formal project specifications. It is expected, however, that the monographs will be used to develop requirements for specific projects and be cited as the applicable documents in mission studies, or in contracts for the design and development of space vehicle systems.

This monograph was prepared under the cognizance of the Goddard Space Flight Center with Scott A. Mills and Mason T. Charak as program coordinators. The principal author was Neil Divine of the Jet Propulsion Laboratory. Valuable contributions were also made by A.J. Beck, C.A. Haudenschild, F.D. Palluconi, and R.A. Schiffer of the Jet Propulsion Laboratory. J.W. Warwick of the University of Colorado served as a consultant in the preparation of the fields and particles sections.

Comments concerning the technical content of these monographs will be welcomed by the National Aeronautics and Space Administration, Goddard Space Flight Center, Systems Reliability Directorate, Greenbelt, Maryland 20771.

December 1971
## CONTENTS

**INTRODUCTION** ................................................................. 1

**STATE OF THE ART** .............................................................. 1

2.1 Physical Properties .......................................................... 1
   2.1.1 Mass
   2.1.2 Dimensions
   2.1.3 Mean Density
   2.1.4 Rotational Quantities

2.2 Gravity Field ................................................................. 5
   2.2.1 Inertial Coordinates
   2.2.2 Corotating Coordinates

2.3 Magnetic Field and Magnetosphere ........................................ 7
   2.3.1 Magnetic Field Strength and Configuration
      2.3.1.1 Dipole Moment
      2.3.1.2 Dipole Inclination
      2.3.1.3 Dipole Displacement
   2.3.2 Magnetosphere

2.4 Electric Fields .............................................................. 8

2.5 Electromagnetic Radiation ................................................ 9
   2.5.1 Light
      2.5.1.1 Solar Radiation
      2.5.1.2 Jupiter Reflected Radiation
      2.5.1.3 Satellites and Planets
      2.5.1.4 Atmospheric Illumination
   2.5.2 Heat
   2.5.3 Radio
      2.5.3.1 Decimetric (UHF) Radiation
      2.5.3.2 Decametric (HF) Radiation

2.6 Satellites and Meteoroids ................................................ 16
   2.6.1 Galilean Satellites
   2.6.2 Smaller Satellites
   2.6.3 Meteoroids

2.7 Charged Particles ............................................................ 18
   2.7.1 Galactic Cosmic Rays
   2.7.2 Solar Particle Events
   2.7.3 Solar Wind
   2.7.4 Trapped Radiation Belts
3.5.1 Above the Tropopause
3.5.2 Below the Tropopause

3.6 Satellites and Meteoroids .............................................. 42

3.7 Charged Particles ....................................................... 44
3.7.1 Particle Types
3.7.2 Calculation of Low Energy Radiation Belt Number Densities
3.7.3 Calculation of Radiation Belt Fluxes

3.8 Transfer Properties in the Magnetosphere and Ionosphere .......... 53

3.9 Atmospheric Structure ................................................. 54

3.10 Cloud Properties ....................................................... 54

3.11 Atmospheric Motions .................................................. 54

3.12 Transfer Properties in the Atmosphere ............................... 64

3.13 Interior ............................................................... 64

REFERENCES .......................................................... 65

APPENDIX A. Symbols ..................................................... 73

APPENDIX B. Dipole Magnetic Field ..................................... 78

APPENDIX C. Relationships among Radio Astronomy Quantities .......... 79

APPENDIX D. Atmosphere and Cloud Relationships ....................... 80

APPENDIX E. Glossary ..................................................... 83

LIST OF NASA SPACE VEHICLE DESIGN CRITERIA MONOGRAPHS ........... 87
THE PLANET JUPITER (1970)

1. INTRODUCTION

Both qualitative and quantitative descriptions of Jupiter are needed to properly design spacecraft for the investigation of that planet and its environment. Although different data are needed for atmospheric probes, orbiters, and flyby missions, the paucity of present data permits inclusion of pertinent information for all three types of missions in one monograph.

This monograph is based primarily on data obtained through 1970 but includes some material published during the first half of 1971 and conclusions of the Jupiter Radiation Belt Workshop held at the Jet Propulsion Laboratory in July 1971. All the information of Jupiter has been derived from data obtained at the angular and spectral resolutions possible with Earth-based instrumentation or with sensors on aircraft, rockets, and balloons. The observations were made primarily in the visible, near visible, infrared, and radio portions of the electromagnetic spectrum. The information was assessed for the potential effects of Jupiter's environmental properties on spacecraft performance. The assessment was done independently for the three types of missions under consideration and formulated for overall spacecraft as well as for subsystem design. Many of the environments presented, including physical properties, gravity field, charged particles, and atmospheric structure, are also important for mission planning.

For flyby spacecraft, the effects of electric and magnetic fields and charged particles on electronic components are important as well as the effects of electromagnetic radiation on light, heat, and radio-wave sensitive subsystems and the hazards of meteoroids on structural integrity and exposed surfaces. An orbiter spacecraft suggests similar environment-design interaction plus the effects of prolonged or repeated exposure on survival capabilities. In the case of atmospheric-entry probes, the effects of static, transport, and dynamic properties of Jupiter's atmosphere on pressure and temperature-sensitive electronic and structural components are important for design consideration. In table I the relative design importance of various aspects of Jupiter's environment on the subsystems of an entry probe are given.

References 1, 2 and 3 describe the electromagnetic, magnetic field, and meteoroid environments which are applicable to the design of spacecraft which will travel to and encounter the environment of Jupiter. Pertinent symbol definitions, mathematical formulations, and glossary are contained in appendices A through E.

2. STATE OF THE ART

2.1 Physical Properties

The physical properties considered here are the planetary mass, radius, shape, mean density, and rotational quantities. Related discussions of the orbital elements are given in Melbourne et al. (ref. 4). Photometric properties are discussed in section 2.5 of this monograph.
## Table I

**Impacts of Jupiter’s Environments on the Design of Atmospheric Probe Subsystems**

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<tr>
<td>Physical Properties</td>
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</tr>
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</table>

*S means strong, probable, or direct design impact.

W means weak, possible, or indirect design impact.
2.1.1 Mass

The mass $M_{jS}$ of the Jupiter system, i.e., planet plus satellites, may be most accurately determined from observations of the angular positions on the sky of certain minor planets and of the outer satellites of Jupiter. In both cases, the Sun and Jupiter furnish the major gravitational forces acting on the smaller bodies. The quantity which is directly derived from the analysis of the varying orbital elements is the “reciprocal mass” $M_\odot/M_{jS}$, the ratio of the mass of the Sun to that of the Jupiter system. Although JPL uses the value $M_\odot/M_{jS} = 1047.3908 \pm 0.0074$ (ref. 4), the associated uncertainty estimate is not broad enough to include other recent values cited by Klepczynski (ref. 5) and Bec (ref. 6). To include these results, the value and uncertainty $M_\odot/M_{jS} = 1047.39 \pm 0.04$ are adopted herein as the basis for the planetary mass values.

To obtain the mass of the planet $M_j$, two additional parameters are needed: the mass of the Sun, $M_\odot = (1.989 \pm 0.002) \times 10^{33}$ g (ref. 7), and the ratio of the mass of the four largest (Galilean) satellites to that of the planet, $(M_{jS} - M_j)/M_j = (1.955 \pm 0.085) \times 10^{-4}$ (ref. 8). The mass of Jupiter is then $M_j = (1.899 \pm 0.002) \times 10^{30}$ grams in which the uncertainty is traced to that in the value of $M_\odot$. Related discussion appears in section 2.2.

2.1.2 Dimensions

The angular diameters of Jupiter have been measured by various direct, visual techniques. In one widely quoted comparative analysis thereof, Rabe (ref. 10) derives 38.09 and 35.76 seconds of arc for the equatorial and polar diameters, respectively, for a mean Jupiter-Earth separation of 5.2028 AU. The uncertainties are large, somewhat less than one second of arc for each value. Dollfus* reports a diameter of $70850 \pm 100$ km for Jupiter from synthesis of double image micrometer observations, but the relationship of this value to the uppermost cloud layer has not been established. A value with reasonable uncertainty results from the thorough analysis of the duration of the eclipses of the Galilean satellites by Sampson (ref. 11) who derived an unweighted average of eight independent determinations (each from several observations) for the equatorial semi-diameter of $18.9275 \pm 0.05$ seconds of arc at 5.2028 AU. This value requires knowledge only of satellite orbit dimensions, period, and eclipse timing. Orbit dimensions whose base is hundreds of seconds of arc come from many visual observations. Consequently, such observations give more precise results than direct diameter measurements. Precision of eclipse timing measurements is excellent. If the modern value for the astronomical unit of $1.49597893 \times 10^8$ km is employed (ref. 4), Sampson’s work yields an equatorial radius $R_j$ of $71422 \pm 200$ km for Jupiter.

The value 0.06117 for the optical flattening obtained from the above angular diameters is probably reliable to only one significant figure. However, the shape of the planetary equipotential surfaces (approximately ellipsoids of revolution) may be derived from the

parameters $J_2$ and $J_4$ of Jupiter's gravitational field (sec. 2.2) and the period of rotation $T_0$ (sec. 2.11.1). Calculation of the polar diameter of that surface which corresponds to the foregoing value of the equatorial radius derived from Sampson (ref. 11) and the values of $J_2$, $J_4$, and $T_0$ (adopted in the sections cited) leads to the value $0.0645 \pm 0.0008$ for the flattening $\varepsilon$. It narrows the range of the optical values cited in the literature (ref. 9).

The conventional expression for the planetary radius $R_s$ as a function of jovicentric latitude $\phi$ is

$$R_s = R_J [1 - \varepsilon (\sin \phi)^2]. \quad (1)$$

Equation (1) is correct to first order in $\varepsilon$, and the retention of higher order terms is not justified because of the large uncertainty in $R_J$. Therefore, equation (1) is adopted here for the specification of the shape of the planet in visible light with the foregoing values of $R_J$ and $\varepsilon$. $R_s$ refers to a level near the clouds in Jupiter's atmosphere and not to a specific solid or liquid planetary surface (secs. 2.9 and 2.13).

### 2.1.3 Mean Density

The mean density $\bar{\rho}$ is given by the following relationship (ref. 9); it is correct only to first order in $\varepsilon$

$$\bar{\rho} = \frac{3 M_J (1 + \varepsilon)}{4 \pi R_J^3}. \quad (2)$$

Equation (2) and the foregoing values for $M_J$, $R_J$, and $\varepsilon$ lead to the value $1.32 \pm 0.01$ g/cm$^3$ for $\bar{\rho}$, a specific gravity for the planetary material only slightly greater than that of water.

### 2.1.4 Rotational Quantities

The periods (near 10 hours) and angular velocities associated with the rotation of Jupiter are discussed in section 2.11.1. The specific values, $T_0 = 9h 55m 29.73 \pm 0.04$ and $\omega_0 = (1.758531 \pm 0.000002) \times 10^{-4}$ radian/sec, are adopted there.

The directions of Jupiter's rotational axis and equatorial plane (determined from measurements of the position angle on the sky of the oval disk's major axis or the banded surface features) suggest that the equatorial plane is identical with the invariant plane of the Galilean satellites. The determination of the equatorial plane by Souillart (ref. 12) has been adopted for this monograph. Accordingly, the celestial coordinates of the north pole of Jupiter's rotation (also the direction of its angular momentum vector with the usual right-hand convention) have been taken as $a_R = 17^h 52^m 00^s.84 + 0^s.247 \quad (t - 1910.0)$ for right ascension and as $\delta_R = +64^\circ 33'34''6 - 0.60'' \quad (t - 1910.0)$ for declination. Here, $t$ is in tropical years, 1910.0 implies noon Greenwich Mean Time, 1 January 1910, and
\( \alpha_R \) and \( \delta_R \) refer to the equator and equinox at time \( t \). The North Pole points into the constellation Draco with the following approximate inclinations: \( 3^\circ \, 07 \) to the north pole of Jupiter's orbit, \( 2^\circ \, 00 \) to the north pole of the ecliptic, and \( 25^\circ \, 45 \) to the north pole of the Earth's rotation.

The moment of inertia \( C \) of Jupiter about its rotational axis may be derived from numerical models of the planet's interior. Various sources (refs. 7, 13, and 14) give \( C \) in the range

\[
C = (0.25 \pm 0.01) M_J R_J^2 = (2.4 \pm 0.1) \times 10^{49} \text{ g cm}^2. \tag{3}
\]

According to DeMarcus (ref. 13), the moment of inertia \( A \) about any equatorial axis may be derived simply in the approximate form

\[
A = C - J_2 M_J R_J^2 = (2.25 \pm 0.1) \times 10^{49} \text{ g cm}^2. \tag{4}
\]

The angular momentum \( C \omega_0 \) of the planet is \((4.3 \pm 0.2) \times 10^{45} \text{ g cm}^2/\text{sec} \) when the foregoing values are employed.

The dimensionless parameter, \( \sigma_R = \omega_0^2 \, R_J^3/G M_J \), is a useful measure of the atmospheric ratio of the centrifugal and gravitational forces and will be required in section 2.2. The value of \( 0.0889 \pm 0.0008 \) may be derived for \( \sigma_R \) from the foregoing adopted values.

2.2 Gravity Field

The gravitational field of a planet is conventionally described in terms of a potential function \( \Psi \), approximated by the first few terms of an infinite series,

\[
\Psi = -\frac{G M_J}{R_J} \left[ \frac{R_J}{R} + \frac{\sigma}{2} \left( \frac{R}{R_J} \cos \phi \right)^2 - J_2 \left( \frac{R_J}{R} \right)^3 P_2(\sin \phi) \right. \\
\left. + J_4 \left( \frac{R_J}{R} \right)^5 P_4(\sin \phi) \right]. \tag{5}
\]

The forces on a test body include those proportional to the gradient of \( \Psi \) to which other relevant forces, e.g., perturbations by the sun and planetary satellites, may be added. The coefficient \( \sigma \) of the term in \( \Psi \) proportional to \( R^2 \) must be set equal to zero if the forces are to be evaluated in an inertial coordinate system. The coefficient \( \sigma \) must be equal to \( \sigma_R \) (sec. 2.1.4) if they are to be evaluated in a coordinate system rotating with the planet, e.g., in the atmosphere.

2.2.1 Inertial Coordinates

An inertial coordinate frame is appropriate for bodies not constrained to move with the planet, e.g., spacecraft beyond Jupiter's atmosphere, and requires that the value \( \sigma = 0 \) be
substituted in expressions for $\Psi$ and its derivatives. In the evaluation of the other constants in equation (5), the mass and radius values adopted in section 2.1 lead to $R_J = 71422 \pm 200$ km and $GM_J/R_J = 1774 \pm 5$ km$^2$/sec$^2$. Values for the harmonic coefficient $J_2$ (which measures the contribution of the planet’s flattening to the gravity field) are derived from analyses of the observed positions of the Galilean satellites. The best determination (ref. 15) is $3/2 J_2 = 0.02206 \pm 0.00022$. A less accurate value for $J_4$ has been obtained from analysis of the observed positions of $JV$, the innermost satellite, as $-\frac{15}{4} J_4 = 0.00253 \pm 0.00141$ (ref. 8). Thus, the maximum contribution of the $J_2$ and $J_4$ terms in $\Psi$ is approximately 1.5 percent, and an appropriate approximate form for the gravitational potential is

$$\Psi = -(1774 \pm 30 \text{ km}^2/\text{sec}^2)(R_J/R). \quad (6)$$

The gravitational forces on a spacecraft additional to those proportional to the gradient of $\Psi$, caused by the gravity of the Sun and the four Galilean satellites, may be calculated from the data given in section 2.6 and the positions tabulated annually in the *American Ephemeris and Nautical Almanac* (ref. 16).

### 2.2.2 Corotating Coordinates

A corotating coordinate frame is appropriate in situations in which the body of interest is constrained to rotate with the planet, and in these cases $\sigma = \sigma_R$ (sec. 2.1.4). Although this frame can be applied to magnetospheric charged particles whose motions may be constrained by rotating magnetic field lines, it is used in this monograph only for regions in which the quantity $\sigma(R/R_J)^2$ is small compared to unity, applicable, for example, to an atmospheric entry vehicle. For convenience, the altitude $z$ (here $z \ll R_J$) above the planetary surface (eq. 1) is defined by the relationship

$$R = R_J[1 - \epsilon(\sin \phi)^2 + z/R_J]. \quad (7)$$

The acceleration of gravity (directed at an angle $\eta_1$ equatorward of the inward radial direction and of magnitude $g$) then can be specified simply by expressions of first order in $z$, $J_2$, $\sigma_R$, and $\epsilon$, namely

$$g = \frac{GM_J}{R_J^2} [1 - \sigma_R + \frac{3}{2} J_2 - \frac{2z}{R_J^2}] [1 + (2\epsilon + \sigma_R - \frac{9}{2} J_2)(\sin \phi)^2] \quad (8)$$

and

$$\eta_1 = \left(\frac{3}{2} J_2 + \sigma_R/2\right)\sin 2\phi. \quad (9)$$

In the evaluation of these expressions, the approximate relationship, $\epsilon = \frac{3}{2} J_2 + \sigma_R/2$
(which is closely satisfied in Jupiter’s case), may be used. With the foregoing values the range of magnitudes of the acceleration of gravity is $2500 \pm 200 \text{ cm/sec}^2$, and $\eta_1$ has maximum deflection of $3^\circ 7 \pm 0^\circ 3$ at jovicentric latitude $\phi$ of $45^\circ$.

### 2.3 Magnetic Field and Magnetosphere

Properties of Jupiter’s magnetic field and magnetosphere are inferred from the radio emissions (sec. 2.5.3) received at Earth from Jupiter. Carr and Gulkis (ref. 17) and Warwick (ref. 18) review the considerations involved. Section 2.7 treats charged particles within the magnetosphere.

#### 2.3.1 Magnetic Field Strength and Configuration

The characteristics of Jupiter’s UHF and HF radiation (sections 2.5.3.1 and 2.5.3.2) permit quantitative descriptions of the magnetic field strength and configuration to be formulated although the uncertainties are considerable. Appendix B includes formulas for calculation of the magnetic field strength and lines of force at any location in Jupiter’s magnetosphere on the basis of the dipole moment, inclination, and displacement adopted in the following sections.

**2.3.1.1 Dipole Moment**

UHF data permit an order-of-magnitude determination of the magnetic field in the radiation belts (sec. 2.7.4); the values cited in the literature are near $1 \text{ gauss}$ about $2 \text{ radii}$ from the dipole (refs. 17, 18, and 19) which correspond to field strengths near $10 \text{ gauss}$ about $1 \text{ radius}$ from the dipole and dipole moment $M_1$ near $3 \times 10^{30} \text{ gauss-cm}^3$. The UHF values at $1 \text{ radius}$ are compatible with the narrower range commonly derived from the HF bursts if the burst frequency is taken to equal the local electron gyrofrequency at the source location. The local magnetic field strengths corresponding to observed HF frequencies near $30 \text{ MHz}$ are near $12 \text{ gauss}$ (refs. 18, 20, and 21). The dipole moment obtained from this field strength value depends sensitively on the relative locations of the dipole and the HF burst generation region. If the dipole is centered on the planet and the emission near $30 \text{ MHz}$ is generated just above Jupiter’s atmosphere in the magnetic equatorial plane, the moment corresponding to the $12 \text{ gauss}$ field strength is approximately $4 \times 10^{30} \text{ gauss-cm}^3$. This value is close to the value of $4.2 \times 10^{30} \text{ gauss-cm}^3$ derived by Warwick (refs. 18 and 22), who completed a detailed analysis of the HF burst structure based on a severely offset dipole (sec. 2.3.1.3). To take into account published field strength and dipole moment values based on analyses of HF and UHF data (refs. 17, 19, 20, 21, and 23), an uncertainty factor of two and a nominal value of $M_1 = 4 \times 10^{30} \text{ gauss-cm}^3$ are adopted for this monograph. The consequent range of field strengths in the atmosphere is from $6 \text{ gauss}$ at the equator to $48 \text{ gauss}$ at the poles if the dipole is centered on Jupiter. The dipole and field have the opposite orientation from the Earth’s as Jupiter’s northern hemisphere contains its north-seeking magnetic pole (refs. 18 and 24).
2.3.1.2 Dipole Inclination

The UHF emission has a linearly polarized component whose position angle is roughly parallel to Jupiter's equator, but the position angle varies regularly by nearly 10° with the period of Jupiter's rotation (sec. 2.11.1 and refs. 25 and 26). This slightly asymmetric variation results from the inclination of Jupiter's magnetic axis to its rotation axis by an angle nearly equal to the above amplitude and the rotation of the field and its sources. This small inclination has a far smaller effect on the field strengths than the large uncertainty in the dipole moment (sec. 2.3.1.1) and should therefore be ignored for design purposes.

2.3.1.3 Dipole Displacement

The asymmetric properties of the UHF polarization and intensity variations and of the HF burst probability and dynamic spectra cannot be explained in terms of an inclined, planet-centered dipole field (ref. 18). Although various authors have contended that significant magnetic field anomalies and quadrupole moment could be responsible for the asymmetries, Warwick (refs. 18 and 22) has claimed that displacement of the magnetic dipole approximately 0.7 radius south of Jupiter's center is responsible. In contrast, many investigators find such a large displacement unreasonable in view of postulated field maintenance processes in Jupiter's interior and of the measurements of the position of Jupiter's UHF emission centroid within 0.4 radius of the geometric center according to reference 27. Berge (ref. 28) has refined the measurement by reference 27 to find a small southward displacement of less than 0.2 radius for the emission centroid. If this displacement corresponds to that of the magnetic dipole, the influence on derived field strengths in the atmosphere or above is everywhere less than a factor two with respect to the centered situation. Thus, it is concluded that for design purposes any possible dipole displacement can be ignored in view of the large uncertainties associated with the dipole moment itself (sec. 2.3.1.1) and with the energetic charged particle fluxes (sec. 2.7.4).

2.3.2 Magnetosphere

Jupiter's magnetosphere is the region within which Jupiter's own magnetic field and charged particles (sec. 2.7) dominate. In the solar direction the magnetosphere boundary is conventionally fixed by equating the solar wind and planetary magnetic field pressures and consists of both a shock front (at which the solar wind velocity decreases abruptly) and a magnetopause (limiting Jupiter's regular magnetic field). The result is 50 ± 20 R_J which includes most estimates of the distance of the sunward boundary from Jupiter (refs. 17, 29, and 30). This distance is adopted here as the minimum radius of the magnetosphere. In other directions additional complex features, including distortion by Jupiter's rotation (sec. 2.11.1) and a more distant magnetosphere boundary, are expected by analogy with the Earth. In the anti-solar direction the length of Jupiter's magnetospheric tail is greater than 100 R_J (refs. 17 and 31).

2.4 Electric Fields

There is very little evidence for static electric fields, either large scale or local, in any observations of Jupiter or in their interpretations.
The local value of the field strength suggested in the HF burst theory of Goldreich and Lynden-Bell (ref. 21) ranges up to approximately 10 volts/meter. The field which would be required to accelerate electrons (charge \( -e \)) electrostatically to the energy \( (E \approx 6 \text{ MeV}, \text{sec. 2.7.4.2}) \) inferred from the UHF observations has the order of magnitude of \( \mathcal{E} = E/eR_J = 0.1 \text{ volt/meter} \). The electric field induced in an inertial coordinate system by the rotation of Jupiter’s magnetic dipole field is perpendicular to the magnetic field lines and of the order of magnitude of \( \mathcal{E} = \nu \mathbf{B} = (\omega_B R)(M_I/R^3) = (14 \text{ volts/meter})(R_J/R)^2 \).

For this monograph, a range between zero and the largest of the foregoing estimates is adopted for the general field within the magnetosphere:

\[
0 \leq \mathcal{E} \leq (14 \text{ volts/meter})(R_J/R)^2.
\]  

(10)

Local, time-varying electric fields could be associated with plasma waves of the disturbances implied by the HF bursts, but their field strengths are unknown.

Within the ionosphere the electrical conductivity of the plasma is probably large enough for the static field strengths already quoted to be ample upper limits. Below the region of maximum ionospheric electron density, the low atmospheric electric conductivity (sec. 2.12.1) is conducive to the build-up of static charge accumulations, particularly in regions of convection and cloud formation. By analogy with the Earth’s atmosphere, randomly oriented electric fields of strength up to the order of \( 3 \times 10^4 \text{ volts/meter} \) might be expected in scales of hours and kilometers. Other fields up to the order of \( 10^6 \text{ volts/meter} \) may occur in scales of seconds and centimeters just before lightning strikes sharp, projecting objects (ref. 32).

### 2.5 Electromagnetic Radiation

The electromagnetic radiation environment in the vicinity of Jupiter is divided into light, heat, and the radio spectrum for presentation in the following sections.

#### 2.5.1 Light

This section treats X-rays, the ultraviolet, visible, and near infrared portions of the electromagnetic spectrum, i.e., wavelengths between about 1 Å and 5 μm.

##### 2.5.1.1 Solar Radiation

Near Jupiter the Sun is the dominant source of light which may be direct, reflected, or scattered. NASA SP-8005 (ref. 1) summarizes the direct solar spectrum at 1 AU and specifies \( P_\lambda \), the power per unit area and per unit wavelength interval for wavelengths between 50 Å and 100 cm. The values from reference 1 and standard relationships for intensity and flux (app. E) lead to the formulas given for the Sun in table VII (sec. 3.5).
2.5.1.2 Jupiter Reflected Radiation

The solar radiation reflected from Jupiter (or its clouds) has been observed only at phase angles less than 12° and is conventionally described in terms of astronomical magnitudes, colors, and albedos. The definition of geometric albedo $p$ (app. E) leads to formulae for the intensity and flux of Jupiter-reflected radiation observed at zero phase angle and jovicentric distance $R$

$$I_\lambda = \frac{pP_\lambda}{\pi r^2}$$  \hspace{1cm} (11)$$

and

$$F_\lambda = \frac{pP_\lambda}{r^2 (R/R_J)^2}$$  \hspace{1cm} (12)$$

where $r^2$ evaluated for Jupiter in (AU)$^2$ has the value $27 \pm 3$. For other observable phase angles ($< 12°$), these quantities are smaller than the results of equations (11) and (12). Therefore, it can be inferred that the zero phase angle formulae constitute correct upper limits at all phase angles.

Photoelectric measurements of the geometric albedo $p$ for Jupiter at wavelengths between 2000 Å and 1.5 μm are summarized in table II. Limiting values from table II are shown in figure 1, and although the data are restricted to one decade in wavelength, the albedo is apparently decreasing as both ends of the observed range are approached. Thus, a range of values for $p$ between 0 and 0.3 at unobserved wavelengths in addition to the range from table II is adopted herein. Table II also lists the result of a photoelectric determination of the bolometric albedo $p_b$. That value and its substitution in the following expression for the integrated reflected flux

$$F = \frac{p_b S}{r^2 (R/R_J)^2},$$  \hspace{1cm} (13)$$

are adopted here.

2.5.1.3 Satellites and Planets

The brightnesses of Jupiter's satellites and other planets are specified by their magnitudes and colors (defined in appendix E) which are summarized in references 7, 9, and 38. Their variations with satellite rotation and phase angle $a$ are limited respectively by ±0.3 magnitude (ref. 38) and probably by $(0.03 \pm 0.02) \alpha$ for $a$ in degrees; the phase angle uncertainty includes all Jupiter satellite observations ($a \leq 12°$) and the complete phase variations for Venus (deep atmosphere) and the Moon (no atmosphere) to within 0.1 magnitude. When
the geometry of satellite illumination and observation is specified in terms of \( a, r, \) and \( \Delta \) (\( r \) and \( \Delta \) in AU only and \( a \) in degrees), the relationship

\[
m_v = (m_0 \pm 0.3) + 5 \log(r\Delta) + (0.03 \pm 0.02)a
\]

(14)

applies. Equation 14 is adopted for use with the magnitudes and colors given in table IX (sec. 3.5) for both Jupiter's satellites and other planets.

### TABLE II

**OBSERVED GEOMETRIC ALBEDOS FOR JUPITER**

<table>
<thead>
<tr>
<th>Wavelength** (Å)</th>
<th>Geometric Albedo ( p )</th>
<th>Source</th>
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<tbody>
<tr>
<td>1730</td>
<td>0.2 ± 0.15</td>
<td>Moos et al. (ref. 33)</td>
</tr>
<tr>
<td>2100</td>
<td>0.2 ± 0.15</td>
<td>Stecher (ref. 34), Jenkins et al. (ref. 35), and Anderson et al. (ref. 36)</td>
</tr>
<tr>
<td>2800</td>
<td>0.23 ± 0.1</td>
<td></td>
</tr>
<tr>
<td>3530 (U)</td>
<td>0.31 ± 0.05</td>
<td>Irvine et al. (ref. 37) and Harris (ref. 38)</td>
</tr>
<tr>
<td>4480 (B)</td>
<td>0.42 ± 0.05</td>
<td></td>
</tr>
<tr>
<td>5540 (V)</td>
<td>0.50 ± 0.05</td>
<td></td>
</tr>
<tr>
<td>6900 (R)</td>
<td>0.47 ± 0.05, 0.35 ± 0.05</td>
<td>Harris (ref. 38)</td>
</tr>
<tr>
<td>8200 (I)</td>
<td>0.31 ± 0.1</td>
<td>Irvine et al. (ref. 37)</td>
</tr>
<tr>
<td>10635</td>
<td>0.31 ± 0.1</td>
<td></td>
</tr>
<tr>
<td>15000</td>
<td>0.15 ± 0.1</td>
<td>Danielson estimate (ref. 39)</td>
</tr>
<tr>
<td>Bolometric Albedo</td>
<td>( p_b = 0.28 ± 0.03 )</td>
<td>Taylor (ref. 40)</td>
</tr>
</tbody>
</table>

*Uncertainties are estimated from comparison with additional published data.

**Letters specify pass bands in the conventional magnitude system.
2.5.1.4 Atmospheric Illumination

Within Jupiter's atmosphere the foregoing intensities, fluxes, and magnitudes most likely represent appropriate upper limits on radiation from the Sun, Jupiter's clouds, its satellites, and other planets if \( r^2 = 27 \pm 3 \text{ (AU)}^2 \) and \( \frac{R}{R_J} = 1 \). Additional sources of light radiation which might be expected are scattering, airglow, aurorae, and lightning. No simple description of any of these phenomena of Jupiter's atmosphere is possible, but the light reflected from Jupiter's clouds is adopted as an upper limit for the intensities and fluxes of the first three. Lightning could presumably be as intense as that in the Earth's atmosphere and thus could require that light-sensitive devices be covered during storms.

![Figure 1. Typical values of geometric albedo \( p \) and disk brightness temperature \( T_{BD} \), taken from table II and references cited in section 2.5.2. (The temperature scale is logarithmic.)](image-url)
2.5.2 Heat

Direct solar radiation and intrinsic, thermal Jupiter radiation are important in the wavelength range between $\mu m$ and 1 cm. The solar contribution is covered in section 2.5.1.

The determination of Jupiter's infrared spectrum is difficult because terrestrial atmospheric absorption distorts data observed even at great altitudes. The difficulty is especially severe at wavelengths near 20 $\mu m$ where Jupiter's thermal radiation is greatest. Even the thorough discussion of data at wavelengths shorter than 14 $\mu m$ by Gillett et al. (ref. 41) may not have completely eliminated the absorption problem from ground-based observations. At longer wavelengths, extending into the microwave radio region (sec. 2.5.3.1), Kellermann (ref. 42) summarizes the data of many observers in terms of the brightness temperature $T_{BD}$ of Jupiter's thermal disk radiation (fig. 1, sec. 2.5.1). For wavelengths between 1 $\mu m$ and 100 cm, the band labelled $T_{BD}$ in figure 7 indicates (with ample uncertainty limits, and reasonable extrapolations into unobserved regions) the range of values reported by these authors. The increase of $T_{BD}$ with wavelength in the centimeter range implies that the radiation emerges from lower, warmer atmospheric levels at longer wavelengths. Ammonia absorption is responsible for the dip in $T_{BD}$ near 1 cm wavelength. The intensity and flux of the thermal disk radiation may be computed readily from the brightness temperature through the Planck function $B_{\lambda}(T)$ or $B_{\nu}(T)$, tabulated, for example, by Allen (ref. 7) and the formulas given in table VII (sec. 3.5).

The effective temperature of Jupiter resulting from data obtained with an aircraft-borne Germanium bolometer is given as $T_e = 134 \pm 4^\circ K$ (ref. 43). However, the uncertainty may be too small because Jupiter was observed on only one date and the 8 to 14 $\mu m$ results commonly fall below 130$^\circ K$ (ref. 44). Thus, it is concluded that for this monograph the uncertainty be increased slightly to $T_e = 134 \pm 6^\circ K$, corresponding to a 20 percent uncertainty in the flux of energy from Jupiter. This energy, given in table VII (sec. 3.5), almost certainly is radiated nearly uniformly over the planet's surface.

2.5.3 Radio

The range considered here includes all wavelengths greater than approximately 1 cm (or, equivalently, frequencies less than 30 000 MHz). Radio observations of Jupiter conventionally are made in two distinct spectral regions and indicate that three distinct source mechanisms are involved.

2.5.3.1 Decimetric (UHF) Radiation

Thermal radiation from Jupiter's disk (sec. 2.5.2) contributes significantly to the UHF radiation environment near Jupiter. It is randomly polarized and continuous in frequency. Its brightness temperature is constant in time and uniform over the surface to the accuracy shown in figure 7. Its radio intensity and flux may be calculated from figure 7 and the relationships given in appendix C.

Non-thermal UHF radiation is observed from a region several Jupiter radii in extent, elongated parallel to Jupiter's magnetic equator. As Jupiter rotates, this radiation varies by
about 15 percent in intensity because of the changing magnetic latitude of the observer, and the position angle of its linearly polarized component (maximum about 30 percent near $\lambda = 30$ cm) moves roughly parallel to the magnetic equator. An observer at a distance of 4.04 AU receives $7 \pm 3$ flux units ($1 \text{ FU} = 10^{-26} \text{ watt/m}^2 \text{ Hz}$) at all wavelengths between 1 and 300 cm with perhaps 50 percent greater uncertainty outside the wavelength range 3 to 100 cm. Dickel et al. (ref. 45) summarize the UHF spectrum and cite some of the many references from which the foregoing brief description was compiled. The non-thermal component synchrotron radiation from relativistic electrons trapped in Jupiter's magnetic field (secs. 2.3.1 and 2.7.4).

Close to Jupiter the specification of the spatial distribution of the synchrotron radiation requires the utilization of data reported by Berge (ref. 24) and Branson (ref. 46) at 10.4 and 21 cm wavelengths. Figures 2 and 3 show their reconstructed emission distributions as seen from the Earth in terms of brightness temperatures at resolutions of the order of Jupiter's radius $R_J$. Although some data suggest that the source may extend to significantly greater distances from Jupiter at longer wavelengths (refs. 47 and 48), this concept does not have general concurrence (ref. 18). To provide a model which yields approximately correct brightness temperatures for the synchrotron source, the following approach is suggested. Figure 8 illustrates the volume enclosed by a sphere of radius $3 R_J$, centered on the magnetic dipole and truncated by two planes parallel to and 1 $R_J$ away from the magnetic equator. Then if $D$ is the path length within the described volume in the direction of observation (extending either to the planetary surface or to infinity), the brightness temperature is given by

$$T_{BS} = \frac{D \lambda^2}{R_J} (0.30 \pm 0.15)^{\circ} K$$

for $\lambda$ in cm. The uncertainty in this expression is large enough so that $D$ can be obtained from a reasonable estimate instead of from a geometric calculation. Thus over the center of the disk, by taking $D = 2 R_J$, the $T_{BS}$ values $66 \pm 33^\circ K$ and $260 \pm 130^\circ K$ are obtained at $\lambda 10.4$ and 21 cm; just beyond the limb, by taking $D = 5 R_J$, the values $165 \pm 82^\circ K$ and $650 \pm 325^\circ K$ are obtained at the same two wavelengths. The results of Berge (ref. 24 and Branson (ref. 46) are included within the uncertainties quoted.

The intensity of the UHF synchrotron radiation is then given in terms of the brightness temperature $T_{BS}$ by the Rayleigh-Jeans law (app. C). The maximum flux is $7 \pm 3$ flux units, independent of $\lambda$, at 4.04 AU, proportional to $R^{-2}$ for $R \gg R_J$ and nearly independent of $R$ within the source, i.e., for $R \lesssim 2 R_J$. An expression satisfying these requirements is

$$F_{\nu} = \frac{(5 \pm 2) \times 10^8 \text{ FU}}{4 + (R/R_J)^2} \quad (16)$$

and its maximum value is consistent with the intensity and brightness temperature formulation at the maximum $D \approx 5 R_J$. The multiplication of equation (16) by the appropriate bandwidth (approximately 3000 MHz) yields the corresponding integrated flux. Table VII (sec. 3.5) summarizes the foregoing conclusions.
Figure 2.—Aperture synthesis map of Jupiter's 10.4 cm radiation. The central circle represents Jupiter's optical disk, the curved lines represent contours of constant brightness temperature, and the bars at lower right indicate the instrumental resolution (from ref. 24).

Figure 3.—Three aperture synthesis maps of Jupiter's 21 cm radiation. The central circles represent Jupiter's optical disk, the curved lines represent contours of constant brightness temperature (interval 47°K), the oval at upper left represents the instrumental resolution, and the figures at lower right represent the orientations of the rotational and magnetic axes (from ref. 46).
2.5.3.2 Decametric (HF) Radiation

HF bursts are observed in the frequency range 5 to 39.5 MHz (wavelengths between 7.6 and 60 meters). Many features of their polarization, directionality, and time and frequency structure display great variety although some properties (notably burst detectability and dynamic spectra) are correlated with the relative orientations of the Earth, Jupiter, and the Galilean satellite, Io. Their characteristics are reviewed by Warwick (ref. 49) and Carr and Gulkis (ref. 17) who give numerous references to original observations and analysis. Because sizes of HF burst sources are small in comparison to planetary dimensions (refs. 50 and others), because bursts are sporadic, and because there is no consensus for any of the proposed generation mechanisms (refs. 18, 20, 21, and 51), it is not appropriate to propose particular characteristics for HF radiation near Jupiter. Because the intensities could be very high (ref. 18), however, the use of spacecraft equipment sensitive to external radiation at wavelengths longer than 100 cm is not recommended near Jupiter.

2.6 Satellites and Meteoroids

The distribution and properties of the solid particles in the Jupiter environment beyond the planet’s atmosphere may be inferred from optical, telescopic observations of the natural satellites of Jupiter. The bodies divide naturally into three categories.

2.6.1 Galilean Satellites

The four largest satellites are visible with even very small telescopes. Analysis of their observed angular positions with respect to Jupiter implies orbits of negligible eccentricity and inclinations to Jupiter’s equatorial plane, separations from Jupiter’s center between 5.9 and 27 R_J, and periods between 42 and 400 hours, i.e., between 4.2 and 40 Jupiter days. Analysis of their mutual perturbations leads to masses between 0.00002 and 0.0001 times that of Jupiter (ref. 9). Visual observations and measurements (ref. 52) indicate that their rotation is synchronous with their orbits and that values for their diameters are of the order of one second of arc as seen from the Earth.

Photometric and spectroscopic observations (refs. 38 and 53) do not furnish much information on the satellite properties. Surface temperatures may be variable within the range from 3°K to 170°K (refs. 54 and 55). With these thermal and appropriate gravitational conditions, the presence of satellite atmospheres is not excluded, but attempted observations thereof have been inconclusive.

Some of the most certain information available for these satellites is presented in table X (sec. 3.6) and references 4, 7, 9, and 56. Photometric information is discussed in section 2.5.1.3, and detailed orbital information is presented in the American Ephemeris and Nautical Almanac (ref. 16).

2.6.2 Smaller Satellites

Eight smaller satellites (identified as J V through XII), each much smaller than the Galilean satellites, have been observed photographically. The orbital elements of J V vary because of
strong perturbations by Jupiter's oblateness, whereas the orbital elements of the others vary because of perturbations by the Sun. The latter divide naturally into the categories J VI, VII, and X with prograde orbits of periods near 260 days and J VIII, IX, XI, and XII with retrograde orbits of periods near 700 days. Because the eight smaller satellites are so faint (magnitudes +13 to +19 as seen from the Earth), their masses (probably less than $10^{+23}$ grams), radii, and physical properties are not known. There is speculation that their physical properties resemble those of the asteroids. Other undiscovered, even smaller satellites of Jupiter also may exist.

Some of the most certain information available for these objects is presented in table X. Additional information and references may be found in table 14 of reference 4.

2.6.3 Meteoroids

There is no direct evidence of meteoroidal debris in the environment of Jupiter. Traditionally, the meteoroid environment near a planet has been inferred from the presumed gravitational enhancement of the nearby interplanetary cometary debris (ref. 3). Although such an enhancement is definitely to be expected, the uncertainties involved in the evaluation of the interplanetary debris distribution and the enhancement characteristics in the case of Jupiter amount to several orders of magnitude.

Another reasonable approach is to consider Jupiter's irregular satellites (J VI through XII) as the observable, large body end of a meteoroid distribution analogous to that expected in the asteroid belt. Such an approach assumes that the smaller satellites of Jupiter have individual properties and a mass distribution comparable to those of the asteroids, an assumption which is neither confirmed nor denied by the meager satellite data available.

The masses $M_n$ in grams and radii $R_n$ in cm for J V through XII presented in table X are derived from the relationships

$$\log M_n = 26 \pm 1 - 0.6 m_0$$

and

$$\log R_n = 8.3 \pm 0.3 - 0.2 m_0$$

which are reasonable for asteroids of absolute magnitude $m_0$, mass density $3.5 \text{ g/cm}^3$, and visual albedo 0.2. If the positions of the outer seven satellites are distributed uniformly within a volume centered on Jupiter of radius $500R_J$ and at latitudes within ±30°, the number per unit volume $N_M$ of particles of mass $M \geq 10^{18.5}$ grams is $10^{-31 \pm 1} \text{ m}^{-3}$. This is almost identical to the peak spatial density at this mass in the asteroid belt (ref. 3). Application of a crude mass distribution similar to that in reference 3 gives

$$\log N_M = -31 \pm 1 - (0.8 \pm 0.1)(\log M - 18.5).$$

Equation (19) requires $N_M$ in particles/m$^3$ and applies for $10^{-9} \leq M \leq 10^{20}$ grams within the foregoing given volume near Jupiter.

*The inner boundary for meteoroids could be $100R_J$, because of sweeping action by the satellites and gravitational forces. However, an inner boundary of $1R_J$ is assumed herein.
Omnidirectional fluxes of particles may be computed by multiplying \( N_M \) by \( v/4 \) where the speed \( v \) of the particles relative to a spacecraft (with speed \( v_s \)) is approximately \( v = \left( \frac{v_s^2}{2} + \frac{(R_J/R)(40 \text{ km/sec})^2}{2} \right)^{1/2} \). The particle density is \( 3.5 \times 10^{11} \text{ g/cm}^3 \). Lack of relevant data makes it impossible to decide whether such a distribution dominates anticipated cometary debris of lower density, but it is recommended that cometary debris be neglected on the basis of the small fluxes suggested in reference 3.

2.7 Charged Particles

Figure 4 schematically shows various charged particle environments near Jupiter. The following sections describe the number density, energy, and flux for each; transfer properties are discussed in section 2.8.

2.7.1 Galactic Cosmic Rays

Near the Earth, galactic cosmic ray intensities are modulated by the interplanetary magnetic field. In general, it is expected that this modulation reduces the intensities more severely at lower energies, closer to the Sun, and during intervals of greater solar activity. Quantitative predictions of the intensities near Jupiter have not been made, however. Jupiter's own magnetic field may reduce the intensities further, particularly at low energies and latitudes. Thus the approach adopted is to specify the fluxes in the energy range observed (0.1 to \( 10^2 \) GeV) as between zero and a spectrum extrapolated from the highest energies observed for the most abundant particle kinds at times near minimum solar activity. This spectrum can be approximated by the following expression for the flux of particles with kinetic energy greater than \( E \)

\[
\Phi_E = A(E + m_0c^2)^{-1.5}
\]  

(20)

where \( m_0c^2 \) is the rest energy of the particle and \( E \) is the particle kinetic energy in GeV (per nucleon for alpha-particles). The summary given by Haffner (ref. 57, fig. 2-3) specifies \( A \approx 10 \text{ cm}^{-2} \text{ sec}^{-1} \) for protons and \( A \approx 1 \text{ cm}^{-2} \text{ sec}^{-1} \) for alpha-particles; Fanselow (ref. 58) specifies \( A \approx 0.2 \text{ cm}^{-2} \text{ sec}^{-1} \) for electrons.

2.7.2 Solar Particle Events

Protons and alpha-particles of energy greater than 1 MeV are emitted sporadically by the Sun and have been detected by ground-based and spacecraft-borne instrumentation. Their fluxes near the Earth vary over several orders of magnitude in time, have both directional and isotropic components, and frequently can be identified with specific solar flares. Properties of sample events observed near the Earth are described in NASA TR R-169 (ref. 59). The recent theoretical model by Englade (ref. 60) cites the important references which treat the complex processes involved. This literature shows that neither the probability of occurrence of an event nor the variation of the particle fluxes with solar distance can be reliably estimated at this time.

*The average particle velocity relative to Jupiter is taken here as the orbital velocity (40 km/sec at \( R/R_J = 1 \)) which is distinct from the escape velocity given in reference 3.
Figure 4.—Schematic of charged particle regions near Jupiter.
2.7.3 Solar Wind

Properties of the solar wind are summarized by Hundhausen (ref. 61). On the basis of observations from Mariner spacecraft at 0.8 to 1.5 AU from the Sun, protons and electrons during quiet solar conditions have concentrations of approximately $5r^2 \text{cm}^{-3}$ (for $r$ in AU) and are streaming radially away from the Sun at speeds near 320 km/sec. Increases in solar activity are accompanied by temporary increases up to factors of 10 in the concentration and 3 in the speed. The applicable theory suggests that the extrapolation of these conditions on the basis of heliocentric distance $r$ to the orbit of Jupiter is justified. The interaction of the solar wind with Jupiter’s magnetic field is responsible for the formation of the outer boundary of Jupiter’s magnetosphere (sec. 2.3.2).

2.7.4 Trapped Radiation Belts

The high energy (several MeV) electron component of Jupiter’s charged particle environment is the only one about which inferences can be made directly from phenomena observed from the Earth. Part of Jupiter’s UHF radiation (sec. 2.5.3.1) is interpreted as synchrotron radiation from electrons trapped in Jupiter’s dipole magnetic field (sec. 2.3.1).

The theoretical relationships which may be used to predict the features of such synchrotron radiation are sufficiently complex that a unique solution for the electron distributions in energy, position, and pitch angle is not possible from observations so far made of the UHF radiation from Jupiter. The complexity increases if the magnetic field is treated as an unknown to be determined from the UHF data analysis. Thus, many electron density estimates are simply order-of-magnitude such as made by Carr and Gulkis (ref. 17) who suggest that a region roughly enclosed by two right circular cylinders centered on Jupiter of radii 1 and 3 R$_J$ and height 2 R$_J$ is characterized by a magnetic field of 1 gauss and contains 14-MeV electrons at a density near $10^{-3}$ cm$^{-3}$ and a flux near $3 \times 10^7$ cm$^{-2}$ sec$^{-1}$.

2.7.4.1 Relativistic Electron Flux

a. Peak Flux

Figure 5 shows the dependence of the total relativistic (kinetic energy $E > 0.5$ MeV, the electron rest energy) electron flux $\Phi$ (in cm$^{-2}$ sec$^{-1}$) on distance $L$ (in Jupiter radii) in the plane of the magnetic equator. Here $L$ is the magnetic shell parameter as in app. B. Near $L = 2$ the existence of a flux peak is a common interpretation of the 10 and 21 cm UHF emission patterns derived by Berge (ref. 24) and Branson (ref. 46) from aperture synthesis and reproduced in figures 2 and 3 (sec. 2.5.3.1). The models of Haffner (ref. 63) and Koepp-Baker (ref. 64) place their values of the flux in this peak (fig. 5) on the assumption that the synchrotron power is radiated isotropically from Jupiter. This approach neglects the strong beaming of the UHF radiation into the magnetic equatorial plane which is required by observations (ref. 65) and theory (ref. 18) and thus overestimates the total power and the electron flux required to produce it.
Figure 5.— Flux of relativistic electrons as a function of the distance from the dipole in the plane of the magnetic equator, for proposed models of Jupiter's trapped radiation belt.
Warwick (ref. 18) derives a flux $1.9 \times 10^7 \text{ cm}^{-2} \text{ sec}^{-1}$ for the peak near $L = 2$ by using the value of the magnetic field strength (2.0 gauss at $L = 1.8$ and $\phi = 0$) derived from the dipole moment quoted in section 2.3.1, an electron energy of 6.2 MeV (sec. 2.7.4.2), and an estimate for the UHF brightness temperature (183°K) above the center of the disk from Branson's map of 21 cm (ref. 46). The order-of-magnitude estimates of Eggen (ref. 62), Branson (ref. 46), and Carr and Gulkis (ref. 17) support the Warwick flux level. As an estimate of the correct flux at the major electron peak in Jupiter's belts, this monograph adopts $2 \times 10^7 \text{ cm}^{-2} \text{ sec}^{-1}$ with an uncertainty factor of three.

b Distribution with Distance

The distributions of radiating electrons as functions of distance from the dipole within Jupiter's belts shown in figure 5 are derived from Earth analogy, analysis of the distribution of the UHF radiation, and theory. Eggen's model (ref. 62) shows several peaks derived entirely from Earth analogy, and although there may be various peaks in Jupiter's electron distribution, it is unlikely that this simply-scaled model locates them correctly. Also, the observations (fig. 2, sec. 2.5.3.1) indicate that significant flux levels occur with $L < 3$.

Koepp-Baker's model (ref. 64) derives its double peak distribution from an interpretation of the Stokes parameters associated with a UHF emission map derived by Berge from 10.4 cm observations and reported by Roberts (ref. 65). Emission maps have since been improved, and Branson's smooth, single peak model, derived directly from his own 21 cm observations (ref. 46), is more reasonable.

McAdam (ref. 47) and Gulkis (ref. 48) report 75 cm wavelength observations which support the existence of a distribution extending to a greater distance from Jupiter than $L = 3$ or possibly additional flux peaks there. McAdam's results, obtained with a Mills cross, are of too low a resolution (more than 2 radii) to warrant acceptance of his conclusion that a significant fraction of the radiation is generated near $L = 6$ (refs. 18 and 48). Gulkis' results (ref. 48), obtained from the observation of a lunar occultation of Jupiter with a single 140-foot radio telescope, suggest a more modest breadth, i.e., $L = 4$ or 5. On this basis it is reasonable to conclude that the uncertainties in the flux levels at $L = 3$ or beyond should be taken as sufficiently larger than those near $L = 2$ to permit the possibility of broad trapped radiation belts.

Warwick (ref. 18) has noted that if conservation of the magnetic moment of the electrons at the peak of the radiation belt is applied in an extrapolation to the magnetic field strengths expected in the solar wind near Jupiter, energies are predicted that are reasonable for electrons which share some of the proton energy as they are trapped by Jupiter's magnetic field from the solar wind. Thus, he suggested L-shell diffusion with magnetic moment conservation as a possible mechanism for the population of Jupiter's magnetosphere with electrons from the solar wind. The non-relativistic treatment of Davis and Chang (ref. 66) for the diffusion has been modified by Davis* to include correct relativistic relations among speed, momentum and energy. The resulting equation for the flux dependence on position in the diffusion region is

$$\Phi = \phi_1 \left( \frac{L_{\text{L}}}{L} \right)^6 \left( \frac{E_1}{m_0 c^2} + 1 \right)^2 \cdot \left\{ \left[ \left( \frac{E_1}{m_0 c^2} + 1 \right)^2 - 1 \right] \frac{L_{\text{L}}^3}{L^3} + 1 \right\}^{1/2}$$

where the flux and energy are $\Phi_1$ and $E_1$ respectively at $L = L_1$. Additionally, for small $L$ values the synchrotron radiation may be responsible for particle loss of energy. Thus, as a schematic representation of the $L$-dependence of the electron flux in Jupiter’s belts, this monograph adopts the values quoted at the end of the section 2.7.4.1a for $L \leq 2$ and the dependence specified by equation (21) for $2 < L < 50$ where $E_1$ and $\Phi_1$ are evaluated at $L_1 = 2$. An uncertainty factor of 3 is applied to the flux $\Phi$ everywhere. These results are consistent with the conclusions reached at the Jupiter Radiation Belt Workshop.** At any given time and location, zero flux is an appropriate lower limit because local particle loss mechanisms could be strong.

c. Distribution with Latitude

Because the radio observations display poorer angular resolutions parallel to Jupiter’s magnetic dipole than perpendicular to it, they provide few details of the electron distribution with latitude. The 10-cm map (fig. 2, sec. 2.5.3.1) of Berge (ref. 24) suggests that the UHF source extends perhaps 1 $R_J$ above and below the magnetic equator (compared to 2 $R_J$ on either side of the dipole along the equatorial plane). This suggests that the characteristic latitude extent of the belts may be of the order of $26^\circ$ ($\tan^{-1} 0.5$) with higher values to $45^\circ$ not excluded. This monograph adopts for the latitude distribution a form proportional to $\exp\left[-\left(\phi/\phi_0\right)^2\right]$ where $\phi_0 \approx 30^\circ$.

d. Directionality

The particle fluxes are probably greatest in directions perpendicular to the magnetic field lines, corresponding to large pitch angles near $90^\circ$, but this is likely to affect vehicle design in the same way as an isotropic flux distribution which accordingly is adopted here. Fluxes are calculated by multiplying the concentration and the corresponding speed; for electrons the speed of light $c$ is appropriate, whereas the protons are subrelativistic and their speed depends on their energy. Electron concentrations and fluxes are summarized by the formulas in table XII (sec. 3.7) in which the specified uncertainty factors are large enough to bound likely unknown variations in position and time.

2.7.4.2 Relativistic Electron Energy

According to the synchrotron theory, a single relativistic electron of kinetic energy $E_1$ and pitch angle $\theta$ in the presence of a magnetic field of strength $B$ emits UHF radiation within the wide frequency bandwidth, $W = (40 \text{ MHz}/\text{MeV}^2 \text{Gauss})E_1^2 B \sin \theta$, at frequencies near and below $W$ itself (ref. 18). However Barber and Gower (ref. 67) quote a constant 2.5 times smaller. To estimate $E_1$, it is appropriate to select the values $\theta = 90^\circ$, $B = 2.0$ gauss at $L = 1.8$ from the dipole moment (sec. 2.3.1), and $W = 3000 \text{ MHz}$ from the UHF observations which extend to wavelengths shorter than 10 cm; the resulting value is $E_1 = 6.2$ MeV (ref. 18). The reasonable uncertainties of factors 2 in $B$, 2.5 in the constant, and 2 in $W$ lead to an uncertainty of a factor 3 in $E_1$. The value and uncertainty thus derived are adopted for the characteristic energy for the relativistic electrons in the flux peak near $L = 1.8$. The uncertainty includes many of the energy values suggested by several authors on the basis of the UHF data, 14 or 15 MeV (refs. 17 and 68), 10 MeV (refs. 48 and 67), and 1 to 30 MeV (ref. 46).

**Held 13-15 July 1971 at Jet Propulsion Laboratory, Pasadena, California.
Because the frequency distribution of the radiation emitted by a single electron is so broad, it would be possible to reproduce many of the characteristics of the observed UHF data from monoenergetic electrons (say 6.2 MeV) trapped on a single set of field lines (say at \( L = 1.8 \)). Such sharp distributions seldom, of course, occur in nature, and the L-shell distribution has already been discussed (sec. 2.7.4.1). In order to compare the energy distributions suggested in the literature, figure 6 shows the ratio of \( N_E \) (the number of electrons with energies greater than \( E \)) to \( N_5 \) (the number with energies greater than 5 MeV, i.e., near the characteristic energy) at the location of the peak flux in each of several models, as a function of electron kinetic energy \( E \). The order-of-magnitude energy distributions of Warwick (ref. 18) and Carr and Gulkis (ref. 17) are monoenergetic and are not intended to be fully realistic descriptions. Because it has been demonstrated on the basis of the theory (refs. 69 and 70) that a distribution in which \( dN \) is proportional to \( E^{-1} dE \) over a wide but limited energy range leads to a flat, i.e., constant flux density, emission spectrum, many authors have adopted such a distribution. Some of these results (refs. 62, 64, and 46) are shown in figure 6. Hauffner (ref. 63) has adopted an exponential energy distribution in which \( dN \) is proportional to \( e^{-E/E_0} dE \) on the basis of his understanding of the energy distributions in the Earth's belts, extending to zero energy. Gulkis (ref. 48) discusses the energy spectra carefully and concludes that the failure of the observed intensity to increase at long wavelengths requires that the low-energy cutoff in the \( E^{-1} \) distribution occur at an energy smaller than, but comparable to, the characteristic energy. The conclusion, supported by reference 18, to be drawn from this survey is that the number of relativistic electrons at energies below perhaps 5 MeV must be small compared to the total, at least in those regions of greatest field strength where the observed UHF emission is generated. Such a requirement is not in agreement with the low-energy portions of the Hauffner (ref. 63) and Koepp-Baker (ref. 64) distributions shown in figure 6. It is also an apparent consensus from references 46 and 48 that the distribution does not contain a significant fraction of electrons at energies greater than about 30 MeV.

The concept of a nearly monoenergetic distribution is supported by the theoretical possibility of an energy-selective acceleration mechanism, whereas that of a broad distribution, extending particularly to low energies, is supported by the likelihood of energy loss mechanisms which could smear out energy peaks. Although these theoretical arguments are inconclusive, the data (ref. 48) favor the adoption of a distribution which does not permit a significant fraction of the electrons to have energies more than a factor two or three different from the characteristic energy for the distribution.

To satisfy the foregoing suggestions, a new electron energy distribution is proposed here as an artificial but reasonable compromise between monoenergetic and exponential energy distributions. The flux \( \Phi_E \) of electrons with energy greater than \( E \) and the differential flux \( d\Phi \) are given for all relativistic values of \( E \) by

\[
\Phi_E = \Phi_0 \left( 1 + \frac{E}{E_0} \right) \exp \left( - \frac{E}{E_0} \right)
\]

\[
d\Phi = -\Phi \frac{E}{E_0^2} \exp \left( - \frac{E}{E_0} \right) dE
\]

\[22\]
Here $\Phi_0$ is the flux parameter*, and $E_0$ is the local characteristic energy. The integrated number ratios corresponding to the distribution adopted (equation 22) have also been plotted in figure 6 for $E_0=6.2$ MeV. The adopted distribution has only 19 percent of its electrons at energies $E<5$ MeV and only 5 percent at energies $E>30$ MeV. The uncertainty in the flux at a given energy can be specified as part of the uncertainty in $\Phi_0$ or $E_0$ without the additional specification of an uncertainty in the foregoing relative energy distribution functions.

The UHF data do not contain sufficient information from which the variation with position in the belts of the characteristic energy $E_0$ might be derived. Gulkis (ref. 48) claims that the reality of a broad emission source at long wavelengths perhaps extending to $L=3$ at 75 cm, compared to $L=2$ at 21 cm (sec. 2.7.4.1), would require greater fluxes of higher energy particles at greater distances from the dipole. He views this as unlikely on intuitive grounds, and Warwick (ref. 18) claims as well that the reality of such outer belts is not required because the supporting data have been obtained at marginal resolution. The theoretical treatment which has been applied (sec. 2.7.4.1) to the variation of the flux

![Figure 6](image)

Figure 6.—Ratio of the number of relativistic electrons with kinetic energy $\geq E$ to the number with energy $\geq 5$ MeV for proposed trapped radiation belt models, evaluated at their innermost flux peaks.

*See Symbols (app. A).
with position requires that the energy dependence reflect the conservation of the magnetic moment. Simple, relativistic considerations* lead to the relationship

\[ E_0 = m_0 c^2 \left[ \left( \frac{E_1}{m_0 c^2} + 1 \right)^2 - 1 \right] \frac{L_1^3}{L^3} + 1 \]  

(23)

where \( E_1 \) may be interpreted as the local characteristic energy at \( L = L_1 \) and \( E_0 \) is the local characteristic energy elsewhere. For \( E_0 \ll m_0 c^2 \) equation (23) implies that \( E_0 L^3 \) is approximately constant (independent of \( L \)). In the Earth’s belts, the variation of energy is not so steep, with \( L \) exponents of 2 and 1.36 being reported as compatible with the data (refs. 18 and 63). For Jupiter, the synchrotron radiation reduces the energy at small \( L \) values. Thus, this monograph adopts a nominal \( E_0 = 6.2 \text{ MeV} \) for \( L \leq 2 \) and the nominal energy elsewhere is joined to this value at \( L_1 = 2 \) by equation (23) for \( 2 \leq L \leq 50 \). An uncertainty factor of 3 yields an appropriate upper limit for \( E_0 \) at \( L \leq 2 \), and the upper limit energy elsewhere is joined to the value \( E_0 = 19 \text{ MeV} \) at \( L_1 = 2 \) by equation (23).

**2.7.4.3 Energetic Protons**

There are no positive Jupiter data from which proton fluxes in Jupiter’s radiation belts can be inferred. Energy-dependent upper limits can be set on the basis of trapping by the magnetic field, but it is not considered likely that the proton population approaches these comparatively extreme upper limits. The three published studies which derive peak proton fluxes on the basis of Earth analogy do so in very different ways. The resulting characteristic energies diverge widely for the innermost belt, 0.1 to 4 MeV (ref. 62), 9.4 GeV (ref. 63), and \( 10^3 \text{ MeV} \) (ref. 64). Various unpublished studies** apply cosmic ray albedo neutron decay theory to Jupiter and also derive widely divergent parameters for the protons.

This monograph adopts the conclusions** of several knowledgeable scientists that an \( L \)-shell diffusion model is appropriate for a nominal description of the protons, such that equations (21) and (23) describe the flux and energy for the entire range \( 1 \leq L \leq 50 \). This assumes that at \( L_1 = 2 \) the flux is \( \Phi_1 = 7 \times 10^6 \text{ cm}^{-2} \text{ sec}^{-1} \) (corresponding to a number density equal to that of the electrons) and the characteristic energy is \( E_1 = 60 \text{ MeV} \) (equal to ten times that of the electrons). Because the uncertainties are so large, a considerably more severe upper limit model is adopted in which the fluxes are limited at intermediate \( L \) values by ion cyclotron instabilities.** The characteristic energy is given by equation (23) where \( E_1 = 100 \text{ MeV} \) at \( L_1 = 2 \), but the dependence of the flux on \( L \) is the complicated function shown in figure 9 and table XII (sec. 3.7.3). Because local proton loss mechanisms such as satellite sweeping could be strong, an appropriate lower limit is zero flux.

**2.7.4.4 Low Energy Particles**

The arguments suggested in the foregoing sections do not apply at energies below about 1 MeV, and few observations have been so interpreted. This situation contrasts with the Earth’s radiation belts in which particle energies in the range 1 keV to 1 MeV are common.

---

**Jupiter Radiation Belt Workshop, held 13-15 July 1971 at the Jet Propulsion Laboratory, Pasadena, Calif.
Warwick (ref. 18) and Goldreich and Lynden-Bell (ref. 21) require electrons of several keV energy in their HF burst generation mechanisms, but such particles would exist just above Jupiter's atmosphere and their fluxes are not intended to apply to the trapped radiation belts. Thus all that can be said about low energy particles is that their concentrations should not exceed the total magnetospheric plasma density (sec. 2.7.5) or the maximum which could be trapped in, and share the rotation of, Jupiter's magnetic field. The latter energy- and L-dependent limit can be simply derived for protons, and, assuming equal concentrations, the following limit also applies to electrons

\[ N_E < \frac{(M_1/R_J^3)^2}{8\pi L^6 \left[ E + (m_p/2)(LR_J \omega_0)^2 \right]} \]  

(24)

If the nominal values for \( M_1 \) and \( R_J \) (secs. 2.3.1 and 2.1.2) are substituted in this expression, the numerical limit

\[ N_E < \frac{3.6 \times 10^{12}}{L^6(E + 3.3L^2)} \]  

(25)

is obtained for \( N_E \) in \( \text{cm}^{-3} \) and \( E \) in eV.

**2.7.5 Magnetospheric Plasma**

Theoretical discussions of the distribution of thermal ions and electrons in Jupiter's magnetosphere are complicated by the strong influences of magnetic, gravitational, and centrifugal forces as well as by uncertainties in the plasma temperature and source mechanism. Melrose (ref. 30) and Carr and Gulkis (ref. 17) present thorough mathematical and qualitative discussions of many of the foregoing considerations, and the importance of possible instability, fragmentation, and ring current formation is stressed. The present, inconclusive state of the art permits a range in electron number density of several orders of magnitude and leaves open whether the plasmasphere is cool and close to the planet or is warm and populates the magnetosphere. Consequently, very broad limits, representing extremes of values suggested by the literature are adopted herein. The plasma temperature ranges from 150°K (typical of ionospheric temperatures, section 2.7.6) to 10^5°K (refs. 17 and 71). The plasma electron density ranges from 0.1 cm\(^{-3}\) (credited to Brice by reference 17) to 10^7 cm\(^{-3}\) (a high ionosphere electron density estimate). For \( L > 4.2 \), the upper limit electron density which can be contained by the magnetic field is given by \( 10^{12}/L^8 \) cm\(^{-3}\) (eq. 25 and ref. 17). The plasma corotates with Jupiter's magnetic field at least as far out as \( L = 7 \) or 8 (ref. 30).

There are few discussions based on observations. Warwick (ref. 18) argues that the total Faraday rotation experienced by an HF wave propagating from its source near \( R = R_J \) to
the observer on Earth is less than one rotation. This reasoning requires that, if the electron (or proton) density $N$ is either uniform or decreases with increasing distance $R$ from Jupiter, the upper limit

$$N < \frac{(40 \text{ cm}^{-3})}{(1 - R_j^2/R^2)}$$

must be satisfied everywhere. The reality of the mode coupling which is basic in deriving this result has been challenged (ref. 17); and until this difference is settled, the foregoing cannot be adopted as a firm upper limit.

2.7.6 Ionosphere

There are two approaches to the description of Jupiter's ionosphere. The first is theoretical in which considerations such as photoionization and solar UV flux values (useful in understanding the Earth's ionosphere) are applied to Jupiter. The other approach is based on inferences from the HF bursts to which various generation mechanisms and requirements are applied. Within the rather large uncertainties inherent in the two approaches, agreement is quite satisfactory.

The theoretical approach is exemplified by the work of Gross and Rasool (ref. 72). Their results are supported by those derived by Shimizu (ref. 73), Hunten (ref. 74), Henry and McElroy (ref. 75), and Capone and Barrow (ref. 76) which consider similar processes in less detail. In reference 72, solar ultraviolet radiation at wavelengths of the order of 1000 Å is responsible for the photoionization of molecular and atomic hydrogen down to those altitudes at which it is fully absorbed. Secondary reactions, e.g., recombination, occur such that the equilibrium abundances of the species $H_2$, He, H, $H^+$, and $e^-$ (listed roughly in order of decreasing concentration) are significant near the maximum of charged particle density, approximately $10^6$ electrons/cm$^3$ and $10^6$ protons/cm$^3$. The local gas pressure at that maximum is approximately 3 dyn/cm$^2$; atmospheric models suggest that this level occurs within the altitude range $z_0 = 250 \pm 150$ km (sec. 2.9.4). Below this altitude the electron density becomes negligible within several kilometers, whereas upward the electron density decreases roughly exponentially with a scale height approximately 100 km. According to the same authors, the ionospheric temperature ranges from 100 to 200°K.

Numerous arguments from mechanisms of HF burst generation have yielded electron density values independent of the foregoing considerations. Warwick (ref. 18) suggests that a maximum electron number density of the order of $10^7$ cm$^{-3}$ may be required to cause reflection of the HF bursts at the ionosphere. Investigations by Ellis (ref. 20) and Papadopoulos and Lerche (ref. 71) use various HF burst mechanisms to derive values within one or two orders of magnitude of the preceding density.

The presence in the Earth's ionosphere of several distinct layers of electron and ion concentration suggests that a one-layer model is probably too simplistic for Jupiter. For this reason and on the basis of the foregoing derived values, it is appropriate to adopt a peak
number density of $10^6 \text{ cm}^{-3}$ for electrons and protons and to estimate the uncertainty as an order of magnitude in either direction.

2.8 Transfer Properties in the Magnetosphere and Ionosphere

The presence of magnetic fields and charged particles in the magnetosphere and ionosphere may lead to values for the electrical and thermal conductivity, opacity, and index of refraction different than those for a vacuum or a neutral gas. In most cases, these parameters must be evaluated by applying appropriate theory to the charged particle environment given in section 2.7 because there are scant data for determining them directly.

2.8.1 Electrical and Thermal Conductivity

The discussion by Hess and Mead (ref. 77) of the electrical conductivity of an ionized gas in a magnetic field can be applied to Jupiter's magnetosphere and ionosphere. The fact that the electron gyro-frequency is comparable to or greater than the collision frequency leads to strongly anisotropic electrical conductivity. In directions parallel to the magnetic field lines, current flow is not affected, and direct conductivity is appropriate. However, if the current is perpendicular to the field lines, it is inhibited by the inability of the electrons to move far from their original field lines and so the Pederson and Hall conductivities apply which are much smaller than the direct conductivity.

In the magnetosphere, the medium (secs. 2.3.2 and 2.7.5) is a highly ionized, low density hydrogen plasma in which Coulomb collisions among the ions and electrons determine the collision frequency. For this case Alfvén and Fälthammar (ref. 78) have graphed the value of the direct conductivity. It is nearly independent of the electron density but depends strongly on the temperature, approximately as $(1.35 \times 10^{-3} \text{ mho/meter})T^{3/2}$ with $T$ in °K. Thus, the temperature uncertainty in section 2.7.5 leads to the range of values $2.5$ to $4.3 \times 10^4 \text{ mho/meter}$ for the direct conductivity. The formulae given by Hess and Mead (ref. 77) lead to Pederson and Hall conductivities which are less than $10^{-6}$ times the direct.

In the ionosphere, the medium (sec. 2.7.6) is a weakly ionized, low density plasma in which collisions with neutral particles determine the collision frequency. In this case the direct conductivity depends weakly on temperature but strongly on electron and neutral particle densities. The formulae of Hess and Mead (ref. 77) lead to the range of values $10^{-2.2}$ mho/meter for the direct conductivity at the peak of Jupiter's ionosphere. Their formulae also indicate that the Pederson and Hall conductivities are less than $10^{-2}$ times the direct. Because the concentration of neutral particles falls off rapidly with increasing altitude, the conductivities increase above the electron density peak.

For design purposes, it does not seem worthwhile to consider the foregoing complexities because it may not be possible to ascertain electron densities and temperatures and spacecraft component orientations in advance. Thus, for this monograph the range for electrical conductivity of $0$ to $4.3 \times 10^4 \text{ mho/meter}$ for altitudes $z > 400 \text{ km}$ is adopted for the magnetosphere; for the ionosphere the range of $0$ to $1 \text{ mho/meter}$ is adopted for altitudes $100 < z < 400 \text{ km}$. 

29
The corresponding thermal conductivities are so small that conduction is negligible by comparison with radiative heat transport processes.

### 2.8.2 Opacity and Index of Refraction

The formulas and discussions provided by Ratcliffe (ref. 79) apply to the evaluation of the opacity and index of refraction. These quantities are related to the local electron gyro-frequency, the local plasma frequency, and the local electron-neutral particle collision frequency. For all electromagnetic radiation at wavelengths shorter than one meter, the wave frequency \( \nu > 300 \text{ Hz} \) is much greater than the frequencies given in section 2.8.1 for the ionosphere and magnetosphere. Therefore, the expressions given by Ratcliffe (1959) for the index of refraction \( n \) reduce to

\[
\frac{n^2}{\nu^2} = \frac{1}{\nu^2} \left( \nu \pm \nu_B \cos \theta \right).
\]  

In equation (27), \( \theta \) is the angle between the direction of propagation and the magnetic field \( B \), \( \nu_B \) is the electron gyro-frequency (given by \( \nu_B = (2.8 \text{ MHz})B \) for \( B \) in gauss), and \( \nu_p \) is the plasma frequency (given by \( \nu_p = (9 \text{ kHz})N_e^{1/2} \) where \( N_e \) is the electron concentration in \( \text{cm}^{-3} \)).

For \( N_e \leq 10^7 \text{ cm}^{-3} \), as given in section 2.7, the range of values for \( n \) is

\[
1 \geq n \geq 1 - (4.5 \times 10^{-7})\lambda^2
\]  

for wavelengths \( \lambda \) in cm. The corresponding opacity and optical thickness are negligible; this conclusion is supported by the observation that the thermal radiation from the atmosphere passes through these regions apparently without attenuation. For lower frequency radio waves (wavelengths longer than one meter), the foregoing frequencies, \( \nu_B \) and \( \nu_p \), can be comparable to the wave frequency. Under such circumstances, there is some expectation that complex refraction, reflection, absorption, and Faraday rotation processes will occur in the ionosphere and magnetosphere. There are insufficient data to predict these processes with confidence, however.

### 2.9 Atmospheric Structure

The variable image of the disk of Jupiter, as observed telescopically in visible light, strongly indicates an active atmosphere of sufficient thickness to obscure any features of a solid planetary surface. Current understanding of the composition, temperature, pressure, density, and other significant physical quantities within this atmosphere and a set of numerical models are given in the following sections.

#### 2.9.1 Composition

Recent papers by Owen (ref. 80) and McElroy (ref. 81) summarize the latest abundance results from visible and infrared spectra in which absorption lines of \( \text{H}_2 \), \( \text{CH}_4 \) and \( \text{NH}_3 \) can be identified as originating in the atmosphere of Jupiter. They conclude that the abundance
ratios are consistent with those of a solar mixture of elements which has evolved into simple hydrogen-bearing molecules with modification by saturation and condensation of some species, particularly NH\textsubscript{3} and H\textsubscript{2}O. This conclusion is adopted herein although it must be stressed that absorption lines of He, a presumably important constituent, and minor constituents such as H\textsubscript{2}O and Ne have not been observed in the spectra. In view of the uncertainty thereby introduced and the uncertain solar ratios of elements (particularly He), it is appropriate to consider the fractions by mass of all molecules other than H\textsubscript{2} uncertain by a factor of two in either direction. Table III shows the nominal and limiting compositions adopted for the atmosphere which is assumed homogeneous below the clouds. The nominal values of the H and He abundances are consistent with the results of Torres-Peimbert et al. (ref. 82), whereas the ratios of H to C, N, O, and Ne are based on Lambert (ref. 83).

### 2.9.2 Lower Atmosphere

The lower atmosphere is thought to be in hydrostatic equilibrium and convective to a very great depth. The variable appearance of the planet in visible light, analogous to weather patterns on the Earth (ref. 84), is attributed to large-scale cloud features which are condensates formed in upward convecting gas. Among the numerous references treating the clouds, the most complete in describing physical processes is that of Lewis (ref. 85) who concludes that the three most significant condensates in upper regions of the lower atmosphere are ammonia ice, NH\textsubscript{4}SH, and water. The water may be ice or liquid with NH\textsubscript{3} in solution. Even in the presence of such condensates, the lapse rate is nearly equal to the “dry”, i.e., condensation free, adiabatic value appropriate to the gas composition adopted.

In addition, a pressure-temperature correspondence at some point is required for a model. Attempts to establish such a correspondence from measurements of spectral line widths, equivalent widths, and their ratios are available but open to serious questions. In particular, “cloud top reflection layer” and ammonia saturation concepts (ref. 84) are not considered satisfactory; and on-going discussions of rotational temperatures and the associated abundance (refs. 86, 87, and 88) are incomplete. An alternate approach by Gillett et al. (ref. 41) combines the 125°K temperature measured near 12\textmu m with the calculated opacity of the dominant absorber H\textsubscript{2} at that wavelength to give an abundance of 12 km-atm H\textsubscript{2} * and a partial pressure of 1/4 atm H\textsubscript{2}. This correspondence and the foregoing composition and lapse rate lead to a reasonable lower atmosphere model, similar to that of Owen (ref. 89), which explains many of the observed features of the reflection and emission spectra.

Observation of infrared (thermal) radiation emitted in the lower atmosphere at wavelengths longer than about five \textmu m leads to two important conclusions: (1) that the disk is uniform within about 20 percent in brightness between 8 and 14 \textmu m (ref. 44), and (2) that the planet radiates considerably more energy than it receives from the Sun (ref. 43). These conclusions suggest that lower atmospheric parameters need not be specifically associated

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*This means that the H\textsubscript{2} on Jupiter above the level at which the radiation is emitted would form a layer 12 km thick if it were of uniform density at standard temperature and pressure (0° C and 1 atm) and no other gases were present.
with planetary latitudes and time of day but that these variations should be included within
the uncertainty specified by a single set of models.

The most comprehensive survey (ref. 43) of Jupiter's infrared emission suggests an effective
temperature near 134°K. Following Owen (ref. 89), it is reasonable to take the upper
boundary of the lower atmosphere as the level at which the temperature reaches \((0.5)^{1/4}(134°K) = 113°K\) and to overlay that level with an isothermal stratosphere extending
upward at least one scale height.

**2.9.3 Upper Atmosphere**

For upper atmosphere strata (above the regions of significant radiation reflection, absorp-
tion, and emission), observational data are meager. Wrixon (ref. 90) attributes narrow emis-
sion lines in the \(\text{NH}_3\) absorption feature near 1.25 cm wavelength to an atmospheric
inversion layer near 145°K temperature and 6.5 mb pressure. The most thorough theoretical
analyses (refs. 72, 74, 75, and 91) rely heavily on calculations of equilibrium among absorp-
tion, photodissociation, photoionization, diffusion, and recombination of \(\text{H}_2\), \(\text{He}\), \(\text{CH}_4\), \(\text{H}\),
\(\text{H}^+\), and \(e^-\). The results have some quantitative uncertainty because the rates are poorly
known and qualitative aspects have been emphasized. The major conclusions are (1) that the
temperatures are probably near 150°K (ref. 92) but because of possible effects of the solar
cycle may range as high as 500°K (ref. 49) and (2) that molecular hydrogen predominates
up to heights of the order of 500 km where pressures are less than \(10^{-6}\) atm.

**2.9.4 Model Atmospheres**

Several numerical models of Jupiter’s atmosphere have been presented in the literature;
those of Kuiper (ref. 84), Gross and Rasool (ref. 72), Trafton (ref. 93), and Lewis (ref. 85)
emphasize distinct, important aspects of the physical processes involved. Except for that of
Smith and Weidner (ref. 94), based on the work of Gross and Rasool (ref. 72), none of the
available models completely tabulates the basic parameters of height, temperature, pressure,
and density. Besides, the models do not incorporate some of the recent observational
information cited in the foregoing sections. Therefore, it has seemed appropriate to derive a
new set of models. These models assume hydrostatic equilibrium and the perfect gas law and
follow smoothly varying adiabatic lapse rates in the troposphere (app. D). Local adjustments
of the troposphere lapse rate for the partial replacement of convective transport by radiative
equilibrium and for the inclusion of “wet” saturation and cloud condensation effects have
not been made because these additional complications were considered unwarranted in view
of other large uncertainties.

The nominal model has been constructed with the constant values of \(T_e = 134°K\),
\(g = 2500\ \text{cm/sec}^2\), and \(u = 2.30\) which correspond to the “nominal” composition and
other quantities shown in table III. The convective troposphere has an adiabatic lapse rate
near \(-2°K/km\), extends indefinitely downward from the tropopause boundary temperature
of 113°K, and includes the correspondence level established at \(P = 0.30\ \text{atm}\) and
\(T = 125°K\). Above the tropopause an isothermal stratosphere extends upward one scale
height (16.3 km). Above the stratosphere an inversion layer of constant \((P/T)(dT/dP)\) is
limited by an uppermost level at which $P = 6.5 \text{ mb}$ and $T = 145^\circ \text{K}$. An isothermal region extends indefinitely upward from the latter level.

The limiting models are "cool" and "warm" in the sense that for a given pressure of the nominal model they provide extremes of temperature and density which are thought to bracket the range of possible Jupiter values. The values for model construction were selected as reasonable limits to the nominal values and are given in table III. The numerical values calculated for these models are based on the atmospheric relationships described in appendix D, and are given in tables XIV, XV, and XVI for the pressure range between $2 \times 10^{-7}$ and $10^3$ atmospheres beyond which various assumptions become insupportable. By analogy with the Earth, zero altitude is set at the level at which the pressure $P$ equals one atmosphere in all three models. This level is near the ammonia clouds (sec. 2.10) and corresponds to the optical surface or limb of the planet. The profiles of pressure, temperature, density, and altitude are shown in figures 10 through 14 (sec. 3.9).

### 2.10 Cloud Properties

#### 2.10.1 Observations

Monitoring of the planet Jupiter visually and photographically has resulted in the accumulation of a great deal of detailed information about the visible markings. Representative extended and brief summaries of such information appear in Peek (ref. 95) and Michaux (ref. 9), respectively. The changeable positions, sizes, shapes, and colors of both the short-lived and permanent markings have led to the conclusion that the visible atmospheric layers are affected by both winds and clouds.

Beyond approximately $\pm 50^\circ$ latitude are the two polar regions which are dark and featureless at the resolution and wavelengths observed from the Earth. At lower latitudes the planet has stripes parallel to the equator, classified as zones and belts. As usually observed from the Earth in visible light, the zones are bright and the belts dark. Their latitudes and relative and absolute brightnesses vary with longitude, time, and wavelength of observation, i.e., color effects are important. The semi-permanent, well documented, generally uniform belts and zones occasionally are divided by thin bands also parallel to the equator and sometimes are blemished by light or dark spots whose motions in latitude and longitude have provided information on planetary rotation and atmospheric currents (sec. 2.11). These spots have characteristic horizontal dimensions ranging from 1000 to 10 000 km and lifetimes ranging from a few hours to weeks. Their resolution is marginal from Earth observations so that very little is known about their actual properties. The spots are comparable in size and possibly in nature to major weather systems in the Earth's atmosphere.

One much larger feature has existed for at least three centuries, namely the Great Red Spot. In a presentation of recent observational results, Solberg (refs. 96 and 97) reports that the Red Spot usually contrasts sharply with its surroundings in light of wavelengths $\lambda < 5000 \text{ Å}$ (very dark). Its oval shape extends horizontally about 13 000 km in latitude and 40 000 km in longitude, and its area varies by about 50 percent. The latitude of its center is usually within $3^\circ$ of $22^\circ$ South. Its longitude oscillates with an amplitude near $0^\circ 8$
TABLE III
COMPOSITIONS AND OTHER PARAMETERS
FOR MODEL ATMOSPHERES OF JUPITER*

<table>
<thead>
<tr>
<th>PARAMETER</th>
<th>COOL MODEL</th>
<th>NOMINAL MODEL</th>
<th>WARM MODEL</th>
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</thead>
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<tr>
<td>Fractions by mass (or weight)</td>
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<td></td>
</tr>
<tr>
<td>$H_2$</td>
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<tr>
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<tr>
<td>Inversion level temperature (°K)</td>
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<td>500</td>
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<tr>
<td>Inversion level pressure (atm)</td>
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<td>$2 \times 10^{-7}$</td>
</tr>
</tbody>
</table>

*The zero of altitude, at the distance $R_e$ from the center (eq. 1), is taken as having a pressure of 1 atm in all three models.
and a period near 90 days, and drifts with respect to longitude System II* at a rate of several degrees/year. The Red Spot’s nature and how it retains its integrity are enigmatic although Hide (ref. 98) has developed the Taylor column hypothesis in which relatively uniform, stagnant fluid overlies some topographical feature in the rotating fluid which forms Jupiter’s deep atmosphere.

2.10.2 Theories

Theories of cloud formation and structure have been applied to the vertical distribution of condensates in Jupiter’s lower atmosphere (sec. 2.9.2). Techniques based on the analysis of Lewis (ref. 85) have been applied to the models given in section 2.9.4 with numerical data for the condensates of ammonia and water from the International Critical Tables (ref. 99) and Lange (ref. 100). Some of the cloud properties (calculated in accordance with app. D) are given in tables XIV, XV, and XVI and figures 10, 11, 13 and 14. Cloud altitudes, masses, and compositions, primarily H₂O and NH₃ ices with a dilute NH₃-H₂O solution replacing H₂O in the cool model, show broad ranges which may be exceeded in the real atmosphere for any of the following reasons: (1) additional condensates, particularly NH₄SH and NH₄Cl, may form clouds of their own or contaminate the ones considered, leading to non-uniform cloud structure and color (ref. 85); (2) convection may carry the condensates to other altitudes; and (3) in downward convecting regions (cooler than upward-moving ones at a given pressure), condensates are less likely to be present so there is a reasonable basis for equating downward convecting regions with Jupiter’s dark belts. Lastly, by Earth analogy, it would not be surprising if storms were occasionally present in regions of high cloud activity, leading to rain, hail, snow, thunder and lightning within the lower atmosphere.

2.11 Atmospheric Motions

The properties of horizontal motions in the troposphere of Jupiter may be derived directly from the variable apparent positions of well-observed spots at latitudes up to 50°. Vertical motions, motions at polar latitudes, motions above the tropopause, and motions on distance scales smaller than about 1000 km can only be inferred from assumptions and theories consistent with the data. Roughly the motions may be classified as rotation and winds.

2.11.1 Rotation

The rotation rate derived from the radio observations is probably the most physically significant of the rates which have so far been deduced. Recent determinations of the rotation rate have been made from the periodicities in the temporal variations of both the UHF data (ref. 26) and the HF burst data (ref. 101). The resulting rotation periods agree within the associated uncertainties given in the foregoing two references; that from reference 101, \( T_0 = 9^h 55^m 29^s 73 \pm 0^s 04 \), is adopted here. Because the rotation rate is that of the planetary dipole magnetic field, it is probably also that of Jupiter’s deep interior.

*Rotational coordinate system, defined in Glossary (app. E).
where the field presumably is generated. This was recognized early in the history of the radio observations, and the then current value of the radio period $T_3 = 9^h 55^m 29^s 37$ was adopted in the definition of longitude in System III (1957.0)* to which Jupiter's radio events are often referred (ref. 9).

Rotation periods derived from the observed positions of spots and other visible features on the planetary disk have a range of several minutes although the uncertainty of the period of any given spot observed for several rotations may be near ±1 second. Averages over such periods resulted in the adoption of two distinct periods, $T_1 = 9^h 50^m 30^s 003$ for latitudes within approximately ±10° of the equator and $T_2 = 9^h 55^m 40^s 632$ for all other latitudes. The longitude specifications for features observable on the planetary disk employ Systems I and II* which correspond to the foregoing periods. Reference 9 provides details and further references. The foregoing derivations of several values for the rotation period constitute substantial evidence for actual variation in rotation rates.

2.11.2 Winds

In accord with Chapman's concept of the horizontal surface motions on Jupiter (ref. 102), it is appropriate to define winds as any motions differing from solid body planetary rotation (assumed to have a period equal to that for Jupiter's radio sources). At latitudes within 50° of the equator, visual and photographic observations of the positions of spots and features on the Jupiter disk imply horizontal winds on scales larger than 1000 km that are zonal, i.e., parallel to the equator, and apparently unceasing. According to Chapman (ref. 102), the major winds can be described by (1) a westerly equatorial current between latitudes ±10° of average speed near 100 m/sec, corresponding to the fast System I* rotation period $T_1$, with local winds up to 10 m/sec in any direction with respect to the prevailing wind, (2) local winds between latitudes 10° and 50°N up to 10 m/sec in any direction with a frequent westerly current (the north temperature current) of speeds up to 100 m/sec within 10° of latitude 25°N, (3) local winds between latitudes 10° and 50°S up to 20 m/sec of which the most prominent and persistent (the circulating current) has an easterly component near 20°S and a westerly component near 25°S, and (4) a wind having counter-clockwise vorticity near the edges of the Great Red Spot of speeds near 10 m/sec. Additional details of the winds may be inferred from the observational analysis of Peek (ref. 95) and the theoretical analyses of Hide (ref. 103) and Ingersoll and Cuzzi (ref. 104). Winds have not been observed on a smaller scale, at latitudes beyond ±50°, or moving in vertical directions; but a reasonable assumption is that for such winds the characteristic speed of 10 m/sec holds. The limits in height and depth to which these winds might extend are unknown, but a reasonable assumption is that the horizontal winds may occur everywhere in the atmosphere, whereas the vertical ones are confined to the troposphere (sec. 2.10). The literature contains no estimate of the magnitude of wind shear.

2.12 Transfer Properties in the Atmosphere

Consideration of the atmospheric models in section 2.9.4 in conjunction with theoretical and empirical gas properties leads to values of transfer properties for which there are no

*Rotational coordinate system, defined in Glossary (app. E).
supporting observations. The following sections separately treat the various transfer properties.

2.12.1 Electromagnetic

At optical wavelengths, \( \lambda < 1 \mu \text{m} \), the major sources of opacity are (1) Rayleigh scattering by molecular hydrogen (ref. 7), (2) absorption and scattering by cloud particles (ref. 105), and (3) molecular absorption bands of \( \text{CH}_4 \) and \( \text{NH}_3 \) (ref. 106). These references lead to an opacity of \( \kappa_R = (1.8 \pm 0.6) \times 10^{-5} / \lambda^4 \text{ cm}^2/\text{g} \) (for \( \lambda \) in \( \mu \text{m} \)) for Rayleigh scattering throughout the atmosphere. Within the clouds the opacity is expected to be considerably larger but cannot be predicted because the cloud masses and particle sizes are uncertain and probably variable. Absorption by the molecules \( \text{CH}_4 \) and \( \text{NH}_3 \) can be inferred from laboratory data but depends strongly on wavelength and weakly on pressure. Because of these complications, the opacity is taken herein as the Rayleigh scattering value above the clouds and zero visibility is adopted near and below the clouds. This simple approximation should be considered only as an interim model until a more rigorous solution becomes available.

The index of refraction \( n \) may be calculated from the information provided in the Smithsonian Physical Tables (ref. 107) and is roughly independent of wavelength. It is given by \( n = 1 + (n_s - 1) P_T / P_s T \) where \( n_s, P_s, \) and \( T_s \) represent values for the gas mixture at standard temperature and pressure. The quantity \( (n_s - 1) \) has values in the range 0.00012±0.00002 for the compositions listed in section 2.9.1 and \( n = 1.00019±0.00003 \) at the level at which \( P = 1 \text{ atm} \) in the model atmospheres of section 2.9.4.

In the infrared, the opacity is that of collision-induced transitions in molecular hydrogen (refs. 93, 108, and 109) and of rotational bands of ammonia (ref. 110). Both opacities depend strongly on pressure, and simple formulae for the opacity cannot be obtained. A similarly complex situation exists in the microwave region where the opacity results not only from the molecular hydrogen absorption but also from the center and wings of the strong ammonia band near 1.3 cm wavelength (refs. 90 and 111). The brightness temperature data (sec. 2.5.2) for the Jupiter disk indicate that the opacity decreases as the wavelength increases beyond this band and that optical depth unity is reached at pressures of the order of magnitude of a few atmospheres near 10 cm wavelength. It is thus reasonable to expect that communications wavelengths greater than 10 cm would encounter less atmospheric absorption and would be able to penetrate deeper layers of the atmosphere. In addition, however, the possibly complex transmission properties of the ionosphere (sec. 2.8) must be considered at wavelengths much longer than 100 cm.

The electrical conductivity of gas mixtures involves complex, unpredictable phenomena, and the important effects of impurities are uncertain (ref. 112). However, Jupiter’s atmosphere is expected to be an excellent insulator like the Earth’s atmosphere.

2.12.2 Mechanical and Thermal

The calculation of viscosity and thermal conductivity is simplified by the graphical procedures in reference 113. The application of these techniques to the gas mixtures of
section 2.9 yields temperature-dependent values of the viscosity $\eta$ and thermal conductivity $\lambda_t$ which can be summarized for the range of the model atmospheres by

$$\log \eta = (0.4 \pm 0.1) + (0.65)\log T \quad (29)$$

and

$$\log \lambda_t = -(5.3 \pm 0.1) + (0.77)\log T \quad (30)$$

where $T$ must be in K, $\eta$ in micropoise, and $\lambda_t$ in cal/cm-sec°K. The uncertainty arises primarily from the range of model compositions and the inaccuracy of the log-linear fit. Neither of the uncertainty factors depends on pressure or density.

The range of specific heat values for molecular hydrogen (ref. 114) is so broad for the range of atmospheric temperatures that the variations added by the inclusion of other constituents and condensates are not significant. That range is specified simply by $C_v = 21 \pm 7$ joule/mole°K and $C_p = 29 \pm 7$ joule/mole°K. The ratio of specific heats has the corresponding range of $\gamma = 1.45 \pm 0.15$. Inclusion of the range of molecular weights in the formula for the speed of sound, $c_s = (\gamma R_g T/\mu)^{1/2}$ where $T$ is in K, yields $c_s = (70 \pm 10$ m/sec)$^{1/2}$.

### 2.13 Interior

The interior of Jupiter is that portion of the planet below the atmospheric altitudes at which radiation of various wavelengths escapes. The only compositions capable of retaining the low mean density of Jupiter (sec. 2.1.3) at the pressures required by its high gravity consist primarily of hydrogen and possibly include helium and other elements per table III. Theoretical models of the interior have been constructed by DeMarcus (ref. 115) and refined by Peebles (ref. 116); thermal effects are included in references 117, 118, and 119. The models suggest the following probable conclusions. (1) The planet is fluid either everywhere or throughout much of its bulk with the metallic phase of hydrogen being likely in the central regions; a small, heavy-element core is not excluded. (2) The pressure, density, and temperature increase in moving inward; the rate of increase is uncertain because the equation of state is poorly known. The temperature reaches values near the center of the order of magnitude of $10^4$°K. (3) Energy is released to the exterior from cooling and gravitational contraction at the luminosity suggested in section 2.5.2. (4) The energy is transported by convection in most layers of the planet, and electron conduction is significant elsewhere. (5) Convection in the rotating, conducting fluid hydrogen interior is responsible for the maintenance of the planetary magnetic field. (The magneto-hydrodynamic processes involved are not well understood.) (6) Although phase transitions may occur even in the probably chemically homogeneous interior, they do not constitute a rigid, solid “surface” at any level but are characterized by density and viscosity graduation over a radial distance scale of kilometers or greater. Therefore, the atmosphere is considered to be an extension of the interior rather than a distinct entity.

*One micropoise equals $10^{-7}$Ns/m$^2$. 38
3. CRITERIA

The following descriptions of the physical properties of the environment of Jupiter are recommended for application to the design of spacecraft intended to function within that environment. When uncertainty estimates or limiting values are given, they represent the extreme range of possible values to be used for design although the real environment probably will provide some surprises. For design purposes the probability that a particular environmental value is appropriate is nearly uniform within range but greatest at the nominal value.

3.1 Physical Properties

Values and uncertainties for Jupiter's mass, radius, figure, moments of inertia, rotation, and angular momentum are presented in table IV. In visible light the planet is axially symmetric about the rotation axis and has radius values \( R_\phi \) which depend on latitude \( \phi \) through the relationship

\[
R_\phi = R_J [1 - \epsilon (\sin \phi)^2].
\]  
(31)

The planet is gaseous to at least 1000 km below the radius specified by equation 31. Jupiter does not have a readily identifiable solid or liquid surface.

3.2 Gravity Field

Table V gives parameters and their uncertainties for Jupiter's gravitational field.

3.3 Magnetic Field and Magnetosphere

The parameters of Jupiter's magnetic field are given in table VI. Appendix B specifies the magnetic field components and lines of force in terms of the parameters. Jupiter's northern hemisphere contains the north-seeking end of the magnetic dipole. The uncertainties in table VI include the field strength ranges expected on the basis of the small dipole inclination and displacement from Jupiter's rotational axis and geometric center. Jupiter's magnetic field dominates within the elongated, teardrop-shaped magnetosphere whose dimensions are given in table VI.

3.4 Electric Fields

Static electric fields of intensity \( 0 \leq \sigma \leq (14 \text{ volts/meter})(R_J/R)^2 \) occur in the magnetosphere and ionosphere, in any direction, but usually perpendicular to the magnetic lines of force.

In the troposphere, field strengths up to \( 3 \times 10^4 \text{ volts/meter} \) on the scale of kilometers and hours may occur in the neighborhood of cloud formations. Field strengths up to \( 10^6 \text{ volts/meter} \) on the scale of centimeters and seconds may occur prior to strokes of lightning.
<table>
<thead>
<tr>
<th>Physical Parameter</th>
<th>Value</th>
</tr>
</thead>
<tbody>
<tr>
<td>Distance from Sun</td>
<td>5.20 ± 0.25 AU</td>
</tr>
<tr>
<td>Period of revolution about Sun</td>
<td>11.862 years</td>
</tr>
<tr>
<td>Mass of planet, excluding satellites</td>
<td>(1.899 ± 0.002) X 10^30 gram</td>
</tr>
<tr>
<td>Equatorial radius</td>
<td>71422 ± 200 km</td>
</tr>
<tr>
<td>Mean density</td>
<td>1.32 ± 0.01 g/cm³</td>
</tr>
<tr>
<td>Flattening (dynamical)</td>
<td>0.0645 ± 0.0008</td>
</tr>
<tr>
<td>Period of rotation</td>
<td>9h 55m 29s 73 ± 0s 04</td>
</tr>
<tr>
<td>Rotational angular velocity</td>
<td>(1.758531 ± 0.000002) X 10^-4 rad/sec</td>
</tr>
<tr>
<td>Angular momentum</td>
<td>(4.3 ± 0.2) X 10^45 g cm²/sec</td>
</tr>
<tr>
<td>Celestial coordinates of North Pole</td>
<td>17h 52m 084 + 0247(t - 1919.0)</td>
</tr>
<tr>
<td></td>
<td>+64° 33′ 34″ 6 - 0″ 60(t - 1910.0)</td>
</tr>
<tr>
<td>Moment of inertia about rotational axis</td>
<td>(2.4 ± 0.1) X 10^49 g cm²</td>
</tr>
<tr>
<td>Moment of inertia about equatorial axes</td>
<td>(2.25 ± 0.1) X 10^49 g cm²</td>
</tr>
</tbody>
</table>
### TABLE V

**GRAVITATIONAL PARAMETERS OF JUPITER**

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Equation/Expression</th>
</tr>
</thead>
<tbody>
<tr>
<td>Gravitational potential in an inertial coordinate system at distance R from Jupiter</td>
<td>$\Psi = -(1774 \pm 30 \text{ km}^2/\text{sec}^2) (R_J/R)$</td>
</tr>
<tr>
<td>Escape velocity at distance R</td>
<td>$(-2\Psi)^{1/6} = (59.6 \pm 0.6 \text{ km/sec})(R_J/R)^{1/6}$</td>
</tr>
<tr>
<td>Period of Keplerian orbit of semi-major axis a</td>
<td>$T_a = (2.95 \text{ hour})(a/R_J)^{3/2}$</td>
</tr>
<tr>
<td>Atmospheric acceleration of gravity – Magnitude (includes rotation)</td>
<td>$g = 2500 \pm 200 \text{ cm/sec}^2$</td>
</tr>
<tr>
<td>Equatorward deflection from jovicentric direction at latitude $\phi$</td>
<td>$\eta_1 = (3^\circ 70 \pm 0^\circ 03) \sin 2\phi$</td>
</tr>
</tbody>
</table>

### TABLE VI

**PARAMETERS OF JUPITER’S MAGNETIC FIELD AND MAGNETOSPHERE**

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Value</th>
</tr>
</thead>
<tbody>
<tr>
<td>Dipole Moment (centered with respect to planet)</td>
<td>$M_1 = (4 \times 2^{\pm1}) \times 10^{30} \text{ gauss-cm}^3$</td>
</tr>
<tr>
<td>Atmospheric field strength – At equator</td>
<td>$M_1/R_J^3 = 12 \times 2^{\pm1} \text{ gauss}$</td>
</tr>
<tr>
<td>At poles</td>
<td>$2M_1/R_J^3 = 24 \times 2^{\pm1} \text{ gauss}$</td>
</tr>
<tr>
<td>Magnetosphere dimensions – Standoff in solar direction</td>
<td>$50 \pm 20 R_J$</td>
</tr>
<tr>
<td>Length in anti-solar direction</td>
<td>$&gt; 100 R_J$</td>
</tr>
</tbody>
</table>
3.5 Electromagnetic Radiation

3.5.1 Above the Tropopause

Beyond Jupiter's atmosphere (here the tropopause at altitude \( z = 35 \pm 15 \text{ km} \) is the appropriate limit), the formulae in table VII specify the ranges of intensities, fluxes, and temperatures associated with the electromagnetic radiation environment under conditions of maximum illumination. Under partial illumination, total shadowing, or eclipse the values for intensities and fluxes range down to zero from the maximum illumination conditions of table VII and temperatures range down to 300K. The wavelength of interest determines which of the radiation sources listed in table VII should be considered. Figure 7 shows the geometric albedo \( p \) and brightness temperature \( T_{BD} \) as functions of the wavelength. In table VII the quantity \( D \) is the approximate path length of the line of sight through the schematic synchrotron emission volume shown in figure 8.

In addition, Jupiter's satellites and the other planets constitute light sources, but their fluxes near Jupiter are minor compared to the sources listed in table VII. Their positions are published annually in the *American Ephemeris and Nautical Almanac* (ref. 16), and their visual magnitudes and colors are specified by

\[
m_v = (m_0 \pm 0.3) + 5 \log r \Delta + (0.03 \pm 0.02)a
\]

and the parameters in table IX where \( a \) is in degrees, and \( r \) and \( \Delta \) in AU. The radiation from stars and other non-solar-system objects is identical to that beyond the Earth's atmosphere.

3.5.2 Below the Tropopause

Below the tropopause, i.e., altitudes \( z<35\pm15 \text{ km} \), the maximum contributions of the Sun and the synchrotron source are identical to those specified in table VII which uses the geometric albedo \( p \) and brightness temperature \( T_{BD} \) given in figure 7. The reflected sunlight intensity and flux in table VII are appropriate but must be considered omnidirectional to account for the maximum scattered light. The thermal radiation is specified by table VIII. In addition, sporadic lightning may require that light-sensitive surfaces be covered when local meteorological conditions are stormy.

Radiation and its effects at wavelengths shorter than 1 Å (γ-rays) should be ignored. The power emitted sporadically in the HF bursts near Jupiter ranges up to \( 4 \times 10^{16} \text{ erg/sec} \) at wavelengths longer than 7.5 meters; intensities at wavelengths longer than 1 meter are unknown but probably high compared to communications intensities.

3.6 Satellites and Meteoroids

Four categories of known satellites of Jupiter are listed in table X which specifies some of their properties. Additional photometric data for the Galilean satellites are given in section 3.5. Detailed ephemerides may be found in the *American Ephemeris and Nautical Almanac* (ref. 16) and the references specified there.

Within the volume centered on Jupiter of radius 500 R\(_J\) and at latitudes within ±30°, asteroidal meteoroids of mass density \( 3.5 \times 2^{\pm1} \text{ g/cm}^3 \) and relative velocity
### TABLE VII

**ELECTROMAGNETIC RADIATION PARAMETERS NEAR JUPITER WITH MAXIMUM ILLUMINATION**

<table>
<thead>
<tr>
<th>Parameters</th>
<th>Direct Sunlight*</th>
<th>Reflected Sunlight*</th>
<th>Thermal Radiation**</th>
<th>Synchrotron UHF Radiation***</th>
</tr>
</thead>
<tbody>
<tr>
<td>Wavelength</td>
<td>$1 , \text{Å} &lt; \lambda &lt; 100 , \text{cm}$</td>
<td>$1 , \text{Å} &lt; \lambda &lt; 10 , \mu\text{m}$</td>
<td>$1 , \mu\text{m} &lt; \lambda &lt; 100 , \text{cm}$</td>
<td>$1 , \text{cm} &lt; \lambda &lt; 100 , \text{cm}$</td>
</tr>
<tr>
<td>Intensity — Power/Area-</td>
<td>$I_\lambda = \frac{P_\lambda}{(0.000068 \text{ sterad})}$</td>
<td>$I_\lambda = \frac{pP_\lambda}{(90 \pm 10)}$</td>
<td>$I_\lambda = B_\lambda(T_{BD})$</td>
<td>$I_\nu = 2kT_{BS}/\lambda^2$</td>
</tr>
<tr>
<td>wavelength-</td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>solid angle</td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>Power/Area-frequency</td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>angle</td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>Flux — Power/Area-wavelength</td>
<td>$F_\lambda = \frac{P_\lambda}{r^2}$</td>
<td>$F_\lambda = \frac{pP_\lambda}{(27 \pm 3)(R/R_J)^2}$</td>
<td>$F_\lambda = \frac{\pi B_\lambda(T_{BD})}{(R/R_J)^2}$</td>
<td>$F_\nu = \frac{(5 \pm 2) \times 10^6}{4 + (R/R_J)^2} F_U$</td>
</tr>
<tr>
<td>Power/Area-frequency</td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>Integrated Flux — Power/Area —</td>
<td>$F = \frac{(1.353 \pm 0.021)}{r^2} \times 10^6 \frac{\text{erg}}{\text{cm}^2\text{sec}}$</td>
<td>$F = \frac{(1.4 \pm 0.2)}{(R/R_J)^2} \times 10^4 \frac{\text{erg}}{\text{cm}^2\text{sec}}$</td>
<td>$F = \frac{(1.8 \pm 0.3)}{(R/R_J)^2} \times 10^4 \frac{\text{erg}}{\text{cm}^2\text{sec}}$</td>
<td>$F = \frac{(15 \pm 7) \times 10^6}{4 + (R/R_J)^2} \frac{\text{erg}}{\text{cm}^2\text{sec}}$</td>
</tr>
<tr>
<td>Brightness Temperature</td>
<td></td>
<td></td>
<td></td>
<td>$T_{BD}$</td>
</tr>
<tr>
<td>Effective Temperature</td>
<td>5800 $\pm 15^\circ\text{K}$</td>
<td></td>
<td></td>
<td>$T_{e} = 134 \pm 6^\circ\text{K}$</td>
</tr>
</tbody>
</table>

* Solar irradiance from NASA SP-8005 (ref. 1), solar distance $r$ in AU only, and geometric albedo $p$ from figure 7.

** Brightness temperature $T_{BD}$ from figure 7 and Planck functions $B_\lambda(T)$ and $B_\nu(T)$ from Allen (ref. 7) or elsewhere.

*** Use $\lambda$ in cm only and estimate $D/R_J$ from figure 8 with $D$ being the approximate path length of line of sight through the volume. Radiation is up to 30 percent linearly polarized.
TABLE VIII
THERMAL RADIATION PARAMETERS BELOW
THE TROPOPAUSE IN JUPITER'S ATMOSPHERE

<table>
<thead>
<tr>
<th>Wavelength</th>
<th>$1 \mu m &lt; \lambda &lt; 100$ cm</th>
</tr>
</thead>
</table>
| **Intensities*** | $I_\lambda = B_\lambda(t)$  
$I_\nu = B_\nu(t)$ |
| **Fluxes*** | $F_\lambda = \pi B_\lambda(t)$  
$F_\nu = \pi B_\nu(t)$ |
| **Integrated Fluxes** | $F = \begin{cases} 
(1.8 \pm 0.3) \times 10^4 \left( \frac{\text{erg}}{\text{cm}^2\text{sec}} \right) & \text{for } T < 134^\circ K \\
\left( \frac{T}{134} \right)^4 (1.8 \pm 0.3) \times 10^4 \left( \frac{\text{erg}}{\text{cm}^2\text{sec}} \right) & \text{for } T > 134^\circ K 
\end{cases}$ |
| **Brightness Temperature** | $t = \text{larger of } T \text{ and } T_{BD}$ |

* $t$ is the larger of the brightness temperature $T_{BD}$ shown in figure 7 and the local tropospheric temperature $T$ (sec. 3.9); Planck functions $B_\lambda(T)$ and $B_\nu(T)$ from Allen (ref. 7) or elsewhere should be used.

$v = \sqrt{v_s^2 + (R_J/R) (40 \text{ km/sec})^2}$ have a concentration of particles of mass greater than $M$ given by

$$
\log N_M = -31 \pm 1 - (0.8 \pm 0.1)(\log M - 18.5)
$$

(33)

where $v_s$ is the spacecraft speed, $N_M$ is in $m^{-3}$, and $M$ is in grams ranging from $10^{-9}$ to $10^{20}$. The corresponding isotropic meteoroid flux is $N_M v/4$. The observed, small satellites of Jupiter form the large-mass end of this distribution.

3.7 Charged Particles

3.7.1 Particle Types

Table XI presents formulas and parameters for galactic cosmic rays, solar protons, solar wind, low energy trapped radiation belts, magnetospheric plasma, and the ionosphere. Table XII
<table>
<thead>
<tr>
<th></th>
<th>m0</th>
<th>U-B</th>
<th>B-V</th>
</tr>
</thead>
<tbody>
<tr>
<td>Jupiter V (Amalthea)</td>
<td>6.3</td>
<td>-</td>
<td>-</td>
</tr>
<tr>
<td>I (Io)</td>
<td>-1.90</td>
<td>1.30</td>
<td>1.17</td>
</tr>
<tr>
<td>II (Europa)</td>
<td>-1.53</td>
<td>0.52</td>
<td>0.87</td>
</tr>
<tr>
<td>III (Ganymede)</td>
<td>-2.16</td>
<td>0.50</td>
<td>0.83</td>
</tr>
<tr>
<td>IV (Callisto)</td>
<td>-1.20</td>
<td>0.55</td>
<td>0.86</td>
</tr>
<tr>
<td>VI-XII</td>
<td>≥7.0</td>
<td>-</td>
<td>-</td>
</tr>
<tr>
<td>Mercury</td>
<td>-0.36</td>
<td>-</td>
<td>0.93</td>
</tr>
<tr>
<td>Venus</td>
<td>-4.29</td>
<td>0.50</td>
<td>0.82</td>
</tr>
<tr>
<td>Earth</td>
<td>-3.87</td>
<td>-</td>
<td>0.2</td>
</tr>
<tr>
<td>Mars</td>
<td>-1.52</td>
<td>0.58</td>
<td>1.36</td>
</tr>
<tr>
<td>Jupiter</td>
<td>-9.25</td>
<td>0.48</td>
<td>0.83</td>
</tr>
<tr>
<td>Saturn</td>
<td>-8.88</td>
<td>0.58</td>
<td>1.04</td>
</tr>
<tr>
<td>Uranus</td>
<td>-7.19</td>
<td>0.28</td>
<td>0.56</td>
</tr>
<tr>
<td>Neptune</td>
<td>-6.87</td>
<td>0.21</td>
<td>0.41</td>
</tr>
<tr>
<td>Pluto</td>
<td>-1.01</td>
<td>0.27</td>
<td>0.80</td>
</tr>
</tbody>
</table>

*Symbols defined in appendix A and parameters obtained from reference 38.
Figure 7 - Ranges of geometric albedo and disk brightness temperature as functions of wavelength.
Figure 8.—Schematic volume for the calculation of the brightness temperature, $T_{BS}$, of Jupiter's synchrotron emission source.
TABLE X
PROPERTIES OF JUPITER'S SATELITES

<table>
<thead>
<tr>
<th>Satellite</th>
<th>Orbital Period (days)</th>
<th>Range of Distance from Jupiter (radii of Jupiter, R_J)</th>
<th>Maximum Jovian-centric Latitude (degrees)</th>
<th>Orbital Speed (km/sec)</th>
<th>Radius R_n (km)</th>
<th>Mass M_n (g)</th>
</tr>
</thead>
<tbody>
<tr>
<td>V (Amalthea)</td>
<td>0.498</td>
<td>2.539 ± 0.08</td>
<td>0.4</td>
<td>26</td>
<td>100 ± 2</td>
<td>10^{22} ± 1</td>
</tr>
<tr>
<td>I (Io)</td>
<td>1.76914</td>
<td>5.905</td>
<td>0</td>
<td>17</td>
<td>1800 ± 160</td>
<td>(7.2 ± 0.6) × 10^{25}</td>
</tr>
<tr>
<td>II (Europa)</td>
<td>3.55118</td>
<td>9.396</td>
<td>0</td>
<td>14</td>
<td>1550 ± 100</td>
<td>(4.7 ± 0.1) × 10^{25}</td>
</tr>
<tr>
<td>III (Ganymede)</td>
<td>7.15455</td>
<td>14.99</td>
<td>0</td>
<td>11</td>
<td>2600 ± 400</td>
<td>(15.5 ± 0.2) × 10^{25}</td>
</tr>
<tr>
<td>IV (Callisto)</td>
<td>16.68902</td>
<td>26.36</td>
<td>0</td>
<td>8</td>
<td>2400 ± 400</td>
<td>(9.6 ± 0.8) × 10^{25}</td>
</tr>
<tr>
<td>VI</td>
<td>251</td>
<td>161 ± 25</td>
<td>29</td>
<td>3</td>
<td>70 ± 2</td>
<td>10^{21.5±1}</td>
</tr>
<tr>
<td>VII</td>
<td>260</td>
<td>164 ± 34</td>
<td>28</td>
<td>3</td>
<td>10 ± 2</td>
<td>10^{19±1}</td>
</tr>
<tr>
<td>X</td>
<td>253</td>
<td>164 ± 23</td>
<td>29</td>
<td>3</td>
<td>7 ± 2</td>
<td>10^{18.5±1}</td>
</tr>
<tr>
<td>VIII</td>
<td>737</td>
<td>300 ± 200</td>
<td>33</td>
<td>2</td>
<td>8 ± 2</td>
<td>10^{19±1}</td>
</tr>
<tr>
<td>IX</td>
<td>758</td>
<td>330 ± 120</td>
<td>23</td>
<td>2</td>
<td>7 ± 2</td>
<td>10^{18.5±1}</td>
</tr>
<tr>
<td>XI</td>
<td>692</td>
<td>313 ± 63</td>
<td>18</td>
<td>2</td>
<td>7 ± 2</td>
<td>10^{18.5±1}</td>
</tr>
<tr>
<td>XII</td>
<td>631</td>
<td>290 ± 50</td>
<td>34</td>
<td>2</td>
<td>7 ± 2</td>
<td>10^{18.5±1}</td>
</tr>
</tbody>
</table>
TABLE XI
PARAMETERS OF CHARGED PARTICLES NEAR JUPITER*

<table>
<thead>
<tr>
<th>Particle</th>
<th>Location</th>
<th>Temperature, Energy or Velocity***</th>
<th>Number Density or Flux***</th>
</tr>
</thead>
<tbody>
<tr>
<td>Galactic Cosmic Rays</td>
<td>Everywhere</td>
<td>$0.1 &lt; E &lt; 10^{10}$ GeV</td>
<td>$\Phi_E = A(E + m_0c^2)^{-1.5}$ for $E$ in units of adjacent column</td>
</tr>
<tr>
<td></td>
<td></td>
<td>(GeV/nucleon for alphas)</td>
<td>Electrons: $0 \leq A \leq 0.02$ cm$^{-2}$sec$^{-1}$</td>
</tr>
<tr>
<td></td>
<td></td>
<td></td>
<td>Protons: $0 \leq A \leq 10$ cm$^{-2}$sec$^{-1}$</td>
</tr>
<tr>
<td></td>
<td></td>
<td></td>
<td>Alphas: $0 \leq A \leq 1$ cm$^{-2}$sec$^{-1}$</td>
</tr>
<tr>
<td>Solar Cosmic Rays (Protons)</td>
<td>Everywhere</td>
<td>$10 &lt; E &lt; 10^4$ MeV</td>
<td>Sporadic, with fluxes between 0.0 and 1.0 times those specified in NASA TR R-169 (ref. 59)</td>
</tr>
<tr>
<td>Solar Wind (electrons and protons)</td>
<td>Beyond Magnetosphere (sec. 3.3)**</td>
<td>$v = 320$ km/sec (up to 960 km/sec, at peak solar activity)</td>
<td>$N_0 = 0.2$ cm$^{-3}$ (up to 2.0 cm$^{-3}$, at peak solar activity)</td>
</tr>
<tr>
<td>Trapped Radiation* (electrons and protons)</td>
<td>Within Magnetosphere (sec. 3.3)**</td>
<td>$10 &lt; E &lt; 10^6$ eV</td>
<td>$0 &lt; N_E &lt; \frac{3.6 \times 10^{12}}{(E+3.3L^2)L^6}$ cm$^{-3}$ or $N_E &lt; 10^7$ cm$^{-3}$ whichever is smaller.</td>
</tr>
<tr>
<td>Magnetospheric Plasma (electrons and protons)</td>
<td>Within Magnetosphere (sec. 3.3)**</td>
<td>$150 &lt; T &lt; 10^{5}$K</td>
<td>$0.1 &lt; N_0 &lt; 10^7$ cm$^{-3}$ for $L \leq 4.2$</td>
</tr>
<tr>
<td></td>
<td></td>
<td>$10^{-2} &lt; E &lt; 10$ eV</td>
<td>$0 &lt; N_0 &lt; 10^{12}/L^8$ for $L \geq 4.2$</td>
</tr>
<tr>
<td>Ionosphere (electrons and protons)</td>
<td>For $z &gt; z_0$ where $z_0 = 250 \pm 150$ km</td>
<td>$T = 150 \pm 50^0$K</td>
<td>$N_0 = (10^{6 \pm 1}$cm$^{-3}) \exp\left(-\frac{z - z_0}{100 \text{ km}}\right)$</td>
</tr>
</tbody>
</table>

*Except trapped radiation at energies $> 1$ MeV to which tables XII and XIII apply.
**Shares the rotation of Jupiter's magnetic field (sec. 3.3).
***As defined in appendix A.
### TABLE XII

**FORMULAS FOR ENERGETIC CHARGED PARTICLES IN JUPITER'S TRAPPED RADIATION BELTS**

<table>
<thead>
<tr>
<th>MODEL</th>
<th>LOCATION</th>
<th>FLUX PARAMETER, $\phi_0$ (cm$^{-2}$ sec$^{-1}$)</th>
<th>CHARACTERISTIC ENERGY, $E_0$ (MeV)</th>
</tr>
</thead>
<tbody>
<tr>
<td><strong>NOMINAL ELECTRONS</strong></td>
<td>$1 &lt; L \leq 2$</td>
<td>$2.0 \times 10^7$</td>
<td>$6.2$</td>
</tr>
<tr>
<td></td>
<td>$2 \leq L &lt; 50$</td>
<td>$\frac{1.7 \times 10^{10}}{L^6} \left( \frac{1377}{L^3} + 1 \right)^{-\frac{1}{2}}$</td>
<td>$(0.51) \left[ \left( \frac{1377}{L^3} + 1 \right)^{\frac{1}{2}} \right] - 1$</td>
</tr>
<tr>
<td><strong>UPPER LIMIT ELECTRONS</strong></td>
<td>$1 &lt; L \leq 2$</td>
<td>$6.0 \times 10^7$</td>
<td>$20.0$</td>
</tr>
<tr>
<td></td>
<td>$2 \leq L &lt; 50$</td>
<td>$\frac{1.5 \times 10^{11}}{L^6} \left( \frac{12930}{L^3} + 1 \right)^{-\frac{1}{2}}$</td>
<td>$(0.51) \left[ \left( \frac{12930}{L^3} + 1 \right)^{\frac{1}{2}} \right] - 1$</td>
</tr>
<tr>
<td><strong>NOMINAL PROTONS</strong></td>
<td>$1 &lt; L &lt; 50$</td>
<td>$\frac{4.7 \times 10^8}{L^6} \left( \frac{1.06}{L^3} + 1 \right)^{-\frac{1}{2}}$</td>
<td>$(938) \left[ \left( \frac{1.06}{L^3} + 1 \right)^{\frac{1}{2}} \right] - 1$</td>
</tr>
<tr>
<td><strong>UPPER LIMIT PROTONS</strong></td>
<td>$1 &lt; L \leq 6$</td>
<td>$\frac{1.0 \times 10^{12}}{L^6} \left( \frac{1.8}{L^3} + 1 \right)^{-\frac{1}{2}}$</td>
<td>$(938) \left[ \left( \frac{1.8}{L^3} + 1 \right)^{\frac{1}{2}} \right] - 1$</td>
</tr>
<tr>
<td></td>
<td>$6 \leq L &lt; 12$</td>
<td>$(2.8 \times 10^{10}) L^{-4} + (5.2 \times 10^7) \exp \left[ -(12 - L)^2 \right]$</td>
<td></td>
</tr>
<tr>
<td></td>
<td>$12 \leq L &lt; 50$</td>
<td>$\frac{1.6 \times 10^{14}}{L^6}$</td>
<td></td>
</tr>
</tbody>
</table>

**DISTRIBUTIONS WITH ENERGY AND LATITUDE**

$$\phi_E = \phi_0 \left(1 + \frac{E}{E_0} \right) \exp \left( -\frac{E}{E_0} - \frac{\phi^2}{1000} \right)$$

* Latitude $\phi$ in degrees only.
### TABLE XIII

ENERGETIC CHARGED PARTICLE FLUXES FOR INDIVIDUAL ENERGY INTERVALS AT THE LOCATION OF THE PEAK ELECTRON FLUXES IN JUPITER'S TRAPPED RADIATION BELTS (L = 2 and $\phi = 0$)

<table>
<thead>
<tr>
<th>Particle Type</th>
<th>Energy Interval (MeV)</th>
<th>Flux, $(\Delta \Phi)_E$ in cm$^{-2}$ sec$^{-1}$</th>
</tr>
</thead>
<tbody>
<tr>
<td></td>
<td></td>
<td>Nominal</td>
</tr>
<tr>
<td>Electrons</td>
<td>1-3</td>
<td>$1.5 \times 10^6$</td>
</tr>
<tr>
<td></td>
<td>3-10</td>
<td>$7.9 \times 10^6$</td>
</tr>
<tr>
<td></td>
<td>10-30</td>
<td>$9.5 \times 10^6$</td>
</tr>
<tr>
<td></td>
<td>30-100</td>
<td>$9.3 \times 10^5$</td>
</tr>
<tr>
<td></td>
<td>100-300</td>
<td>34</td>
</tr>
<tr>
<td></td>
<td>300-1000</td>
<td>0.0</td>
</tr>
<tr>
<td>Protons</td>
<td>1-3</td>
<td>$7.4 \times 10^3$</td>
</tr>
<tr>
<td></td>
<td>3-10</td>
<td>$7.8 \times 10^4$</td>
</tr>
<tr>
<td></td>
<td>10-30</td>
<td>$5.4 \times 10^5$</td>
</tr>
<tr>
<td></td>
<td>30-100</td>
<td>$2.8 \times 10^6$</td>
</tr>
<tr>
<td></td>
<td>100-300</td>
<td>$3.2 \times 10^6$</td>
</tr>
<tr>
<td></td>
<td>300-1000</td>
<td>$2.8 \times 10^5$</td>
</tr>
<tr>
<td></td>
<td>1000-3000</td>
<td>7.0</td>
</tr>
<tr>
<td></td>
<td>3000-10000</td>
<td>0.0</td>
</tr>
</tbody>
</table>
Figure 9. - Fluxes of charged particles in Jupiter's trapped radiation belts as functions of distance from the magnetic dipole in the magnetic equatorial plane.
presents formulas and parameters for relativistic electrons and energetic protons in the trapped radiation belts, and table XIII evaluates their nominal and upper limit fluxes in several energy intervals at the location of the peak electron fluxes in the belts (L=2 and \( \phi=0 \)).

### 3.7.2 Calculation of Low Energy Radiation Belt Number Densities

Upper limits to the number densities for trapped electrons and protons having energies less than 1 MeV are specified by the formulas given in table XI. The pertinent entry in table XI specifies the upper limit to the concentration \( N_E \) (in \( \text{cm}^{-3} \)) of particles of energy greater than or equal to \( E \) (in MeV) as a function of magnetic shell parameter \( L \), subject to the condition that \( N_E \) does not exceed \( 10^7 \) \( \text{cm}^{-3} \). In case a higher value for \( N_E \) is obtained, the upper limit is taken as \( 10^7 \) \( \text{cm}^{-3} \). Although the upper limit specifies the maximum concentration at any energy, the concentration may not equal this upper limit at more than one energy at any given time and location.

### 3.7.3 Calculation of Radiation Belt Fluxes

To calculate particle fluxes from the formulas in table XII for a specific position (specified by distance \( R \) and latitude \( \phi \) with respect to Jupiter) and energy interval (between \( E \) and \( E+\Delta E \)), the following procedures should be employed.

1. Calculate the magnetic shell parameter \( L \) from \( L=R/R_J(\cos\phi)^2 \).  
2. Evaluate the flux parameter \( \Phi_0 \) (shown in figure 9) and the characteristic energy \( E_0 \) from the first several rows in the table, chosen according to the value of \( L \), the particle type, and the model required.  
3. With the last formula in table XII, calculate \( \Phi_E \), the particle flux for energy greater than \( E \), for both end-points of the energy interval and take the difference to obtain \( (\Delta\Phi)_E \), the flux in the interval of interest.

At any given time and location, the minimum, or lower limit flux is zero.

### 3.8 Transfer Properties in the Magnetosphere and Ionosphere

The range of values for the anisotropic electrical conductivity in the magnetosphere (the altitude range between 400 km and \( 50 \pm 20 R_J \)) is 0 to \( 4.3 \times 10^4 \) mho/meter. In the ionosphere, between 100 and 400 km altitude, the anisotropic conductivity range is 0 to 1 mho/meter. In both regions, thermal conduction is negligible compared to heat transport by radiative processes.

For electromagnetic radiation at wavelengths shorter than one meter (frequency \( \nu > 300 \) MHz), the index of refraction \( n \) is bounded by

\[
1 \geq n \geq 1 - (4.5 \times 10^{-7})\lambda^2
\]  

where \( \lambda \) is the wavelength in cm. At these wavelengths the opacity and optical thickness of the ionosphere and magnetosphere are neglected. At wavelengths greater than one meter, unpredictable refraction, reflection, absorption, and Faraday rotation effects occur.
3.9 Atmospheric Structure

The nominal model atmosphere of Jupiter given in table XIV contains values for the temperature $T$, density $\rho$, altitude $z$, pressure and density scale heights $H_p$ and $H_\rho$, and cloud mass $w$ as functions of pressure $P$ (for $0.2 \text{ mb} \leq P \leq 1000 \text{ atm}$). Tables XV and XVI give the "cool" and "warm" model atmospheres which encompass observational uncertainty and variations expected with planetary latitude, time of day, and local features. The models represent cool and warm extremes at given pressures not altitudes. The compositions for the three models are given in table XVII which does not take into account condensation effects. For all models, the zero of altitude is arbitrarily set at a pressure of one Earth atmosphere ($1.01325 \times 10^6 \text{ dyn/cm}^2$) and corresponds to the distance from Jupiter's center of $R_s$, given by

$$R_s = R_J \left[1 - \epsilon \sin^2 \phi \right]$$

(35)

where $R_J = 71422 \pm 200 \text{ km}$, $\epsilon = 0.0645 \pm 0.0008$, $\phi$ is the latitude, and $R_s$ specifies the level of the optical disk of the planet but not a real liquid or solid surface. Interpolation and extrapolation beyond the tabulated pressure range should be carried out linearly in $\log P$, $\log T$, $\log \rho$, and $z$. Figures 10 through 14 present the three models graphically.

3.10 Cloud Properties

The ranges of altitude and mass of the clouds and their composition (primarily water and ammonia ices and solution with additional unknown coloring agents and minor compounds) are specified in tables XIV, XV, and XVI and figures 10, 11, 13 and 14. Condensates are expected to be common in the bright zones (rising gas) and infrequent in the dark belts (descending gas), but Jupiter's clouds like Earth's are associated with complex and hard-to-predict weather phenomena. Hazards to spacecraft from corrosive properties of the clouds and from occasional water and ammonia rain, hail, snow, thunder, and lightning are anticipated. Among the highly variable cloud features, the most distinctive and permanent is the Great Red Spot which is dark at wavelengths less than 5000 Å and drifts slowly in longitude but remains near 22° South latitude.

3.11 Atmospheric Motions

The planetary rotation underlying the atmosphere is prograde at an angular velocity $\omega_0 = (1.758531 \pm 0.000002) \times 10^{-4} \text{ radian/sec}$, corresponding to the radio period $T_0 = 9^h 55^m 29^s 73 \pm 0^s 04$. Within the convective troposphere (altitudes below $35 \pm 15 \text{ km}$), there exist subsonic prevailing zonal winds, westerly at approximately 100 m/sec within $10^\circ$ of the equator and variable at $10 \pm 10 \text{ m/sec}$ easterly or westerly elsewhere. In the same region turbulence (local horizontal and vertical winds) of order of magnitude $10 \text{ m/sec}$ occurs in random directions with respect to the prevailing winds. Within the stratosphere and beyond (altitudes above $35 \pm 15 \text{ km}$), turbulence of the same magnitude occurs horizontally but not vertically. The strength of possible wind shear has not been estimated.
<table>
<thead>
<tr>
<th>( \rho ) (atm)</th>
<th>( T ) (°K)</th>
<th>( \rho ) (g/cm³)</th>
<th>( z^* ) (km)</th>
<th>( H_p ) (km)</th>
<th>( H_p ) (km)</th>
<th>( w ) (mg/liter)</th>
<th>Remarks</th>
</tr>
</thead>
<tbody>
<tr>
<td>2.00 ( \times 10^{-7} )</td>
<td>145.0</td>
<td>( 3.86 \times 10^{-11} )</td>
<td>313.4</td>
<td>21.0</td>
<td>21.0</td>
<td></td>
<td></td>
</tr>
<tr>
<td>3.00 ( \times 10^{-7} )</td>
<td>145.0</td>
<td>( 5.80 \times 10^{-11} )</td>
<td>304.9</td>
<td>21.0</td>
<td>21.0</td>
<td></td>
<td></td>
</tr>
<tr>
<td>1.00 ( \times 10^{-6} )</td>
<td>145.0</td>
<td>( 1.93 \times 10^{-10} )</td>
<td>279.7</td>
<td>21.0</td>
<td>21.0</td>
<td></td>
<td></td>
</tr>
<tr>
<td>3.00 ( \times 10^{-6} )</td>
<td>145.0</td>
<td>( 5.80 \times 10^{-10} )</td>
<td>256.6</td>
<td>21.0</td>
<td>21.0</td>
<td></td>
<td></td>
</tr>
<tr>
<td>1.00 ( \times 10^{-5} )</td>
<td>145.0</td>
<td>( 1.93 \times 10^{-9} )</td>
<td>231.4</td>
<td>21.0</td>
<td>21.0</td>
<td></td>
<td></td>
</tr>
<tr>
<td>3.00 ( \times 10^{-5} )</td>
<td>145.0</td>
<td>( 5.80 \times 10^{-9} )</td>
<td>208.3</td>
<td>21.0</td>
<td>21.0</td>
<td></td>
<td></td>
</tr>
<tr>
<td>1.00 ( \times 10^{-4} )</td>
<td>145.0</td>
<td>( 1.93 \times 10^{-8} )</td>
<td>183.1</td>
<td>21.0</td>
<td>21.0</td>
<td></td>
<td></td>
</tr>
<tr>
<td>3.00 ( \times 10^{-4} )</td>
<td>145.0</td>
<td>( 5.80 \times 10^{-8} )</td>
<td>160.0</td>
<td>21.0</td>
<td>21.0</td>
<td></td>
<td></td>
</tr>
<tr>
<td>0.00100</td>
<td>145.0</td>
<td>( 1.93 \times 10^{-7} )</td>
<td>134.8</td>
<td>21.0</td>
<td>21.0</td>
<td></td>
<td></td>
</tr>
<tr>
<td>0.00300</td>
<td>145.0</td>
<td>( 5.80 \times 10^{-7} )</td>
<td>111.7</td>
<td>21.0</td>
<td>21.0</td>
<td></td>
<td></td>
</tr>
<tr>
<td>0.00650</td>
<td>145.0</td>
<td>( 1.26 \times 10^{-6} )</td>
<td>95.5</td>
<td>21.0</td>
<td>21.0</td>
<td></td>
<td></td>
</tr>
<tr>
<td>0.0100</td>
<td>139.0</td>
<td>( 2.02 \times 10^{-6} )</td>
<td>86.7</td>
<td>20.1</td>
<td>18.3</td>
<td></td>
<td></td>
</tr>
<tr>
<td>0.0300</td>
<td>124.8</td>
<td>( 6.73 \times 10^{-6} )</td>
<td>65.7</td>
<td>18.1</td>
<td>16.4</td>
<td></td>
<td></td>
</tr>
<tr>
<td>0.0829</td>
<td>113.0</td>
<td>( 2.06 \times 10^{-5} )</td>
<td>48.3</td>
<td>16.3</td>
<td>16.3</td>
<td></td>
<td></td>
</tr>
<tr>
<td>0.100</td>
<td>113.0</td>
<td>( 2.48 \times 10^{-5} )</td>
<td>45.2</td>
<td>16.3</td>
<td>16.3</td>
<td></td>
<td></td>
</tr>
<tr>
<td>0.225</td>
<td>113.0</td>
<td>( 5.59 \times 10^{-5} )</td>
<td>31.9</td>
<td>16.3</td>
<td>16.3</td>
<td></td>
<td></td>
</tr>
<tr>
<td>0.267</td>
<td>120.0</td>
<td>( 6.24 \times 10^{-5} )</td>
<td>29.1</td>
<td>17.4</td>
<td>26.8</td>
<td></td>
<td></td>
</tr>
<tr>
<td>0.300</td>
<td>125.0</td>
<td>( 6.72 \times 10^{-5} )</td>
<td>27.0</td>
<td>18.0</td>
<td>27.9</td>
<td>0.00119</td>
<td></td>
</tr>
<tr>
<td>0.350</td>
<td>132.0</td>
<td>( 7.44 \times 10^{-5} )</td>
<td>24.1</td>
<td>19.1</td>
<td>29.3</td>
<td>0.00583</td>
<td></td>
</tr>
<tr>
<td>0.406</td>
<td>139.0</td>
<td>( 8.20 \times 10^{-5} )</td>
<td>21.2</td>
<td>20.1</td>
<td>30.8</td>
<td>0.0243</td>
<td></td>
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*The zero of altitude is specified by equation 35 (sec. 3.9).*
# TABLE XV

**COOL MODEL ATMOSPHERE OF JUPITER**

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<th>$P$ (atm)</th>
<th>$T$ (°K)</th>
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<th>$z^*$ (km)</th>
<th>$H_p$ (km)</th>
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<th>$w$ (mg/liter)</th>
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*The zero of altitude is specified by equation 35 (sec. 3.9).*

Remarks:
- Tropopause
- Correspondence Level
- Zero of altitude
- NH$_3$ ice cloud base
- Solution cloud base (NH$_3$·H$_2$O)
### TABLE XVI

WARM MODEL ATMOSPHERE OF JUPITER

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<th>$P$ (atm)</th>
<th>$T$ (°K)</th>
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<th>$z^*$ (km)</th>
<th>$H_\rho$ (km)</th>
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*The zero of altitude is specified by equation 35 (sec. 3.9).*
### TABLE XVII

**COMPOSITIONS OF MODELS OF JUPITER'S ATMOSPHERE**

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<th>Nominal Model</th>
<th>Warm Model</th>
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<td>0.00429</td>
<td>0.00214</td>
</tr>
<tr>
<td>NH₃</td>
<td>0.00219</td>
<td>0.00109</td>
<td>0.00055</td>
</tr>
<tr>
<td>H₂O</td>
<td>0.01601</td>
<td>0.00800</td>
<td>0.00400</td>
</tr>
<tr>
<td>Ne</td>
<td>0.00229</td>
<td>0.00115</td>
<td>0.00057</td>
</tr>
<tr>
<td>others</td>
<td>0.00398</td>
<td>0.00199</td>
<td>0.00100</td>
</tr>
<tr>
<td>Fractions by number (or volume)</td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>H₂</td>
<td>0.68454</td>
<td>0.86578</td>
<td>0.93754</td>
</tr>
<tr>
<td>He</td>
<td>0.31057</td>
<td>0.13214</td>
<td>0.06149</td>
</tr>
<tr>
<td>CH₄</td>
<td>0.00145</td>
<td>0.00062</td>
<td>0.00028</td>
</tr>
<tr>
<td>NH₃</td>
<td>0.00035</td>
<td>0.00015</td>
<td>0.00007</td>
</tr>
<tr>
<td>H₂O</td>
<td>0.00240</td>
<td>0.00102</td>
<td>0.00048</td>
</tr>
<tr>
<td>Ne</td>
<td>0.00031</td>
<td>0.00013</td>
<td>0.00006</td>
</tr>
<tr>
<td>others</td>
<td>0.00038</td>
<td>0.00016</td>
<td>0.00008</td>
</tr>
<tr>
<td>Mean molecular weight (u) (grams/mole)</td>
<td>2.70</td>
<td>2.30</td>
<td>2.14</td>
</tr>
</tbody>
</table>
Figure 10.—Pressure vs temperature for the Jupiter model atmospheres.
Figure 11.—Pressure vs density and cloud masses at cloud bases for the Jupiter model atmospheres.
Figure 12.—Pressure vs altitude for the Jupiter model atmospheres. The zero of altitude is specified by equation 35 (sec. 3.9).
Figure 13.—Temperature vs altitude for the Jupiter model atmospheres. The zero of altitude is specified by equation 35 (sec. 3.9).
Figure 14.—Density vs altitude and cloud masses at cloud bases for the Jupiter model atmospheres. The zero of altitude is specified by equation 35 (sec. 3.9).
3.12 Transfer Properties in the Atmosphere

Table XVIII provides expressions for electromagnetic, thermal, and mechanical transfer properties in the atmosphere.

### TABLE XVIII

TRANSFER PROPERTIES IN JUPITER'S ATMOSPHERE

<table>
<thead>
<tr>
<th>Property</th>
<th>Expression</th>
</tr>
</thead>
<tbody>
<tr>
<td>Opacity $\kappa$ (cm$^2$/gram)*</td>
<td>$\kappa = (1.8 \pm 0.6) \times 10^{-5}/\lambda^4$</td>
</tr>
<tr>
<td>Index of refraction $n$**</td>
<td>$n = 1.0 + (0.033 \pm 0.005)P/T$</td>
</tr>
<tr>
<td>Viscosity $\eta$ (micropoise)***</td>
<td>$\log \eta = (0.4 \pm 0.1) + (0.65)\log T$</td>
</tr>
<tr>
<td>Thermal conductivity $\lambda_t$*** (cal/cm sec °K)</td>
<td>$\log \lambda_t = (-5.3 \pm 0.1) + (0.77)\log T$</td>
</tr>
</tbody>
</table>
| Specific heats $C_v$, $C_p$ (joule/mole °K) | $C_v = 21 \pm 7$  
$C_p = 29 \pm 7$ |
| Ratio of specific heats $\gamma$ | $\gamma = C_p/C_v = 1.45 \pm 0.15$ |
| Speed of sound $c_s$ (m/sec)*** | $c_s = (70 \pm 10) T^{\frac{1}{2}}$ |

*For $\lambda < 1$ μm and altitude $z > 35 \pm 15$ km; elsewhere opacity cannot be specified.

**For $P$ in atm and $T$ in °K only.

***For $T$ in °K only. One micropoise equals $10^{-9}$Ns/m$^2$.

3.13 Interior

Below the deepest levels specified in section 3.9, the interior pressure, density, and temperature of the convective, hydrogen-rich interior increase inward toward the center of the planet, but their relationships (the equation of state) cannot be anticipated with reliability. No solid or liquid surface is expected to exist at any depth within the atmosphere or interior.
REFERENCES


APPENDIX A
SYMBOLS*

A coefficient in cosmic-ray flux formula (sec 2.7.1)
A moment of inertia about equatorial axis (sec. 2.1.4)
A_j parameter of condensible gas j (app. D)
a semi-major axis of orbit about Jupiter (table V)
B blue magnitude (table IX)
B magnetic field strength (app. B)
B_R component of B in radial direction (app. B)
B_\phi component of B in latitudinal direction (app. B)
B_\lambda(T) Planck function, per unit wavelength interval (sec. 2.5.2)
B_\nu(T) Planck function, per unit frequency interval (sec. 2.5.2)
C moment of inertia about polar axis (sec. 2.1.4)
C_p specific heat at constant pressure (sec. 2.12.2, app. D)
C_v specific heat at constant volume (sec. 2.12.2, app. D)
c speed of light, $3 \times 10^{10}$ cm/sec (sec. 2.7)
c_s speed of sound (sec. 2.12.2)
D path length within synchrotron radiation volume (sec. 2.5.3.1)
E charged particle kinetic energy (sec. 2.7)
E_0 local characteristic energy (sec. 2.7.4.2)
E_1 characteristic energy at flux peak (sec. 2.7.4.2)
\xi electric field strength, in volts/m (sec. 2.4)
e absolute value of electron charge, $4.8 \times 10^{-10}$ e.s.u. (sec. 2.4)
F integrated flux of electromagnetic radiation (sec. 2.5.1.2)
F_\lambda flux of electromagnetic radiation, per unit wavelength interval (sec. 2.5.1.2)
F_\nu flux of electromagnetic radiation, per unit frequency interval (app. C)
G universal constant of gravitation, $6.668 \times 10^{-8}$ dyne/cm^2/grams^2 (sec. 2.2)
g local acceleration of gravity (sec. 2.2.2)

*Words in bold face are defined in Glossary (app. E).
<table>
<thead>
<tr>
<th>Symbol</th>
<th>Description</th>
</tr>
</thead>
<tbody>
<tr>
<td>$H_p$</td>
<td>local pressure scale height (app. D)</td>
</tr>
<tr>
<td>$H'_p$</td>
<td>local density scale height (app. D)</td>
</tr>
<tr>
<td>$I_\lambda$</td>
<td>intensity per unit wavelength interval (sec. 2.5.1.2)</td>
</tr>
<tr>
<td>$I_\nu$</td>
<td>intensity per unit frequency interval (sec. 2.5.1.2 and app. C)</td>
</tr>
<tr>
<td>$J_2$</td>
<td>coefficient in gravitational potential (sec. 2.2.1)</td>
</tr>
<tr>
<td>$J_4$</td>
<td>coefficient in gravitational potential (sec. 2.2.1)</td>
</tr>
<tr>
<td>$K_1$</td>
<td>parameter of adiabatic gradient (app. D)</td>
</tr>
<tr>
<td>$K_2$</td>
<td>parameter of adiabatic gradient (app. D)</td>
</tr>
<tr>
<td>$k$</td>
<td>Boltzmann's constant, $1.38 \times 10^{-16}$ erg/°K (app. C)</td>
</tr>
<tr>
<td>$L$</td>
<td>magnetic shell parameter (app. B)</td>
</tr>
<tr>
<td>$L_0$</td>
<td>$L$ evaluated at inner boundary of radiation belts (sec. 2.7.4.1 b)</td>
</tr>
<tr>
<td>$L_1$</td>
<td>$L$ evaluated at peak of radiation belts (sec. 2.7.4.1 b)</td>
</tr>
<tr>
<td>$L_2$</td>
<td>$L$ evaluated at magnetopause (sec. 2.7.4.1 b)</td>
</tr>
<tr>
<td>$M$</td>
<td>meteoroid mass, in grams (sec. 2.6.3)</td>
</tr>
<tr>
<td>$M_1$</td>
<td>magnetic dipole moment (app. B)</td>
</tr>
<tr>
<td>$M_J$</td>
<td>mass of Jupiter, excluding satellites (sec. 2.1.1)</td>
</tr>
<tr>
<td>$M_{JS}$</td>
<td>mass of Jupiter system, including satellites (sec. 2.1.1)</td>
</tr>
<tr>
<td>$M_n$</td>
<td>mass of satellite $n$ (sec. 2.6.3)</td>
</tr>
<tr>
<td>$M_\odot$</td>
<td>mass of the Sun, $1.989 \times 10^{33}$ grams (sec. 2.1.1)</td>
</tr>
<tr>
<td>$m_0$</td>
<td>absolute magnitude (sec. 2.5.1.3)</td>
</tr>
<tr>
<td>$m_0$</td>
<td>rest mass of cosmic ray particle (sec. 2.7.1)</td>
</tr>
<tr>
<td>$m_p$</td>
<td>rest mass of proton, $1.67 \times 10^{-24}$ gram (sec. 2.7.4.4)</td>
</tr>
<tr>
<td>$m_v$</td>
<td>apparent visual magnitude (sec. 2.5.1.3)</td>
</tr>
<tr>
<td>$N$</td>
<td>charged particle concentration, in cm$^{-3}$ (sec. 2.7.5)</td>
</tr>
<tr>
<td>$N_0$</td>
<td>local radiation belt concentration parameter (sec. 2.7.4.1 b)</td>
</tr>
<tr>
<td>$N_1$</td>
<td>$N_0$ evaluated at flux peak (sec. 2.7.4.1 a)</td>
</tr>
<tr>
<td>$N_5$</td>
<td>$N$ evaluated for electrons with $E &gt; 5$ MeV (sec. 2.7.4.2)</td>
</tr>
<tr>
<td>$N_E$</td>
<td>$N$ evaluated for particles with kinetic energy $&gt; E$ (sec. 2.7.4.2)</td>
</tr>
<tr>
<td>$N_e$</td>
<td>electron concentration in the ionosphere, in cm$^{-3}$ (sec. 2.8.2)</td>
</tr>
</tbody>
</table>
$N_M$ concentration of meteoroids of mass $> M$ (sec. 2.6.3)

$n$ index of refraction (secs. 2.8.2 and 2.12.1)

$n_s$ $n$ evaluated at standard temperature and pressure (sec. 2.12.1)

$P$ local gas pressure in the atmosphere (app. D)

$P_2$ Legendre polynomial (sec. 2.2)

$P_4$ Legendre polynomial (sec. 2.2)

$P_A$ $P$ evaluated at some specific atmospheric level (app. D)

$P_{js}$ saturation vapor pressure of condensible gas $j$ (app. D)

$P_s$ standard pressure, 1 atm = 1.013 $\times$ 10$^6$ dynes/cm$^2$ (sec. 2.12.1)

$P_{\lambda}$ solar spectral irradiance (sec. 2.5.1.1 and ref. 1)

$p$ geometric albedo, function of wavelength (sec. 2.5.1.2)

$p_b$ bolometric geometric albedo (sec. 2.5.1.2)

$R$ separation from center of Jupiter (sec. 2.2)

$R_g$ universal gas constant (app. D)

$R_J$ equatorial radius of Jupiter's visible disk (sec. 2.1.2)

$R_n$ characteristic semi-diameter of satellite $n$ (sec. 2.6.3)

$R_s$ $R$ evaluated for Jupiter's visible disk (sec. 2.1.2)

$r$ distance from the Sun, in AU (sec. 2.5.1.2)

$S$ solar constant, 1.353 $\times$ 10$^6$ erg/cm$^2$ sec (sec. 2.5.1.2)

$T$ period of Jupiter's revolution about the Sun (table IV)

$T$ local temperature in atmosphere or ionosphere (app. D)

$T_0$ period of rotation of Jupiter's core and magnetic field (sec. 2.11.1)

$T_1$ period of rotation of System I (sec. 2.11.1)

$T_2$ period of rotation of System II (sec. 2.11.1)

$T_3$ period of rotation of System III (sec. 2.11.1)

$T_A$ $T$ evaluated at some specific atmospheric level (app. D)

$T_a$ period of revolution in orbit about Jupiter (table V)

$T_B$ brightness temperature (app. C)

$T_{BD}$ $T_B$ evaluated for Jupiter's disk (sec. 2.5.2)

$T_{BS}$ $T_B$ evaluated for Jupiter's synchrotron radiation source (sec. 2.5.3.1)
\( T_e \) effective temperature of Jupiter (sec. 2.5.2)
\( T_s \) standard temperature, 273.16 °K (sec. 2.12.1)
\( U \) ultraviolet magnitude (table IX)
\( u \) mean molecular weight, grams/mole (app. D)
\( u_j \) molecular weight of condensible gas j (app. D)
\( V \) visual magnitude (table IX)
\( v \) meteoroid speed relative to spacecraft (sec. 2.6.3)
\( v_s \) spacecraft speed relative to Jupiter (sec. 2.6.3)
\( W \) bandwidth of Jupiter's synchrotron radiation (sec. 2.7.4.2)
\( w \) representative cloud mass, per unit volume of gas (app. D)
\( w_j \) w evaluated for condensible gas j (app. D)
\( z \) altitude above Jupiter's visible disk (sec. 2.2.2)
\( z_0 \) z evaluated at base of Jupiter's ionosphere (sec. 2.7.6)
\( z_A \) z evaluated at some specific atmospheric level (app. D)
\( a \) phase angle (sec. 2.5.1.3)
\( a_R \) right ascension of Jupiter's north rotational pole (sec. 2.1.4)
\( \beta \) adiabatic gradient, \( \frac{d \log T}{d \log P} \) (app. D)
\( \beta_0 \) parameter of adiabatic gradient (app. D)
\( \gamma \) ratio of specific heats, \( C_p/C_v \) (app. D)
\( \Delta \) distance from reflecting surface to observer, in AU (sec. 2.5.1.3)
\( (\Delta \Phi)_E \) \( \Phi \) evaluated for energies between \( E \) and \( E+\Delta E \) (sec. 3.7.2)
\( \delta_R \) declination of Jupiter's north rotational pole (sec. 2.1.4)
\( \epsilon \) flattening of Jupiter's disk (sec. 2.1.2)
\( \eta \) coefficient of viscosity, in micropoise (sec. 2.12.2)
\( \eta_1 \) angle of equatorward deflection of gravity (sec. 2.2.2)
\( \theta \) pitch angle of electron trajectory (sec. 2.7.4.2)
\( \theta \) angle between propagation and magnetic field directions (sec. 2.8.2)
\( \kappa \) opacity in Jupiter's atmosphere (table XVIII)
\( \kappa_R \) \( \kappa \) evaluated for Rayleigh scattering (sec. 2.12.1)
\( \lambda \) wavelength of electromagnetic radiation (sec. 2.5)
\( \lambda_j \) parameter of condensible gas j (app. D)
\( \lambda_t \)  
thermal conductivity (sec. 2.12.2)

\( \nu \)  
frequency of electromagnetic radiation (sec. 2.8.2)

\( \nu_B \)  
electron gyrofrequency (sec. 2.8.2)

\( \nu_p \)  
plasma frequency (sec. 2.8.2)

\( \rho \)  
local gas density in the atmosphere (app. D)

\( \bar{\rho} \)  
mean density of Jupiter (sec. 2.1.3)

\( \sigma \)  
coefficient of rotation term in \( \Psi \) (sec. 2.2)

\( \sigma_R \)  
\( \sigma \) evaluated for rotation with Jupiter (secs. 2.1.4 and 2.2.2)

\( \tau \)  
optical depth (app. E)

\( \Phi \)  
flux of charged particles (sec. 2.7)

\( \Phi_E \)  
\( \Phi \) evaluated for particles with kinetic energy > \( E \) [\( E \neq 0 \)] (sec. 2.7)

\( \Phi_0 \)  
local radiation belt flux parameter (sec. 2.7.4.2)

\( \phi \)  
latitude, with respect to Jupiter's rotational equator (sec. 2.1.2)

\( \phi_0 \)  
parameter of radiation belt latitude distribution (sec. 2.7.4.1 c)

\( \Psi \)  
gravitational potential function (sec. 2.2)

\( \Omega \)  
solid angle of radio source (app. C)

\( \omega_0 \)  
angular velocity of Jupiter's core and magnetic field (sec. 2.1.4)
APPENDIX B
DIPOLE MAGNETIC FIELD

At distance R from Jupiter’s center and latitude $\phi$, the components of the magnetic field of the dipole whose magnetic moment is $M_1$ are

$$B_R = \frac{2M_1}{R^3} \sin \phi$$  \hspace{1cm} (B1)

and

$$B_\phi = -\frac{M_1}{R^3} \cos \phi.$$  \hspace{1cm} (B2)

The magnetic field has the magnitude

$$B = \frac{M_1}{R^3} \left[1 + 3 (\sin \phi)^2\right]^{1/2}.$$  \hspace{1cm} (B3)

It is directed along field lines in that plane which contains both the dipole and the point of interest; the shape of these field lines is specified by

$$R = L R_J (\cos \phi)^2$$  \hspace{1cm} (B4)

where the quantity $L$, the magnetic shell parameter, represents the distance from the magnetic dipole at which the magnetic line of force (on which the test point lies) crosses the magnetic equator, in units of $R_J$. 

78
APPENDIX C
RELATIONSHIPS AMONG RADIO ASTRONOMY QUANTITIES

At UHF (microwave) frequencies, corresponding to wavelengths from a few cm to a few meters, the results of radio astronomical observations are expressed in terms of either flux \( F_\nu \) or brightness temperature \( T_B \) (app. E) which can be related by a form of the Rayleigh-Jeans law

\[
F_\nu = 2k T_B \Omega / \lambda^2
\]  

where \( k \) is Boltzmann’s constant, \( 1.38 \times 10^{-23} \text{ joule/}^\circ \text{K} \), \( \lambda \) is the wavelength of observation, and \( \Omega \) is the solid angle of the source in steradians (\( \Omega << \pi \) is required). The flux density \( F_\nu \) is commonly expressed in flux units FU where 1 FU = \( 10^{-26} \text{ watt/m}^2 \text{ Hz} \). The intensity \( I_\nu \) is then correctly given by \( F_\nu / \Omega \) for all values of \( \Omega \) and is independent of distance from the source.
APPENDIX D
ATMOSPHERE AND CLOUD RELATIONSHIPS

In terms of the symbols defined in appendix A, the model atmospheres of sections 2.9 and 3.9 are governed, for each atmospheric region, by

(1) hydrostatic equilibrium
\[ \frac{dP}{dz} = - \rho g \]  
(D1)
(2) the perfect gas law
\[ \rho = \frac{uP}{R,T}, \]  
(D2)
(3) a temperature-dependent gradient
\[ \frac{d \log T}{d \log P} = \beta = \beta_0 \frac{T + K_1}{T + K_2} \]  
(D3)

where values of the constants $K_1$ and $K_2$ are restricted by $K_2 \geq 0$ in general, by $K_1 = K_2 > 0$ for regions of constant $\beta = \beta_0 \neq 0$, and by $K_1 = -T_A$ and $K_2 = 0$ for regions of $\beta = 0$, which implies constant temperature $T = T_A$. Values of $K_1$ and $K_2$ can be chosen to ensure that $\beta$ approximates the adiabatic value $(\gamma - 1)/\gamma$ for a real gas mixture whose specific heats are related by $\gamma = C_p/C_v$. The inequality $K_1 > K_2 > 0$ results if $C_p$ and $C_v$ increase with temperature.

The solution of equations D1 through D3 requires that $T$ and $z$ at any value of $P$ be related to those at $T_A$, $z_A$, and $P_A$ in the same region by

\[ \left(rac{P}{P_A}\right)^{\beta_0} = \left(rac{T}{T_A}\right)^{K_2/K_1} \left( \frac{T + K_1}{T_A + K_1} \right)^{(K_1 - K_2)/K_1} \]  
(D4)

and

\[ z = z_A - \frac{R_g}{u_g} \left( \frac{T - T_A}{\beta_0} - K_1 \ln \frac{P}{P_A} + \frac{K_2}{\beta_0} \ln \frac{T}{T_A} \right). \]  
(D5)

The lapse rate and scale heights are given by

\[ \frac{dT}{dz} = -\frac{\beta u_g R_g}{P} \]  
(D6)
\[ H_p = \frac{R_g T}{u_g} \]  
(D7)

and

\[ H_p = H_p/(1 - \beta). \]  
(D8)

It is assumed that clouds are formed by condensation of a species $j$ in convective regions moving upward to lower temperature and pressure when the saturation vapor pressure $P_{j, s}$. 

80
would be exceeded by the partial pressure (both for the condensible constituent) if condensation were not to occur. \( P_{js} \) is conventionally described by a form of the Clausius-Clapeyron equation

\[
P_{js} = \exp \left[ A_j - \left( \frac{\lambda_j}{R_g T} \right) \right].
\]

Here \( A_j \) and \( \lambda_j \) are constants, and \( \lambda_j \) is the latent heat of condensation. In this case \( P_{js} \) also specifies the partial pressure of the condensible gas in the cloud region above the lowest level of condensation, and the mass of the cloud per unit volume of gas is approximated by

\[
w_j = \frac{\beta \lambda_j}{R_g T} \left( \frac{\beta \lambda_j - \frac{R_g T}{\beta \lambda_j}}{\beta \lambda_j} \right)
\]

where \( \beta \) is given by equation D3 and \( u_j \) is the mean molecular weight of the condensate.

Table D-1 lists the three species actually used in computing cloud properties and the appropriate constants; for \( \text{NH}_3-\text{H}_2\text{O} \) solutions as dilute as those likely, the liquid water constants were considered satisfactory approximations.

**TABLE D-1**

**PARAMETERS OF CONDENSATES**

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Liquid H\textsubscript{2}O</th>
<th>Solid H\textsubscript{2}O Ice</th>
<th>Solid NH\textsubscript{3} Ice</th>
</tr>
</thead>
<tbody>
<tr>
<td>( u_j ) (grams/mole)</td>
<td>18</td>
<td>18</td>
<td>17</td>
</tr>
<tr>
<td>( A_j ) (for ( P_{js} ) in mm Hg)</td>
<td>20.592</td>
<td>24.055</td>
<td>23.744</td>
</tr>
<tr>
<td>( \lambda_j ) (cal/mole)</td>
<td>10 350.67</td>
<td>12 230.27</td>
<td>7 740.64</td>
</tr>
</tbody>
</table>
APPENDIX E
GLOSSARY*

Adiabatic – Characteristic of processes in which heat is not transferred across system boundaries; in an atmosphere such a system is a hypothetically rising or falling parcel of gas, and the adiabatic requirement must be satisfied when the parcel achieves the local pressure and density.

Astronomical Unit (AU) – A unit of length approximately equal to the semi-major-axis of the Earth’s orbit about the Sun. A modern value cited by Melbourne et al. (ref. 4) is $1 \text{ AU} = 1.49597893 \times 10^8 \pm 5 \text{ km}$.

Bandwidth (W) – The range of frequencies (or wavelengths) within which electromagnetic radiation is emitted or detected; the power or response distributions need not be uniform within this range.

Bolometric – Characteristic of an infinite bandwidth, and including electromagnetic radiation at all frequencies (or wavelengths) and polarizations.

Brightness Temperature ($T_B$) – The temperature at which a blackbody would radiate an intensity of electromagnetic radiation identical to that of the source for the bandwidth and polarization considered.

Color – For a given light source, the difference in magnitude for two bandwidths centered on different wavelengths.

Conductivity, Electrical – The ratio of current density to the impressed electric field strength. In the presence of both electric and magnetic fields, the direct conductivity is the ratio of the current to the electric field component parallel to the magnetic field; in the plane perpendicular to the magnetic field, the Pederson conductivity is the ratio of the current component, parallel to the electric field component, to that electric field component; and the Hall conductivity is the ratio of the current component, perpendicular to both fields, to the electric field component in the plane perpendicular to the magnetic field.

Conductivity, Thermal ($\lambda_t$) – The ratio of the thermal power per unit area crossing an imaginary plane surface perpendicular to a temperature gradient, to that gradient.

Decametric – Characteristic of electromagnetic radiation at those radio wavelengths between 10 and 100 meters; used here synonymously with high frequency (HF) to cover a broader range extending to about 7 meters.

Decimetric – Characteristic of electromagnetic radiation at those microwave wavelengths between 10 and 100 cm; used here synonymously with ultra high frequency (UHF) to cover a broader range extending to about 1 cm.

Declination – The celestial coordinate equal to the angle (north taken positive) between the direction of the item considered and the plane of the Earth’s equator; the precession of the latter implies a slow variation of the declination even of fixed directions.

*Cross references within the glossary are indicated by bold face.
Effective Temperature ($T_e$) – The temperature at which a blackbody would radiate a bolometric intensity of electromagnetic radiation identical to that of the source.

Flattening (e) – The positive difference between unity and the ratio of the polar to the equatorial diameter of a planetary disk (optical), or the value for the same quantity which would be derived on the basis of hydrodynamic theory and the gravitational potential inferred from observed satellite motions (dynamical).

Flux, of Electromagnetic Radiation ($F, F_\nu$, or $F_\lambda$) – The power per unit area crossing an imaginary plane surface from one side to the other, either per unit bandwidth or integrated over all frequencies.

Flux, of Particles ($\Phi$ or $\Phi_E$) – The number of particles per unit area and per unit time crossing an imaginary plane surface with positive or negative (but not both) velocity components perpendicular to the surface.

Geometric Albedo ($p$) – The ratio of the reflected electromagnetic flux from an astronomical object (observed at distance $\Delta$, zero phase angle, and zero optical depth) to the quotient of the solar power intercepted by the object divided by $\pi\Delta^2$, in the bandwidth considered. Here the flux, the power, and $\Delta$ must be expressed in consistent units, and $\Delta$ must be large compared to the dimensions of the object.

Gyrofrequency – The frequency of circling of a charged particle (mass $m$, charge $e$) in a uniform magnetic field $B$, given by $eB/2\pi m$ in general, and by $(2.8 \text{ MHz})B$ for electrons when $B$ is given in gauss.

Hall Conductivity – See conductivity, electrical.

High Frequency (HF) – Characteristic of electromagnetic radiation at those radio frequencies between 3 and 30 MHz, used here synonymously with decametric to cover a broader range extending to about 40 MHz.

Inertial Coördinates – Any system of coördinates characterized by zero or constant vector velocity with respect to the “fixed stars”.

Intensity ($I_\nu$ or $I_\lambda$) – The flux of electromagnetic radiation per unit solid angle of the source for a defining imaginary surface whose normal intersects the source; intensity is independent of the source-surface separation.

Ionosphere – The atmospheric layer which includes the major maxima of electron and ion concentration.

Magnetopause – The magnetosphere’s outer boundary, where a planet’s field interacts with external magnetic field and charged particle environments, particularly the solar wind.

Magnetosphere – The region surrounding a planet in which the local magnetic field is dominated by planet-associated fields rather than by external environments.

Magnitude ($m, m_0$) – Five-halves times the common logarithm (base ten) of the ratio of the power received per unit area within some bandwidth for a standard object to that for an astronomical object of interest. The base of this logarithmic scale is $x = 2.512$ so that an increase of one magnitude corresponds to a decrease in power by a factor $x^{-1}$ or an increase in distance by a factor $x^{1/2}$. Apparent magnitudes ($m$) are those observed. Absolute magnitudes ($m_0$) are those for which absorption and scattering effects have
been removed, and observation is assumed to occur in a standard geometrical configuration. For solar system objects, the standard configuration is sun-object distance $r = 1 \text{ AU}$, object-observer distance $\Delta = 1 \text{ AU}$, and phase angle $\alpha = 0$.

**Megahertz (MHz)** – The frequency unit $10^6$ cycles per second.

**Opacity ($\kappa$)** – The cross-section area per unit mass of absorber or scatterer.

**Optical Depth ($\tau$)** – The integral of the opacity $\kappa$ and the mass density $\rho$ along the absorption or scattering path length $ds$. The ratio of the intensity of a source after to that before absorption is $e^{-\tau}$ where $\tau = \int \kappa \rho \, ds$.

**Pederson Conductivity** – See conductivity, electrical.

**Phase Angle ($\alpha$)** – The angle formed by the Sun, the point of reflection, and the observer.

**Pitch Angle ($\theta$)** – The angle between the particle velocity vector and the external magnetic field vector evaluated at the position of the particle.

**Plasma** – A gas in which the concentration of charged particles has more than negligible effects on the properties of the gas.

**Position Angle** – The angle which describes the direction from one point to another on the sky. It is conventionally taken as zero when the direction of the line between the points is north; the angle increases as the line's direction moves east.

**Precession** – The change in the directions of the Earth’s rotational axis and vernal equinox with respect to the “fixed stars”. It is approximately periodic with period of 26 000 years.

**Prograde** – The direction of rotation and revolution common in the solar system in which the motion is counter-clockwise as viewed from the north.

**Rayleigh Scattering** – Scattering of electromagnetic radiation by particles (in Jupiter’s case H$_2$ molecules) whose characteristic size is small compared to the wavelength $\lambda$. The scattering center may be approximated by an oscillating electric dipole whose scattering cross-section and opacity are proportional to $\lambda^{-4}$.

**Relativistic** – Characteristic of a particle whose kinetic energy is comparable to or greater than its rest energy and whose speed is only slightly less than that of light.

**Right Ascension** – The celestial coordinate equal to the angle (east taken positive) between the projection of the direction of the item considered on the plane of the Earth's equator and that of the vernal equinox, expressed as fractions of 24 hours rather than 360°; the precession of the Earth's rotation axis implies a slow variation of the right ascension even of fixed directions.

**Scale Height ($H_p$, $H_p$)** – A measure of the vertical gradient of a quantity $x$, e.g., pressure, electron concentration, such that if $H = -x(dx/dz)^{-1}$ is constant with altitude $z$, the quantity $x$ changes by a factor $e$ within the altitude interval $H$.

**Stratopause** – The upper boundary of the stratosphere, characterized by a near-discontinuity in the temperature gradient.
Stratosphere — The atmospheric layer directly above the troposphere, within which the temperature is constant or increases with altitude.

Synchrotron Radiation — Electromagnetic radiation emitted by a relativistic charged particle as it is accelerated, i.e., follows a curved path, in a magnetic field.

System I — The artificial Jupiter longitude measurement system in which the meridian at 47°31 is directed toward the Earth at 0° Universal Time on 14 July 1897 and has prograde rotation with angular velocity of 877.90 degrees/day. It is applied to the analysis of optical observations at equatorial latitudes.

System II — The artificial Jupiter longitude measurement system in which the meridian at 96°58 is directed toward the Earth at 0° Universal Time on 14 July 1897 and has prograde rotation with angular velocity of 870.27 degrees/day. It is applied to the analysis of optical observations at mid-latitudes.

System III (1957.0) — The artificial Jupiter longitude measurement system in which the meridian at 108°02 is directed toward the Earth at 0° Universal Time on 1 January 1957 and has prograde rotation with period of 9h 55m 29s 37. It is applied to the analysis of radio observations.

Trapped Radiation — Energetic charged particles whose trajectories in a planetary magnetic field are confined within certain boundaries. A particle travels nearly along the field lines, “mirrors” at equal north and south magnetic latitudes, and drifts in longitude.

Tropopause — The upper boundary of the troposphere and the lower boundary of the stratosphere, characterized by a near-discontinuity in the temperature gradient.

Troposphere — The atmospheric layer within which major weather phenomena occur, characterized by temperatures which decrease as altitude increases.

Ultra High Frequency (UHF) — Characteristic of electromagnetic radiation at those microwave frequencies between 300 and 3000 megahertz (MHz); used here synonymously with decimetric to cover a broader range extending perhaps to 30 000 MHz.

Vernal Equinox — The direction from the center of the Earth to the center of the Sun at the time when the latter lies in the plane of the Earth's equator in March of each year.

Viscosity (η) — The ratio of the tangential component (perpendicular to the velocity gradient) of the force per unit area to the velocity gradient in a moving fluid.

Zenith — The direction opposite to that of the local acceleration of gravity (including the centrifugal terms) and perpendicular to the local horizon.

Zenith Angle — The angle between the directions to the zenith and to an object observed.
LIST OF NASA SPACE VEHICLE DESIGN CRITERIA MONOGRAPHS

SP-8001 (Structures) Buffeting During Atmospheric Ascent, Revised November 1970
SP-8002 (Structures) Flight-Loads Measurements During Launch and Exit, December 1964
SP-8003 (Structures) Flutter, Buzz, and Divergence, July 1964
SP-8004 (Structures) Panel Flutter, July 1964
SP-8005 (Environment) Solar Electromagnetic Radiation, Revised, May 1971
SP-8006 (Structures) Local Steady Aerodynamic Loads During Launch and Exit, May 1965
SP-8007 (Structures) Buckling of Thin-Walled Circular Cylinders, revised August 1968
SP-8008 (Structures) Prelaunch Ground Wind Loads, November 1965
SP-8009 (Structures) Propellant Slosh Loads, August 1968
SP-8010 (Environment) Models of Mars Atmosphere (1967), May 1968
SP-8011 (Environment) Models of Venus Atmosphere (1968), December 1968
SP-8012 (Structures) Natural Vibration Modal Analyses, September 1968
SP-8014 (Structures) Entry Thermal Protection, August 1968
SP-8015 (Guidance and Control) Guidance and Navigation for Entry Vehicles, November 1968
SP-8016 (Guidance and Control) Effects of Structural Flexibility on Spacecraft Control Systems, April 1969
SP-8017 (Environment) Magnetic Fields – Earth and Extraterrestrial, March 1969
SP-8018 (Guidance and Control) Spacecraft Magnetic Torques, March 1969
SP-8019 (Structures) Buckling of Thin-Walled Truncated Cones, September 1968
SP-8020 (Environment) Mars Surface Models (1968), May 1969
SP-8021 (Environment) Models of Earth's Atmosphere (120 to 1000 km), May 1969
SP-8022 (Structures) Staging Loads, February 1969
SP-8023 (Environment) Lunar Surface Models, May 1969
| SP-8024 | (Guidance and Control) | Spacecraft Gravitational Torques, May 1969 |
| SP-8025 | (Chemical Propulsion) | Solid Rocket Motor Metal Cases, April 1970 |
| SP-8026 | (Guidance and Control) | Spacecraft Star Trackers, July 1970 |
| SP-8027 | (Guidance and Control) | Spacecraft Radiation Torques, October 1969 |
| SP-8028 | (Guidance and Control) | Entry Vehicle Control, November 1969 |
| SP-8029 | (Structures) | Aerodynamic and Rocket-Exhaust Heating During Launch and Ascent, May 1969 |
| SP-8030 | (Structures) | Transient Loads From Thrust Excitation, February 1969 |
| SP-8031 | (Structures) | Slosh Suppression, May 1969 |
| SP-8032 | (Structures) | Buckling of Thin-Walled Doubly Curved Shells, August 1969 |
| SP-8033 | (Guidance and Control) | Spacecraft Earth Horizon Sensors, December 1969 |
| SP-8034 | (Guidance and Control) | Spacecraft Mass Expulsion Torques, December 1969 |
| SP-8035 | (Structures) | Wind Loads During Ascent, June 1970 |
| SP-8036 | (Guidance and Control) | Effects of Structural Flexibility on Launch Vehicle Control Systems, February 1970 |
| SP-8037 | (Environment) | Assessment and Control of Spacecraft Magnetic Fields, September 1970 |
| SP-8040 | (Structures) | Fracture Control of Metallic Pressure Vessels, May 1970 |
| SP-8042 | (Structures) | Meteoroid Damage Assessment, May 1970 |
| SP-8043 | (Structures) | Design-Development Testing, May 1970 |
| SP-8044 | (Structures) | Qualification Testing, May 1970 |
| SP-8045 | (Structures) | Acceptance Testing, April 1970 |
| SP-8046 | (Structures) | Landing Impact Attenuation For Non-Surface-Planing Landers, April 1970 |
SP-8047 (Guidance and Control)  Spacecraft Sun Sensors, June 1970

SP-8048 (Chemical Propulsion) Liquid Rocket Engine Turbopump Bearings, March 1971

SP-8049 (Environment) The Earth’s Ionosphere, March 1971

SP-8050 (Structures) Structural Vibration Prediction, June 1970

SP-8051 (Chemical Propulsion) Solid Rocket Motor Igniters, March 1971

SP-8052 (Chemical Propulsion) Liquid Rocket Engine Turbopump Inducers, May 1971

SP-8053 (Structures) Nuclear and Space Radiation Effects on Materials, June 1970

SP-8054 (Structures) Space Radiation Protection, June 1970

SP-8055 (Structures) Prevention of Coupled Structure-Propulsion Instability (Pogo), October 1970

SP-8056 (Structures) Flight Separation Mechanisms, October 1970

SP-8057 (Structures) Structural Design Criteria Applicable to a Space Shuttle, January 1971

SP-8058 (Guidance and Control) Spacecraft Aerodynamic Torques, January 1971

SP-8059 (Guidance and Control) Spacecraft Attitude Control During Thrusting Maneuvers, February 1971

SP-8060 (Structures) Compartment Venting, November 1970

SP-8061 (Structures) Interaction With Umbilicals and Launch Stand, August 1970

SP-8062 (Structures) Entry Gasdynamic Heating, January 1971

SP-8063 (Structures) Lubrication, Friction, and Wear June 1971
| SP-8065  | (Guidance and Control) | Tubular Spacecraft Booms (Extendible, Reel Stored), February 1971 |
| SP-8066  | (Structures)           | Deployable Aerodynamic Deceleration Systems, June 1971 |
| SP-8067  | (Environment)          | Earth Albedo and Emitted Radiation, July 1971 |
| SP-8068  | (Structures)           | Buckling Strength of Structural Plates, June 1971 |
| SP-8070  | (Guidance and Control) | Spaceborne Digital Computer Systems, March 1971 |
| SP-8071  | (Guidance and Control) | Passive Gravity-Gradient Libration Dampers, February 1971 |
| SP-8072  | (Structures)           | Acoustic Loads Generated by the Propulsion System, June 1971 |
| SP-8074  | (Guidance and Control) | Spacecraft Solar Cell Arrays, May 1971 |