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PARTICLES, ENVIRONMENTS AND POSSIBLE ECOLOGIES

IN THE JOVIAN ATMOSPHERE

(Revised Version)

Carl Sagan and E. E. Salpeter



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Abstract

The eddy diffusion coefficient is estimated as a function of altitude, separately for the Jovian troposphere and mesosphere. The growth-rate and motion of particles is estimated for various substances: the water clouds are probably nucleated by NH_4Cl , and sodium compounds are likely to be absent at and above the levels of the water clouds. Complex organic molecules produced by the Ly α photolysis of methane may possibly be the absorbers in the lower mesosphere which account for the low reflectivity of Jupiter in the near ultraviolet. The optical frequency chromophores are localized at or just below the Jovian tropopause. Candidate chromophore molecules must satisfy the condition that they are produced sufficiently rapidly that convective pyrolysis maintains the observed chromophore optical depth. Organic molecules and polymeric sulfur produced through H_2S photolysis at $\lambda > 2300 \text{ \AA}$ probably fail this test, even if a slow, deep circulation pattern, driven by latent heat, is present. The condition may be satisfied if complex organic chromophores are produced with high quantum yield by NH_3 photolysis at $\lambda < 2300 \text{ \AA}$. However, Jovian photoautotrophs in the upper troposphere satisfy this condition well, even with fast circulation, assuming only biochemical properties of comparable terrestrial organisms. Unless buoyancy can be achieved, a hypothetical organism drifts downwards and is pyrolyzed. An organism in the form of a thin, gas-filled balloon can grow fast enough to replicate if (i) it can survive at the low mesospheric temperatures, or if (ii) photosynthesis occurs in the troposphere. If hypothetical organisms are capable of slow, powered locomotion and coalescence, they can grow large enough to achieve buoyancy. Ecological niches for sinkers, floaters and hunters appear to exist in the Jovian atmosphere.

I. INTRODUCTION

The contemporary Jovian atmosphere has some similarities to the primitive terrestrial atmosphere, where organic molecules were produced readily; and experiments in which a wide variety of organic molecules are produced under conditions similar to those of the Jovian atmosphere have been carried out (Sagan and Miller, 1960; Sagan et al., 1967; Woeller and Ponnampereuma, 1969; Rabinowitz et al., 1969; Sagan and Khare, 1971a, 1971b; Khare and Sagan, 1973; Ferris and Chen, 1975). This has led to the hypothesis that there may be a Jovian biology (Sagan, 1961; Shklovskii and Sagan, 1966). Observations in the visible and in the ultraviolet have implied the presence of particulate matter in the Jovian atmosphere, which may (but need not necessarily) be connected with questions of Jovian organic chemistry and biology. Both chemical and biological issues are affected by fluid dynamics in the Jovian atmosphere. We are here concerned largely with these effects.

Modern models of the structure of Jupiter (for references, see, e.g., Stevenson and Salpeter, 1976) predict that the bulk constituent, a hydrogen-helium mixture, is fluid throughout with temperatures increasing inwards to above 10^4 °K. Plausible organisms require temperatures well below 10^3 °K. One principle problem which must be faced by such hypothetical organisms is descent to pyrolytic depths due either to the acceleration of gravity or to convective downdrafts. The second half of the present paper addresses (1) whether such organisms unde 30

significant growth and replication in less than the timescale for drift to pyrolytic depths; or (2) whether other adaptations to avoid pyrolysis are hydrodynamically and biologically feasible. We shall see that (1) becomes easier for smaller organisms and (2) for larger ones.

That plausible and internally consistent Jovian ecologies can be described cannot, by itself, demonstrate the likelihood of life on Jupiter. A tenable argument for life on Jupiter must also show that the origin of life on that planet was possible. Since many steps, particularly the later ones, which led to the origin of life on Earth are only incompletely understood (see, e.g., Miller and Orgel, 1974) a thorough discussion of the origin of life on Jupiter cannot be undertaken at the present time. Most skepticism on this subject seems to arise from the contention that synthesized molecules will be carried convectively to pyrolytic depths before the origin of life can occur. The unspoken premise in such an argument is that pyrolysis would not have occurred on the primitive Earth. However, the measured Arrhenius rate constants imply through a regression analysis that many essential organic molecules, for example a number of biological amino acids, suffer thermal degradation at typical terrestrial surface temperatures in geologically short time periods. The very limited available experimental data show, for example, that the half life at 30° C of the amino acid serine is $\sim 10^4$ yrs and of threonine is $\lesssim 10^5$ yrs; while at the same temperature the half life of the simplest amino acid alanine is $> 10^{10}$ yrs (Vallentyne, 1964). Because of the exponential Arrhenius kinetics, the highest temperatures on a temperature-variable planet dominate the net molecular degradations; also there is evidence that aqueous solution and alkaline pH's -- both expected on the

primitive Earth -- enhance the decomposition rates (Vallentyne, 1964). It does not seem at all improbable that some essential precursor molecules for the origin of life had half lives on the primitive Earth of 10^2 to 10^3 yrs, not greatly dissimilar to the lifetimes against convective pyrolysis of organic molecules on contemporary Jupiter, as calculated below.

Many workers on the origin of life have suggested that the rate-limiting steps were stochastic in nature. The volume available for natural molecular experiments on Jupiter is approximately 10^3 times larger than on Earth. If phase interfaces play a critical role in the origin of life (see, e.g., Bernal, 1973) Jupiter -- with a great concentration of solid and liquid atmospheric aerosol particles -- would be particularly favorable.

The origin of life on Earth took a period < 0.1 the lifetime of the Earth, and possibly much less (Sagan, 1975). On Jupiter, we have much longer periods of time available for chemical experiments, a larger interaction volume, and a much larger area for phase interface chemistry. On the other hand the thermal degradation timescales are shorter for many molecules, and there is a much greater excess of molecular hydrogen than on the primitive Earth, where H_2 was rapidly depleted during planetary formation and in subsequent Jeans escape. The effect of molecular hydrogen on pre-biological organic chemistry is inadequately understood, but is briefly discussed below. Its effect on quenching gas phase organic reactions is less than linear, at least for some reactions. We conclude that the absolute a priori likelihood of life on Jupiter, as on Earth, cannot be estimated reliably at the present time, but -- as on Earth -- can

hardly be ruled out. The relative a priori probabilities of the origin of life, in a comparison of Jupiter and Earth, depend on which reactions and molecular precursors are rate-limiting. If, for example, the alanine abundance is rate-limiting, the Earth is strongly favored; if serine is rate-limiting, Jupiter may be favored. It is also possible that there exist a variety of different paths to the origin of life, for some of which the Earth is favored and for others of which Jupiter is favored.

We devote the remainder of the present paper to a characterization of the relevant physical environment of Jupiter; to a discussion of the chromophores responsible for the observed coloration of the planet; and to an analysis of some permissible ecological niches of hypothetical Jovian organisms.

Values for the eddy diffusion coefficient K play an important role in such arguments. Estimates are given in Section II, separately for the convective troposphere and for the stabler mesosphere. Analytic approximations are also given there for abundance as a function of height for particles which are produced at some upper level and are transported downwards towards pyrolysis.

We are concerned with two sorts of particulate matter: particles that condense out under conditions of thermal equilibrium and particles that require an investment of free energy for their production. Equilibrium condensation has been studied previously for H_2O , NH_4SH and NH_3 ; we investigate it further in Section III for a number of substances with higher condensation temperatures. We estimate growth-rates for these condensates and discuss whether $NaCl$ or other particles can be present at the level of the water clouds to act as nucleation agents. The production of chromophores and of non-equilibrium organic molecules is treated in Section IV. We obtain a lower limit to

the production rate of ethane from methane photolysis by solar Lyman-alpha; and treat organic chromophores produced by various wavelengths of light, as well as polymeric sulfur chromophores.

The motion of hypothetical organisms is discussed in Section V; the diffusion of metabolites and the consequent growth of organisms in Section VI. We consider four different cases depending on whether organisms can or cannot survive at the low temperatures of the mesosphere and whether they can or cannot utilize visible light for photosynthesis. In Section VII we consider the more rapid growth attainable if organisms can control their motions.

II. EDDY DIFFUSION IN THE JOVIAN ATMOSPHERE

The total infrared flux emerging from Jupiter (Aumann, et al., 1969; Ingersoll, et al., 1975) is $F_{\text{tot}} \equiv \sigma T_{\text{eff}}^4 \sim 1.5 \times 10^4 \text{ erg cm}^{-2} \text{ s}^{-1}$ ($T_{\text{eff}} \sim 127^\circ \text{K}$), with about half of this flux coming from the deep interior ($F_{\text{int}} \sim 0.8 \times 10^4 \text{ erg cm}^{-2} \text{ s}^{-1}$). The radiative opacity increases fairly rapidly with increasing pressure, P , and radiative-convective model calculations predict a fairly rapid change across the boundary level defined here as the tropopause. Above the tropopause we have radiative (subadiabatic) heat transport in the mesosphere; below we have convective heat transport (with the temperature gradient only slightly superadiabatic) in the troposphere. Model atmospheres (Trafton and Munch, 1969; Hogan, et al., 1969; Wallace, et al., 1974) and spectroscopic observations (Gillett, et al., 1969; Aitken and Jones, 1972; Lacy, et al., 1975) suggest a temperature $T \sim 175^\circ \text{K}$ at the pressure level of $P = 1 \text{ bar}$ and place the tropopause somewhere slightly above this. We shall adopt the level ($T \sim 145^\circ$, $P \sim 0.5 \text{ bar}$) for the tropopause. These numbers are consistent with Pioneer 10 infrared radiometer data (Orton, 1975a) and with ground-based infrared radiometric measures (Orton 1975b) and possibly also with Pioneer 10 and 11 S-band occultation data (Kliore and Woiceshyn 1976). We assume a helium abundance of about 10 percent by number, consistent with separate measurements from the occultation of Beta Scorpii (Elliot, et al., 1974); from Pioneer 10 and 11 ultraviolet photometry (Carlson and

Judge, 1974); and from the pressure broadening of the $J = 0$ and $J = 1$ hydrogen lines (Houck, et al., 1975.) Assuming a helium-abundance of about 10% (by number), it then follows that $\gamma/(\gamma - 1) \approx 3$, where γ is the ratio of specific heats, and that the mean molecular weight $\mu \approx 2.2$. For pressure P , pressure scale height $H = P (dP/dz)^{-1}$, density ρ and kinematic viscosity ν in the troposphere we adopt

$$P \approx \left(\frac{T}{1750\text{K}}\right)^3 \text{ bar}; H \approx \left(\frac{T}{1750\text{K}}\right) 27\text{km}; \quad (1)$$

$$\rho \approx \left(\frac{T}{1750\text{K}}\right)^2 1.6 \times 10^{-4} \text{ gm cm}^{-3}.$$

$$\nu \approx \left(\frac{T}{1750\text{K}}\right)^{-1.5} 0.5 \text{ cm}^2 \text{ s}^{-1}$$

In the mesosphere, the temperature gradient decreases rapidly (for $P \lesssim 0.2$ bar, say), and the minimum temperature is about 120°K followed by a temperature inversion whose structure is not yet well-known (possibly $T_{\text{max}} \approx 150^\circ\text{K}$ at $P \sim 10^{-3}$ bar).

For a general convective circulation pattern the vertical convective heat-flux F_c is given by

$$F_c = c_p \rho \Delta T w_c \quad (2)$$

where c_p is the specific heat at constant pressure, ΔT is the typical vertical excess in the temperature drop and w_c the typical vertical convection speed. For vertical distances larger than the vertical extent ℓ of the circulation pattern one has an eddy diffusion coefficient $K_c \equiv w_c \ell$ and we shall denote by t_c the eddy diffusion time, H^2/K_c , for diffusion through a single pressure scale height H . In the simplest form of Prandtl mixing-length theory (see, e.g., Schwarzschild, 1958) the circulation is characterized by two parameters, the mixing-length ℓ and the superadiabaticity with w_c^2 and $\Delta T/T$ both

proportional to the superadiabaticity. We replace F_c by $1.4 \times 10^4 \text{ erg cm}^{-2} \text{ s}^{-1}$, a typical value for the Jovian troposphere, employ eq. (1) and write K_c , w_c and t_c in the form

$$\begin{aligned} K_c &= \xi (T/175^\circ\text{K})^{1/3} \times 10^7 \text{ cm}^2 \text{ s}^{-1}, \\ w_c &= \xi^{1/4} (T/175^\circ\text{K})^{-2/3} \times 76 \text{ cm s}^{-1}, \\ t_c &= \xi^{-1} (T/175^\circ\text{K})^{5/3} \times 7.5 \times 10^5 \text{ s}. \end{aligned} \quad (3)$$

In the notation of mixing-length theory our parameter ξ equals $(\ell/0.05H)^{4/3}$. The Jovian tropospheric lapse rate is very close to adiabatic ($\Delta T/T \sim 10^{-6}$), as is true by the same argument (Sagan, 1960) for all other planetary tropospheres.

In the simplest version of mixing-length theory $\ell \sim H$ and $\xi \sim 100$. The actual circulation patterns in the troposphere are likely to be much more complex (Hunten 1975). In particular one should consider the effects of rotation at great depths and of latent heat release in the cloud layers near the tropopause (Gierasch, 1973). For a constant value of the heat-flux F_c in eq. (2), these effects usually increase ΔT and decrease w_c considerably, leading to smaller values of K_c and ξ in eq. (3). These considerations suggest, as a simple working hypothesis, $\xi = 1$ which should be accurate to within a factor of about 200.

Gierasch (1976) has suggested a specific large-scale circulation pattern which extends all the way from the bottom of the water-cloud layer ($T \sim 280^\circ\text{K}$) to the tropopause. This pattern is driven by the latent heat of water which leads to a relatively large value of $\Delta T/T \sim 10^{-2}$. This corresponds to $\xi \sim 0.01$ with timescales for complete turnover of the region of up to a few years, even though some cloud patterns (above the water clouds) are observed to change in shorter times (Peek, 1958). On Gierasch's model, the eddy diffusion coefficient increases by a very large factor just below the cloud-bottom to $\xi \sim 10$.

In the mesosphere the bulk of the heat flux F_{tot} is carried by radiation, but some fraction of the flux could drive a circulation pattern working against a subadiabatic temperature gradient. The vertical speed w_m of such a circulation pattern must satisfy an inequality, based on the fact that the adiabatic cooling rate cannot exceed the total available heating rate:

$$H\rho w_m [c_p (dT/dz) + g] \leq F_{\text{tot}}. \quad (4)$$

Except in the immediate vicinity of the tropopause, the temperature gradient in eq. (4) can be neglected for the troposphere. Putting $T \sim 150^\circ\text{K}$ and assuming that the vertical length-scale for the circulation pattern is of order H , the eddy diffusion coefficient K_m is of order $w_m H$, so that

$$w_m \leq (1 \text{ bar}/P) 0.016 \text{ cm s}^{-1}; K_m \leq (1 \text{ bar}/P) 7 \times 10^4 \text{ cm}^2 \text{ s}^{-1}. \quad (5)$$

Thus, typical velocities just below the tropopause are many thousands of times larger than just above the tropopause. Analysis of the observed Ly α albedo of the upper mesosphere (Wallace and Hunten, 1973, Carlson and Judge, 1974) indicates $K_m \sim 3 \times 10^{8 \pm 1} \text{ cm}^2 \text{ s}^{-1}$ at levels where $P \sim 3 \times 10^{-7}$ bar; analysis of β Sco occultation measurements (Sagan et al., 1974) indicates $K_m > 10^6 \text{ cm}^2 \text{ s}^{-1}$ at similar levels. It is theoretically plausible (French and Gierasch, 1974) that gravity wave propagation upwards in the mesosphere contributes to the eddy diffusion. For a single wave mode, K_m has a step-function behavior, but overall it seems plausible to assume some constant fraction of the maximum expression in eq. (5). The value given by the Ly α albedo then suggests

$$K_m \sim (1 \text{ bar}/P) 100 \text{ cm}^2 \text{ s}^{-1}, \quad (6)$$

uncertain by factors of order 10 or 100. Note that $K_m \propto e^{z/H}$. Temperature T and $K(z)$ are plotted against pressure, schematically, in Figure 1. [The dashed curve for K corresponds to Gierasch's (1976) model in which the latent heat of water condensation is taken into account.]

We shall need continuity equations for $n_j(z)$, the number density as a function of height z of some complex molecule (of molecular weight $\mu_j \gg 2.2$) or of some particle. Let $w_d(z)$ be the steady-state vertical drift velocity of the molecule or droplet falling under gravity through an atmosphere with pressure scale height H , eddy diffusion coefficient K and mean molecular weight 2.2. In standard notation, the continuity equation (see, e.g., Strobel, 1973) reads

$$(K + D_j) \frac{dn_j}{dz} + \left(\frac{K}{H} + \frac{D_j}{H_j}\right) n_j = -\Phi_j,$$

where D_j is the effective molecular (or particle) diffusion coefficient and $H_j = H (2.2/\mu_j) \ll H$ is the effective scale height for the j th particles alone. Φ_j is the vertical flux of particles of species j . H_j and the drift-velocity w_d for free fall under gravity are related by $D_j = w_d H_j$. We shall be interested only in levels well below the turbopause where $D_j \ll K$ and give only a simplified version of the standard treatment (Strobel 1973; Prinn 1973). The steady-state continuity equation then reduces to

$$K(z) \frac{dn_j(z)}{dz} + \left[\frac{K(z)}{H} + w_d(z)\right] n_j(z) = -\Phi_j(z). \quad (7)$$

In Section IV we shall apply eq. (7) to the number density $n(z)$ of organic molecules or other chromophore particles for which the drift-velocity can be neglected ($w_j \ll K/H$). Assume further that the particles are mainly produced at some higher level $z \approx Z_0 > 0$ and destroyed at some lower level $z \equiv 0$, so that the flux Φ equals

the total production rate ($\text{cm}^{-2} \text{ s}^{-1}$) for $0 < z < Z_0$ and $\Phi = 0$ outside that range. If the production level Z_0 is in the troposphere, the eddy diffusion coefficient is given by K_c in eq. (3), which varies quite slowly over a scale-height H . An approximate solution of eq. (7), which neglects terms in dK_c/dz , is then

$$\begin{aligned} n(z) &\approx (\Phi H / K_c) (1 - e^{-z/H}) \quad \text{for } 0 < z < Z_0, \\ &\approx (\Phi H / K_c) [e^{-(z - Z_0)/H} - e^{-Z_0/H}] \quad \text{for } Z_0 < z. \end{aligned} \quad (8)$$

Thus, particle number densities decline exponentially with altitude above the production level and decrease more slowly with depth below the production level. Mathematically, the solution in eq. (8) depends on the boundary condition that $n(z) = 0$ for $z \leq 0$ and that Φ is already constant for infinitesimally small positive z . Physically, the requirement is merely that the destruction rate change from very small values (at positive z) to very large values (at negative z) over less than a scale height. This should be a good approximation since pyrolysis rates are very temperature sensitive. (But note that the level $z = 0$ refers to different altitudes for different molecular species.)

The situation is more complicated if the production level Z_0 is in the mesosphere and the destruction level ($z \equiv 0$) in the troposphere, with $z = Z_p < Z_0$ representing the tropopause. However, neglecting again derivatives of $K = K_c$ (for $z < Z_p$) and assuming that $K = K_m$ (for $Z_p < z$) is proportional to $e^{z/H}$ as in eq. (6), the continuity eq. (7) can again be solved analytically. With the fact that $Z_p \gg H$ and $K_c(Z_p) \gg K_m(Z_p)$ for cases of practical interest, the solution can be approximated by

$$\begin{aligned}
 n(z) &\approx (\Phi H / K_c) (1 - e^{-z/H}) \quad \text{for } 0 < z < Z_p, \\
 &\approx \frac{\Phi}{K_m(z)} \left[z - Z_p + H \frac{K_m(Z_p)}{K_c(Z_p)} \right] \quad \text{for } Z_p < z < Z_o, \\
 &\approx [\Phi(Z_o - Z_p) / K_m(Z_o)] e^{-(z-Z_o)/H} \quad \text{for } Z_o < z.
 \end{aligned}
 \tag{9}$$

Thus, particle number densities increase slowly with altitude above the destruction level until the tropopause; decline slower than exponentially ($\propto z e^{-z/H}$) between the tropopause and the production level; and decline exponentially above the production level. For mesospheric production and tropospheric destruction, we therefore expect maximum number densities just above the tropopause. Again, no particles exist at $Z < 0$. Note that the mixing-ratio ($\propto n/\rho_{\text{gas}}$) remains constant above the production level whether this level is in the mesosphere or troposphere. In reality, this mixing ratio would decrease slightly with increasing height z if some additional destruction takes place above the production level or if the drift-velocity w_d is not entirely negligible.

III. THERMAL CONDENSATION IN DEEP CLOUD LAYERS

The thermal condensation of droplets to form cloud layers in the Jovian atmosphere has been discussed (Weidenschilling and Lewis, 1973; Stauffer and Kiang, 1974) only for the upper levels of the troposphere ($T \lesssim 300^{\circ}\text{K}$), where direct observations are possible. Molecules involving H, C, N, O, P and S have been considered, which lead especially to condensation of H_2O ($T \lesssim 275^{\circ}\text{K}$), NH_4SH ($T \lesssim 220^{\circ}\text{K}$) and NH_3 ($T \lesssim 160^{\circ}\text{K}$). The condensation of droplets of more refractory materials in deeper layers (higher temperature, T) is also of some interest, especially for compounds containing Na and/or Cl. If such small grains or droplets can be convected to the upper troposphere, they can act as seed-nuclei for the condensation of H_2O -drops (at $T \approx 280^{\circ}\text{K}$). Their abundance is also relevant for any hypothetical Jovian biology and for the question of the replication of terrestrial microbial contaminants in the Jovian clouds. The following crude estimates and conjectures should serve at least for purposes of orientation.

In the very deep interior of Jupiter where the density exceeds 0.2 gm cm^{-3} ($P > 10^5 \text{ bar}$, $T \gtrsim 10^4 \text{ K}$), incomplete solubility of various substances in hydrogen could lead to substantial chemical separation. This possibility, discussed so far only for He (Salpeter, 1973; Smoluchowski, 1973), should be investigated further; however, we assume provisionally cosmic abundances of all elements (Allen, 1973)

at $P \sim 10^{3.5}$ bar ($T \sim 2,500^\circ\text{K}$), say. At this level practically all substances are gaseous. Refractory molecules begin to condense at slightly lower temperatures.

For each material, with f the cosmic abundance by number (mixing ratio) of the least abundant element in the compound, we must evaluate a number of quantities:

- (1) The equilibrium condensation temperature T_{con} appropriate to the Jovian troposphere, i.e., that temperature at which the saturation vapor pressure equals f times the pressure in eq. (1). Values for T_{con} have already been given by Lewis (1969b).
- (2) Assuming the absence of any external seed-nuclei, we calculate the amount of supercooling ΔT_{cr} required from homogeneous nucleation theory. For this purpose we need a parameter η , defined in eq. (7) of Salpeter (1974), which is essentially the number of times appropriate molecules of a particular species stick to a given surface site on a droplet during a typical flow-time. In that definition for η we use $s \sim 0.5$ for the surface sticking coefficient, eq. (1) for the pressure and replace the flow-velocity v by the typical convection speed w_c in eq. (3), which gives $\eta \sim f(T_{\text{con}}/175^\circ\text{K})^{4.2} \times 3 \times 10^{13}$. Values of f , and T_{con} (taken from JANAF 1971) are given in Table 1 for a number of compounds and η exceeds 4×10^8 in all cases. The ratio $\Delta T_{\text{cr}}/T_{\text{con}}$ is given by eq. (9) of Salpeter (1974), but multiplied by a factor $6^{-1/2}$ (correcting an error kindly pointed out by B. Draine). This ratio is fairly small in the cases considered and not of direct importance for our crude estimates. But we do need the final radius a_f of typical drops which would result at the end of purely homogeneous nucleation.

At first sight one would expect a_f to be of order η times the intermolecular spacing in the liquid, but there are a number of correction factors. The most important correction is due to the fact that the drop-radius a becomes large compared with the mean free path for a molecule in the gas, $\lambda_m \sim (T/175^\circ\text{K})^{-2} \times 6 \times 10^{-5}$ cm. The droplet's growth is controlled by the diffusion speed rather than the gas-kinetic speed and the rate is multiplied by a correction factor $(\lambda_m/a) \ll 1$. Including some further logarithmic factors we find, as a very rough approximation,

$$a_f \sim (175^\circ\text{K}/T)\eta^{1/2}(\log_{10}\eta)^{-3/4} \times 10^{-6} \text{ cm} \quad (10)$$

Our values for a_f are given in Table 1 and have to be compared with a number of other length scales:

- (3) Let a_{Re} be a critical radius for a spherical droplet such that the hydrodynamic drag-coefficient C_D is unity. For $a \ll a_{\text{Re}}$ the Reynolds number is small and the drag-coefficient is given by the Stokes formula $C_D = (12\nu/aw_d)$, where w_d is the Stokes expression for the terminal vertical drift-velocity of the sphere. Using eq. (1) we have

$$a_{\text{Re}} = 0.007 \text{ cm} \left(\frac{2}{\rho_{\text{dr}}} \frac{175^\circ\text{K}}{T} \right)^{1/3}$$

$$w_d(a) = \frac{2}{9} \frac{a^2 g}{\nu} \frac{\rho_{\text{dr}}}{\rho_g} \sim \left(0.15 \text{ cm s}^{-1} \right) \left(\frac{T}{175^\circ\text{K}} \right)^{-1/2} \left(\frac{\rho_{\text{dr}}}{2} \right) \left(\frac{a}{10^{-4} \text{ cm}} \right)^2 \quad (11)$$

where ρ_{dr} is the internal density (in gm cm^{-3}) of a liquid drop.

Another critical length is a_d , the droplet radius for which the drift-velocity w_d equals the typical convection speed w_c in eq. (3). For the cases of interest ($T \lesssim 2,500^\circ\text{K}$) we have $a_d < a_{\text{Re}}$, so that eq. (11)

holds. With $\rho_{dr} \approx 2$ and $\xi \approx 1$ in eq. (3) we have

$$a_d \approx 0.002 \text{ cm } (T/175^\circ\text{K})^{-1/12} \quad (12)$$

Large drops with $a \gg a_d$ "rain out" rapidly, whereas small droplets with $a \ll a_d$ are carried along by the convection pattern and can move up to higher levels.

For all cases in Table 1 (except possibly for NH_4Cl) we have $a_f > a_d$, so that homogeneous nucleation leads to large enough drops for most of the matter to be rained out. However, if a very large number of nucleation seeds are present originally, a large number of much smaller droplets condense out and the droplets must grow before they can fall out as rain. Let a_{coa} be that radius for which $C_D = \rho_{dr}/\rho_g$,

$$a_{coa} \sim \left(\frac{\rho_{dr}}{2}\right)^{-2/3} \left(\frac{T}{175^\circ\text{K}}\right)^{1/3} \times 3 \times 10^{-4} \text{ cm} \quad (13)$$

Once $a \gtrsim a_{coa}$, droplets usually can grow very quickly by coalescence, but for $a < a_{coa}$ growth must occur by the slower process of smaller drops evaporating and the excess vapor diffusing to larger drops (Langmuir, 1948; Mason, 1972; Rossow, 1976). An important parameter η_{vap} is the growth rate from this process (from radius $0.5 a_{coa}$ to a_{coa} , say) expressed in units of H^2/K_c , a typical convection time. With an uncertainty of at least a factor of 10, we have

$$\eta_{vap} \sim 4 \times 10^4 f \left(\frac{T}{175}\right)^{7/6} \left(\frac{\rho_{dr}}{2}\right)^2 \quad (14)$$

If $\eta_{vap} \gg 1$ for a particular material, most of it will form into large "rain-drops" in a cloud layer at levels close to the condensation temperature, T_{con} , and only small amounts of the material will be found at higher, cooler levels. If $\eta_{vap} \ll 1$

one might expect only small droplets of the material to form, so that $w_d \ll w_c$ and upward convection can maintain an appreciable mixing ratio of the droplets at higher levels. In practice, however, various complications usually tend towards more rapid formation of larger droplets (Rossow, 1976). For $\eta_{\text{vap}} \sim 10^{-1}$ or 10^{-2} the circumstances are probably mixed: Some rain near the T_{con} level, some small droplets a few scale heights higher up, but with a mixing ratio which decreases slightly with increasing height. Values for η_{vap} from eq. (14) are given in Table 1. Except possibly for NH_4Cl , $\eta_{\text{vap}} \gg 10^{-2}$ as well as $a_f > a_d$ and most of the material rains out and cannot move up to higher levels. However, in upward moving convection cells the saturation vapor pressure continues to fall; droplets then continue to condense out, but with decreasing values of the mixing ratio f and hence of the parameter η . Let f_m be that (decreased) value of the mixing ratio, for which the droplet size a_f (due to homogeneous nucleation) in eq. (10) equals a_d . Values of f_m are also given in Table 1. The mixing ratio will not drop appreciably below f_m in upward moving convection cells, since droplets are now too small to rain out rapidly enough (and further growth of droplets is also too slow).

For the cases with $T_{\text{con}} \gtrsim 1600^\circ\text{K}$ the radii satisfy $a_f > a_{\text{coa}}$, coalescence is rapid and large drops are rained out immediately. At the high pressures in Jupiter's interior, liquid Mg_2SiO_4 (as well as Ca- and Al- silicates and metallic Fe) condenses before MgSiO_3 and before SiO_2 . When liquid SiO_2 begins condensing near the 1800°K -level, most of the metallic elements have already been rained out, but sodium has not. For pure Na_2SiO_5 our estimates indicate a slightly lower condensation temperature than for SiO_2 , which introduces an uncertainty: It is conceivable that all the silicon has been rained out before sodium condenses, in which case gaseous NaCl and NaOH (since the elemental abundance of Na exceeds

that of Cl) survive to higher, cooler levels. It is more likely that much of the sodium condenses out between 1800°K and 1600°K in the form of some sodium silicate-silica solution, in which case gaseous HCl and some NaCl survive to higher levels. In either case, more of the surviving NaCl condenses out near 1000°K ; in the more likely case, NH_4Cl condenses near 400°K ; in the less likely case Na_2S condenses near 1000°K .

We have some tentative implications for levels with $T \lesssim 700^{\circ}\text{K}$, where Jovian biology (or at least organic chemistry) might be possible: Sodium (or other alkali metals and alkaline earths) and chlorine (or other halogens) cannot be present simultaneously. However, either sodium sulfide or (more likely) ammonium chloride is present. For the more likely case, some fraction of the NH_4Cl - crystals are convected upward and act as seed-nuclei for the condensation of water droplets, so that essentially no supercooling is present in the base level of the water clouds (Stauffer and Kiang, 1974). For the less likely case of Na_2S crystals convection over a larger height difference is required for nucleating the water, and f_m is slightly smaller. However, the chromophores to be discussed in Section IV, as well as Na_2S and NH_4Cl , could probably also act as seed-nuclei, and it seems generally safe to assume no supersaturation for water. We estimate below the amount of exogenous sodium which may arrive at the cloud level from micrometeoritic infall.

Since both sulfur and sodium have been detected in toroidal nebulae about Jupiter -- probably arising from the surface of Io -- these clouds are a potential source of S and Na in the clouds of Jupiter.

From observations of forbidden SII circumjovian line emission, Brown (1976) calculates a total sulfur flux from the surface of Io of $10^{8.6 \pm 0.7} \text{ cm}^{-2} \text{ sec}^{-1}$, equivalent to roughly $10^4 \text{ cm}^{-2} \text{ sec}^{-1}$ at the Jovian clouds. This is some eight orders of magnitude less than that expected from H_2S photodissociation (see Section IV), and therefore is entirely negligible. The cosmic abundance of Na is at least an order of magnitude less than that of S, and so only very small quantities of sodium could be present at the Jovian clouds due to leakage from the Io Na toroid.

Approximate influxes of micrometeoritic sodium and sulfur, assuming the fluxes are the same on Jupiter as on Earth and that micrometeorites are primarily cometary ices with cosmic abundances of other atoms, are $\sim 10^5 \text{ cm}^{-2} \text{ sec}^{-1}$ for Na, and $\sim 10^6 \text{ cm}^{-2} \text{ sec}^{-1}$ for S. Thus, the micrometeoritic infall rate of these atoms is likely to exceed that from the Io-associated toroidal clouds. In the upper troposphere with $T \sim 175^\circ\text{K}$, eqs. (3) and (9) then give $n \sim 10^4 \text{ Na atoms cm}^{-3}$, corresponding to a mixing ratio $f \lesssim 10^{-15}$, and even smaller at the water cloud levels. These values are small compared with the values of f_m in Table 1.

TABLE 1

The equilibrium condensation temperature, T_{con} , and estimates of the relevant mixing ratio, f , the radius a_f , the parameter η_{vap} and the minimum mixing ratio f_m for a number of substances above their condensation levels.

| | Fe | Mg ₂ SiO ₄ | SiO ₂ | Na ₂ Si ₂ O ₅ | NaCl | NH ₄ Cl | (Na ₂ S) |
|------------------------------------|-----|----------------------------------|------------------|--|------|--------------------|---------------------|
| $-\log_{10} f$ | 4.1 | 4.6 | 4.4 | 5.7 | 6.3 | 6.3 | (5.7) |
| $T_{\text{con}}/100^\circ\text{K}$ | 25 | 23 | 18 | 16 | 10 | 4 | 10 |
| $a_f/10^{-3}$ cm | 120 | 60 | 40 | 20 | 4 | 2 | 8 |
| η_{vap} | 60 | 20 | 2 | 1 | 0.1 | 0.03 | 0.4 |
| $-\log_{10} f_m$ | 7.9 | 7.7 | 7.2 | 7.8 | 7.0 | 6.4 | 7.0 |

IV. PHOTOPRODUCTION OF MOLECULES AND CHROMOPHORES

(a) Far UV:

At thermal equilibrium methane is by far the most abundant carbon-compound in the Jovian atmosphere, where it remains in gaseous form. Methane absorbs UV photons at all wavelengths shorter than 1600\AA and the total flux of such photons at the top of the Jovian atmosphere (mainly from solar Ly α at $\lambda = 1216\text{\AA}$) is $\sim 5 \times 10^9$ photons $\text{cm}^{-2} \text{s}^{-1}$ (global diurnal average). Very little of the Ly α photons are scattered by neutral hydrogen (Carlson and Judge 1974) and about one-third of the photons process CH_4 into other hydrocarbons (Strobel 1973a, 1975). The dominant product is ethane with a predicted production rate of $R \sim 1 \times 10^9 \text{ C}_2\text{H}_6 \text{ cm}^{-2} \text{s}^{-1} \sim 5 \times 10^{-14} \text{ gm cm}^{-2} \text{s}^{-1}$. The ethane produced is rather stable and is destroyed mainly by eddy diffusion into the troposphere, followed by pyrolysis in deeper, hotter layers. For an assumed mesosphere eddy diffusion coefficient $K_m(z)$, accurate numerical calculations for the number density $n(z)$ of C_2H_6 can be carried out (Strobel 1975). However, eqn. (9) is sufficient for purposes of orientation (the largest uncertainty comes from uncertainties in K_m). Assuming the tropopause level Z_p to occur at $P \sim 0.5 \text{ bar}$ (with $P \sim 10^{-6}$ to 10^{-7} bar at the production level Z_o), and with $K_m(z)$ given by eqn. (6), the maximum ethane number density occurs one scale height H above the tropopause:

$$n_{\text{max}} \sim \Phi H / K(Z_p + H) \sim 5 \times 10^{12} \text{ cm}^{-3} \sim 5 \times 10^{-7} n_{\text{H}_2} \quad (15)$$

where $K(Z_p + H)$ is the eddy diffusion coefficient evaluated at altitude $Z_p + H$.

The maximum value of the mixing ratio occurs just below the production level and the total predicted column density N of ethane molecules is

$$N \sim \Phi H^2 / K_m(Z_p) \sim 5 \times 10^{19} \text{ cm}^{-2} \sim 3 \times 10^{-3} \text{ gm cm}^{-2} . \quad (16)$$

Ethane (and smaller amounts of acetylene) have indeed been detected in the Jovian mesosphere by means of the emission spectrum at $\lambda \sim 10\mu$ (Ridgway 1974). High-resolution spectra are not yet available and the column density N of C_2H_6 from the present spectra depends strongly on the distribution of gas temperatures, which is not well-determined by low-resolution spectra: If the temperature in the mesospheric inversion layer turns out to have a maximum value of $150^\circ K$, the observations indicate (Ridgway 1974) $\sim 0.02 \text{ gm cm}^{-2}$ of ethane in this high temperature region. This would be in very serious conflict with the predictions, since only a small fraction of the theoretical column density in eqn. (16) refers to the hotter region many scale heights above the tropopause. Presumably one is dealing with some combination of (a) higher mesospheric temperatures; (b) smaller diffusion rates than in eqn. (6); or (c) some additional source of ethane raising the production rate R above $10^9 \text{ cm}^{-2} \text{ sec}^{-1}$.

Production rates from the same Ly α radiation in the mesosphere for larger, more complex organic molecules has been estimated (Strobel, 1975) to be factors of ~ 10 or 100 (by number) less than ethane, an estimate consistent with the data presented by Noyes and Leighton (1941) who remark that the ultraviolet irradiation of ethane leads to "a surprising variety of products," including unsaturated polymeric hydrocarbons. Assuming one-tenth the rate (by mass), the mass-production rate R_{org} is

$\sim 5 \times 10^{-15} \text{ gm cm}^{-2} \text{ s}^{-1}$ and eqn. (16) would give a column density for complex organics of $N_{\text{org}} \sim 3 \times 10^{-4} \text{ gm cm}^{-2}$.

Complex organic molecules may or may not be the mesospheric chromophores which have been inferred from the UV-albedos: The measured albedo of Jupiter (Wallace et al., 1972; Savage and Caldwell, 1974) to solar photons in the near UV is quite low with a minimum ~ 0.25 near $\lambda \sim 2800 \text{ \AA}$. Unit optical depth for pure (conservative) Rayleigh scattering by H_2 for 2800 \AA photons is reached at a level with $P \sim 0.3$ bar in the lower mesosphere. No simple molecules present near or above this level absorb in this wavelength region (NH_3 is effective only for $\lambda < 2,300 \text{ \AA}$) and complex molecules or very small absorbing dust grains have been invoked to provide the absorption (Axel 1972). Absorption in this region can be provided by a range of not implausible organic molecules (Sagan, 1968; Khare and Sagan, 1973). A mass absorption coefficient near 2800 \AA of $\sim (2N_{\text{org}})^{-1} \sim 10^3 \text{ cm}^2 \text{ gm}^{-1}$ would be required if organic molecules (beyond ethane) produced by Ly α photons are responsible. This is within the range of measured absorption coefficients for complex organic molecules produced by ultraviolet light in crude Jupiter simulation experiments (see, e.g., Khare and Sagan, 1973).

(b) Cloud models

The total thermal infrared flux emerging from Jupiter gives $T_{\text{eff}} \sim 127^\circ\text{K}$. Most of this flux emerges at wavelengths $\lambda \sim 20 \text{ }\mu\text{m}$ to $60 \text{ }\mu\text{m}$; the opacity source is mainly H_2 (with a small contribution from gaseous NH_3); and most of the radiation emanates from levels with $P \sim 0.7$ bar to 0.2 bar ($T \sim 150^\circ\text{K}$ to $\sim 115^\circ\text{K}$). Only a small fraction of the thermal flux emerges at wavelengths $\lambda < 15 \text{ }\mu\text{m}$, but the opacities of the main gaseous constituents of the Jovian atmosphere are low at some of these shorter infrared wavelengths. In the absence of any particulate matter (or complex

molecules), radiation with $\lambda \sim 8.2 \mu\text{m}$ to $9.5 \mu\text{m}$ would emerge from levels with $T \sim 155^\circ\text{K}$, and radiation with $\lambda \sim 5 \mu\text{m}$ from even deeper levels with temperatures up to $T \sim 300^\circ\text{K}$.

Assuming solar abundances of all elements in the Jovian interior and assuming thermochemical equilibrium, one can calculate condensation temperatures T_{con} for various species (cf. Lewis 1969a; Weidenschilling and Lewis 1973; Prinn and Owen 1976). T_{con} determines the level of each "cloud bottom," and above this level the gas-phase abundance is given by the saturation vapor pressure. The abundance and size-distribution of droplets (or crystals) in the upper layers of a cloud cannot be calculated without detailed knowledge of the dynamics (see Sect. III). The cloud bottoms for ice (and aqueous ammonia) are at levels with $T \sim 280^\circ\text{K}$, for NH_4SH at $T \sim 200^\circ\text{K}$, and for NH_3 at $T \sim 150^\circ\text{K}$ ($P \sim 0.7$ bar).

Measurements of spectral reflectivities for solar radiation in the visible range are now available separately for the bright zones and for the darker belts in the Jovian equatorial and temperate latitudes (Orton 1975c). The fraction of the sunlight that is absorbed varies from only ~ 0.2 in zones to ~ 0.3 in belts for red light, but varies from ~ 0.3 in zones to ~ 0.5 in belts for blue light. The data at $\lambda \sim 20\mu$ and 45μ , separately for zones and belts (Orton 1975c), are compatible with a model which includes an optically opaque cloudtop at the level ($T \sim 145^\circ\text{K}$, $P \sim 0.6$ bar) for zones, but not for belts. This model is not unique, but the postulated cloudtop occurs close to the level at which solid NH_3 crystals are expected to precipitate out if the zones represent the upward draft in the overall circulation. The spectral data for $\lambda \sim 8$ to $14 \mu\text{m}$

(averaged over some zones and belts), together with spatially resolved data at $\lambda = 8.11 \mu\text{m}$ and $8.45 \mu\text{m}$ (Orton 1975b), indicate the presence of some obscuring matter at altitudes above the $T = 150^\circ\text{K}$ level. The absorption spectrum of solid NH_3 crystals fits this data quite well, with more absorbing material over zones than over belts suggested by the data. Nevertheless, a small amount of absorption seems to be required over the belts as well, with an extent of only a few kms at levels with $T \sim 145^\circ\text{K}$ to 150°K indicated for this thin haze.

One question of primary concern is the nature and location of the red chromophores which are responsible for the absorption of visible sunlight over the belts (or at least for the preferential absorption in the blue). Upper limits to the location altitude can be inferred from the common observation that the red chromophores are sometimes overlain by time-variable white clouds. The relative heights are apparent because the interface between white and red material is frequently convex outward from the white clouds. As discussed, the highest lying clouds which are possible on Jupiter are ammonia cirrus clouds at levels with $T \sim 145$ to 150°K . The red chromophores must therefore reside mainly at deeper levels with $T > 145^\circ$ and $P > 0.6$ bar. Emission at $\lambda \sim 5\mu$ reveals infrared hotspots in only some regions of the belts (but in none of the zones) where the radiation emerges from deep levels with temperatures up to $T \sim 300^\circ\text{K}$. The red coloration is anticorrelated with these infrared hotspots (Keay et al., 1973; Westphal et al., 1974), which implies either that (a) the red chromophores reside above the level of some intermittent clouds which are opaque to 5μ radiation and which reside somewhere above the $T \sim 300^\circ\text{K}$ level; or (b) that the red chromophores are themselves intermittent and opaque at 5μ . The infrared hotspots are associated with visually dark

patches of bluish hue, probably connected with Rayleigh scattering rather than with blue chromophores, placing the blue regions at pressures of several bars, and $T \lesssim 280^\circ\text{K}$ (see Sagan 1971). To summarize: The red chromophores reside mainly in the troposphere ($T > 145^\circ\text{K}$), extending downwards from the tropopause ($T \sim 145^\circ\text{K}$) to at most the $T \sim 300^\circ\text{K}$ level.

The condensation temperature of solid NH_3 is only slightly higher than the temperature ($T \sim 145^\circ\text{K}$) expected for the tropopause, so that the solid NH_3 clouds reside in the upper layers of the troposphere. The pattern of patches of white clouds overlying the red chromophores changes on a time-scale of a few weeks. This is comparable with the total tropospheric eddy diffusion time t_c in eq. (3) if $\xi \sim 1$, but is also compatible with the longer total circulation time t_{cir} in Gierasch's model, since time-scales for patches of horizontal extent $d_p \ll D_{zb}$ (the separation between zones and belts) are only of order $(d_p/D_{zb})t_{\text{cir}}$. The postulated intermittent 5μ clouds underlying the red chromophores could in principle reside anywhere between the $T \sim 145^\circ\text{K}$ and $T \sim 300^\circ\text{K}$ levels, but on Gierasch's model these levels in belt-regions represent a continuous downdraft of warming material which could not form NH_4SH or H_2O crystals. Solid NH_3 clouds can be present in belts to varying degrees, because the upper horizontal circulation path from zones to belts can include layers just below the tropopause which carry NH_3 crystals. On this picture the red coloration would have to be provided by red chromophores which more or less coexist with NH_3 clouds ($P \sim 0.5$ to 0.7 bar), but with varying amounts of (i) high-level clouds with small NH_3 crystals, (ii) red chromophores, and (iii) clouds at lower levels with larger NH_3 crystals.

(c) Radiative transfer excursion

Besides the red chromophores we are also interested in inferred near UV-chromophores: The measured albedo of Jupiter (Wallace et al., 1972;

Savage and Caldwell 1974) to solar photons in the near UV ($\lambda \sim 2000 \text{ \AA}$ to 3500 \AA , say) is quite low, with a minimum albedo ~ 0.25 near $\lambda \sim 2800 \text{ \AA}$ (after rising slowly towards shorter λ , the albedo again seems to be lower for $\lambda < 2300 \text{ \AA}$). Absorption by H_2S , residing in the NH_4SH clouds at lower levels, has also been considered (Prinn 1970; Lewis and Prinn 1970; Sagan and Khare 1971b; Lewis 1976). A computational problem is often presented by absorbers whose abundance decreases very rapidly with increasing altitude (because of condensation); we briefly review the relevant radiative transfer theory:

Consider plane parallel geometry and assume the scattering is isotropic. Let τ be the total extinction optical depth at some level (scattering plus absorption), $\tilde{\omega}_0$ the single scattering albedo (ratio of scattering to extinction cross-section) at the same level, and $\kappa \equiv [3(1-\tilde{\omega}_0)]^{1/2}$. If κ is independent of τ , the Eddington approximation is known even for slabs of finite thickness (Chandrasekhar 1960, Irvine 1975) and is discussed in our Appendix. The mean intensity $J(\tau)$ as a function of depth is particularly simple for a semi-infinite slab of constant κ (see eqs. A3 to A5 and eq. A17). Unfortunately, κ can be a very steep function of τ as in the example in Fig. 2, which corresponds very roughly to absorption by H_2S and Rayleigh scattering by H_2 of 2500 \AA wavelength photons in the work by Prinn (1970). Prinn used eq. (A3) for $J(\tau)$ in the integrand of an integral which evaluates the absorbed energy layer by layer. Unfortunately, this use of eq. (A3) can lead to gross errors when κ is varying rapidly, and a more accurate numerical evaluation is given in our Appendix for the example in Fig. 2. Fortunately, the correct result for the fraction f of the incident energy which is absorbed is fairly close to the following simple approximation, f_0 : Let τ_0 be the total optical depth at the level where κ equals some predetermined

number, say $\kappa = 0.5$ (in our example $\tau_0 \approx 7.5$). Let f_0 be the fractional energy which would be absorbed according to the Eddington approximation if one had no absorption at all for $\tau < \tau_0$ and a perfect absorber at τ_0 . For diffuse incident illumination (instead of averaging over different angles of incidence) one has $f_0 = [0.75 \tau_0 + 1]^{-1}$ if $\tau_0 \gg 1$ (see eq. A16). For the numerical example in the Appendix the value of f_0 corresponding to the choice $\kappa = 0.5$ for τ_0 fits the correct f quite well.

(d) Near UV:

Solar photons in the near UV ($\lambda \approx 2000 \text{ \AA}$ to 3000 \AA , say) are much more abundant than Ly α -photons. It is of interest to speculate about materials which absorb in the near UV to see (i) if their absorption can explain the low observed albedo in the near UV; and/or (ii) if the end-products of reactions initiated by such absorption can provide the red chromophores for the troposphere; and (iii) how much complex organic material can be so produced.

The absorption of solar radiation with $\lambda \approx 2300 \text{ \AA}$ to 2700 \AA by H_2S in the upper layers of the troposphere has been discussed by the authors listed above (for $\lambda > 2700 \text{ \AA}$ the absorption by H_2S is too weak and photons with $\lambda < 2300 \text{ \AA}$ are already absorbed by NH_3 at higher altitudes). The total flux in this wavelength range (diurnal mean) is about 1×10^{13} photons $\text{cm}^{-2} \text{ s}^{-1}$ and some fraction f_s of these photons are absorbed by H_2S . We first present estimates if an ad hoc value of $f_s = 0.1$ is assumed [and $\xi = 1$ in eq. (3)].

Laboratory simulation experiments by near-ultraviolet irradiation of CH_4 , C_2H_6 , NH_3 , H_2O and H_2S mixtures at ~ 1 bar pressures have been performed, but without initial excess hydrogen (Sagan 1968; Sagan and Khare 1971 a,b; Khare and Sagan 1973, 1975). Hydrogen sulfide is the primary photon acceptor in these experiments. Amino acids were produced in these experiments with a quantum yield of $\sim 10^{-4}$ for molecular weights ~ 100 ($\sim 10^{-2}$ amu/photon). The quantum yield of all organics in such experiments is of the order of 100 times larger (of the order of 1 amu/photon). Some red chromophores were abundantly produced, consisting mainly of polymeric sulfur, but including a substantial fraction of organic compounds. The overall mass absorption coefficient for blue light of these chromophores is $k \sim 10^3 \text{ gm cm}^{-2}$. If such particles are produced in Jupiter at the rate of one S-bond per absorbed photon ($R \sim 10^{12} \text{ S cm}^{-2} \text{ s}^{-1} \sim 5 \times 10^{-11} \text{ gm cm}^{-2} \text{ s}^{-1}$), their column density, N , over a few scale heights down to the thick water clouds, is needed. According to eqns. (3) and (8), $N \sim 2RH^2/K_c \sim 10^{-4.5} \text{ gm cm}^{-2}$ and the optical depth to absorption in the blue is $\tau = kN \sim 0.03$. This is almost two orders of magnitude smaller than the observed optical depth. Moreover the reaction of 8 sulfhydryl radicals to form S_8 is poisoned by competing reactions, and the S_8 quantum yield must be significantly < 1 .

In the laboratory experiments cited, organic molecules, but not polymeric sulfur, were produced from superthermal hydrogen atoms. In the presence of a thousand times more H_2 (by number), the quantum yield might be expected to decrease by about the same factor, if energy loss of the fast atoms is mainly by elastic collisions, to $\sim 10^{-3}$ amu/photon. The total production rate would then be quite small, $R \sim 5 \sim 10^{-14} \text{ gm cm}^{-2} \text{ s}^{-1}$ (which is still ten times larger than our estimated R_{org} for the mesosphere). The corresponding optical depth for organic chromophores is, then, many

orders of magnitude below the observed value. In the laboratory experiments cited, large quantities of molecular hydrogen are produced during the irradiation because the end products are unsaturated and it may be that the correction factor from laboratory to Jovian situation is closer to 10^{-2} than to 10^{-3} . It is also possible that somewhat longer wavelength ultraviolet light is still effective in these syntheses, perhaps as long as 2900 \AA . Recent experiments (Ferris and Chen, 1975) have shown that 1849 \AA irradiation of mixtures of CH_4 and NH_3 in a 10:1 excess of H_2 produce a variety of organic molecules in very high yields, presumably by the generation of hot hydrogen atoms from ammonia photodissociation. These results suggest that the effects of H_2 dilution on poisoning organic photochemistry may be slower than linear. Hence the contribution of organic polymers to near-UV photoproduced optical frequency absorbers on Jupiter may be significant. Polymeric sulfur is not produced by hot hydrogen atoms in such experiments and so the scaling from laboratory to Jupiter is more direct in this case. However, even here, and even under the excessively optimistic assumption that every H_2S photodissociation event leads to polymeric S, the implied optical depth falls short by two orders of magnitude from matching the observed values; moreover, pure polymeric sulfur fits the observed optical properties of the Jovian red chromophores only poorly (Rages and Sagan, in preparation).

We have seen that a slow, deep circulation pattern is likely to extend all the way from the tropopause down to the level with $T \sim 300^\circ\text{K}$ (Gierasch 1976) with $\xi \sim 0.01$ in eq. (3). Abundances of reaction products are proportional to ξ^{-1} and also to $(f_s/0.1)$ which was claimed to be as large as 4 by Prinn (1970). One might then expect H_2S photolysis to dominate chromophore production, but this still seems unlikely for three reasons: (1) As discussed in Section IV (c) and in our Appendix, Prinn's (1970) estimate of $f_s \sim 0.4$ is

likely to be an overestimate since $f_o \equiv [0.75 \tau_o + 1]^{-1}$ is smaller. Thus $f_s \lesssim 0.2$ is more likely even in the absence of any scatterers (other than H_2) or other absorbers above the H_2S levels. The smallness of f_s is also suggested by the fact that the observed albedo of Jupiter does not increase from $\lambda \sim 2800 \text{ \AA}$ (where H_2S is ineffective) to $\lambda \sim 2600 \text{ \AA}$ (where H_2S absorption should already be strong). (2) Since H_2S cannot itself be the cause for the low albedo in this wavelength range, whatever absorbers are the cause must lower f_s further. (3) Most important, the systematic deep circulation pattern mentioned before depresses H_2S photolysis both in zones and belts: In the updraft, represented by zones, H_2S is abundant in the NH_4SH clouds but NH_3 clouds are above these layers and prevent most of the solar radiation from penetrating down. In the downdraft, represented by belts, the NH_3 clouds are absent (or at least diminished) but the material moving downward is practically devoid of H_2O and H_2S (which condensed out on the previous upward journey). H_2O and H_2S are again abundant in deep levels with $T \gtrsim 300^\circ K$ but little of the near UV penetrates this deep.

NH_3 absorbs solar photons of wavelength $\lambda < 2300 \text{ \AA}$ which are somewhat less abundant than photons between 2300 \AA and 2700 \AA . On the other hand, NH_3 should not suffer the shielding difficulties experienced by H_2S because it resides in the highest cloud layer which extends up to the tropopause. NH_3 should therefore be abundant in both zones and belts and even flows into the lower mesosphere at a non-negligible rate. It is therefore likely that most of the photons with $\lambda \sim 2300 \text{ \AA}$ are absorbed by ammonia (Strobel 1973b), which is compatible with the observed decline in albedo from $\lambda > 2300 \text{ \AA}$ to $\lambda < 2300 \text{ \AA}$. Hydrazine particles are also produced by ammonia photolysis and may provide a suitable explanation for the low albedo near $\lambda \sim 2800 \text{ \AA}$ (Prinn 1974). Moreover, organic molecules

produced by NH_3 photodissociation at $\lambda < 2300 \text{ \AA}$ could conceivably be the principal source of red chromophores, if the quantum yields of such reactions with $10^3:1 \text{ H}_2$ dilutions are as high as the experiments of Ferris and Chen (1975) suggest.

(e) Optical frequency photoproduction of chromophores

We here considered two possibilities: If the circulation in the upper part of the troposphere is as slow as Gierasch (1976) suggests, then $\xi \lesssim 0.01$ in eq. (3) and column densities obtained from eq. (8) are increased by two orders of magnitude; thus polymeric sulfur or complex organics might provide the red chromophores. But, as we have seen, there are serious problems with the contention that these molecules are produced by H_2S photolysis. These problems are avoided if shorter wavelength photolysis of NH_3 is an adequate source of complex organic chromophores, but this is a questionable proposition considering the lower photon flux at $\lambda < 2300 \text{ \AA}$, and whatever is the correct value of the poisoning of organic photochemical synthesis in a $10^3:1 \text{ H}_2$ dilution. If, on the other hand, ξ turns out to be of order unity, then ultraviolet photochemistry cannot in any case provide the observed chromophores at a sufficient rate. Thus, in both cases we are led to the possibility that photochemistry at optical frequencies, where more total energy is available, is involved in chromophore production. But no chemical bonds of even moderately abundant molecules can be broken by photons of wavelengths longer than 3500 \AA on Jupiter, from which it follows that optical frequency photochemistry must be at least a two photon process. However, such processes appear to be exclusively biological. We therefore find ourselves led unexpectedly to the hypothesis that the Jovian chromophores are biological in origin and that there is an abundant biota in the Jovian clouds.

Let us assume that the optical chromophores are contained primarily within Jovian organisms in the upper troposphere -- organisms driven by visible light photosynthesis, perhaps primarily in the blue, and utilizing the abundant methane, water and ammonia. Terrestrial algae (using CO_2 instead of CH_4) under laboratory conditions (Lehninger, 1971) have a quantum yield for synthesizing glucose of about 4 amu per photon (in the visible). The equivalent maximum production rate of biological material in the Jovian troposphere would be about $R = 3 \times 10^{-9} \text{ gm cm}^{-2} \text{ s}^{-1}$. With a mass absorption coefficient of $k \sim 10^3 \text{ cm}^{-2}$, such organisms would yield an optical depth of $\tau \sim 2$, slightly more than required by the observations. If the parameter ξ turns out to be of order 10^{-2} , and the red chromophores are due to Jovian photosynthetic organisms, R and/or K would have to be smaller.

Thus Jovian organisms having metabolic and photosynthetic parameters typical of terrestrial algae are able to account for the optical chromophores on Jupiter. Such organisms must however be adapted to the Jovian environment, and in particular must maintain a steady state population in the face of convective pyrolysis. Likely ecological niches for such organisms, as determined by the hydrodynamics on Jupiter, are discussed in the following sections. The following discussion can, however, be treated independently of the hypothesis that Jovian chromophores are biogenic.

V. SINKERS AND FLOATERS

If there is a Jovian biology, we would expect it to fill a rich variety of ecological niches. The best terrestrial analogy seems to be the surface of the sea. Oceanic phytoplankton inhabit a euphotic zone near the ocean surface where photosynthesis is possible. They are slightly denser than seawater and passively sink out of the euphotic zone and die. But such organisms reproduce as they sink, return some daughter cells to the euphotic zone through turbulent mixing, and in this way maintain a steady state population. (The early stages of sinking move the phytoplankton from a region of depleted resources into a region of abundant nutrients, thus stimulating replication).

A more elaborate adaptation is provided by fish and other organisms with float bladders which use metabolic energy to maintain a habitat at suitable pressure levels. They are generally not photosynthetic autotrophs, but rather heterotrophs living off organic molecules produced by autotrophs. A third ecological niche is filled by marine predators, one step further up the food chain, which hunt heterotrophs.

In the following discussion we will consider three comparable ecological niches on Jupiter: The primary photosynthetic autotrophs, which must replicate before they are pyrolyzed, will be described as sinkers. A second category of larger organism which may be either autotrophs or heterotrophs but which actively maintain their pressure level will be described as floaters. A third category of organisms actively seek out

other organisms; we call these hunters, although, as we shall see, the distinction between hunting and mating under these conditions is not sharp. Finally, there is a category of organisms which live almost at pyrolytic depth. They are scavengers, metabolizing the products of thermal degradation of other organisms. For our purposes, these pyrolytic scavengers are identical with floaters. In the following two sections we estimate the growth time t_{gr} required for an organism to double its mass.

We first consider passive organisms with positive excess density $\Delta\rho_a$ above the gas density ρ_g . Such organisms falling under gravity through the atmosphere rapidly acquire a vertical drift velocity, w_d , which depends on location and the size of the organism and increases with increasing $\Delta\rho_a$. There is a premium on keeping $\Delta\rho_a$ and hence w_d small and we consider only organisms in the shape of thin, gas-filled balloons. Such organisms have been proposed briefly before both for Jupiter (Shklovskii and Sagan, 1966) and for Venus (Morowitz and Sagan, 1967).

For simplicity we treat a spherical shell of outer radius a and skin-thickness $d \ll a$. In general we consider organisms filled with gas of the same temperature and mean molecular weight (2.2) as the surrounding ambient atmosphere. The lack of buoyancy is then controlled by the excess averaged density due to the skin, $\Delta\rho_a = (3d/a)\rho_g$ where we assume the biological skin material to have unit density (1 gm cm^{-3}).

For sufficiently small radius a of an organism, the Stokes formula in eqn. (11) holds (with $\rho_{dr} \rightarrow 3d/a$) for the drift velocity w_d and the drag

coefficient $C_D \sim 12\nu/aw_d$ is large. Let a_{Re} be a critical radius for which this relation gives $C_D = 1$ (Reynolds number = 24), so that

$$\left. \begin{aligned} w_d &= \frac{2}{3} \frac{gad}{\nu\rho_g} \text{ for } a \ll a_{Re} , \\ a_{Re} &= (6\nu^2\rho_g/gd)^{1/2} \end{aligned} \right\} \quad (17)$$

For large organisms with $a \gg a_{Re}$ (large Reynolds number) the flow around the body becomes turbulent and the drag coefficient C_D (\equiv drag force/ $0.5\rho_g w_d^2$) is close to 0.5 (see, e.g., Batchelor, 1970). In this regime

$$w_d \sim 4(gd/\rho_g)^{1/2} \quad (18)$$

We shall have to consider separately organisms living below and above the tropopause. The troposphere provides conditions closest to terrestrial ones (similar temperatures and abundant water in liquid or vapor form). We assume that the organisms can survive there over a range L of vertical height ($L \sim 5H$, say), centered approximately on the level with $T \sim 300^\circ\text{K}$. An important timescale is then a nominal drift-time $t_d \equiv L/w_d$. We consider the skin-thickness d as an unknown parameter with $d \sim 10^{-4}$ cm possibly a typical value (biological unit membranes on Earth have $d < 10^{-5}$ cm, but mechanical strength is required for the skin). Using eqns. (1) and (11) and evaluating coefficients at the level with $T = 300^\circ\text{K}$, we find

$$\left. \begin{aligned} t_d &\sim \frac{L}{5H} \left(\frac{T}{300^\circ\text{K}} \right)^{1.5} \frac{1 \text{ cm}}{a} \frac{10^{-4} \text{ cm}}{d} \quad 1.3 \times 10^4 \text{ sec} \\ \text{for } a < a_{Re} &\sim \left(\frac{10^{-4} \text{ cm}}{d} \frac{300 \text{ K}}{T} \right)^{1/2} \quad 0.039 \text{ cm} \end{aligned} \right\} \quad (19)$$

For $a > a_{Re}$, on the other hand, eqn. (18) gives a drift-time which depends on d but not directly on radius a ,

$$t_d \sim \frac{L}{5H} \left(\frac{T}{300^\circ K} \right)^2 \left(\frac{10^{-4} \text{ cm}}{d} \right)^{1/2} \approx 2.5 \times 10^5 \text{ sec} \quad (20)$$

The eddy diffusion coefficient K_c in the convective troposphere is fairly large and we also have to consider the contribution of eddy diffusion to the organism's downward motion over a vertical distance L . We are interested in cases where L is larger than H and the effective circulation time t_{cir} for downward diffusion is then not given by L^2/K_c : With $L \gg H$ and K_c a slowly-varying function of height, the first derivative term in eqn. (7) can be neglected and eddy diffusion has the same effect as increasing w_d to $(w_d + H^{-1}K)$. The relevant time scale is then linear in L and is approximated by the smaller of t_d and the quantity

$$t_{cir} = \frac{L}{H} t_c = \frac{LH}{K_c} \sim \frac{L}{5H} \left(\frac{T}{300^\circ K} \right)^{5/3} \approx 3.7 \times 10^6 \text{ sec}, \quad (21)$$

where we have used eqn. (3) and assumed $\xi \sim 1$. If $d \sim 10^{-4}$ cm, then eddy diffusion competes with gravitational fall only for microorganisms with $a \lesssim 30 \mu\text{m}$. Thus, Jovian organisms the size of small terrestrial protozoa and prokaryotes have typical times for falling through the troposphere to pyrolytic depths of one to two months. To maintain a steady state population, they must only replicate in that time scale.

In the Jovian mesosphere the temperature is low (we assume $T \sim 150^\circ\text{K}$) and water is absent, but some biologies have nevertheless been envisioned for such environments (see, e.g., Pimentel et al., 1966; Sagan, 1970). We shall see that the lower, denser layers of the mesosphere are the most advantageous for the growth of such organisms; eddy diffusion is negligibly slow here compared with the downward drift velocity w_d . For organisms operating between a lower pressure P and the tropopause (pressure ~ 1 bar), the drift-time L/w_d becomes

$$\left. \begin{aligned} t_d &\sim \ln \left(\frac{1 \text{ bar}}{P} \right) \frac{1 \text{ cm}}{a} \frac{10^{-4} \text{ cm}}{d} \times 930 \text{ sec for} \\ a < a_{\text{Re}} &\sim \left(\frac{10^{-4} \text{ cm}}{d} \frac{1 \text{ bar}}{P} \right)^{1/2} 0.044 \text{ cm} \end{aligned} \right\} \quad (22)$$

For larger organisms, $a > a_{\text{Re}}$, the expression for large Reynolds number gives

$$t_d \sim \ln \left(\frac{1 \text{ bar}}{P} \right) \left(\frac{10^{-4} \text{ cm}}{d} \frac{P}{1 \text{ bar}} \right)^{1/2} \times 1.6 \times 10^4 \text{ sec} \quad (23)$$

To obtain the pyrolysis time scale this number must be added to the appropriate tropospheric time scale given by eqs. (19), (20) and (21).

The expressions given above for t_d hold only for balloon organisms with ambient atmosphere inside. For organisms capable of pumping gas, buoyancy can be achieved while maintaining pressure equilibrium by keeping the interior gas pure hydrogen with molecular weight $\mu = 2.0$ instead of the ambient hydrogen-helium mixture with $\mu = 2.2$. A pumped organism can thus float if $\Delta\rho_a = (3d/a)\rho_g < 0.1\rho_g$. This requires radii a larger than a

threshold-radius a_{fl} for floating, given by

$$a_{fl} \sim \left(\frac{d}{10^{-4} \text{ cm}} \right) \left(\frac{T}{150^\circ \text{K}} \right) \left(\frac{1 \text{ bar}}{P} \right) \times 16 \text{ cm} \quad (24)$$

Instead of maintaining a hydrogen interior, an organism could use a fraction of its metabolic energy release to heat its interior gas. The temperature difference ΔT between interior and exterior is determined by the metabolic rate and the thermal diffusion coefficient of the gas. Except at very high pressures, the thermal diffusion is fast enough to keep $\Delta T/T \ll 0.1$, so that smaller values of $\Delta \rho_a$ and larger radii than a_{fl} are required. Buoyancy from a hot interior is thus likely to be less advantageous than pumping out helium. Note that Jovian floaters are macroscopic organisms (eq. (24)). Specialized floaters, with a range of organ systems, might have an effective value of $d \sim 1 \text{ cm}$ as terrestrial reptiles and mammals do; in that case, floaters would have kilometer dimensions, a size within the resolution capability of the Mariner Jupiter/Saturn flyby imaging system. The existence of very large floaters is, however, limited by metabolic constraints described in the next section.

VI. GROWTH OF PASSIVE SINKERS AND FLOATERS

We defer for the moment further questions of buoyancy and of powered locomotion, and consider a range of metabolic niches. Organisms -- sinkers or floaters -- may be photosynthetic autotrophs, obtaining their free energy directly from sunlight. They may also be heterotrophs of which we distinguish two classes: (i) those which feed on organic molecules of non-biological origin (e.g., ethane, an ultraviolet photoproduct), which reach the organism by molecular diffusion; and (ii) organic matter in smaller organisms or their fragments which reach the larger organism because of the difference in vertical drift velocity w_d between predator and prey. Active hunters described in the following section are a related subcase. In all cases we must estimate the growth-time $t_{gr} \equiv [2 d \ln a / dt]^{-1}$ and compare it with the drift-time t_d (or with t_{cir} if that should be smaller). We shall see that the ratio t_{gr} / t_d increases with the radius a of the organism; we are interested in the maximum radius a_{max} for which this ratio is unity. If a_{max} is appreciably larger than the minimum size a_{min} for a particular life-style, then a biology can be maintained by an organism fragmenting into many smaller organisms or dispersules before it reaches size a_{max} , i.e., before it has drifted downwards to a pyrolytic death. If $t_d \gg t_{cir}$ for the dispersules, they can be circulated upwards, grow to a_{max} , and complete the life-cycle.

We saw in Section IV that the most favorable region for direct photosynthesis is probably the upper troposphere, just above the water clouds (with $T \sim 300^\circ K$,

$P \sim 5$ bar). Methane, ammonia and water building blocks are abundantly available there to a photosynthesizing organism and the growth-rate is controlled by the deposition of free energy from solar photons. Furthermore, the observed optical chromophores can be explained by such organisms at such a level. We saw that a value for the production rate of $R \sim 3 \times 10^{-9} \text{ gm cm}^{-2} \text{ s}^{-1}$ seems reasonable and we find for the growth-time

$$t_{gr} \equiv \frac{a/2}{da/dt} = \frac{4d}{R} = \frac{d}{10^{-4} \text{ cm}} \frac{3 \times 10^{-9}}{R} \times 1.3 \times 10^5 \text{ sec} \quad (25)$$

Equating this time with t_d in eq. (19) then gives for a_{max} , the maximum radius to which an organism can grow

$$a_{max} \sim (10^{-4} \text{ cm/d})^2 0.1 \text{ cm} \quad (26)$$

Values of $a_{min} \lesssim 10^{-3}$ seem perfectly feasible, so that one organism can produce more than 10^4 dispersules for which $t_d \ll 4 \times 10^6 \text{ sec} \sim t_{cir}$. The steady-state population of such organisms would then be controlled only by competition for sunlight, and not by gravitational fallout to pyrolytic depths.

Consider next the diffusion onto the skin of an organism of some organic molecule with molecular weight μ and abundance in the surrounding atmosphere of ρ_0 (in gms cm^{-3} of atmosphere). Ethane is the lowest mass ($\mu = 30$) organic molecule which carries free energy (relative to methane, the prevalent carbon-molecule at thermal equilibrium). The diffusion coefficient D of ethane in the Jovian H-He mixture is about 0.4 times the viscosity ν of this mixture (Strobel, 1973) and for heavier molecules D scales reciprocally as the cross-section of the molecule. For approximately spherical organic molecules we adopt

$$D \sim 0.4\nu (\mu/30)^{-2/3}, \quad (27)$$

with v given by eq. (11). If $v > D \gtrsim (aw_d/12)$, the diffusion is almost independent of the motion of the sphere and the steady-state rate for diffusion onto the surface is $4\pi a^2 (\rho_0 D/a)$. Let ϵ be the ratio of increased body-mass of the organism to the mass of ingested metabolites. Assuming unit density for the skin-material the growth-time is then

$$t_{gr} = ad/2\epsilon \rho_0 D. \quad (28)$$

At a mesospheric level with pressure P (in bars), for molecules produced at a mass rate of $R = \mu\Phi$ at some higher level, eqs. (6) and (9) give $\rho_0 \sim (\mu\Phi H/100) P \ln P^{-1}$. The growth-time in the mesosphere is then

$$t_{gr} \sim \frac{a}{1 \text{ cm}} \frac{d}{10^{-4} \text{ cm}} \frac{(\mu/30)^{2/3}}{\ln(1 \text{ bar}/P)} \left[\frac{5 \times 10^{-14}}{\Phi} \right] \frac{0.03}{\epsilon} \times 1.8 \times 10^6 \text{ sec.} \quad (29)$$

For ethane an efficiency factor $\epsilon \sim 0.03$ seems reasonable and $R \sim 5 \times 10^{-14} \text{ gm cm}^{-2} \text{ s}^{-1}$ has already been discussed; for UV photoproduction of more complex organic molecules we may leave $\epsilon \sim 0.3$ and $R \sim 5 \times 10^{-15} \text{ gm cm}^{-2} \text{ s}^{-1}$, so that $R\epsilon$ is the same but the factor $\mu^{2/3}$ makes heavier molecules less favorable. Replacing the slowly-varying factor $\ln(1/P)$ by 3 and $R\epsilon$ by 1.5×10^{-15} we find from eqs. (22) and (29),

$$a_{max} \sim (10^{-4} \text{ cm/d}) (\mu/30)^{-1/3} \times 0.07 \text{ cm} \quad (30)$$

If $\Phi\epsilon$ is particularly large and/or d particularly small, an organism can grow beyond size a_{Re} and the diffusion rate now depends on the drift-speed w_d , given by eq. (18). We omit the complicated transition region $a \sim a_{Re}$ and make only order of magnitude estimates: There is a boundary layer in the flow around the sphere of thickness $b \sim a (4\nu/a w_d)^{1/2} \ll a$, such that the flow-speed a normal distance y

from the body is $\sim (y/2b) w_d$. With $D < v$ the typical stand-off distance for diffusion onto the body-surface is $b_D \sim 2b (D/v)^{1/3}$ and the diffusion rate is $\sim 4\pi a^2 (\rho_0 D/b_D)$, so that

$$t_{gr} \sim (2d/\epsilon\rho_0) (a/Dw_d)^{1/2} (v/D)^{1/6}. \quad (31)$$

For the mesosphere this gives

$$t_{gr} \sim \left(\frac{a}{1 \text{ cm}}\right)^{1/2} \left(\frac{d}{10^{-4} \text{ cm}}\right)^{3/4} \left(\frac{1 \text{ bar}}{P}\right)^{1/4} \frac{(\mu/30)^{4/9}}{\ln(1/P)} \left[\frac{5 \times 10^{-14}}{\Phi} \frac{0.03}{\epsilon}\right] \times 8 \times 10^4 \text{ sec} \quad (32)$$

The same basic formulae apply in the troposphere, but ρ_0 is now much smaller for the same value of Φ because of the larger eddy diffusion coefficient, K_c . On the other hand, we have seen that photosynthesis may be going on there and this may release complex organic molecules at a faster rate Φ . Substituting $\rho_0 = RH/K_c$ and using eq. (3), we find (for $a < a_{Re}$)

$$t_{gr} \sim \frac{a}{1 \text{ cm}} \frac{d}{10^{-4} \text{ cm}} \left(\frac{T}{300^\circ\text{K}} \frac{\mu}{30}\right)^{2/3} \left[\frac{3 \times 10^{-9}}{\Phi} \frac{0.3}{\epsilon}\right] 8 \times 10^7 \text{ sec}. \quad (33)$$

Comparison with eq. (19) shows a_{max} to be of order $(10^{-4} \text{ cm/d}) (30/\mu)^{1/3} 0.013 \text{ cm}$, which is just macroscopic.

If small tropospheric photosynthetic organisms have radii $a_s \lesssim 0.01 \text{ cm}$, their drift speeds are even smaller than the eddy circulation speed and their abundance is still given by $\rho_0 = RH/K_c$. A larger organism of radius a , drifting downwards with speed w_d , can grow by coalescing with such small organisms if the encounter satisfies certain conditions. These conditions are similar to those governing the coalescence of raindrops (Mason, 1972; Rossow, 1975). In analogy with eq. (13) one finds, independent of a_s , and provided $d \ll a_s < a$, the requirement that $a > (10^{-4} \text{ cm/d}) (T/300^\circ\text{K}) 0.002 \text{ cm}$. If these conditions are satisfied, the larger organisms incorporate mass from the smaller ones at a rate of $4\pi a^2 w_d \rho_0$ so that

$t_{gr} \sim (4d/R) (K_c/Hw_d)$. If $a \gtrsim a_{Re} \sim 0.04$ cm, then

$$t_{gr} \sim \left(\frac{d}{10^{-4} \text{ cm}}\right)^{1/2} \left(\frac{T}{300}\right)^{1/3} \left[\frac{3 \times 10^{-9}}{\phi} \frac{0.3}{\epsilon}\right] 2.5 \times 10^4 \text{ sec.} \quad (34)$$

Neither t_{gr} nor t_d in eq. (20) depend explicitly on radius a . An organism can then grow indefinitely as long as $(d/10^{-4} \text{ cm}) \lesssim (T/300^\circ\text{K})^{5/3}$.

Thus, provided the membrane thickness stays smaller than about $1 \mu\text{m}$, such passive sinkers can grow by coalescence with smaller organisms to sizes larger than a_{fl} as given by eq. (24). This provides one evolutionary pathway from sinkers to floaters. However, the condition on membrane thickness is very restrictive. Floaters of this sort are unlikely to be much more stable than soap bubbles and the organ systems which large organisms require could not be accommodated with the thin skin thickness. However, as we shall see in the following section, the hunting lifestyle provides a way out of this dilemma.

To summarize the situation so far: Sinkers can grow and produce reproductive dispersules in a stable steady state life cycle if they either (a) are mesospheric heterotrophs or photoautotrophs, or (b) are tropospheric photoautotrophs. Passive floaters exist only if (1) their skin can be made particularly thin, and (2) a significant fraction of the smaller photoautotrophs is available as food.

VII. POWERED LOCOMOTION AND COALESCENCE

We saw that the efficiency of utilization of metabolites, diffusing onto the surface of a freely falling organism, decreases with increasing mass of the molecule. However, for even more massive metabolites and organisms approaching each other with some relative drift velocity, growth by coalescence is again moderately efficient. The growth-rate by coalescence can be increased greatly if organisms are capable of powered locomotion and can steer toward each other. We shall not discuss the predator-prey relationships of a strict hunting doctrine, but rather only an ecology in which an organism approaches another, coalesces into one larger organism which approaches other organisms, etc. It is clear that this hierarchical mating doctrine can also serve the traditional terrestrial function of mating, namely, the exchange of genetic material. A strict hunting doctrine in which one of the partners in such a coalescence is non-cooperative will be less efficient than the doctrine discussed here. For convenience, however, we will describe both doctrines as hunting. In the following discussion we will assume that only one partner in such a mating event is active in the search; for this reason mating, hunting prey, and hunting abiological organic molecules are not extremely different. The doctrine in which both organisms are active in seeking partners would be more efficient than the ones described below.

For simplicity we assume the presence with number-density n , of organisms with radii of order a_1 at some upper injection-level $z = Z_1$. The organisms coalesce and grow as they drift downwards, so that the size distribution changes with height z . Let $n(a, z)$ and $\rho(a, z) = (4\pi a^2 d) n(a, z)$ be their steady-state number-density

and mass-density, respectively, per logarithmic interval of radius a . In principle we have a relation like eq. (7) separately for each value of a , with the right-hand side representing the flux from one size range to another. We shall make only order of magnitude estimates for $\bar{a}(z)$, the mean radius of organisms at level z , and for $n(z) \equiv \int d\ln a n(a, z)$.

We assume that the organisms are sufficiently large that $w_d(z)$, the drift-velocity for radius $\bar{a}(z)$ at height z , is much larger than K/H in eq. (7). Multiplying by mass (assuming as usual unit density for the skin-material) and integrating over sizes, we find

$$[4\pi a^{-2}(z) d] w_d(z) n(z) = \epsilon(z) R \quad (35)$$

In this relation R is the constant rate in $\text{gm cm}^{-2} \text{s}^{-1}$ at which mass in the form of organisms enters the upper level $z = Z_1$; $\epsilon(z)$ is a slowly-varying factor (slightly smaller than unity) which takes account of any inefficiency in the utilization of biological material, due to waste or expenditure of free energy. We shall have to evaluate the growth-time $t_{gr}(z) \equiv [2 d\ln a/dt]^{-1}$ for organisms with $a \sim \bar{a}(z)$ at level z . The function $\bar{a}(z)$ is then given in terms of t_{gr} by

$$d\ln \bar{a}/dz = [2 t_{gr} w_d(z)]^{-1} \quad (36)$$

For the integrated column-density $N(a) \equiv \int dz n(a, z)$ we also find $N(a) \sim 2 t_{gr} w_d n(\bar{z})$, where \bar{z} is such that $\bar{a}(\bar{z}) = a$. The total mass contained in organisms in different size ranges is then given by

$$(4\pi a^2 d) N(a) \sim 2 t_{gr}(a) \epsilon(z) R. \quad (37)$$

Assume that an organism, while drifting downward with speed w_d , can expend enough mechanical energy to acquire a horizontal velocity component of up to θw_d where θ is some angle less than 1 radian. Its vertical speed relative to other

similar organisms is about $1/2 w_d$ and in time t a volume $(\theta^2/6) (w_d t)^3$ is accessible to a searching organism. When $n(z)$ times this volume is unity, the organism has a good probability of coalescing with a mate and doubling in mass. With $n(z)$ given by eq. (35), the growth-time is then

$$t_{gr} \sim \frac{4}{\ln 2} \left(\frac{3\pi \bar{a}^2 d}{\epsilon R \theta^2 w_d^2} \right)^{1/3} \quad (38)$$

If $\bar{a} > a_{Re}$, we have, independent of d both for the troposphere and the mesosphere,

$$t_{gr} \sim \left[\frac{1}{\theta^2} \left(\frac{a}{1 \text{ cm}} \right)^2 \frac{175^\circ \text{K}}{T} \frac{P}{1 \text{ bar}} \frac{3 \times 10^{-9}}{R} \frac{0.3}{\epsilon} \right]^{1/3} 20 \text{ sec.} \quad (39)$$

For such organisms the troposphere is quite favorable even in the absence of photoautotrophs as prey ($R \sim 5 \times 10^{-14} \text{ g cm}^{-2} \text{ s}^{-1}$); for instance, with $T \sim 300^\circ \text{K}$ and $d \sim 10^{-4} \text{ cm}$, we have $a_{fl} \sim 6 \text{ cm}$ and $a_{max} \sim (\theta/0.002) 6 \text{ cm}$, so that buoyancy can be achieved with quite small values of θ . For $R \sim 3 \times 10^{-9} \text{ g cm}^{-2} \text{ s}^{-1}$,

$a_{fl} \sim 6 \text{ cm}$ requires only $a_{max} \sim (\theta/3 \times 10^{-6}) 6 \text{ cm}$; microradian maneuverability suffices.

The maximum sizes of organisms embracing the hunter doctrine as a function of the maneuverability, θ , and the source of food is given in Table II. We see that very large hunters are permitted.

To derive an explicit relation between typical radius a and vertical level \bar{z} (a), we must integrate eq. (36) with t_{gr} given by eq. (39). The result depends on whether θ is constant or a function of a . Neglecting the variations of θ , w_d and ϵ gives (for a much larger than the original size a_1 at level Z_1)

$$Z_1 - \bar{z}(a) = 3 w_d(\bar{z}) t_{gr}(a) \quad (40)$$

We still have to calculate the mechanical energy expended by an organism in maintaining its horizontal hunting speed θw_d . In pure downward drift the rate of energy dissipation equals the rate of gravitational energy release $mg w_d$. The

TABLE II

Maximum Hunter Sizes

| | Source of Metabolites | |
|----------------------------|-------------------------------|-----------------|
| | Abiological organic matter | Photoautotrophs |
| θ | | |
| 2×10^{-3} radians | 6 cm | 40 m |
| 10^{-1} | 3 m | 2 km |
| 1 | 30 m | 20 km |

equivalent rate for maintaining the horizontal speed (with $\theta < 1$) is at most θ mg w_d (with judicious sailing and searing it could be appreciably less). The relevant quality factor is the ratio $\Delta E/\Delta m$ of mechanical energy expended in finding a partner to the increase in mass of biological material. (Biochemical energy storage is of order 10^{11} erg/gm and a mechanical quality factor $\sim 10^9$ erg/gm is a very modest assumption.) The heat of combustion (full oxidation) of molecular hydrogen is about ten times that of glucose and more than three times that of such fats as stearic acid. It is not the unavailability of reduced compounds but rather the unavailability of oxidized compounds on Jupiter which probably sets limits on the efficiency of internal metabolism of hypothetical Jovian organisms. We have

$$\frac{\Delta E}{\Delta m} \lesssim \theta g w_d t_{gr}, \quad (41)$$

which depends quite weakly on temperature and pressure. For the upper troposphere

$$\frac{\Delta E}{\Delta m} \lesssim \left(\frac{a}{1 \text{ cm}}\right)^{2/3} \left(\frac{d}{10^{-4} \text{ cm}}\right)^{1/2} \left[\theta \frac{3 \times 10^{-9}}{R} \frac{0.3}{\epsilon}\right]^{1/3} 5 \times 10^6 \frac{\text{erg}}{\text{gm}} \quad (42)$$

represents quite a modest power expenditure except for the largest organisms. A 100 m radius hunter of photoautotrophs with $\theta = 0.01$ radian excursion capability and a 1 cm skin thickness requires a quality factor of 10^{11} erg/gm. The same quality factor is required by a 1 m radius hunter of abiological organic matter with $\theta = 0.03$ radians and $d = 1$ cm.

Even smaller power expenditures will be required if, as would be likely, evolution selects organisms with sensory systems which increase the efficiency of hunting food -- such as, e.g., optical sensors for detecting chromophores. Likewise, acceleration sensors capable of detecting convective "thermals" would ease the restrictions on floater sizes and energy expenditures; and habitats near the boundaries between upward and downward convective regimes would improve the cost of Jovian biology.

The path from sinkers to hunters to floaters is an evolutionary path. It provides a possible sequence for the evolution of floaters which, as we have seen, would

be difficult to understand directly from sinkers. Once floaters are established, they may be able to replicate without passing through the sinker stage at any point in their life cycles. For this to occur, of course, the minimum radius of their dispersules must be larger than a_{f1} .

While we have distinguished passive sinkers, hunters, and floaters in this discussion, it is clear that a single highly evolved Jovian organism might partake qualities of all of these lifestyles -- for example, at different stages in the life cycle. Indeed, it may be that only after the hunting doctrine is adopted by sinkers can growth to sizes large enough for the evolution of floaters become possible. We have in this discussion made no distinction among various locales on Jupiter; but it is clear that some locales -- the Great Red Spot, for example -- may be more favored than others because of higher abundances of organic molecules, prevailing updrafts or other reasons.

Among other objects in the outer solar system which possess atmospheres, only Titan and Saturn exhibit colorations similar to those on Jupiter. Because of the absence of convective pyrolysis on Titan, the chromophores there can quite readily be ultraviolet photoproducted organic molecules (Sagan, 1973). However, the physical environment of Saturn is very similar to that of Jupiter in the respects relevant to the arguments of this paper and we, therefore, tentatively postulate an airborne biota on Saturn as well.

The test of these ideas is, of course, observational. Flyby, orbital, and entry probe spacecraft each can perform significant tests of these ideas. For example, entry gas chromatograph/mass spectrometers for the near future are anticipated to have detectivities of organic matter as good as about 10^{-10} gm cm⁻³. The predicted steady state density of complex organic molecules on Jupiter is RH/K_c and is roughly estimated, using the numerical values of this paper, both for the lower mesosphere and upper troposphere and for abiological and biological sources of organic matter in Table III. While most biological organic matter will be in organisms and

TABLE III

Steady State Density of Organic Molecules (g cm^{-3})

| | Abiological | Biological |
|-------------------|---------------------|---------------------|
| Lower Mesosphere | 2×10^{-11} | 9×10^{-6} |
| Upper Troposphere | 2×10^{-15} | 9×10^{-10} |

not in free molecules, the densities predicted in Table III are sufficiently large that a high sensitivity entry mass spectrometer may be a significant test of the biological hypotheses of the present paper. Because of the slow convective velocities of the mesosphere, such instruments should be capable of working at such pressure levels.

The possible existence of indigenous Jovian organisms is also relevant to the question of sterilization of spacecraft intended for entry into the atmosphere of Jupiter. Even if the ambient environment is inconsistent with the growth of microbial contaminants from Earth, the internal environment of Jovian organisms may be much more hospitable. The replication of terrestrial contaminants in the Jovian clouds also depends on the availability of trace elements. We have seen (Section III) that indigenous Na is probably missing from the upper troposphere; and that indigenous Mg must certainly be missing. Magnesium ions are essential for nucleic acid replication and a wide range of other biological functions; if Mg is missing from the clouds, the likelihood of terrestrial biological contamination of Jupiter becomes nil. However, exogenous sources of sodium, magnesium and other trace elements, particularly from micrometeorites, while small, cannot be neglected (Section III). Consequently, it seems judicious not to exclude prematurely the possibility of biological contamination of Jupiter by terrestrial microorganisms.

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APPENDIX

Consider a plane parallel slab of total optical thickness τ_0 and let $I(\tau, \mu)$ be the intensity of radiation at depth τ with μ the cosine of the angle between the propagation direction and the forward normal. Assume isotropic scattering and let $\tilde{\omega}_0$ be the single-scattering albedo, which may be a function of τ . The equation of radiative transfer (Chandrasekhar 1960; Irvine 1975) then reads

$$\mu \frac{d}{d\tau} I'(\tau, \mu) = -I'(\tau, \mu) + \tilde{\omega}_0 \frac{1}{2} \int_{-1}^1 d\mu' I'(\tau, \mu') + \tilde{\omega}_0 e^{-\tau/\mu_0} \quad (A1)$$

for an incident beam of unit intensity with direction cosine μ_0 , where I' is the intensity of the diffuse radiation and $e^{-\tau/\mu_0}$ that of the attenuated incident beam. The Eddington approximation can be obtained by substituting a function of a single variable $I'_+(\tau)$ for $I'(\tau, \mu)$ when $0 < \mu < 1$ and another function $I'_-(\tau)$ when $-1 < \mu < 0$. Let

$$\kappa \equiv [3(1 - \tilde{\omega}_0)]^{\frac{1}{2}} \quad (A2)$$

and let $J'(\tau)$ be the average over all μ of $I(\tau, \mu)$. For constant κ and a semi-infinite medium ($\tau_0 \rightarrow \infty$) one finds

$$\left. \begin{aligned} J'(\tau) &= ae^{-\kappa\tau} - be^{-\mu_0^{-1}\tau}, \\ a &= \frac{1 + (2/3)\mu_0^{-1}}{1 + (2/3)\kappa} \quad b, \quad b = \frac{3\tilde{\omega}_0}{\mu_0^{-2} - \kappa^2} \end{aligned} \right\} \quad (A3)$$

The effective reflection coefficient $R \equiv 1 - f$, where f is the fraction

of the energy flux that is absorbed, is given by

$$R = \frac{\tilde{\omega}_0}{(1 + \kappa\mu_0)[1 + (2/3)\kappa]} \quad (\text{A4})$$

Generalizations for non-isotropic scattering are discussed by Chandrasekhar (1960), Prinn (1970) and Irvine (1975). For isotropic scattering Prinn used eq. (A3) for normal incidence ($\mu_0 = 1$) in his numerical work. For general μ_0 but with $\kappa \ll 1$ eq. (A4) reduces to

$$f(\mu_0) \equiv 1 - R(\mu_0) \approx \kappa(2/3 + \mu_0) \quad (\text{A5})$$

For many purposes one is interested only in results averaged over a whole daylight-hemisphere; i.e., averaged over μ_0 from 0 to 1. Such averages are represented by the single problem we shall consider here, where the incident intensity $I_{\text{inc}}(\mu)$ equals unity for all μ between 0 and 1. In the spirit of the Eddington approximation we replace the problem by one for intensity $I(\tau, \mu)$ without explicitly isolating the attenuated incident beam: In analogy with eq. (A1) we have

$$\mu \frac{d}{d\tau} I(\tau, \mu) = -I(\tau, \mu) + \tilde{\omega}_0 \frac{1}{2} \int_{-1}^1 d\mu' I(\tau, \mu'), \quad (\text{A6})$$

which leads to the coupled equations

$$\left. \begin{aligned} \frac{dF(\tau)}{d\tau} &= -J(\tau)[1 - \tilde{\omega}_0(\tau)], \\ \frac{dK(\tau)}{d\tau} &= -F(\tau) \end{aligned} \right\} \quad (\text{A7})$$

where

$$\left. \begin{aligned} J(\tau) &\equiv \frac{1}{2} \int_{-1}^1 d\mu I(\tau, \mu), & F(\tau) &\equiv \frac{1}{2} \int_{-1}^1 d\mu \mu I(\tau, \mu), \\ K(\tau) &\equiv \frac{1}{2} \int_{-1}^1 d\mu \mu^2 I(\tau, \mu) \end{aligned} \right\} \quad (A8)$$

We again make the approximation (but here for all the radiation) that $I(\tau, \mu) = I_+(\tau)$ for $\mu > 0$ and $I(\tau, \mu) = I_-(\tau)$ for $\mu < 0$, so that

$$J = 3K = \frac{1}{2}(I_+ + I_-), \quad F = \frac{1}{2}(I_+ - I_-) \quad (A9)$$

In this problem, any explicit consideration of the incident beam is simply replaced by the boundary condition that $I_+(\tau = 0) = 1$ and the effective reflection coefficient $R = 1 - f$ is simply given by $I_-(\tau = 0)$.

For general $\tilde{\omega}_0(\tau)$ and with $\kappa(\tau)$ defined by eq. (A2), the use of the Eddington approximation (A9) in eq. (A7) leads to the eikonal equation

$$\left[\frac{d^2}{d\tau^2} - \kappa^2(\tau) \right] I_{\pm}(\tau) = 0 \quad (A10)$$

and to the further relation

$$2\kappa^2(I_+ + I_-) = 3\frac{d}{d\tau}(I_- - I_+) \quad (A11)$$

We first solve these equations for the case of constant κ for a slab of total thickness τ_0 : We assume that the effective reflection coefficient $R_0 \equiv I_-(\tau_0)/I_+(\tau_0)$ is given as one boundary condition, the other boundary condition for the incident beam being $I_+(0) = 1$. The general solution of the coupled equations (A10) and (A11) is

$$I_{\pm}(\tau) = a \left[(3 \pm 2\kappa)e^{-\kappa\tau} - \gamma(3 \mp 2\kappa)e^{\kappa(\tau - 2\tau_0)} \right] \quad (\text{A12})$$

with a and γ disposable constants. The requirement of $I_-/I_+ = R_0$ at the second boundary ($\tau = \tau_0$) leads to

$$\gamma = \frac{(3 - 2\kappa) - R_0(3 + 2\kappa)}{(3 - 2\kappa) - R_0(3 - 2\kappa)} \quad (\text{A13})$$

and the condition $I_+(0) = 1$ then gives a . The final expression for the intensity J (averaged over direction cosines) is then

$$J(\tau) = \frac{3 \left[e^{-\kappa\tau} - \gamma e^{\kappa(\tau - 2\tau_0)} \right]}{(3 + 2\kappa) - (3 - 2\kappa) \gamma e^{-2\kappa\tau_0}} \quad (\text{A14})$$

Another quantity of interest is the effective reflection coefficient R_i of the whole slab at the surface of incidence, which is $\propto I_-(0)$,

$$R_i = \frac{(3 - 2\kappa) - \gamma(3 + 2\kappa)e^{-2\kappa\tau_0}}{(3 + 2\kappa) - \gamma(3 - 2\kappa)e^{-2\kappa\tau_0}} \quad (\text{A15})$$

Two limiting cases are of interest for the solutions given by eqs. (A12) to (A15). Consider first the case with $R_0 = 0$ in the limit of $\kappa \rightarrow 0$ with τ_0 kept finite. In this case eqs. (A14) and (A15) reduce to

$$f \equiv 1 - R_i = \frac{4}{3\tau_0 + 4}, \quad J(\tau) = \frac{3(\tau_0 - \tau) + 2}{3\tau_0 + 4} \quad (\text{A16})$$

The other case is the limiting case of $\tau_0 \rightarrow \infty$ for any constant, non-zero κ . Independently of the value of R_0 , this case gives

$$f \equiv 1 - R_i = \frac{4\kappa}{3 + 2\kappa}, \quad J(\tau) = \frac{3e^{-\kappa\tau}}{3 + 2\kappa} \quad (\text{A17})$$

When $\kappa \ll 1$ the value of f in eq. (A17) agrees with the average over $0 < \mu_0 < 1$ of eq. (A5) which is $\bar{f} = (4\kappa/3)$.

No exact solutions of the eikonal equation (A10) exist for general $\kappa(\tau)$, but if κ varies sufficiently slowly, one is tempted to use the WKB approximation. For a semi-infinite medium, only one of the two WKB solutions survives and we have

$$J(\tau) \propto [\kappa(\tau)]^{-\frac{1}{2}} \exp\left[-\int_0^{\tau} d\tau' \kappa(\tau')\right] \quad (\text{A18})$$

Note that substituting a varying $\kappa(\tau)$ into eq. (A17) (or into eq. (5) of Prinn 1970, which is the equivalent of our eq. (A3) for normal incidence instead of diffuse illumination) is not a good approximation to eq. (A18) if κ varies appreciably even if the variation is slow. Furthermore, the requirement for the validity of the WKB approximation is that $(d\kappa/d\tau) \ll \kappa^2$, which is not satisfied by practical cases for $\kappa(\tau)$, such as the example in Fig. 2. An inelegant but reliable numerical method proceeds as follows: Subdivide the total slab into a large but finite number of thin slabs (with $\Delta\tau \sim 0.3$, say) and replace $\kappa(\tau)$ for each thin slab by its average value over this slab. Start with the last thin slab, assume $R_0 = (3 - 2\kappa)/(3 + 2\kappa)$ and evaluate eqs. (A13) to (A15) with the slab-thickness τ_0 replaced by $\Delta\tau$. The value for R_i obtained from eq. (A15) for the last slab is then used as R_0 for the penultimate thin slab and J and R_i are again evaluated, and so on iteratively until the first thin slab. The solid curve labeled J is the result of such a numerical solution for $J(\tau)$ and the effective reflection coefficient R_i for the whole medium (obtained at $\tau = 0$) is $R_i \approx 0.85$. The dashed curve is the result for $J(\tau)$ in a different problem -- no absorption for $\tau < \tau_0$ and complete absorption for $\tau > \tau_0$, with $\tau_0 \approx 7.5$ (the value

of τ for which $\kappa = 0.5$ in the actual problem). Eq. (A16) gives the solution for $J(\tau)$ for this problem and also yields $R_i \approx 0.85$. Eq. (A16) for R_i should be a fairly good solution for any $\kappa(\tau)$, with τ_0 the value where $\kappa \approx 0.5$, as long as $\tau_0 \gg 1$ and κ is small for τ appreciably less than τ_0 .

The accuracy of the numerical solution described above is limited (apart from computational errors) by errors inherent in the Eddington approximation, including the use of a constant coefficient in eq. (A11). For a single slab this approximation is excellent when $\kappa \ll 1$, although other two-stream approximations are better (Sagan and Pollack 1967; Irvine 1975) when $\tilde{\omega}_0 \ll 1$. For the problems with $\tau_0 \gg 1$, discussed here, κ is small near the incident surface and increases only slowly with τ (at first). The results for $J(\tau)$ and R_i are most sensitive to layers with moderately small τ and are insensitive to inaccuracies in the Eddington approximation at large τ .

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Fig. 1: A schematic plot of temperature, T , and eddy diffusion coefficient, K , versus pressure, P for the Jovian atmosphere. The dashed portion of the curve for K corresponds to Gierasch's (1976) model for a deep circulation pattern driven by latent heat effects.

Fig. 2: A model radiative transfer calculation for an assumed form of the parameter, κ , as a function of total optical depth, τ . Numerical results for the mean intensity J are given in the solid curve. The dashed curve is an analytic approximation, discussed in the text.

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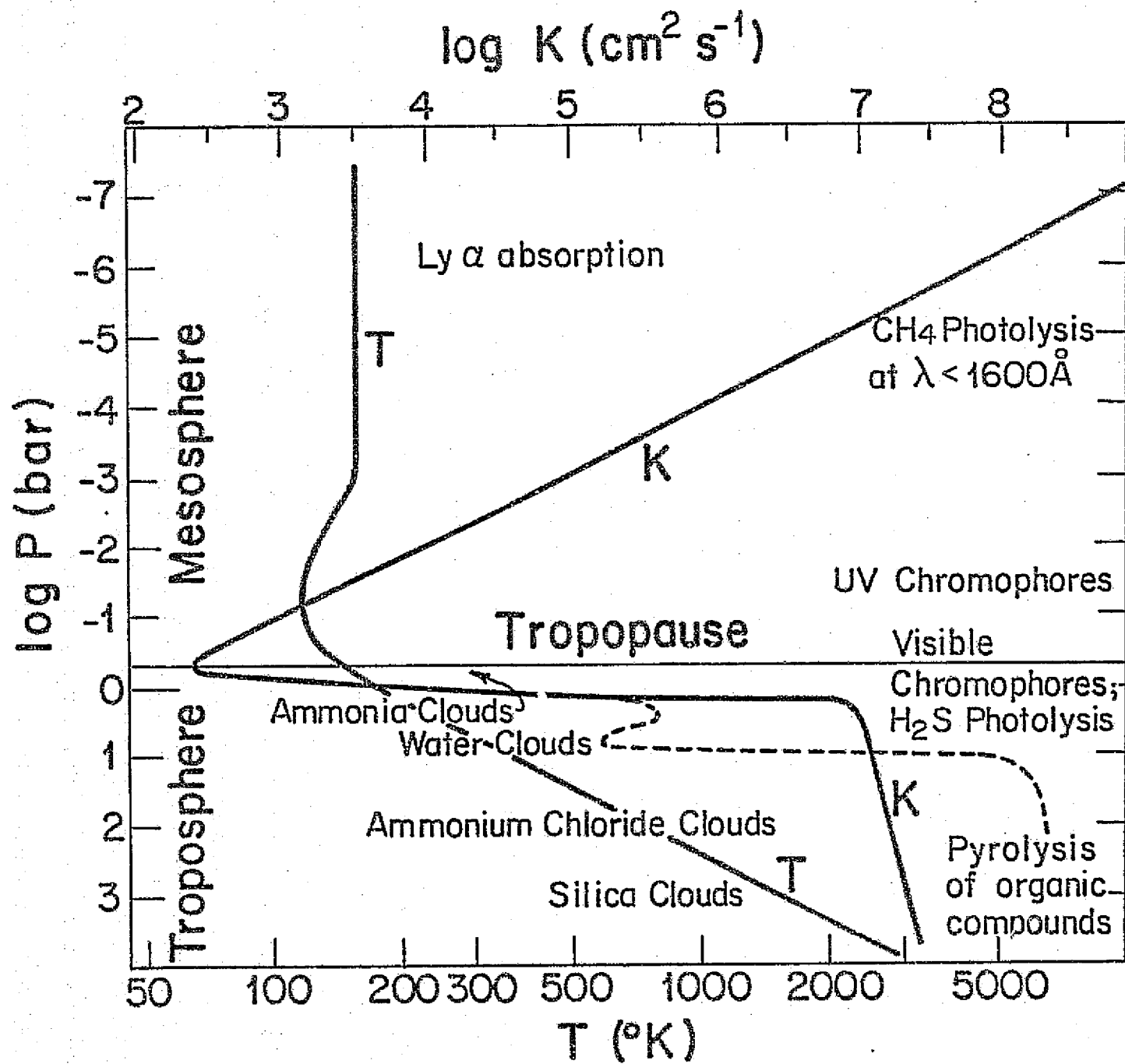
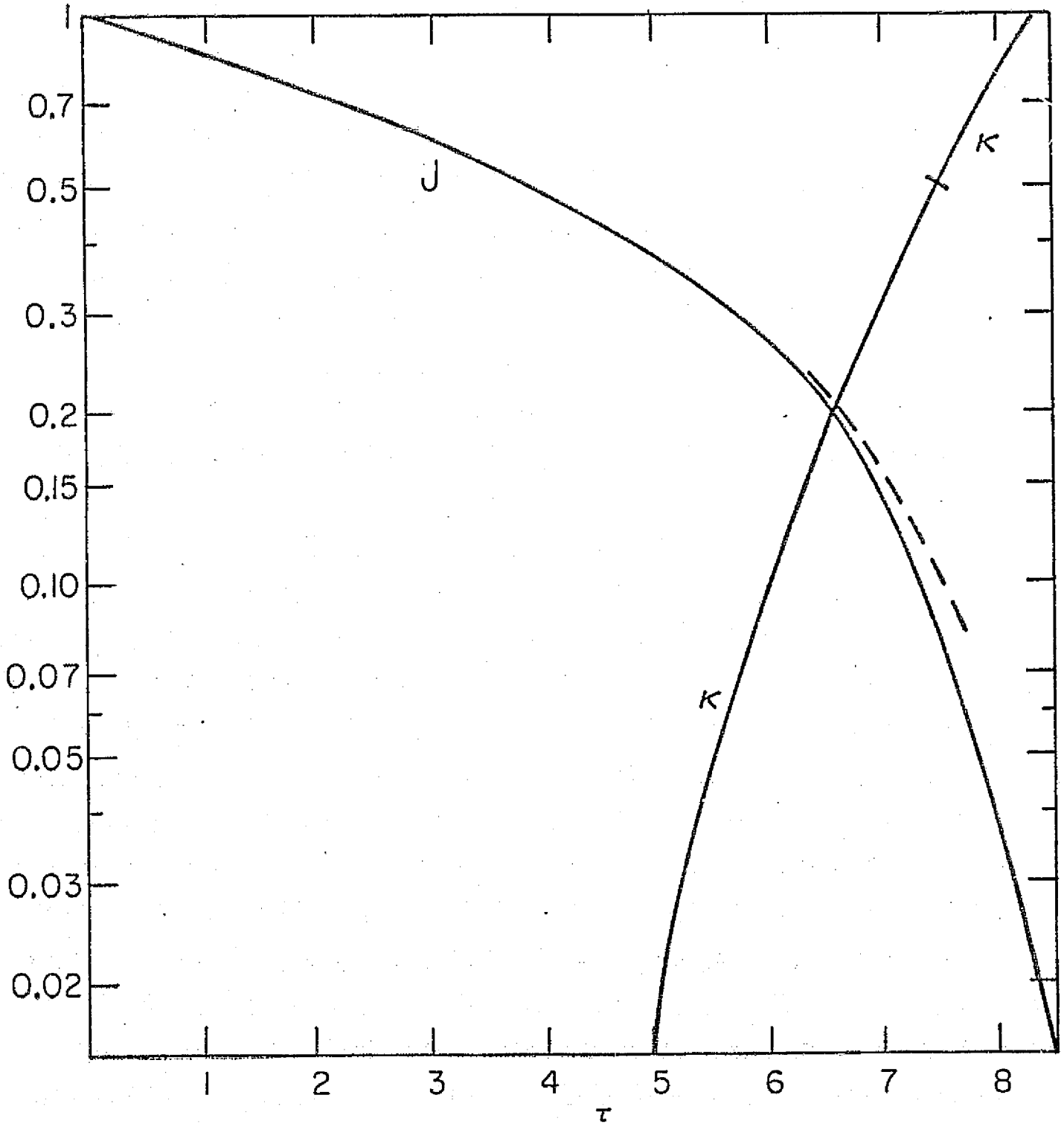


Figure 1.



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Figure 2.