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19850004527

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Paul M. Harvey, Bruce A. Wilking,
Marshall Joy, and Daniel F. Lester

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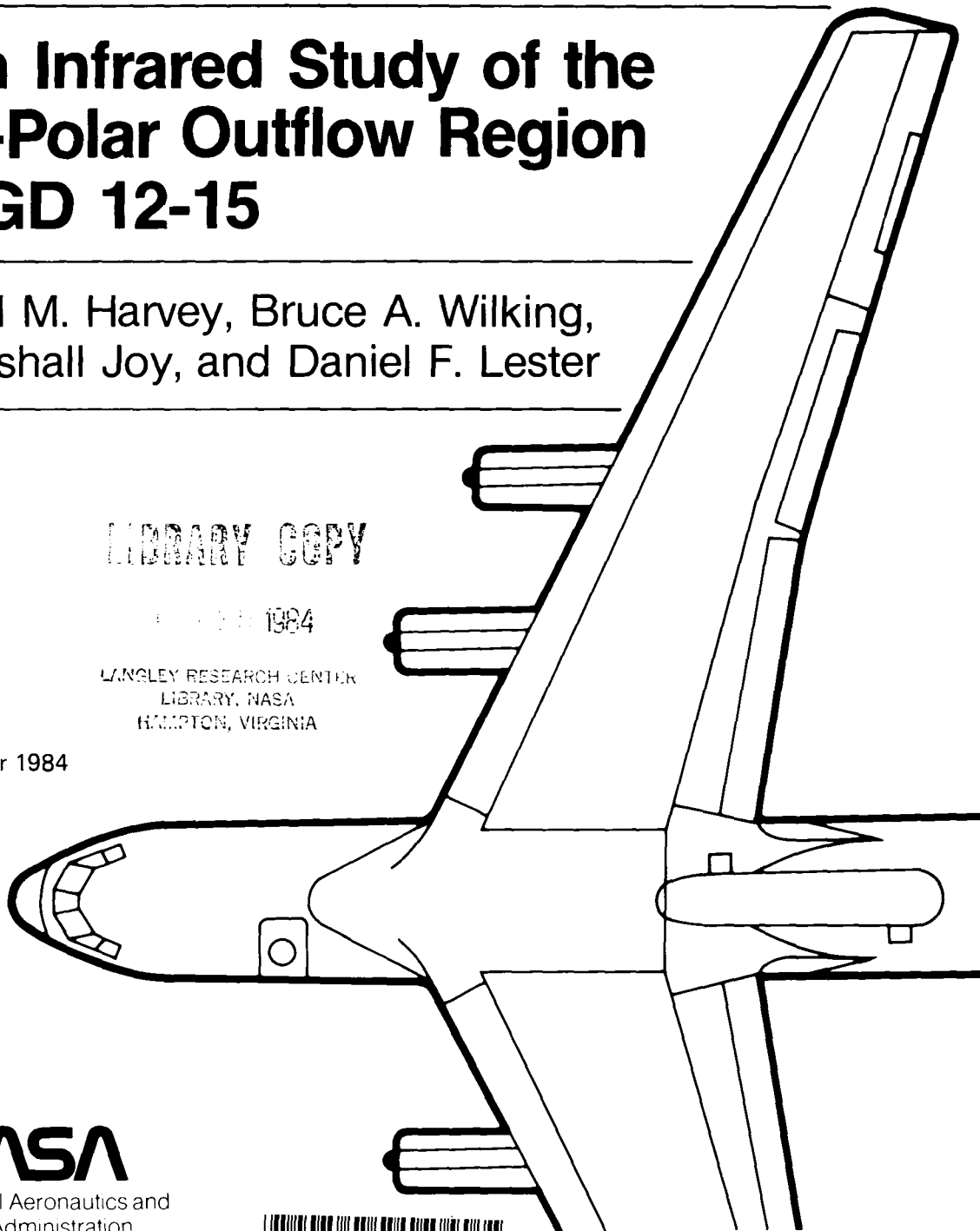
October 1984

NASA

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NF00635



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Paul M. Harvey
Bruce A. Wilking
Marshall Joy
Daniel F. Lester, University of Texas at Austin, Austin, Texas



National Aeronautics and
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Ames Research Center
Moffett Field, California 94035

An Infrared Study of the Bi-Polar
Outflow Region GGD 12-15

by

Paul M. Harvey, Bruce A. Wilking, Marshall Joy, and Daniel F. Lester

Astronomy Department

University of Texas at Austin

and

Visiting Astronomer at the Infrared Telescope Facility
which is operated by the University of Hawaii under contract
from the National Aeronautics and Space Administration.

Accepted (August 1984) for publication in
THE ASTROPHYSICAL JOURNAL

1185-12835 #

Abstract

Infrared observations from 1-100 μm are presented for the region associated with a bipolar CO outflow source near the nebulous objects GGD 12-15. We find a luminous far-infrared source associated with a radio-continuum source in the area. This object appears to be a compact H II region around a nearly main-sequence B0 star. We also find a faint 20 μm source at the position of an H₂O maser 30" northwest of the HII region. This object appears to be associated with but not coincident with a 2 μm reflection nebula. We interpret this structure as evidence for a non-spherically symmetric, possibly disk-like dust distribution around the exciting star for the maser. This object probably powers the bi-polar CO outflow although its luminosity is less than 10% that of the star which excites the compact HII region. A number of other 2 μm sources found in the area are probably members of a recently formed cluster.

I. INTRODUCTION

The region around GGD 12-15 (Gyulbudaghian, Glushkov and Denisyuk 1978) was identified as a likely site of recent star formation in the thorough radio wavelength study of Rodriguez et al. (1980) (RMHG). Subsequently, Rodriguez et al. (1982) found a bipolar outflow in ^{12}CO emission similar to a number of other regions discussed by Bally and Lada (1983). The general area in which the source of the outflow is centered contains a number of faint optical objects, a radio-continuum source, and an H_2O maser 30" away from the continuum source (RMHG). Reipurth and Wamsteker (1983) found several reddened near infrared sources in the region but were unable to identify unambiguously the sources of excitation of the maser and/or radio continuum source. At a distance of 1 kpc (Racine and Van den Bergh 1970) this active region is only 1 pc in diameter.

In order to understand better the nature of the obscured stars in this region we undertook a variety of infrared observations at wavelengths from 1 to 100 μm . These data permit us to identify a group of embedded stars which probably represent a compact, newly-formed cluster in the region. The wide range of wavelengths observed permits us to estimate the total luminosities of the stars and to determine with much more certainty the sources of excitation of the maser and radio continuum source. The results of these observations show that the most luminous object, the exciting source for the radio continuum emission, appears to have a much smaller dynamical effect on the surrounding material than does the much lower luminosity exciting star associated with the maser.

II. OBSERVATIONS

We observed the GGD 12-15 region in the near-infrared from 1-20 μm using the NASA IRTF on Mauna Kea in 1982 December, 1983 January, and 1984 February. We used standard facility photometers with the filters and aperture sizes listed in Table 1. Chopper spacings and direction were chosen to assure that "off"-positions had surface brightnesses less than or equal to the lowest contour level shown in the maps below, typically 80" at 1 - 4 μm and 50" at 8 - 20 μm . The longer wavelength 50 and 100 μm observations (Table 1) were made in 1983 October and 1984 February on NASA's Kuiper Airborne Observatory using the 3 spatial by 2 spectral element array photometer described by Wilking et al. (1984) with a 9' chopper throw roughly east-west. The 1-20 μm observations were calibrated relative to IRTF standard stars; the 50/100 μm flux densities were calibrated relative to S140-IRS (Harvey, Wilking, and Joy 1984).

The IRTF observations consist of: (1) a 2.2 μm map made with a 3".6 beam and 1" spacing between samples (Figure 1), (2) a 20 μm map of the brightest 20 μm emission region made with a 3".8 beam and 1" spacing (Figure 4); (3) a small 20 μm map of the area near the H₂O maser (Figure 2); (4) polarimetry of a near-infrared source near the maser with a 3".6 beam (Table 3); and (5) photometry at several wavelengths between 1 and 20 μm of a number of the sources found (Figure 3 and Tables 2 and 4). All IRTF positions were offset from SAO #132924 and are believed accurate to $\pm 1''$ absolute. Relative positions of nearby sources are probably accurate to $\pm 0.5''$. The KAO observations consist of maps at 50 and 100 μm made with 30" beams and 20" sample spacing over a 5' x 5' area and with 5" spacing in the central 3'x3' as shown in Figure 5. The positions are believed accurate to $\pm 7''$.

III. RESULTS

a) The Maser Region

Figure 1 shows a local $2.2\ \mu\text{m}$ flux maximum (Source 9E) about $4''$ east and slightly south of the H_2O maser position (RMHG). The small area, $20\ \mu\text{m}$ map of this region (Figure 2), however, shows a clear $20\ \mu\text{m}$ flux maximum at the maser position (Source 9M) within the relative positional uncertainties. There is little or no $20\ \mu\text{m}$ emission from the nearby $2\ \mu\text{m}$ peak position. This morphology is similar to that seen in NGC 6334-V by Harvey and Wilking (1984). They proposed that such structure was consistent with a picture in which the $20\ \mu\text{m}$ source/ H_2O maser represents the position of a single, obscured exciting star which is surrounded by a dense dust cloud with non-spherically symmetric geometry. In such a picture the $2\ \mu\text{m}$ emission to the southeast of the GCD 12-15 $20\ \mu\text{m}$ /maser source represents scattered light from the central source which emerges along a line of lower extinction than our line of sight to the central source.

This picture is consistent with two additional observations shown in Figure 3 and Table 3. Figure 3 shows the results of HKL photometry at a number of points along a line parallel to the CO outflow through the maser position. These data show that the near-infrared "nebulosity" becomes progressively bluer from a point $2''$ east of the maser to $6''$ east as would be expected for reflection nebulosity viewed through decreasing amounts of extinction. The K-L color measured at the maser position is not as red as at the point $2''$ east, but this could be due to contamination of the K measurement by the very extended $2.2\ \mu\text{m}$ emission discussed in Sec. III-d if that emission is relatively blue. Table 3 shows the results of polarimetric measurements made at the position of peak $2.2\ \mu\text{m}$ surface brightness southeast of the maser position. The detection of

significant polarization with the E vector along a N-S line is consistent with the interpretation of this source as reflected light from an embedded object to the west, i.e. possibly the maser/20 μ m exciting star.

Figure 3 also shows that there is a second, redder and weaker near-infrared source 12" northwest of the maser which is colinear with the maser and the southeastern 2 μ m peak. This might represent the northwestern lobe of a bipolar reflection nebula analogous to that seen in NGC 6334-V (Harvey and Wilking 1984), if it were on the far side of the exciting star and viewed through a higher column density of material. In that case the circumstellar dust cloud might have a disk-like geometry as proposed for NGC 6334-V (Harvey and Wilking 1984).

An upper limit to the total luminosity of the exciting star for the maser/20 μ m source may be estimated from the shape of the low-level 50 and 100 μ m contours at the position of the maser relative to the point source response as shown in the map of Figure 5. This, together with a lower limit estimated from the 1-20 μ m photometry imply a total luminosity of $10^3 L_{\odot} \gtrsim L \gtrsim 200 L_{\odot}$. A zero-age main-sequence star of luminosity $5 \times 10^2 L_{\odot}$ would have a spectral type of B5. Its total 2.2 μ m brightness would be $K \sim 9.5$ mag. Since the integrated 2.2 μ m brightness in the supposed bi-polar reflection nebula is only $K \sim 11.0$ mag., a B5 ZAMS star could just power the observed reflection nebulae if the reflection efficiency were reasonably high. If the central star has a 2 μ m excess as appears commonly in many other young objects (see sec. III-d), then the required reflection efficiency and extinction requirements are quite reasonable.

b) The Radio Continuum Source

The peak 20-100 μm emission in the GGD 12-15 region is coincident with the position of the 6 cm radio continuum source (RMHG) as shown in Figures 4 and 5. Therefore, since virtually all the luminosity of this region is radiated in the infrared, this object is the dominant luminosity source. The total infrared luminosity in a 30" (0.15 pc) diameter area centered on the 6 cm source is $L \sim 8.0 \pm 2 \times 10^3 L_{\odot}$. RMHG estimated that a $10^4 L_{\odot}$, B0.5 zero-age main-sequence star could produce the H II region inferred from the 6 cm flux. Since that luminosity is nearly equal to the observed infrared luminosity, this object appears to be simply a highly obscured, young H II region around a main-sequence or nearly main-sequence star.

The 8-20 μm energy distribution of the core of this object (Table 4) indicates a relatively cool [T_{d} (8.7-12.5 μm) ~ 150 K] dust temperature and silicate optical depth $\tau_{9.7 \mu\text{m}} \gtrsim 1.0$. Using a ratio of $A_V/\tau_{\text{sil}} \sim 17$ (Rieke and Lebofsky 1984), this implies $A_V \gtrsim 17$ mag. There is no obvious compact feature in the 2.2 μm map (Figure 1) at the position of the 20 μm peak to a level of $K \sim 14.0$ mag, although there is some low-level diffuse emission. A ZAMS B0.5 star would have an apparent K magnitude of 7.8 mag if there were no obscuration. Therefore the observed lack of a point-like exciting source for the H II region implies an extinction at 2.2 μm $A_K \gtrsim 6$ mag, suggesting an even larger $A_V \gtrsim 55$ mag. Another limit may be placed on the extinction to the H II region by comparing the expected 2 μm free-free emission from the H II region to the observed 2 μm flux. From the observed radio flux of 120 mJy (RMHG) we expect $K \sim 10.3$ mag due to free-free, bound-free, and line emission. This implies $A_V \gtrsim 40$ mag.

c) The Other 2 μ m Sources

There are ten other local flux maxima in Figure 1 above the level of 2.4 mJy/beam ($K \leq 13.5$) in addition to the source associated with the maser (IRS 9). Sources 3, 5, 6, and 10 appear essentially pointlike relative to our 3.5" beam down to the 25% contour or below. Sources 1, 2, and possibly 8 may be double at the 25 - 50% contour level with component separations of 2"-4". Sources 7, 11, and 12 are somewhat broader than a point source below the 50% contour. These may be either point sources superimposed on a diffuse background or genuinely resolved sources.

Although some or all of these sources could represent reflection nebulae excited by the star which excites the H II region or by other obscured luminosity sources, it seems more likely that at least the point-like 2 μ m sources probably represent the positions of discrete stellar objects. This is because, first, a highly obscured star which is luminous enough to produce a bright, compact reflection nebula would also be likely to manifest itself as a nearby, detectable longer wavelength source. Secondly, if many of the 2 μ m sources were reflection nebulae powered by the H II region source, the total observed 2 μ m flux density in the region and reasonable reflection efficiencies would require that the H II exciting star have a strong 2 μ m excess. Although the stellar atmospheres of highly obscured young stars have not been well-studied, this seems unlikely for a star whose total luminosity and radio-inferred Lyman continuum luminosity agree so well at a spectral type near B0.5. The best method to answer this question with certainty in the future is probably a combination of polarimetric mapping of the entire region and spectroscopy of the individual sources.

Based on the above arguments, we tentatively interpret the eastern peak of source 1, and sources 3, 5, 6, and 10 as obscured stars because of the point-like contours within a couple of arcseconds of their centers. Sources 7, 8, 11, and 12 are probably individual stars, but could be 2 μ m reflection or emission nebulae illuminated by their own obscured exciting stars or by the exciting star for the obscured H II region. Source 2 appears nearly unresolved along one axis and may represent either such a nebulosity or a partially resolved double star. On the basis of the 1-20 μ m photometry and the lack of more than one local flux maximum in the 50/100 μ m maps none of these sources can have a luminosity significantly greater than $10^3 L_{\odot}$. Table 2 lists the extinction, A_V , derived by assuming that the J and H flux densities represent emission from the reddened photosphere of a hot, star-like object. A ratio of $A_V/E_{J-H} = 11$ was used (Becklin et al. 1978). The de-reddened absolute K magnitudes are also listed in Table 2. For many of the sources these magnitudes are too bright to represent emission from the photospheres of ZAMS stars within the luminosity limit $L \lesssim 10^3 L_{\odot}$. This appears to be a common property of pre-main-sequence stars (Lada and Wilking 1984; Harvey, Wilking, and Joy 1984); the excess may be due to circumstellar gas or dust emission or photospheres which are cooler than appropriate for the zero-age-main-sequence.

d) Large Scale Structure - The Diffuse 2 μ m and 100 μ m Radiation

Both the 2.2 μ m and 100 μ m maps show evidence for diffuse emission over much of the mapped area. At 100 μ m this emission is probably due to dust within the extended molecular cloud which is heated both centrally by the H II region and by fainter nearby stars. A visible reflection nebula 1' southwest is probably responsible for the extended contours in that direction. Because of the strong concentration of the 50 μ m emission around the H II region, it

appears that the visible reflection nebula probably represents only a minor total source of luminosity.

The diffuse 2 μm emission is unlikely to be due to thermal dust emission from typical interstellar dust since dust temperatures above 10^3 K would be implied. Two other possibilities are: (1) reflected emission from the central exciting star for the H II region and possibly some of the other sources, or (2) anomalous thermal emission similar to that found by Sellgren, Werner, and Dinerstein (1983) near visible reflection nebulae, due possibly to non-equilibrium heating of very small grains, or (3) the combined background of a large number of fainter stars. The total diffuse 2.2 μm emission over the mapped area amounts to ~ 50 mJy ($K \sim 10$ mag). This could be supplied by the central star of the H II region if of order 10% of the 2 μm luminosity could be converted into reflection nebulosity. This would imply a highly anisotropic distribution of dust around the central star since the extinction along the line of sight is greater than 5 mag at 2.2 μm . Without further observations it will be difficult to determine the source of this diffuse emission with certainty.

IV. DISCUSSION

a) The Source of the CO Outflow

This region provides an interesting comparison between two young objects, (1) a compact H II region/far-infrared source, and (2) a much lower luminosity H_2O maser/IR source. Our infrared data show, in fact, that the luminosity of the object associated with the H_2O maser is at least a factor of 8 less than that of the exciting star for the H II region. Three pieces of evidence suggest, however, that the CO outflow in this region is associated with the

maser source rather than with the H II region. First, for outflow sources in general (Bally and Lada 1983) there is a very strong statistical correlation between the presence of an outflow and the presence of an H₂O maser. Second, geometrically the H₂O maser/IR source in the GGD 12-15 area is significantly closer to the center of the outflow than is the H II region/far-ir source (Rodriguez et al. 1982). Third, the agreement in alignment between the CO outflow axis and the displacement of the 2 μ m source IRS 9E from the 20 μ m source IRS 9M at the maser position suggests a physical relationship between the infrared source/maser and the larger scale CO outflow. Taken together, these three arguments provide strong circumstantial evidence that the maser source powers the outflow. Higher angular resolution radio observations of the region near the maser are probably needed to be absolutely certain of this conclusion.

Rodriguez et al. (1982) and Bally and Lada (1983) have argued that the bipolar outflow occurs in a large fraction of young stars at some time in their evolution. If that is true, and if the absence of an observable outflow associated with the compact H II region is real and not simply due to an effect such as inauspicious orientation, this suggests that its exciting star has not yet passed through the outflow stage or has already completed it. Since the H II region is resolved (RMHG) it seems more likely that its exciting star has evolved beyond the outflow stage since most outflow sources of similar luminosity have unresolved or non-existent H II regions (Bally and Lada 1983; Wynn-Williams 1982). This result is consistent with the idea that the star which excites the H II region and that which excites the maser formed coevally and the more luminous star, which excites the H II region, evolved more rapidly.

b) The Associated Infrared Cluster

There are at least five and perhaps ten more 2 μ m sources within the mapped area (Figure 1) which probably represent obscured stars. This cluster is thus comparable in density of stars to that associated with the bi-polar infrared nebula NGC 6334-V-IRS 4 (Harvey and Wilking 1984) at equal luminosity limits. Unlike that cluster, however, the most luminous young star is not the object associated with a high velocity outflow.

As noted earlier, the de-reddened absolute K magnitudes of virtually all the sources are too bright to be due to photospheric emission from ZAMS stars with low enough luminosities to be consistent with the far-infrared observations. This suggests that most of these stars are young and possibly pre-main-sequence. It is interesting then that there appears to be only a single "outflow" source in this cluster. This is consistent with the idea (Bally and Lada 1983, and references therein) that the outflow stage is a fairly short lived, transient event in the lives of young stars.

Acknowledgements

We thank the ground and flight crew of the Kuiper Observatory and the staff of the Infrared Telescope Facility for their assistance with these observations and for their efforts which provided such excellent telescope performance. This research was supported in part by NASA grant NAG 2-67.

References

- Bally, J., and Lada, C.J. 1983, Ap. J., 265, 824.
- Becklin, E.E., Matthews, K., Neugebauer, G., and Willner, S.P. 1978, Ap. J., 220, 831.
- Gillett, F.C., Forrest, W.F., Merrill, K.M., Capps, R.W., and Soifer, B.T. 1975, Ap. J., 200, 609.
- Gyulbudaghian, A.L., Glushkov, Yu. I., and Denisyuk, E.K. 1978, Ap. J. (Letters), 224, L137.
- Harvey, P.M., and Wilking, B.A. 1984, Ap. J. (Letters), in press.
- Harvey, P.M., Wilking, B.A., and Joy, M. 1984, Nature, in press.
- Racine, R., and Van den Bergh, S. 1970, in The Spiral Structure of Our Galaxy, ed. W. Becker and G. Contopoulos (Dordrecht: Reidel).
- Reipurth, B., and Wamsteker, W. 1983, Astron. Ap., 119, 14.
- Rieke, G.H., and Lebofsky, M.J. 1984, preprint.
- Rodriguez, L.F., Carral, P., Ho, P.T.P., and Moran, J.M. 1982, Ap. J., 260, 635.
- Rodriguez, L.F., Moran, J.M., Ho, P.T.P., and Gottlieb, E.W. 1980, Ap. J., 235, 845.
- Sellgren, K., Werner, M.W., and Dinerstein, H.L. 1983, Ap. J. (Letters), 271, L13.
- Wilking, B.A., Harvey, P.M., Lada, C.J., Joy, M., and Doering, C.R. 1984, Ap. J., in press.
- Wynn-Williams, C.G. 1982, Annual Reviews of Astronomy and Astrophysics, 20, 587.

TABLE 1
Instrumental Parameters

Filter	λ_{eff} (μm)	Beam Size	Telescope
J	1.25	3.6"	IRTF
H	1.65	"	"
K	2.2	"	"
L	3.5	"	"
8.7	8.7	3.8"	"
9.7	9.7	"	"
N	10.2	"	"
10.3	10.3	"	"
12.6	12.6	"	"
Q	20.0	"	"
50	47	27	KAO
100	95	29	"

TABLE 2
1-20 μ m Photometry

Source #	R.A. (1950) 06 ^h 08 ^m	Dec. -06 ^o	K mag.	J-H	H-K	K-L	N	Q	Derived Av	M _k
1	24 ^s .01	11'21"5	11.57	1.11	0.68	0.62	6.9 \pm .25	2.7 \pm .28	20	0.2
2	23.84	11'15"	11.01	1.65	1.57	2.02	2.58	-1.47	18	-1.0
3	24.25	11'12	12.04	1.75	1.16	1.23	4.72	-0.86	19	-0.1
4	23.98	11'07"	14.28	1.3 \pm 0.2	1.14	3.24	2.46	-2.82	14	2.7
5	23.35	11'02"5	11.40	1.29	1.11	1.14	6.2 \pm .25	2.8 \pm .25	14	-0.2
6	22.96	10'59"	11.68	\geq 3	2.34	1.89	6.2 \pm .25	3.2 \pm .3	33	-2.0
7	24.31	10'57"	12.19	2.7 \pm 0.2	2.23	1.80		\geq 3.0	30	-1.1
8	24.09	10'53"	12.43	2.35	1.83	1.36		\geq 2.8	26	-0.4
9E	25.95	10'51"	12.03	1.98	1.12	1.11	6.6 \pm .3	\geq 3.0	30	-1.5
9M	25.75	10'49"5	14.40	\geq 1.0	1.89	1.3	5.9 \pm .25	1.85		
10	23.96	10'36"5	9.57	2.45	1.97	1.50	5.6 \pm .2	\geq 3.0	27	-3.4
11	25.10	10'53"0	13.51	2.5	1.76	1.32	>6.5	>2.0	27	0.5
12	22.75	10'46"7	12.15	1.63	0.61	0.48			18	0.2

Observational uncertainties are \pm 0.1 mag at J, H, K, L, \pm 0.15 mag at N, and \pm 0.2 mag at Q unless otherwise noted.

TABLE 3
Polarimetry of IRS 9-SE

	H(1.65 μm)	K(2.2 μm)
P	$4.8 \pm 1.0\%$	$2.5 \pm 0.5\%$
θ	$170^\circ \pm 15$	$178^\circ \pm 12$

TABLE 4
Narrow-Band 10 μ m Photometry

Wavelength (μ m)	IRS 2 (Jy)	IRS 4 (Jy)
8.7	2.27	1.36
9.7	2.53	1.21
10.3	3.31	2.33
11.6	5.94	7.71
12.5	7.92	12.9
20 (broad)	39	134

Uncertainties as in Table 2.

Figure Captions

Figure 1 - Map of the GGD 12-15 area at $2.2\ \mu\text{m}$ made with $3''.6$ beam (FWHM).

Contours are at levels 1, 1.5, 2, 4, 8, 16, 32, 64, and 128 x 0.6 mJy per beam. The total region which was mapped includes most of the area shown, with the exception of small areas in the northeast and northwest corners. The various sources found are labeled with the numbers used in Table 2. The dashed square around IRS-9M indicates the area of the $20\ \mu\text{m}$ map shown in Figure 2. Likewise for the dashed rectangle around IRS 1-4 and Figure 4. The dashed line through IRS-9M and 9E shows the scan line along which the observations in Figure 3 were made. The "H" and surrounding dashed circle show the position and size of the compact H II region found by RMHG which is also coincident with IRS 4 of Table 2. In this map and those that follow, the lowest contour is at a level where the signal-to-noise ratio was of order 3:1.

Figure 2 - Map with a $3''.8$ beam at $20\ \mu\text{m}$ centered on the position of the H_2O maser (RMHG) indicated by the dashed cross near the center.

Fluxes are in Jy per beam. Contours are sketched at 0.5, 1.0 and 1.5 Jy per beam.

Figure 3 - Scans at 1.65, 2.2, and $3.5\ \mu\text{m}$ made through maser position (0,0) along CO outflow axis with a $3''.6$ beam. Offsets shown are R.A. and Dec (+ east, + north) in arcseconds.

Figure 4 - Map made with $3''.8$ beam at $20\ \mu\text{m}$ of the brightest $20\ \mu\text{m}$ sources found in the GGD 12-15 region, IRS-2, 3, 4. Contours are at levels 10, 20, 40, 60, 80, 100, 120, and 135 Jy per beam. The dashed circle and plus sign indicate the size and central position of the H II region found by

RMHG.

Figure 5 - Maps made at 50 and 100 μm with 28" beams. Contours are at levels 90, 80, 60, 50, 40, 20, 10, 5, and (100 μm only) 2.5% of the peak fluxes which were 2.2×10^3 and 2.1×10^3 Jy per beam respectively at 50 and 100 μm . The plus sign indicates the position of the compact H II region and the cross shows the position of the H₂O maser (RMHG).

Authors' Postal Addresses

P. M. Harvey, M. Joy, D. F. Lester, and B. A. Wilking

Astronomy Department

University of Texas

Austin, TX 78712

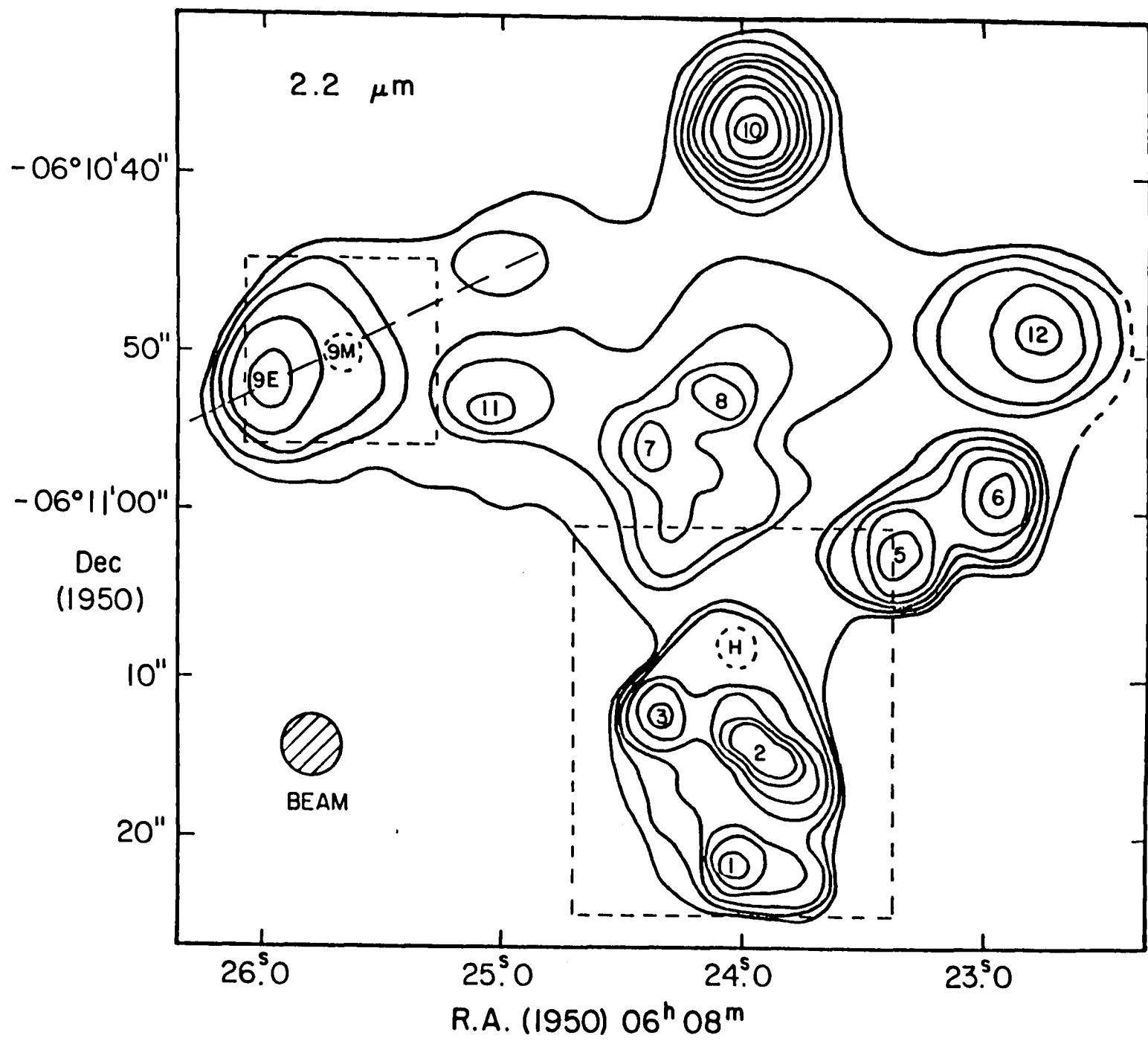
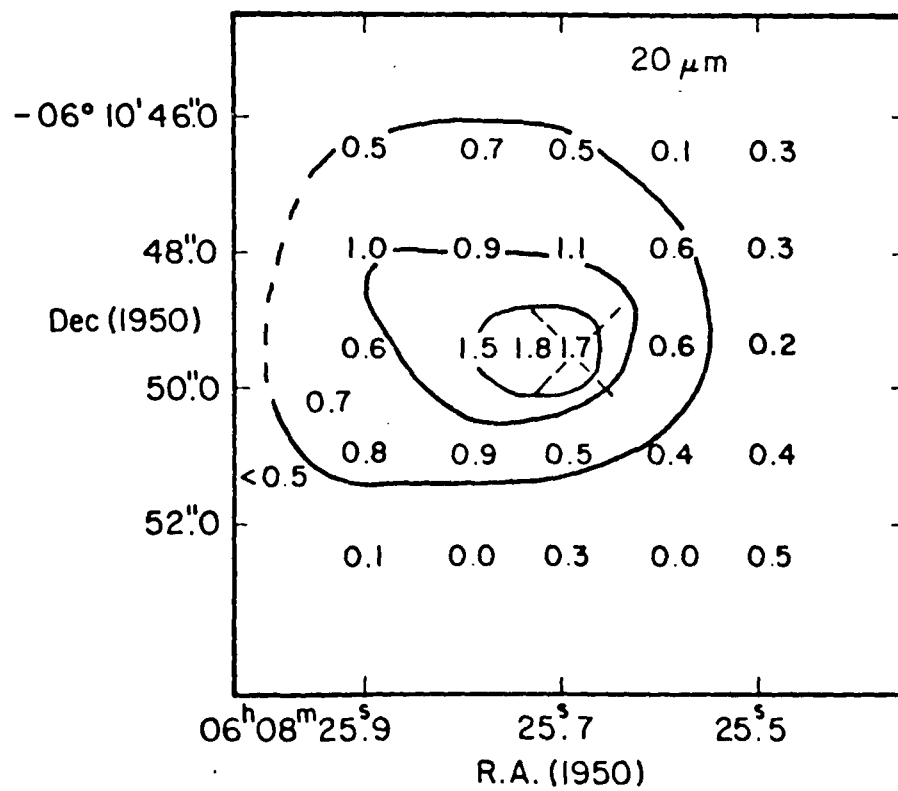


FIGURE 1

FIGURE 2



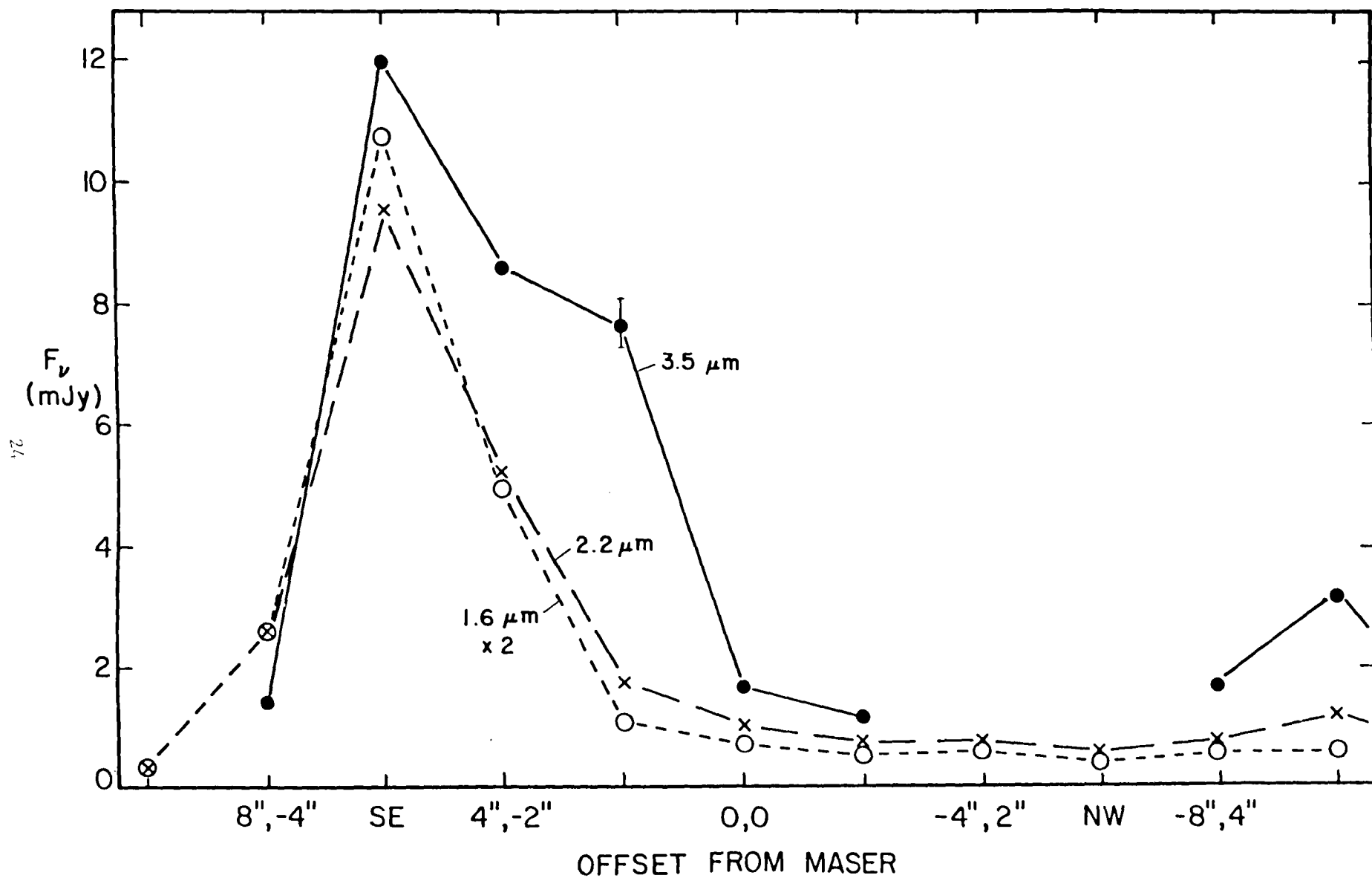


FIGURE 3

FIGURE 4

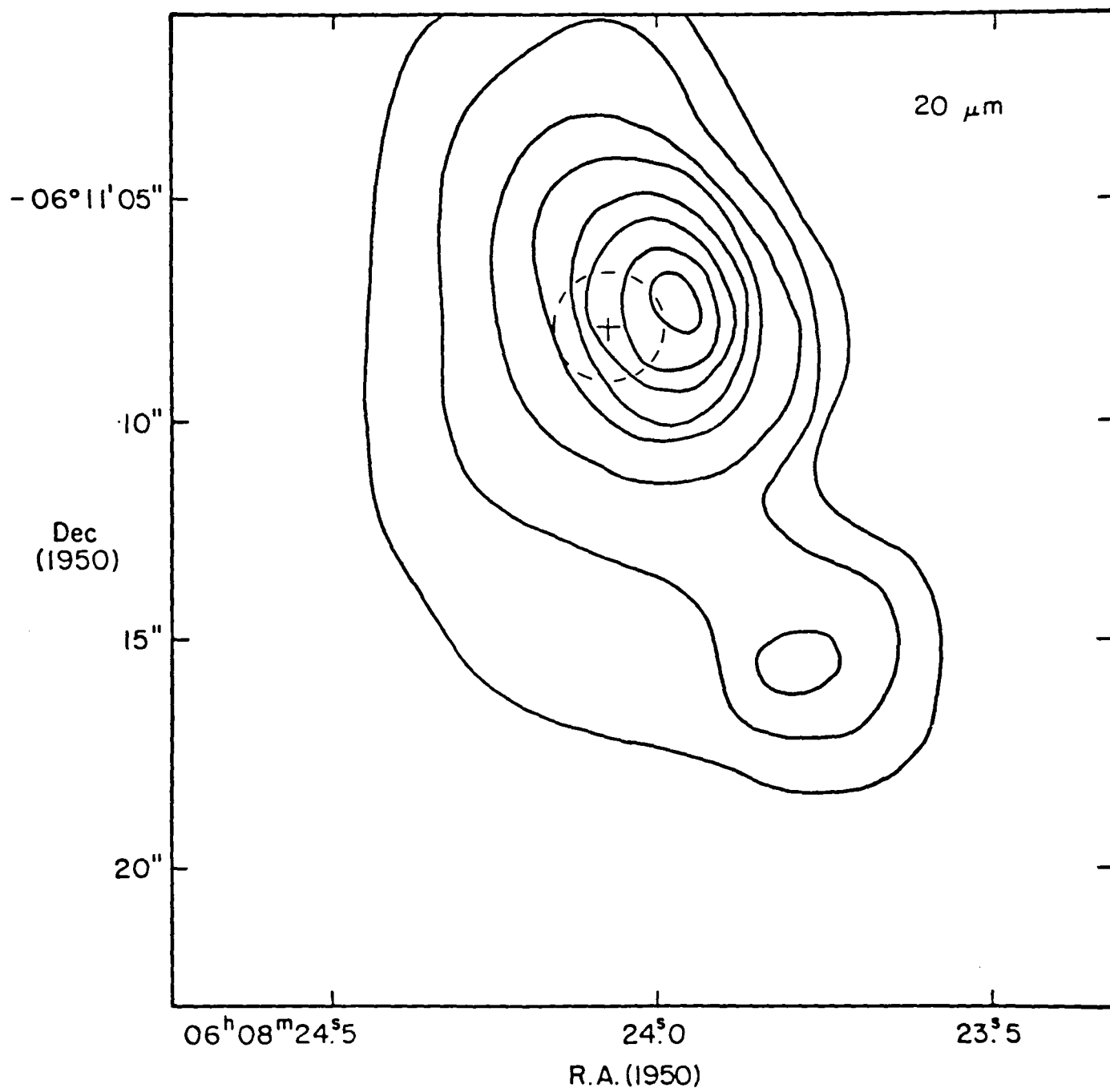
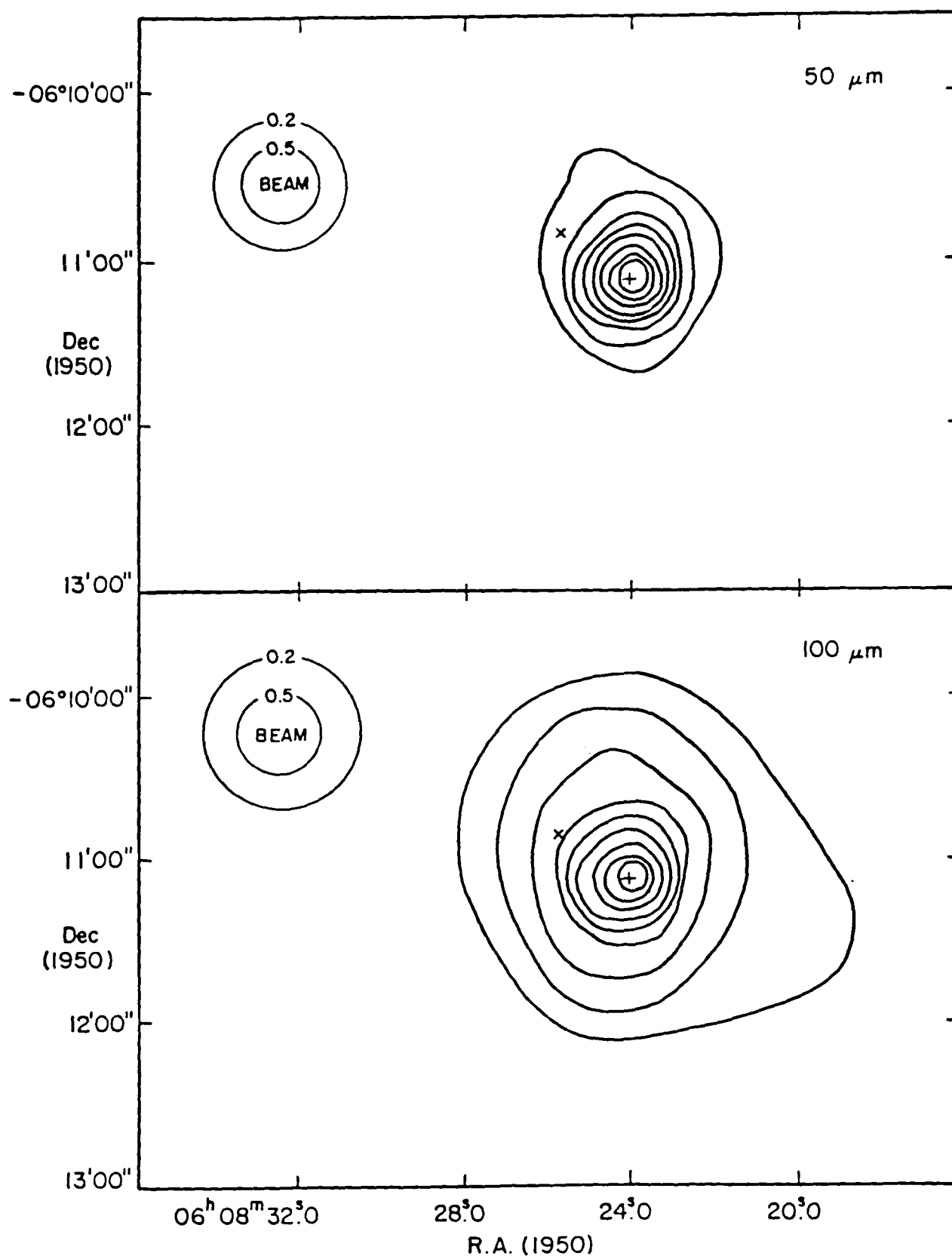


FIGURE 5



1. Report No. NASA TM-86655	2. Government Accession No.	3. Recipient's Catalog No.	
4. Title and Subtitle AN INFRARED STUDY OF THE BI-POLAR OUTFLOW REGION GGD 12-15		5. Report Date November 1984	
		6. Performing Organization Code	
7. Author(s) Paul M. Harvey, Bruce A. Wilking, Marshall Joy, and Daniel F. Lester		8. Performing Organization Report No. 85017	
9. Performing Organization Name and Address University of Texas at Austin Austin, TX 78712		10. Work Unit No.	
		11. Contract or Grant No.	
12. Sponsoring Agency Name and Address National Aeronautics and Space Administration Washington, DC 20546		13. Type of Report and Period Covered Technical Memorandum	
		14. Sponsoring Agency Code 352-02-03	
15. Supplementary Notes Preprint Series #24. Supported by NASA grants. Point of Contact: L. C. Haughney, Ames Research Center, M/S 211-12, Moffett Field, CA 94035, (415) 965-5539, FTS 448-5339.			
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17. Key Words (Suggested by Author(s)) Nebulae: CO outflow Nebulae: far-infrared Nebulae: dust emission		18. Distribution Statement Unlimited Subject category - 89	
19. Security Classif. (of this report) Unclassified	20. Security Classif. (of this page) Unclassified	21. No. of Pages 28	22. Price* A03

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