

STUDY OF NON-THERMAL PHOTON PRODUCTION UNDER DIFFERENT SCENARIOS IN SOLAR FLARES. II. THE COMPTON INVERSE AND BREMSSTRAHLUNG MODELS AND FITTINGS.

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1. Introduction. Energy spectra of photons emitted from Bremsstrahlung (BR) of energetic electrons with matter, is obtained from the deconvolution of the electron energy spectra derived in SH 1.2.-2. According to Kane and Anderson (1970), the differential photon flux at the earth level is

$$\frac{dJ(h\nu)}{d(h\nu)} = \frac{3.53 \times 10^{-28} E.M.}{n} \int_E^{E_m} \frac{d\sigma(E, h\nu)}{dE} N(E) dE \quad (\text{photons/s KeV cm}^2) \quad [1]$$

where  $d\sigma(E, h\nu)/dE$  is the differential cross-section for electron-proton (BR), given in the non-relativistic range by the Bethe-Heitler formula (Jackson, 1962) and in the relativistic domain by the Koch and Motz formula (Bai and Ramaty, 1976);  $E.M. = n^2V$  is the emission measure, with  $n$  and  $V$  the number density and volume respectively. For Inverse Compton (IC) we followed Sheng (1972), introducing  $\sigma(W, h\nu) = \sigma_e \delta[h\nu - (4/3)\bar{\epsilon} (W/m_e c^2)^2]$  in eq.(1) for  $W < (m_e c^2)^2 / 4\bar{\epsilon} \approx 290$  MeV, where  $W$  = total electron energy,  $\sigma_e = 6.65 \times 10^{-25}$  cm<sup>2</sup> = Thompson cross-section and  $\bar{\epsilon} = 2.7kT$  = mean thermal photon energy, and we obtained at 1 Astronomical Unit.

$$dJ(h\nu)/d(h\nu) = (4.39 \times 10^{-31} E.M. \omega_{ph}(h\nu)^{1/2} / n^2 T^{3/2}) N(E = m_e c^2 \{(3h\nu/4\bar{\epsilon})^{1/2} - 1\}) \quad (\text{photons/s KeV cm}^2) \quad [2]$$

For the evaluation of the electron energy spectra we have explored different combinations of the source physical parameter in the scenarios displayed in SH 1.2-2:  $n = 10^{10} - 10^{13}$  cm<sup>-3</sup> with  $T = 10^5 - 10^7$  °K in scenario (a),  $n = 10^{10} - 10^{12}$  cm<sup>-3</sup> with  $T = 10^6 - 10^7$  °K in scenario (b). For scenario (c) we used the combination of the two previous parameter sets. Values of  $B$  were delimited from the thermal flux  $N_0$  in SH 1.2-2, for every couple  $(n, T)$ , by normalization of  $N_0$  with the point of maximum flux and minimum energy in the observational photon spectra. For the evaluation of the spectrum [10] of scenario (d) we swept  $n = 10^{10} - 10^{13}$  cm<sup>-3</sup>,  $L = 10^8 - 10^9$  cm,  $B = 10^2 - 10^3$  gauss and  $\epsilon = 10^{-3} - 10$  V/cm; in this case two assumptions were worked out for the transport and photon emission regions respectively, first  $n = 10^{10} - 10^{11}$  cm<sup>-3</sup>,  $T = 10^5 - 10^6$  °K with  $n = 10^9 - 10^{10}$  cm<sup>-3</sup>,  $T = 10^6 - 10^8$  °K, and on the other hand  $n = 10^{10} - 10^{12}$  cm<sup>-3</sup>,  $T = 10^5 - 10^6$  °K with  $n = 10^{12} - 10^{13}$  cm<sup>-3</sup>,  $T = 10^4 - 10^5$  °K. For the acceleration efficiencies of the Fermi, Betatron and electric field acceleration processes in scenarios (a)-(c) we have required that the net energy change rate  $dE/dt > 0$  in eqs. [7] and [8] of paper SH 1.2-2. The mean acceleration time  $\tau$  in eq. [6] is the free parameter of our analysis, however, we have restricted it to physically reasonable values quoted in the literature. Similarly we have proceeded for the selection of the three characteristic times of scenario (d): the characteristic time  $\tau^i$  of the injection rate in eq. [5] of SH 1.2-2, and the mean remain times  $\tau_1$  and  $\tau_2$  of particles in the trans-

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port and emission regions respectively. It should be noted that, although electric field acceleration in scenario (a) and neutral sheet acceleration in scenario (d) is basically the same process, in the former case we are adopting a thick geometry, while in the later energy losses in the acceleration volume are neglected (thin geometry), and the spectrum is derived in a quite a different form, by following particle trajectories in the electromagnetic field of the sheet is diffusion region. For the photon field we swept from  $\omega_{ph} = 10^{12}-10^{18}$  eV/cm<sup>3</sup>.

2. Results. Low energy events and the impulsive phase of high energy events are better described within the frame of scenario (a) rather than with (d) (Figs.1-3). Acceleration in those events is better described by impulsive electric field acceleration in the context of thick geometries, while the dominant radiation process is the (IC) effect, even within the frame of the thin geometry of scenario (d). The range of accelerating electric fields falls between  $6.5 \times 10^{-3}-10^{-2}$  V/cm, whereas in scenario (d) higher fields are required (1-15) V/cm. Similarly, the adequate magnetic field strengths are 50-100 gauss, whereas in scenario (d) it is needed 400-500 gauss. Typical number densities are  $10^{11}$  cm<sup>-3</sup>, but  $10^{10}$  cm<sup>-3</sup> in the acceleration region with  $10^{13}$  cm<sup>-3</sup> in the emission region in scenario (d). Temperatures of  $10^6$  °K and photon fields  $> 10^{17}$  eV/cm<sup>3</sup> prevail in the source. Characteristic acceleration times are 0.03-0.06 s and much higher (30-50)s in scenario (d), where corresponding emission times are ~80-100 s. Non-impulsive energy spectra of low and high energy events (usually associated to 2nd acceleration stage) are better described by stochastic acceleration from thermal energies, scenario (b), and radiation from (IC) than with injection from a preliminar acceleration phase within the volume of secondary stochastic acceleration scenario (c), in which case radiation appears from (BR). The acceleration efficiency in this kind of events must be very high ( $\alpha=10-20$  s<sup>-1</sup>), while the corresponding acceleration times are quite shorter ~ 0.002 s, but if the contribution of a first acceleration step is considered the acceleration efficiency turns to be lower. Typical parameters involved in these events are  $n=10^{11}$  cm<sup>-3</sup>,  $T \sim 10^6$  °K,  $B \sim 50$  gauss, while  $\omega_{ph}$  values ( $10^{16}-10^{17}$ ) eV/cm<sup>3</sup> are lower than in impulsive events. On the other hand, the global description of energy spectra composed of two different components (usually associated with two acceleration phases) is better assuming scenario (c) than with (b), in which case radiation from (IC) is dominant, with  $\omega_{ph} \geq 10^{17}$  eV/cm<sup>3</sup>,  $\alpha \sim 1$  s<sup>-1</sup>,  $\tau=0.02$ s,  $n=10^{11}$  cm<sup>-3</sup>,  $T=10^6$  °K and  $B=50$  gauss. Optimum fits to the observational spectra are summarized in table 1.

3. Conclusions. From this study it can be inferred that the scenario for the production of (X-γ) rays continuum in solar flares may vary from event to event, however, it is possible in many cases to associate low energy events to impulsive acceleration, and the high energy phase of some events to stochastic acceleration. In both cases, flare particles seems to be strongly modulated by local energy losses. Electric field acceleration, associated for instance to neutral current sheets is a suitable candidat for impulsive acceleration. Finally we claim that the predominant radiation process of this radiation is the (IC) effect due to the local flare photon field.

TABLE 1

SH 1.2-3

EVENT	BEST SCENARIO	ACCEL. & LOSSES PROCESSES	RAD. PROC.	$\omega_{ph}$ (eV/cm <sup>3</sup> )	ACCEL. PARAMETERS			FLARE PARAMETERS			BEST FITTING (FIGS.1-6)
					$\epsilon$ (V/cm) $\alpha$ (s <sup>-1</sup> )	$\tau$ (s) $\tau', \tau_1, \tau_2$ (s)	$n$ (cm <sup>-3</sup> )	$T$ (°K)	$B$ (gauss)		
1-111- 970 X-RAYS 35-160 KeV ONE PHASE	(a) (a) (d)	EFA=electric field acceleration BETA= betatron NSA=neutral sheet acceleration	IC IC BR	$5 \times 10^{17}$ $3 \times 10^{16}$ -	$\epsilon=10^{-2}$ $\alpha=1$ $\epsilon=10$	$6 \times 10^{-3}$ 0.3 50,500,100	$10^{11}$ $10^{11}$ $10^{10}-10^{13}$	$10^6$ $10^7$ $10^6$	90 100 500	EFA-IC	
30-111-1969 X-RAYS 28-254 KeV ONE PHASE	(a) (a) (d)	EFA EFA NSA	IC BR IC	$10^{17}$ - $3 \times 10^{17}$	$\epsilon=6 \times 10^{-3}$ $\epsilon=2$ $\epsilon=15$	$6 \times 10^{-3}$ 10 50,2000,100	$10^{11}$ $10^{13}$ $10^{10}-10^{13}$	$10^6$ $2 \times 10^5$ $10^6$	50 100 400	EFA-IC	
4-111-1972 γ-RAYS 0.4-0.7 MeV 1st PHASE	(a) (a) (a) (d)	EFA BETA BETA NSA	IC IC BR IC	$10^{17}$ $10^{16}$ - $2.5 \times 10^{17}$	$\epsilon=5 \times 10^{-2}$ $\alpha=1$ $\alpha=10$ $\epsilon=1$	$3 \times 10^{-2}$ 0.2 1. 30,400,80	$10^{11}$ $10^{11}$ $10^{11}$ $10^{10}-10^{12}$	$10^6$ $10^7$ $10^6$ $10^6$	50 120 50 400	EFA-IC	
4-111-1972 γ-RAYS 0.8-7MeV 2nd. PHASE	(b) (b) (c)	FERMI FERMI (BETA-IC)-FERMI	IC BR BR	$10^{17}$ - $10^{16}$	$\alpha=20$ $\alpha=1$ $\alpha=1$	$2 \times 10^{-3}$ 1. $5 \times 10^{-2}$	$10^{11}$ $10^{11}$ $10^{11}$	$10^6$ $10^6$ $10^7$	50 50 160	FERMI-IC ~ ~	
30-111-1969 X-RAYS 28-254 KeV 2nd. PHASE	(b) (b) (c)	FERMI FERMI (EFA-IC)-FERMI	IC BR BR	$10^{16}$ - $10^{17}$	$\alpha=10$ $\alpha=0.5$ $\epsilon=10$	$2 \times 10^{-3}$ 0.5 $2 \times 10^{-2}$	$10^{11}$ $10^{11}$ $10^{12}$	$10^6$ $10^6$ $10^6$	45 40 100	FERMI-IC ~ ~	
4-111-1972 γ-RAYS 0.4-7 MeV 1st. & 2nd. PHASE	(b) (b) (c) (c) (c)	FERMI FERMI (EFA-IC)-FER (BETA-IC)-FERMI (BETA-IC)-FERMI	IC BR IC IC BR	$10^{16}$ - $10^{17}$ $2 \times 10^{17}$ $3 \times 10^{16}$	$\alpha=10$ $\alpha=1$ $5 \times 10^{-2}$ 1. $\alpha=10$	$1.5 \times 10^{-3}$ 0.3 $2.5 \times 10^{-2}$ 0.02 $7 \times 10^{-2}$	$10^9$ $10^{10}$ $10^{11}$ $10^{11}$ $10^{11}$	$10^7$ $10^6$ $10^6$ $10^6$ $10^7$	45 50 50 50 100	~ ~ ~ (BETA-IC)-FERMI-IC	

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