

TDA Progress Report 42-111

1993.0.1720

495086

11

November 15, 1992

Correlated Flux Densities From VLBI Observations With the DSN

R. F. Coker
Tracking Systems and Applications Section

Correlated flux densities of extragalactic radio sources in the very long baseline interferometry (VLBI) astrometric catalog are required for the VLBI tracking of Galileo, Mars Observer, and future missions. A system to produce correlated and total flux density catalogs has been developed to meet these requirements. A correlated flux density catalog of 274 sources, accurate to about 20 percent, has been derived from more than 5000 DSN VLBI observations at 2.3 GHz (S-band) and 8.4 GHz (X-band) using 43 VLBI radio reference frame experiments during the period 1989-1992. Various consistency checks have been carried out to ensure the accuracy of the correlated flux densities. All observations were made on the California-Spain and California-Australia DSN baselines using the Mark III wideband data acquisition system. A total flux density catalog, accurate to about 20 percent, with data on 150 sources has also been created. Together, these catalogs can be used to predict

baseline. Nominally, the correlated flux density detection limits, using the DSN's operational 250-kHz bandwidth navigation VLBI system (the Block I system), are about 0.2 Jy (1 Jy is 10^{-26} W/m²/Hz) for two 70-m antennas, 0.4 Jy for one 70-m and one 34-m antenna, and 0.8 Jy for two 34-m antennas.² These limits apply at 8.4 GHz (X-band); at 2.3 GHz (S-band) the limits are somewhat higher at 0.2, 0.5, and 1.0 Jy, respectively, due to higher system temperatures. Since the correlated flux density of a given quasar can change from month to month by as much as 50 percent in extreme cases (especially at higher observing frequencies, such as X-band) [3], it is necessary to make correlated flux density estimates a routine part of reference frame data processing. A method has been developed which easily and quickly produces a correlated and total flux density catalog that is accurate enough (~ 20 percent) to determine which antenna pairs will be able to detect specified reference sources.

Flux density monitoring will also assist in quickly identifying sources that are inappropriate for reference frame determinations due to large variations in correlated flux density with baseline angle (as a result of extensive structure). It will also provide valuable time and source structure variability information. Prior to this work, no method existed for producing correlated flux density data from DSN Mark III Block II observations (see [4] for Mark III system details). In addition, the production of correlated flux density data from DSN Mark II observations was tedious and the results were often highly uncertain.³ The new flux density estimates (with a detection limit of approximately 30 mJy), as well as the developed methodology, will be useful to anyone using the DSN to observe natural radio sources.

This article is organized in the following manner: Section II presents the expression for correlated flux density and gives details of the production and contents of the flux density catalogs. Section III describes the consistency checks and integrity tests carried out to verify the flux density results. The error budget is discussed in Section IV. Section V summarizes the results.

² J. S. Border, "An Error Analysis for Magellan Differential Delay Rate Measurements," JPL Engineering Memorandum 335-96 (internal document), Jet Propulsion Laboratory, Pasadena, California, February 23, 1987.

³ I. Chen and J. S. Ulvestad, "Determination of Correlated Flux Densities from Mark II VLBI Experiments," JPL Interoffice Memorandum 335.3-88-102 (internal document), Jet Propulsion Laboratory, Pasadena, California, September 22, 1988.

II. ESTIMATING CORRELATED FLUX DENSITIES AND TOTAL FLUX DENSITIES

For a radio source with nonzero angular size, the correlated flux density (i.e., that measured with an interferometer) will always be less than the total flux density. Roughly speaking, the correlated flux density is the strength of that part of the source which is angularly small as compared with an interference fringe on the sky. The angular width of an interference fringe is λ/B , where λ is the observing wavelength and B is the projected length of the interferometer baseline in the plane perpendicular to the source direction. The correlated flux density of a resolved source will vary as a function of baseline length and orientation because source structure is two-dimensional. In the absence of a complete, current description of the structure of a source, there is no way to specify a priori the source visibility (defined as the ratio of correlated flux density to total flux density) [5].

The correlated flux density, S_{cor} , of a source is given by⁴

$$S_{cor} = \frac{\sin(\rho\pi/2)}{b} \sqrt{\frac{T_{sys_1}T_{sys_2}}{G_1G_2}} \quad (1)$$

where ρ is the correlation amplitude for data recorded with one-bit sampling, b is a normalization factor due to correlator and bandpass effects (with this definition, b for DSN Mark III data is greater than 1; see Section III.A for details), T_{sys_i} is the total on-source system temperature at station i in kelvins (K), and G_i is the antenna gain at station i in K/Jy [6]. In addition, the total flux density, S_{tot} , as measured by a single antenna with gain G , can be written as

$$S_{tot} = \frac{T_{on} - T_{off}}{G} = \frac{T_{ant}}{G} \quad (2)$$

Here T_{on} and T_{off} are on-source and off-source system temperatures, respectively, and T_{ant} is the antenna temperature.

Using the new procedure reported here, correlated flux densities are obtained from DSN Mark III VLBI data that have been correlated at the JPL/CIT Block II correlator, with observable (e.g., delay and amplitude) extraction by

⁴ R. F. Coker, "Mark III Correlated Flux Densities on the DSN," JPL Interoffice Memorandum, 335.6-91-028 (internal document), Jet Propulsion Laboratory, Pasadena, California, December 31, 1991.

the program FIT [7]. A master program takes the schedule file for the DSN antennas, correlation amplitudes, and observing frequencies from FIT, the measured system temperatures from antenna log files, and the modeled antenna gain and produces a correlated flux density and a total

$$T_{sys} = T_{zen} \frac{10(P/10)}{10(P_{zen}/10)} \quad (3)$$

where P is a power reading (in dBm) for that observation, P_{zen} is a power reading (in dBm) taken at zenith on

T_{FO} , were incorrect before April 1991.⁶ This introduced an error of a few K in T_{zen} . Formal precision measurement errors for T_{zen} appear small (~ 0.1 K), but the stability of T_{LNA} , T_{FO} , and T_{amb} values (as well as operator errors and weather effects) suggest that the overall accuracy of T_{zen} , as calculated above, is ~ 5 percent.^{7,8}

The method for converting observed power to temperature, as described by Eqs. (3) and (5), implicitly assumes that for a given bandpass, power is directly proportional to the system temperature. For very large values of total power (corresponding to system temperatures of ~ 300 K), this assumption of detector linearity breaks down (at the 1- to 2-percent level).

C. Antenna Gains

The antenna gain for each station is modeled using experimentally determined functions; only a summary of the gain estimation procedure is given here (see [8] for details). The desired quantity is gain, G , in units of K/Jy (to use in Eq. (1)):

$$G = \frac{\eta \text{ Area}}{2 k} \quad (6)$$

where k is Boltzmann's constant, η is the antenna aperture efficiency, and *Area* is the antenna's total physical collecting area. Variations in η across a given band are assumed to be negligible. The antenna aperture efficiency depends on the deformation of the antenna from its optimal shape, on weather conditions, and on the basic characteristics of the antenna surface. Thus, it can be written as

$$\eta = \eta_{max} \eta_{def} \eta_{atm} \quad (7)$$

where η_{max} is the aperture efficiency in a vacuum with no deviations from a perfectly shaped surface; η_{def} accounts for efficiency losses due to antenna gravitational deformation; and η_{atm} accounts for losses due to atmospheric attenuation. The value η_{max} is constant for a given antenna while η_{def} and η_{atm} depend on antenna elevation and weather conditions. The overall accuracy of

⁶ Personal communication from T. Anderson, TDA Mission Support and DSN Operations Section, Jet Propulsion Laboratory, Pasadena, California, to C. Jacobs, member of technical staff, Tracking Systems and Applications Section, Jet Propulsion Laboratory, Pasadena, California, April 1, 1991.

the antenna gains modeled in this fashion (exclusive of antenna pointing effects; see Subsection III.C) is estimated at ~ 5 percent.

III. Consistency Checks and Integrity Testing

A number of tests were carried out to verify the validity of the flux density catalogs. Since standard flux monitoring techniques (such as antenna nodding, repeated zenith calibrations, and standard source comparisons) are impractical for reference frame experiments due to time constraints, the internal and absolute accuracy of the DSN Mark III flux densities had to be determined in other ways.

A. Correlation Amplitudes and Their Normalization

The b -factor in Eq. (1) normalizes the correlation amplitude so that a signal with 100-percent correlation produces an amplitude of unity. Assuming identical bandpasses and a flat power spectrum and ignoring Doppler and polarization effects, one has, for some frequency ω [11],

$$b_\omega = a_1 \frac{(BW_0) B^2}{D_N} \quad (8)$$

where

$a_1 = 1.176$ for a three-level lobe rotator, one path [6]

BW_0 = nominal or Nyquist bandwidth (2 MHz for a single Mark III channel)

B^2 = bandpass amplitude at ω

$D_N = \int_0^\infty B^2 d\omega$

Note that the correction for single-bit quantization ($\approx 2/\pi$) is included in the expression for S_{cor} in Eq. (1).

The value B^2 is included in FIT's calculation of the correlation amplitude, ρ . This results in a b -factor that is independent of frequency. The value of B^2 was empirically calculated from over 200 Mark III and Mark II observations; the results, which agree reasonably well with theoretical bandpass shapes, are listed in Table 1. These values are used in FIT to weight the amplitude of each bin to produce one composite amplitude for each band. Also included in Table 1 are values for D_N . The value D_N is

Variable Doppler shifting due to Earth rotation causes the two station bandpasses to be mismatched by up to 20 kHz at X-band; this has been ignored. The frequencies of the local oscillators are not adjusted to correct for this Doppler shift. In addition, there will be some small changes in b due to mismatched bandpasses as a result of polarization effects (the actual size of this effect is unknown but DSN specifications suggest it is ~1 percent). These effects, combined with aliasing of frequencies above 2 MHz, correlation of neighboring bits due to oversampling, and incomplete sideband rejection imply that b will vary at the ~2-percent level. These effects, plus errors in modeling, give an estimated maximum error of approximately 5 percent in the b -factor.

Two errors in the standard data path were uncovered during this work. The first involved the fractional bit-shift (FBS) correction at the correlator. If the FBS correction is not done, the upper and lower 10 percent of the bandpass lose ~10 to 15 percent in amplitude [12]. This causes a composite loss in band amplitude of ~4 percent. Before June 10, 1991, the Block II correlator was not doing this correction properly [13].

The second error occurred in FIT. Instead of summing the correlator sine and cosine counts, FIT averaged the amplitudes over time and frequency; this resulted in too large an amplitude for observations with low SNR. This has been corrected. FIT now also makes a correction for a nonzero residual fringe rate.⁹ FIT amplitudes are normalized with respect to the number of lags, the number of time points, and the number of frequency bins containing good data.

To verify the above calculations for correlated flux densities, a short baseline experiment was conducted. On a baseline of 500 m, the correlated flux densities of the compact extragalactic radio sources used for VLBI astrometry are almost the same (within 10 percent) as the total flux densities. Figure 1 is a plot of the experiment's final visibilities (S_{cor}/S_{total}) versus scan number. Fitting a constant to the visibilities gave a ratio of 0.92. This suggests one or more of the following: The sources are all partially resolved, the antenna gains used in the data analysis are too high, the system temperatures are too low, or the b -factor is too large. However, the chi-square of a fit to unity is only 0.72, so changing the b -factor or antenna gain models is unjustified. *If the sources are resolved by a few percent on average, as expected, the b-factor is accurate to about 5 percent.*

⁹ S. C. Unwin, "Amplitude Errors in AIPS Global Fringe-Fitting,"

B. System Temperatures

The power meter program, QMVLBI, has a zenith system temperature calibration procedure, as described above. The calibration is performed at the start of an experiment. As time passes, the accuracy of the calibration degrades due to weather changes and receiver gain variations. The receiver gain was checked for variability over a 24-hour experiment and was found to be extremely stable (less than 1-percent variance in the power-to-temperature conversion factor). However, changes in the outside temperature over 24 hours can cause a significant (up to 5-percent) change in the zenith system temperature. Weather changes (such as a passing cloud) can cause a more substantial change in system temperature at low elevations; these are measured by the power meter.

The power meters measure power across the entire high electron mobility transistor (HEMT) bandwidth (400 MHz at X-band and 100 MHz at S-band). The 14 recorded 2-MHz channels (which are not contiguous) are scattered throughout this range. Therefore, any radio interference falling outside the observed channels (but within the HEMT bandwidth) will cause a spuriously high system temperature reading. In the next version of station software, it will be possible to take system temperature measurements for each channel,¹⁰ thereby eliminating this potentially serious problem. Until then, it is assumed that variations in system temperature (as well as receiver gain) across an entire band are negligible.

System temperature measurements using the power meters, the strip charts, and results from modeling all agree within the expected errors. Figure 2 is an example of temperature versus air mass for each observation for an experiment with measurements by all three methods. Note that even in the absence of any station system temperature measurements during an experiment, models can be used to get system temperatures that agree with the power meter values to about 20 percent.

On- and off-source system temperature measurements are estimated in a somewhat subjective fashion. For strong sources, inspection of the power meter data clearly reveals when the antenna gets on source and when it moves off. The magnitudes of the observed sharp rise and fall of the system temperature values are used to estimate the antenna temperature (that part of the system temperature due to the source). For the case of weak sources, it is not always possible to determine the antenna temperature. In addition, the error in antenna temperature is

¹⁰ DSN Systems Requirements Document 820-13, Module VLB-13-5.

difficult to determine since the bulk of the error in system temperature is not statistical in nature. Repeatability suggests an error of about 15 percent in antenna temperature (down to the detection limit of 0.1 K). Note that variations in antenna temperature across the 400-MHz bandwidth spanned by the wideband amplifier at X-band due to a nonzero spectral index will be less than 3 percent for a typical VLBI source. Also, the next version of station software will be able to take power data with finer time and power resolution than the present power meters; this will lower the error in total flux density for weak sources.

C. Antenna Pointing and Gain

Total flux density measurements are determined only from antenna temperature and antenna gain. Therefore, using a strong source (so as to minimize antenna temperature errors) with a known total flux density, it is possible to estimate the antenna gain. Some observations of the Very Large Array (VLA) calibrator source 3C 286 were used to compare the modeled antenna gains with gains estimated in this fashion. The resulting total flux densities were consistent at the 20-percent level.

An error in antenna pointing will yield reduced antenna temperatures and correlation amplitudes, giving spuriously low values of both total and correlated flux densities. Reference frame experiments use blind pointing; pointing correction tables, squint corrections, and refraction correction calculations are estimated to yield a blind pointing accuracy of ~ 10 mdeg for the 34-m antennas.¹¹ The refraction correction model used at the antennas was verified

D. Comparison of Results With a Known Structure

The source 4C 39.25 has been a fairly well-studied source at high frequencies (see, for example, [14]). It is known to be extended in the east-west direction with knots in its structure. At X-band it is also highly time-variable [15]. Therefore, X-band correlated flux densities on two baselines on a monthly basis can give little source structure detail; indeed, it is found that measurements vary by more than a factor of 2 from month to month. The picture at S-band is much clearer; Fig. 3 is a plot of measured correlated flux density versus the projected east-west baseline length. The smoothness of the curve implies a smooth underlying source structure (to the limit of the resolution of the interferometer) and little time variability. The fact that the correlated flux density depends solely on the absolute value of the east-west projected baseline length suggests that the source is not resolved in the north-south direction. This is completely consistent with maps of 4C 39.25 that are in the literature.

IV. Error Budget

The variance, σ^2 , in measured correlated flux density (for $\rho\pi/2 \ll 1$), derived from Eq. (1), is:

$$\begin{aligned} \sigma^2 = S_{cor}^2 & \left(\frac{\sigma_\rho^2}{\rho^2} + \frac{\sigma_b^2}{b^2} + \frac{\sigma_{T_1}^2}{4 T_1^2} \right. \\ & \left. + \frac{\sigma_{T_2}^2}{4 T_2^2} + \frac{\sigma_{G_1}^2}{4 G_1^2} + \frac{\sigma_{G_2}^2}{4 G_2^2} \right) \end{aligned} \quad (9)$$

V. Summary

With the methods described here, it is possible to provide correlated and total flux density measurements accurate to ~ 20 percent. This accuracy is sufficient to de-

termine which antennas are necessary to observe a given source on a given baseline at a given time. A complete correlated and total flux density catalog can now be built. However, to get better than 20-percent accuracy, the pointing problems of the DSN antennas must be corrected.

Acknowledgments

Dave Fort was of great help in explaining the Block II correlator. Steve Lowe assisted with many discussions on correlated amplitudes. Roger Linfield provided invaluable comments and critiques.

References

- [1] O. J. Sovers, C. D. Edwards, C. S. Jacobs, G. E. Lanyi, K. M. Liewer, and R. N. Treuhaft, "Astrometric Results of 1978–1985 Deep Space Network Radio Interferometry: The JPL 1987-1 Extragalactic Source Catalog," *Astron. J.*, vol. 95, pp. 1647–1658, 1988.
- [2] R. N. Treuhaft and S. T. Lowe, "A Nanoradian Differential VLBI Tracking Demonstration," *TDA Progress Report 42-109*, vol. January–March 1992, Jet Propulsion Laboratory, Pasadena, California, pp. 1–15, May 15, 1992.
- [3] T. F. Haddock, H. D. Aller, and M. F. Aller, "Frequent Observations of Extragalactic Compact Sources at 24 GHz," *Astron. J.*, vol. 93, pp. 1356–1367, June 10, 1987.
- [4] A. E. Rogers and T. A. Clark, *Mark III Data Acquisition Terminal*, Crustal Dynamics Project, Goddard Space Flight Center, Greenbelt, Maryland, January 1982.
- [5] J. S. Ulvestad and R. P. Linfield, "The Search for Reference Sources for Δ VLBI Navigation of the Galileo Spacecraft," *TDA Progress Report 42-84*, vol. October–December 1985, Jet Propulsion Laboratory, Pasadena, California, pp. 152–163, February 15, 1986.
- [6] A. R. Thompson, J. M. Moran, and G. W. Swenson, Jr., *Interferometry and Synthesis in Radio Astronomy*, New York: John Wiley and Sons, pp. 10, 88, 248, and 305, 1986.
- [7] S. T. Lowe, *Theory of Post Block II VLBI Observable Extraction*, JPL Publication 92-7, Jet Propulsion Laboratory, Pasadena, California, July 15, 1992.
- [8] P. H. Richter and S. D. Slobin, "DSN 70-Meter Antenna X- and S-Band Calibration Part I: Gain Measurements," *TDA Progress Report 42-97*, vol. January–March 1989, Jet Propulsion Laboratory, Pasadena, California, pp. 314–346, May 15, 1989.

- [9] S. D. Slobin and P. H. Richter, "DSN 70-Meter Antenna X- and S-Band Calibration Part II: System Noise Temperature Measurements and Telecommunications Link Evaluation," *TDA Progress Report 42-97*, vol. January–March 1989, Jet Propulsion Laboratory, Pasadena, California, pp. 352–366, May 15, 1989.
- [10] C. C. Chao, *The Tropospheric Calibration Model for Mariner Mars 1971*, Technical Report 32-1587, Jet Propulsion Laboratory, Pasadena, California, pp. 61–76, March 1974.
- [11] J. B. Thomas, *Interferometry Theory for the Block II Processor*, JPL Publication 87-29, Jet Propulsion Laboratory, Pasadena, California, October 15, 1987.
- [12] M. L. Meeks, ed., *Methods of Experimental Physics*, vol. 12, Part C, Astrophysics, Radio Observations, New York: Academic Press, 1989.
- [13] S. C. Unwin, ed., *JPL/Caltech Block II VLBI Correlator Newsletter*, vol. 2, no. 3, California Institute of Technology, Pasadena, California, July 1991.
- [14] D. Shaffer and A. Marscher, "4C 39.25: Superluminal Motion Between Stationary Components," in *Superluminal Radio Sources*, edited by J. A. Zensus and T. J. Pearson, New York: Cambridge University Press, pp. 67–71, 1987.
- [15] J. M.Marcaide, A. Alberdi, P. Elosegui, A. P. Marscher, Y. F. Zhang, D. B. Shaffer, C. J. Schalinski, A. Witzel, N. Jackson, and G. Sandell, "Detection of a New Component in the Peculiar Superluminal Quasar 4C 39.25," in *Parsec-scale Radio Jets: Proceedings of a Workshop Held at the National Radio Astronomy Observatory, Socorro, New Mexico, October 17–18, 1989*, edited by J. A. Zensus and T. J. Pearson, New York: Cambridge University Press, pp. 59–65, 1990.

Table 1. *b*-factor bandpass parameters for the Block II correlator.

	Mark III				Mark II			
	8 lags		16 lags		8 lags		16 lags	
	Empirical	Theory	Empirical	Theory	Empirical	Theory	Empirical	Theory
B^2 for bin 1	0.840	0.899	0.830	0.889	—	0.907	1.050	0.892
bin 2	1.002	1.002	0.992	1.015	—	0.993	1.122	1.015
bin 3	1.023	0.997	0.987	0.994	—	1.019	1.020	0.994
bin 4	0.643	0.604	1.003	1.003	—	0.523	1.004	1.004
bin 5	—	—	1.011	0.997	—	—	0.978	0.996
bin 6	—	—	1.001	0.977	—	—	0.935	0.999
bin 7	—	—	0.877	0.813	—	—	0.913	0.804
bin 8	—	—	0.467	0.408	—	—	0.337	0.244
D_N	0.877	0.876	0.896	0.887	—	0.871	0.921	0.869

Notes:

B^2 and D_N are relative to a square bandpass of amplitude unity.

No S- or X-band Mark II 8-lag data were available for analysis.

In 8-lag mode, bins 5 through 8 contain no signal.

Empirical fits were made relative to bin 2 (for 8 lags) or bin 4 (16 lags).

Mark II theoretical calculations based on an 11-pole Butterworth filter.

Table 2. Error sources.

Error source	Resulting fractional error, percent	
	X-band	S-band
Gain (each station)		
Weather and modeling	5	5
Pointing errors	0 to 50	0 to 10
<i>b</i> -factor	5	5
Amplitudes		
Signal to noise	0 to 15	0 to 15
FIT approximations	5	5
System temperatures (each station)		
Using model	10 to 15	15 to 20
Using strip chart	10	15
Using power meters	5	10
Estimated total (using power meters with T_{zen} at both stations)		
Clear weather, high elevation, high SNR source	10	10
Poor weather, low elevation, low SNR source	30	20

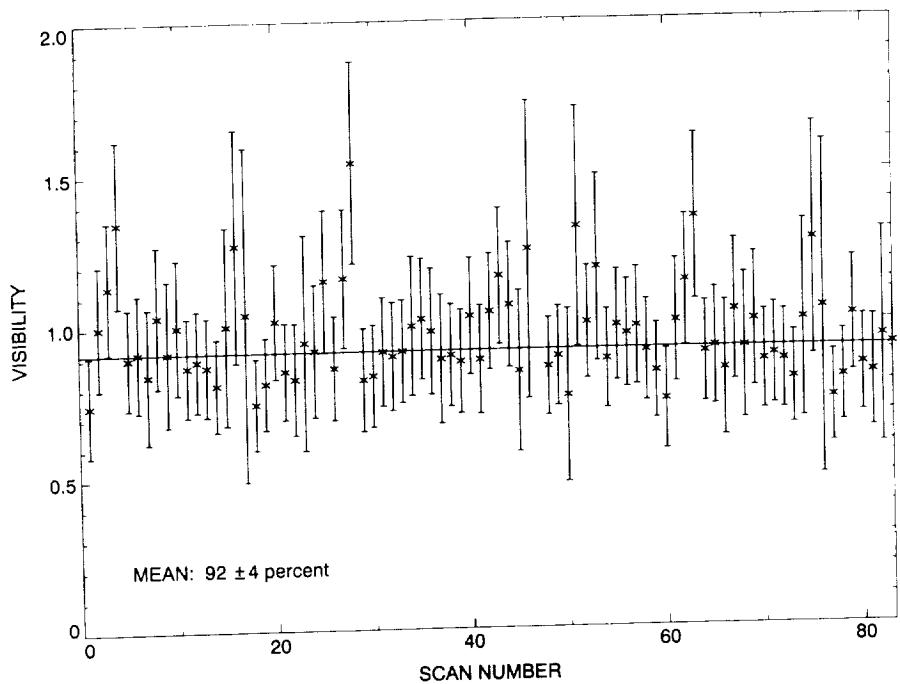


Fig. 1. The final visibilities (S_{cor}/S_{total}) for the short-baseline experiment 92CC124 using DSS 14 and DSS 15. The values for the total flux density, S_{total} , were taken from antenna temperature measurements at DSS 15. A weighted fit of zero slope was made to the data. The error bars are calculated from each observation's error in total flux density and in correlated flux density. For some scans, the antenna temperature was comparable to the resolution limit of 0.1 K; hence the large error bars.

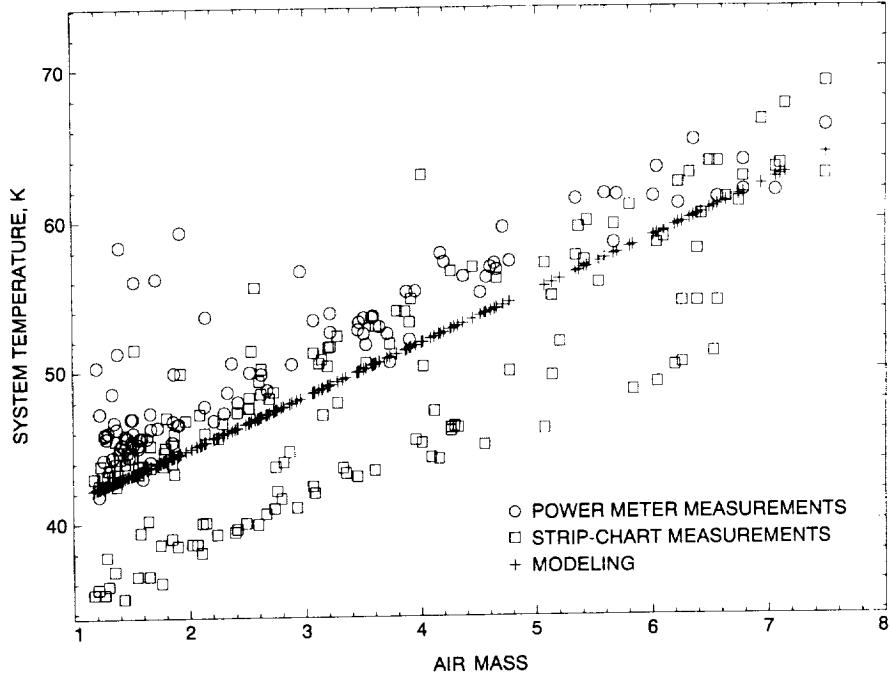


Fig. 2. A plot of X-band (8.4-GHz) system temperature versus air mass ($1/\sin(\text{elevation})$) for the three different methods for the experiment 90CA049. In most cases, agreement of the model with the power meter data is within 20 percent. The strip-chart data show an unexplained (and probably unreal) drop in system temperature halfway through the experiment.

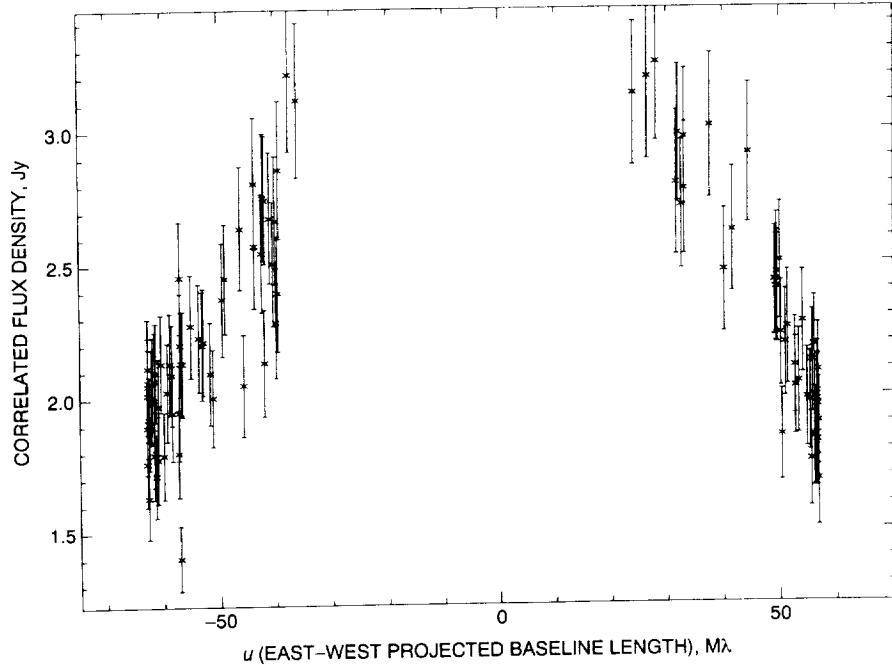


Fig. 3. A plot of correlated flux density versus u for the source 4C 39.25 at 2.3 GHz. The smoothness of the curve suggests smooth source structure. The dependence of correlated flux density solely on the absolute value of u suggests that the source is not resolved in the north-south direction but that it is extended in the east-west direction.