

information from fusion of complex images acquired by different sensors such as SAR, visible, and infrared [3].

References: [1] Kher A. and Mitra S. (1992) *Proc. SPIE*. [2] Maragos P. (1989) *IEEE Trans. Pattern Anal. Mach. Intellig.*, 11. [3] Mitra S. and Kher A. (1992) Paper presented at the International Space Year Conference at JPL, Pasadena, California, 10-13 February, 1992.

A UNIQUE PHOTON BOMBARDMENT SYSTEM FOR SPACE APPLICATIONS. E. J. Klein, KET Canada Inc./Sol-RF Energy Systems Inc., Box 2550, Winnipeg, Canada, R3C 4B3.

The innovative (patents pending) Electromagnetic Radiation Collection and Concentration System (EMRCCS) described here is the foundation for the development of a multiplicity of space and terrestrial system formats. The system capability allows its use in the visual, infrared, and ultraviolet ranges of the spectrum for EM collection, concentration, source/receptor tracking, and targeting.

The nonimaging modular optical system uses a physically static position aperture for EM radiation collection. Folded optics provide the concentration of the radiation and source autotracking. The collected and concentrated electromagnetic radiation is utilized in many applications, e.g., solar spectrum in thermal and associative photon bombardment applications for hazardous waste management, water purification, metal hardening, hydrogen generation, photovoltaics, etc., in both space and terrestrial segment utilization. Additionally, at the high end of the concentration capability range, i.e., 60,000+, a solar-pulsed laser system is possible.

The system outputs the concentrated flux, orthogonal (normally incident) to the input plane of an output port. The orthogonality remains constant regardless of the radiation input angle to the collection aperture, allowing simplification of radiation receptor design and highly efficient utilization of the concentrated radiation. The system configuration is arrayed for extremely high levels of flux concentration in windowing and targeting applications. Other system design formats provide power generation and thermal processes for heating and absorption cooling.

Fixed portable and mobile (space and terrestrial) applications include designs that incorporate a phased RF and/or the system array for purposes of radiation source acquisition/tracking and data derivation. The data is utilized in source acquisition (array capture angle of $\pm 75^\circ$ in the orthogonal E and H planes), source autotracking in the same angular intervals, and, subsequent to source and receptor acquisition, control of direction and magnitude of the output concentrated radiation at a given target range. In addition, the phased array can provide EM channel voice or data capability.

DETECTION OF OTHER PLANETARY SYSTEMS USING PHOTOMETRY. D. Koch¹, W. Borucki¹, and H. Reitsem², ¹Mail Stop 245-6, NASA Ames Research Center, Moffett Field CA 94035, USA, ²Ball Aerospace Systems Group, P.O. Box 1062, Boulder CO 80306, USA.

Detection of extrasolar short-period planets, particularly if they are in the liquid-water zone, would be one of the most exciting discoveries of our lifetime. A well-planned space mission has the capability of making this discovery using the photometric method.

An Earth-sized planet transiting a Sun-like star will cause a decrease in the apparent luminosity of the star by one part in 10,000 with a duration of about 12 hours and a period of about one year. Given a random orientation of orbital plane alignments with the line-of-sight to a star, and assuming our solar system to be typical, one would expect 1% of the stars monitored to exhibit planetary transits. A null result would also be significant and indicate that Earth-sized planets are rare.

For the mission to be successful one needs a sensor system that can simultaneously monitor many thousands of stars (F, G, and K dwarfs) with a photometric precision of one part in 30,000 per hour of integration. The stellar magnitude, integration time, and desired photometric precision determine the aperture size. The field of view and limiting stellar magnitude determine the number of stars that can be monitored. A 1.5-m telescope is required to attain the photometric precision for 12.5 mag stars. An 8° field of view will yield many thousands of stars and several transit detections per month. Confirmation of a detection will involve detection of a second transit that will yield a period and predict the time for a third and subsequent transits.

The technology issues that need to be addressed are twofold: One is for an appropriate optical design; the other is for a detector system with the necessary photometric precision. Two candidates for the detector system are silicon diodes and CCDs. It has been demonstrated that discrete silicon diodes have the required precision. However, the technology for building them into arrays with readouts needs development. The other approach is to use silicon CCDs. These already exist as arrays. However, the required photometric precision technology has yet to be demonstrated. Data processing complexity can be reduced by using the local-area-readout technique to obtain the flux for a few hundred stars per CCD.

AN INTEGRATED XRF/XRD INSTRUMENT FOR MARS EXOBIOLOGY AND GEOLOGY EXPERIMENTS. L. N. Koppel¹, E. D. Franco¹, J. A. Kerner¹, M. L. Fonda², D. E. Schwartz², and J. R. Marshall², ¹ARACOR, 425 Lakeside Drive, Sunnyvale CA 94086-4701, USA, ²Mail Stop 239-12, NASA Ames Research Center, Moffett Field CA 94035-1000, USA.

By employing an integrated X-ray instrument on a future Mars mission, data obtained will greatly augment those returned by Viking; details characterizing the past and present environment on Mars and those relevant to the possibility of the origin and evolution of life will be acquired. A combined XRF/XRD instrument has been breadboarded and demonstrated to accommodate important exobiology and geology experiment objectives outlined for MESUR and future Mars missions. Among others, primary objectives for the exploration of Mars include the intense study of local areas on Mars to "establish the chemical, mineralogical, and petrological character of different components of the surface material; to determine the distribution, abundance, and sources and sinks of volatile materials, including an assessment of the biologic potential, now and during past epochs; and to establish the global chemical and physical characteristics of the martian surface" [1].

The XRF/XRD breadboard instrument identifies and quantifies soil surface elemental, mineralogical, and petrological characteristics and acquires data necessary to address questions on volatile abundance and distribution. Additionally, the breadboard is able to