

The Carnegie Hubble Program: The Infrared Leavitt Law in IC 1613

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ABSTRACT

We have observed the dwarf galaxy IC 1613 at multiple epochs in the mid-infrared using *Spitzer* and the in the near-infrared using the new FourStar near-IR camera on Magellan. We have constructed Cepheid period–luminosity relations in the J , H , K_s , [3.6] and [4.5] bands and have used the run of their apparent distance moduli as a function of wavelength to derive the line of sight reddening and distance to IC 1613. Using a nine-band fit, we find $E(B-V) = 0.05 \pm 0.01$ mag and an extinction-corrected distance modulus of $\mu_0 = 24.29 \pm 0.03_{\text{statistical}} \pm 0.03_{\text{systematic}}$ mag. By comparing our multi-band and [3.6] distance moduli to results from the tip of the red giant branch and red clump distance indicators, we find that metallicity has no measurable effect on Cepheid distances at $3.6 \mu\text{m}$ in the metallicity range $-1.0 \leq [Fe/H] \leq 0.2$, hence derivations of the Hubble constant at this wavelength require no correction for metallicity.

Subject headings: Galaxies: distances and redshifts — Galaxies: individual: IC 1613 — Infrared: Galaxies — Infrared: stars — Stars: variables: Cepheids

1. Introduction

The stated aim of the Carnegie Hubble Program (CHP) is to measure the Hubble constant to an accuracy of 2% using data from the Warm *Spitzer* mission, and future data from *JWST* and Gaia (see Freedman et al. 2011, for a summary). The calibration of the CHP distance scale is based on mid-infrared observations of Cepheids in the Milky Way and Local Group galaxies. The distances to individual galaxies are measured by comparing the period–luminosity (PL) relations of their Cepheid populations to that of the Milky Way. In this paper we establish a precise and accurate distance to the Local Group dwarf galaxy IC 1613 using its known population of Cepheids.

Although the slope of the PL relation at a given wavelength is not observed to vary from galaxy to galaxy (Pietrzyński et al. 2006), the sensitivity of the zero–point to various factors is still being debated (e.g. Storm et al. 2011). For example, there has been much discussion over the last few decades regarding the sensitivity of the PL zero–point to metallicity (see Romaniello et al. 2008, for an overview of recent measurements). To test for such an effect we can either observe a galaxy with a known metallicity gradient (as first suggested by Freedman & Madore 1990 using M31, and later applied to M101 by Kennicutt et al. 1998 and M33 by Scowcroft et al. 2009), or look at a selection of galaxies of different metallicities with an independent distance indicator (e.g. Lee et al. 1993; Sakai et al. 2004).

With $[Fe/H] \simeq -1$ (Dolphin et al. 2001) IC 1613 is more metal–poor than the SMC, making it extremely useful in setting the the low–metallicity calibration of the PL relation. In this paper we compare the distance obtained from the PL relation in the near- and mid-infrared with that measured from the tip of the red giant branch (TRGB) method. Theory suggests that the effect of metallicity in the mid-infrared will be negligible (e.g. McGonegal et al. 1982); however, this has yet to be demonstrated conclusively. The uncertainty in the effect of metallicity on Cepheid magnitudes is one of the dominant systematics still remaining in the extragalactic distance scale. The present test with IC 1613, along with the other metallicity tests described in Freedman et al. (2011) allow us to measure the size and sense of the effect.

IC 1613 was discovered by Wolf (1906). It is a type IB(s)m dwarf galaxy in the Local Group (Sandage 1971; de Vaucouleurs et al. 1991), situated at high galactic latitude at a consensus distance of 736 ± 49 kpc¹. This converts to a true distance modulus of $\mu = 24.33 \pm 0.07$ mag, which is slightly closer than M31. IC 1613 is highly resolved and its position above the plane of the Milky Way results in low foreground extinction. As a dwarf galaxy, IC 1613 should have low internal extinction (see below), it is an ideal system for which to measure a distance, as well as to test and compare different distance indicators.

Studies of the Cepheids in IC 1613 began with Sandage (1971), who completed the unpublished work begun by Baade over forty years before. Baade chose to observe IC 1613 for the same reasons

¹Average distance based on 57 measurements in NED: <http://ned.ipac.caltech.edu/cgi-bin/nDistance?name=IC+1613>

as we do today: a resolved stellar population and low internal extinction. Baade deduced the latter from the fact that many background galaxies were visible through the main body of IC 1613. This fact meant that any dispersion in the period–luminosity relation would most probably be due to effects intrinsic to the Cepheids themselves, rather than differential reddening, say. Sandage (and Baade) found an anomalously and significantly shallower slope of the period–luminosity (PL) relation, -1.52 versus -2.85 in the blue, the latter established for Local Group galaxies. However, the observations were compromised by calibration issues and especially small–number statistics. Later work by Freedman (1988b) resolved the issue, showing that the slope of the PL relation in the visible did not change between galaxies.

This is not the first test of the Cepheid metallicity effect using IC 1613. Lee et al. (1993) compared TRGB distances of several local group galaxies to those from Cepheids and RR Lyrae stars and found no significant trend in $\Delta\mu$ with $[Fe/H]$ in the I band. Udalski et al. (2001) found 134 Cepheids during the OGLE survey of this galaxy. They compared the V , I and Wesenheit parameter W_{VI} Cepheid distances to the TRGB distance and again found no metallicity effect at those wavelengths. We are repeating their test much further redwards with FourStar (J , H and K_S) and IRAC (3.6 and 4.5 μm). The dispersion of the Leavitt law at this wavelength is minimised, not just by the reduction in differential reddening, but because the amplitudes of the Cepheid light curves (and hence the width of the instability strip) are minimised (Madore & Freedman 2012)

A preliminary measure of the mid–infrared PL relation in IC 1613 was made by Freedman et al. (2009, hereafter F09) who searched for Cepheids in archival *Spitzer* data. They found six Cepheids in the cold–mission data and presented PL relations in the 3.6 and 4.5 μm bands. We compare our results to previous work in Section 4.

2. Observations, Data Reduction, and Calibration

2.1. FourStar: J , H , and K_s

IC 1613 was observed on three nights using the recently commissioned FourStar wide field camera (Persson et al. 2013) on the Magellan Baade 6.5 m telescope at Las Campanas. Table 1 contains the dates and exposure details. IC 1613 virtually fills the 10.9×10.9 arcmin field of view so the sky background was determined by imaging an adjacent (sparse) field before and after each IC 1613 dither sequence. The same individual exposure times were used but a different number of co-adds and dithers were used to save on overhead. Sources were detected and masked in both the IC 1613 and sky frames. For each IC 1613 frame the nearest nine (in time) sky frames were combined using the unmasked region in common, then scaled to and subtracted from the on–target frames. The IC 1613 frames were then combined using an average with sigma-clipping and input rejection masks. This procedure is not expected to produce either flat or sky subtracted frames with zero offset. However, because the stellar photometry uses local sky measurements, the systematic errors due to offset levels should be negligible.

The photometry of the FourStar data was performed with DAOPHOT (Stetson 1987). For each field we identified stars to a signal-to-noise ratio $S/N \sim 3$. We generated a model of the PSF across the field from 100 isolated bright stars. We then performed PSF fitting photometry with ALLSTAR allowing the PSF model to vary linearly with x and y across the field.

To determine the precision of the photometry we added 650,000 artificial stars across each mosaic image. Stars were laid down in a grid with spatial offsets between stars of 40 pixels (10,000 stars at a time) so as to not increase the crowding. The process was repeated with different grid locations until the library of artificial stars was accumulated. We recovered the positions and photometry of the artificial stars by rerunning the same DAOPHOT routines used for the actual photometry. We achieved better than 10 % precision for stars brighter than $J \sim 21.0$, $H \sim 19.5$, and $K_s \sim 18.5$. At the high luminosity end the precision is better than 3%.

Photometric zero-points were determined for each epoch by matching 18 stars in the field of IC 1613 to the Two Micron All Sky Survey (Skrutskie et al. 2006). The uncertainties in the zero-points were found to be 0.015 mag or better. However, the scatter around the zero-point for the third epoch of data was about 4 times larger than for the first two epochs. The larger third epoch scatter was determined to have been caused by excellent seeing. Stellar profiles were $\sim 0.3''$ FWHM, causing the PSF to be under-sampled. We applied a smoothing kernel of 1.5 pixels across the mosaics and reran the photometry. This reduced the scatter between 2MASS and FourStar for that epoch to levels comparable to the first epoch. It also reduced the scatter between the FourStar photometry across all three epochs. It did not, however, change the zero-point significantly, viz., less than 0.005 mag.

2.2. Spitzer IRAC: 3.6 and 4.5 μm

The mid-infrared observations presented here were taken as part of the Warm Spitzer Program PID 61001. We observed in both the 3.6 μm and 4.5 μm channels with a frame time of 30 s. The galaxy was observed twelve times over fifteen months between 2010 January 26 and 2011 March 5. The observations were split into three blocks of roughly one month each and were spaced approximately evenly over that time. The dates of the observations are given in Table 2.

The IRAC camera on *Spitzer* has two operating channels: the 3.6 μm detector observes one field and the 4.5 μm detector simultaneously observes a closely adjacent field. At any given time one of the detectors will be centered on the target position and the other will be offset. When IRAC takes an exposure it can record both fields simultaneously. This is an advantage in programs such as ours, where a whole galaxy is to be surveyed, as it can cut down the total observation time needed to cover the field in both bands.

However, as the year progresses, *Spitzer* rotates about its axis so the position of the off-target field rotates around the on-target field. This means that in two thirds of our observations, the 3.6 μm off-target observations are to the SE of the main field and in one third they are to the NW

(and vice versa for the $4.5\ \mu\text{m}$ observations). Therefore, any objects that were not in the “main” field were imaged in only four or eight epochs, rather than all twelve.

Each observation produced two offset maps, containing the galaxy and its surrounding area, covering approximately 0.15×0.10 degrees, with a subset roughly 0.10×0.10 degrees covered by both channels.

2.2.1. IRAC Mosaic Creation

The data were analyzed in two ways: a time-resolved analysis using one mosaic per channel per epoch of observation, and an averaged analysis using a single mosaic per channel, comprising all of the time domain data for that wavelength. In addition to these mosaics, a “master” mosaic was built from all the frames in both channels (1944 frames) and used to determine accurate positions for all the stars in the field. The procedures are described in Section 2.2.5.

2.2.2. Time-averaged mosaics

The data from all epochs were stacked into two mosaics (one per channel) using MOPEX (Makovoz & Khan 2005). Each mosaic comprises 960 individual basic calibrated data frames (BCDs). Because of the rotation of the telescope, as described in Section 2, the exposure time coverage within the mosaic is not uniform. The region containing the galaxy has an average coverage of approximately 45 minutes per pixel in each channel, and has three times the coverage of the cold mission data described in F09. These images will be referred to as the “science mosaics”. The $3.6\ \mu\text{m}$ science mosaic is shown in Figure 1; the $4.5\ \mu\text{m}$ mosaic covers the same area. Approximately the central third of the image has the full twelve-epoch coverage, with the outer thirds having coverage at either four or eight epochs.

Due to the large ($1.2''$) pixel scale of IRAC, stellar profiles are badly under-sampled in single IRAC frames. However, making use of the large number of observations at each spatial position (960) allowed us to resample the images to achieve higher resolution. The mosaics were created using a pixel scale of $0.75''$. Other resolutions were tested but $0.75''$ pixels delivered the best sampled PSF and thus smooth, well-sampled profiles for the stars.

Finally, the science mosaics were converted from MJy sr^{-1} to data counts using the conversion factors and exposure times in the image headers. The conversion was performed so that ALLFRAME could give a correct estimation of the magnitude uncertainties.

When creating mosaics, MOPEX preserves the fluxes in the original pixels. This means that the variations in the Cepheid fluxes are not truly lost, they are just averaged over. Although the amplitudes of Cepheids in the mid-infrared can reach 0.6 mag, the average of twelve phase points drawn randomly from the light curve will give a good approximation of the mean flux. For example,

if we consider a Cepheid with a mid-infrared amplitude of 0.4 mag (typical of the Cepheids in our study), then the average of 12 random observations will have an uncertainty of 0.03 mag (details of this calculation are in the appendix of Scowcroft et al. 2011). It makes no difference to the Cepheids whether we average these points before or after photometry so long as flux is conserved. However, if we stack the twelve images first then we achieve a higher signal-to-noise ratio image and can detect fainter stars than if we examined single images. Hence, mosaicking the time-resolved images in a way that preserves flux will give us a good value for the average flux of the Cepheid.

2.2.3. *Single-epoch mosaics*

The single-epoch mosaics were created similarly to the time-averaged mosaics. Each one was made from 81 images and was resampled to a pixel scale of $0.6''^2$. The images were converted from MJy sr^{-1} to counts using the conversion factors and exposure times in the image headers. Again, the dither pattern meant that the mosaics were not uniformly exposed, but now had typical integration times of 5 minutes per pixel. Note that the single-epoch mosaics are shallower than the data used by F09. Consequently, the time-averaged mosaics were used for the final photometry, while the time-resolved data were used only to confirm that the stars identified as Cepheids were truly variable.

2.2.4. *Correction mosaics*

In addition to the science mosaics, MOPEX was used to make “correction” mosaics. These are made identically to the science mosaics, mosaicking the location-dependent correction images provided by the Spitzer Science Center³ in the same geometrical pattern used for the science mosaics. The correction mosaics are necessary as IRAC is not uniformly sensitive over its entire field of view. The non-uniform coverage depth of our science mosaics can further exacerbate the problem. Inspection of the correction mosaics showed that the residual location-dependent effect had mean values of approximately 2% in both channels, and would reach as high as 7% at $3.6 \mu\text{m}$ and 10% at $4.5 \mu\text{m}$ in particularly non-uniform regions, if left uncorrected.

²As the single-epoch mosaics were shallow we did not perform the same tests to find the optimum image scale, instead we used the default value.

³Details of the location-dependent photometric correction are given in section 4.5 of the IRAC instrument handbook.

2.2.5. DAOPHOT and ALLFRAME reduction

The photometry was performed using the DAOPHOT and ALLFRAME packages (Stetson 1987, 1994). The mosaics were run through DAOPHOT to detect the stars and create a point spread function (PSF) model. The detected objects were subtracted and the resulting frame processed again to detect the remaining objects. The PSF model was created using ~ 100 stars in each mosaic. The final photometry was done using ALLFRAME. ALLFRAME is preferred over ALLSTAR in studies where there are multiple frames of the same field as it will produce a master detection list. When the photometry is done using the positions in this list, rather than remeasuring the positions from each frame, the number of free parameters in the PSF fit are reduced. This significantly reduces the uncertainty in the final magnitude, and allows better de-blending of close sources. A detailed description of the process can be found in Stetson (1994). The master detection list was generated by running DAOPHOT on the master frame. The coordinate transformations between this frame and all the other mosaics were determined, then were input to ALLFRAME to produce the instrumental photometry.

2.2.6. Calibration

All the *Spitzer* photometry in the Carnegie Hubble Program is set to be on the standard system defined by Reach et al. (2005, hereafter R05). As we used mosaics for our photometry rather than single images we did not need to apply a pixel-phase correction. Any effects due to pixel phase should be adequately averaged over by the dithering and mosaicking. Placing the instrumental magnitudes on the R05 system was achieved by using the PSF stars in the science mosaics as local standards.

The PSF stars in the flux-units versions of the science mosaics (i.e. after mosaicking but before conversion to counts) were each measured using the PHOT aperture photometry routine in IRAF⁴. The *zmag* parameter in PHOT was set for each channel such that the procedure would output calibrated magnitudes.

Each of the stars was measured in a 3.6" radius aperture and sky annulus from 3.6" to 8.4", corresponding to the standard aperture set of 3, 3, 7 native IRAC pixels. To convert this to the 12" radius (10 native pixels) standard aperture used by R05, the latest warm mission aperture corrections of 1.128 and 1.127 in channels one and two, respectively (Spitzer Science Center, Private Communication, 2012) were applied to the measured fluxes.

The ALLFRAME photometry was calibrated by finding the offset between it and the corrected aperture magnitudes described above and correcting each star in the field accordingly. Finally, the

⁴IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy (AURA) under cooperative agreement with the National Science Foundation.

location–dependent correction was applied by measuring the pixel values in the correction mosaic at the positions of each star in our catalog and multiplying the flux of each star by this value.

3. The Cepheid Population of IC 1613

Cepheids were identified by matching to the catalog from the OGLE study (Udalski et al. 2001). We also identified V22 from Sandage (1971), a long period Cepheid which was not included in the OGLE catalog, but was included by F09. The Cepheids were initially identified by their positions derived from the science mosaics. The science mosaics were then visually inspected to check for possible nearby contaminants. The light curves generated from the single–epoch mosaics were inspected to check for variability, but this became increasingly difficult at periods below ten days as the uncertainties on the individual points were comparable to the amplitudes of the light curves.

Thirty-one stars measured at 3.6 and 4.5 μm were matched with the OGLE catalog, and of these twenty-two were measured at J, H , and K_s . This is just under 25% of the original OGLE sample. Unfortunately, the majority of Cepheids in IC 1613 have periods below 6 days. Short period Cepheids are naturally fainter and the majority of the OGLE sample fell below our detection limit.

Several of the detected Cepheids appeared anomalously bright for their known periods. After visual inspection of the images the stars were deemed to be blended with unresolved companions and were excluded from all further analysis. After blends were removed, twenty Cepheids remain. The photometry of the final sample of Cepheids is given in Table 3. The near–infrared time series data is given in Tables 4, 5 and 6.

3.1. Near–Infrared Period–Luminosity Relations

Period–luminosity relations were obtained in the J, H , and K_s bands. In this case we had three epochs that were reduced individually. The magnitudes are weighted means of the three observations of each Cepheid, and the uncertainties are the errors of the weighted mean. The systematic uncertainties in the near–infrared photometric zero–points are 0.017, 0.020 and 0.021 mag in J, H , and K_s , respectively.

The period–luminosity relations for the near–infrared J, H , and K_s bands are shown in Figure 2. Prior to fitting, the magnitudes were converted from the 2MASS system to the LCO photometric system, using the transformations described in Section 4.6 of Carpenter (2001). The LMC PL relations from Table 6 of Persson et al. (2004) are adopted as fiducial and are rewritten in the form

$$M = a(\log P - 1.0) + b. \tag{1}$$

We fixed the slopes (a) to the LMC values and used an unweighted least squares fit to find the zero–points (b). The resulting fits are listed in Table 7. The fits assume $\mu_{0,LMC} = 18.48 \pm 0.03$ mag,

as derived in Monson et al. (2012) and Freedman et al. (2012); this puts them on the same scale as the mid-infrared values. The apparent distance moduli derived from the near-infrared PL relations are 24.36 ± 0.05 , 24.31 ± 0.04 and 24.34 ± 0.05 mag in J , H , and K_s respectively.

3.2. Mid-Infrared Period-Luminosity Relations

The mid-infrared Leavitt laws take the same form as the near-infrared laws given in Equation 1, but in this case the slope, a , was taken from the Scowcroft et al. (2011) LMC results and the zero-points, b were derived using an unweighted least-squares fit.

The first comparison we make is with the earlier results of F09. They measured single-phase magnitudes for six IC 1613 Cepheids in data obtained from the *Spitzer* archive. The resulting PL relations are plotted in Figure 3, along with the CHP time-averaged magnitudes for the same six Cepheids.

We do not expect the values to be exactly the same; the amplitude of each Cepheid’s light curve will change its position relative to the ridge line of the PL relation in the F09 values (and to a much lesser extent in the CHP values depending on the dispersion of the phase points throughout the pulsation cycle). Hence it is not necessarily useful to compare the magnitudes of individual stars. However, *on average* the results should agree, such that we should get the same result when we fit the PL relation to either set of data.

The PL was found by fixing the slope to the values derived from the LMC (Scowcroft et al. 2011) and making an unweighted least-squares fit to the Cepheids with $6 \leq P \leq 60$ days. This period range was chosen to match the Milky Way and LMC samples used to define the IRAC PL relations. The fits to the F09 and CHP data are indistinguishable when plotted; the zero-points differ by only 0.02 mag, which is much smaller than the formal 1σ error on the fitted zero-points (0.15 and 0.12 mag in [3.6] and [4.5], respectively).

Finally, we re-fit the PL using the whole sample of Cepheids in the period range 6 to 60 days. A montage of the $3.6\mu\text{m}$ images of these Cepheids can be found in Figure 4. The results are given in Table 7, and the relations are plotted in Figure 5. By comparing the calculated zero-points to those in the Milky Way (MW) PL relations given in Monson et al. (2012) we find apparent distance moduli of IC 1613 of 24.31 ± 0.09 and 24.26 ± 0.08 mag using the [3.6] and [4.5] PL relations, respectively. The quoted uncertainties include the error estimates from both the MW and IC 1613 fits.

3.3. Reddening Corrected Distance Modulus

We can make use of the broad wavelength coverage of archival observations to derive the total line-of-sight reddening and extinction to the Cepheids in IC 1613. Figure 6 demonstrates the

technique. The distance moduli are plotted as a function of inverse wavelength in microns. Three extinction laws — the optical and near-infrared laws from Cardelli et al. (1989) and the mid-infrared law from Indebetouw et al. (2005)⁵ — are combined to fit to the data, assuming $R_V = 3.1$. The best fit $E(B - V)$ value was found by minimizing the dispersion of the distance moduli around the scaled and shifted extinction law, and was found to be $E(B - V) = 0.05 \pm 0.01$ mag. The top panel in Figure 6 shows the scaled, shifted extinction law with the apparent distance moduli at nine wavelengths; the dashed lines show how the fit changes if $E(B - V)$ is changed by 1σ . Applying this scaled correction to each of the apparent distance moduli and taking a weighted average results in an absolute distance modulus of $\langle\mu_0\rangle = 24.29 \pm 0.03_{stat} \pm 0.03_{sys}$ mag. The deviations of the extinction-corrected distances around this value are shown in the bottom panel of Figure 6. The systematic uncertainty comes from the LMC distance we adopt: $\mu_{LMC} = 18.48 \pm 0.03$ mag.

4. Independent Distance Comparisons

The nine-band fit presented in Figure 6 produces a reddening-corrected Cepheid distance modulus of $24.29 \pm 0.03_{stat} \pm 0.03_{sys}$ mag. In this section we compare our result to other recent measurements also using Cepheids, and then compared to other independent distance indicators.

4.1. Cepheid Comparison

The Auracaria project is using near-infrared observations of Cepheid populations to determine distances to nearby galaxies. We compare our results with their study of IC 1613 (Pietrzyński et al. 2006), in which the authors use the template fitting method of Soszyński et al. (2005) to obtain mean-light magnitudes of Cepheids in J and K from single-epoch observations. They observed 39 Cepheids in the galaxy, the majority of which are also observed in our study. Adopting the PL slopes from Persson et al. (2004) and adopting an LMC distance modulus of 18.50 ± 0.10 mag they find $\mu_J = 24.385 \pm 0.040$ mag and $\mu_K = 24.306 \pm 0.045$ mag. Both of these values agree with our FourStar results within the 1σ uncertainties.

Pietrzyński et al. (2006) combine their J and K distance moduli with the V and I values derived in OGLE II by Udalski et al. (2001) to derive a multi-wavelength fit for $E(B - V)$ and the extinction-corrected distance modulus. Using their four-band fit they measure $E(B - V) = 0.090 \pm 0.019$ mag and derive $\mu_0 = 24.291 \pm 0.035$ mag. Their derived reddening is slightly higher than our value of 0.06 ± 0.01 mag, but is barely outside the respective 1σ error bars. Their de-reddened distance, however, is in complete agreement with our value of 24.29 ± 0.03 mag.

More recently, Bernard et al. (2010) (B10) observed IC 1613 at optical wavelengths using the

⁵Use of the Indebetouw et al. (2005) law was justified by Monson et al. (2012).

Advanced Camera for Surveys (ACS) on the *Hubble Space Telescope* (HST). Their study looked at the fainter variable stars in the galaxy and found 44 Cepheids pulsating in either the fundamental or overtone modes. These Cepheids have short periods (the majority have $\log P < 0.5$) and were not detected in the CHP observations.

B10 derive the Cepheid distance using the Wesenheit index W_{VI} (Madore 1976). The W_{VI} index is reddening-free by design, hence is only affected by the choice of reddening law, and not at all by the total amount of reddening. Adopting an LMC distance modulus of 18.515 ± 0.085 (from Clementini et al. 2003), they find $\mu_{0,W} = 25.50 \pm 0.11$ using only the fundamental mode Cepheids. This value is significantly higher than the distance we find in Section 3.3. However, it is possible that this is due to the lack of a metallicity correction on $\mu_{0,W}$.

B10 make no correction for metallicity in their Cepheid analysis, and note that they do not believe a correction is necessary at the low metallicity of IC 1613. It may be the case that the metallicity correction required on $\mu_{0,W}$ decreases as we move to low $[Fe/H]$ populations, but this has not been proved conclusively. To this end we take the correction on $\mu_{0,W_{VI}}$ from Scowcroft et al. (2009) of $\gamma_{W_{VI}} = -0.29 \pm 0.11$ mag dex $^{-1}$ and apply it to the distance derived by B10. Assuming $12 + \log(O/H)$ of 8.34 (Sakai et al. 2004) and 7.90 (Bresolin et al. 2007) for the LMC and IC 1613 respectively, and now adopting the CHP LMC distance modulus of 18.48 mag, we find a metallicity corrected distance modulus of $\mu_{0,W,Z} = 24.33 \pm 0.14$ mag. This is now consistent with our value derived in Section 3.3, but its error bar is much larger than the original B10 result. This should be taken purely as an indication that the metallicity effect on the Leavitt Law zero-point in the optical bands requires more study, and that its effect at the lowest metallicities is not yet conclusively ruled out.

4.2. RR Lyrae Comparison

The Bernard et al. (2010) study not only found Cepheids in IC 1613 but also RR Lyrae stars. RR Lyrae stars obey a luminosity-metallicity relation at optical wavelengths. Adopting a mean metallicity of $[Fe/H] = -1.6 \pm 0.2$ they calculate the absolute magnitude of the horizontal branch in IC 1613 to be $M_V = +0.52 \pm 0.12$ mag, and derive a reddening corrected RR Lyrae distance modulus of 24.39 ± 0.12 mag. This is larger than, but still within 1σ of our multi-band Cepheid fit presented above.

4.3. Tip of the Red Giant Branch Comparison

An independent measure of the distance to IC 1613 can be obtained using the tip of the red giant branch (TRGB). The absolute magnitude of the TRGB is a physical property of the stellar population and does not depend on any measurements further down the distance ladder. Moreover, in the I band its absolute magnitude is relatively insensitive to metallicity and age effects, making

it a robust measure of distance for old resolved stellar populations, and an independent check of the Cepheid distance moduli we have presented in the previous section.

We compare our result to the work of Dolphin et al. (2001, hereafter D01), who derived the TRGB distance to IC 1613 using V and I band photometry from WFPC2 on *Hubble*. They provide two estimates of the TRGB apparent magnitude — 20.40 ± 0.09 from their own data and 20.35 ± 0.07 from a re-reduction of the data from Cole et al. (1999). The second value is more robust as it is measured from a region with higher stellar density, hence more stars on the red giant branch. They assume the absolute magnitude of the TRGB is $M_I = -4.02 \pm 0.05$, and an extinction of $A_I = 0.05 \pm 0.02$, resulting in an extinction-corrected distance modulus of $\mu_0 = 24.32 \pm 0.09$ mag.

The true distance modulus derived from our multi-band fit is consistent with the TRGB distance from D01. This suggests that metallicity effects are not significantly affecting the Cepheid distance modulus derived here; our PL relations were all calibrated to the MW and LMC which have much higher average metallicities than IC 1613. Note, however, that this result applies to the Cepheid distance modulus derived from a multi-band fit. It does not tell us anything about the effect of metallicity on an individual PL relation.

4.4. Dispersion of Independent Measurements

In Figure 7 we make a graphical comparison of our newly determined Cepheid distance to IC 1613 with the published record of prior distance determinations as found in the December 2012 release of the compilation of redshift-independent distances in NED-D. No attempt has been made to put any of these distances onto a common zero point; the data therefore reflect a variety of adopted reddenings, zero points and wavelengths. We have however, subdivided the data down to a comparison of three major methods: the tip of the red giant branch (TRGB) method, the RR Lyraes, red clump stars and previous determinations using Cepheids. The distance modulus determined here is shown as a vertical line in each of the plots. Individual determinations are shown as unit-area gaussians whose width is the published error on the mean. The solid line is the renormalized (Frequentist) sum of those individual probability density distributions (PDD); its peak is the mode of the published distribution. The filled circle (with horizontal error bars) is the median value of the PDD. The larger error bars capture 68% of the density around the median. The smaller error bar is the error on the mean.

On average, the (Population II) RR Lyrae variables are seen to give a slightly lower distance moduli than our (Population I) Cepheids. On the other hand, the (Population II) TRGB Method appears to give, on average, slightly larger moduli than our Cepheid modulus, although specific studies can be selected that agree exactly. The red clump moduli are few in number and widely scattered, although they do broadly agree with our Cepheid distance. The largest number of distance determinations come from previously published studies of the Cepheids themselves. Here we compare 31 previous determinations and remark that the mode of this distribution is in good agreement

with the latest value, although the range of values accumulated over the years is considerable.

In their figure 16, B10 plot the distance moduli in the literature when corrected to a common LMC distance and common $E(B - V)$. The dispersion is somewhat reduced, but we note that their adopted μ_{LMC} is slightly higher than the CHP value, and their adopted $E(B - V) = 0.025$ mag is approximately half the value we derive from our multi-band fit. It is clear from this that reddening and the LMC distance are the dominant systematics in the determination of the distance to IC 1613.

Fortunately, reducing these two systematics is entirely the domain of the CHP. Our distance ladder is tied to parallax measurements of MW Cepheids, and we virtually eliminate reddening in the mid-infrared. Our result, $\mu_0 = 24.29 \pm 0.03_{stat} \pm 0.03_{sys}$ reflects the reduced systematic uncertainty in these values, showing the power of moving to the mid-infrared for Cepheid distance studies.

4.5. Metallicity Effects in the Mid-Infrared

IC 1613 has a metallicity of $[Fe/H] \approx -1$ (D01), significantly lower than the MW and LMC which were the two calibration galaxies for the CHP PL relations. This makes it an ideal test-bed for searching for metallicity effects in the Cepheid PL. Initial tests for metallicity sensitivity in the mid-infrared were presented in Freedman et al. (2012), where we plotted the residual from the PL relation against spectroscopic metallicity for individual Cepheids. We found no significant effect at $-0.6 \leq [Fe/H] \leq 0.2$. Including IC 1613 in our studies increases the metallicity range of CHP Cepheids by a factor of two, so if a significant metallicity effect were present we should be able to detect it somewhere in this range.

To test for a metallicity effect in IC 1613 we must use a different tactic from the one we applied in the MW and LMC. As we do not have metallicity measurements of the individual stars we must treat them as an ensemble. We assume that there is no effect on the PL slope and that any difference would manifest itself in the zero-point. Therefore, if composition does have an effect, we should find a different distance modulus than with an independent measure such as the TRGB or red clump.

Correcting our [3.6] and [4.5] distances for extinction using $E(B - V) = 0.05 \pm 0.03$ mag we derive $\mu_{0,[3.6]} = 24.30 \pm 0.09$ mag and $\mu_{0,[4.5]} = 24.25 \pm 0.08$ mag. Both of these values are in excellent agreement with the TRGB distance from D01. The 4.5 μm distance modulus is slightly (0.05 mag) brighter than $\mu_{0,[3.6]}$ but still agrees to within one σ . As has been discussed previously (Freedman et al. 2011; Scowcroft et al. 2011; Monson et al. 2012), we believe that the [4.5] band is unsuitable for distance measurements as it is contaminated by the temperature and metallicity sensitive CO band-head at 4.6 μm . Henceforth, all references to our mid-infrared Cepheid distance pertain solely to the 3.6 μm measurement.

The consistency of our [3.6] distance modulus with the TRGB, red clump and multi-band distances shows again that metallicity is not significantly impacting the distance measurements over the range $-1.0 \leq [Fe/H] \leq 0.2$. Therefore we conclude that there is no effect of metallicity on the 3.6 μm Cepheid PL relation zero-point. The mid-infrared measurement of the Hubble constant needs no adjustment for metallicity effects.

5. Summary and Conclusions

We performed a multi-epoch survey of IC 1613 using *Spitzer* and the new FourStar camera on Magellan in the mid- and near-infrared. The photometric catalogs were matched to the OGLE Cepheid catalog to locate the Cepheids. Approximations of mean light were obtained for each star and PL relations were constructed in the J , H , K_s , [3.6] and [4.5] bands and distance moduli were derived. Using the 3.6 μm PL relation, where the effects of reddening are minimized, we measure the true distance modulus of IC 1613 as $\mu_{0,3.6} = 24.30 \pm 0.09_{stat} \pm 0.03_{sys}$ mag. This is entirely consistent with the independent TRGB and red clump distance moduli derived in Dolphin et al. (2001).

In addition to the single-band mid-infrared distance we have used near-infrared data from FourStar and archival optical data (corrected to an LMC distance of $\mu_{LMC} = 18.48 \pm 0.03$ mag) to derive a nine-band fit to measure the reddening and distance modulus of IC 1613. We find $E(B - V) = 0.05 \pm 0.01$ and $\mu_0 = 24.29 \pm 0.03_{stat} \pm 0.03_{sys}$ mag.

Finally, we have shown that as the mid-infrared Cepheid distance agrees with the TRGB distance, there must be no significant metallicity effect on the PL relation in the range $-1.0 \leq [Fe/H] \leq 0.2$. This removes any uncertainty in the CHP distance scale due to metallicity effects in the Cepheid calibration, significantly reducing the uncertainty in the CHP mid-infrared determination of the Hubble constant.

Facilities: Spitzer, Magellan:Baade

A. Fitting Using Template Lightcurves

The amplitudes of Cepheid light curves in the near-infrared are much smaller than at optical wavelengths (approximately 1/3 to 2/3 of the V or I band amplitudes — see Table 2 of Soszyński et al. 2005). However, they still reach levels around 0.5 mag which, combined with non-uniform sampling, can significantly affect the mean-light magnitude derived from a straight average. The effect of small numbers of non-uniform observations can be negated using template fitting, as was first shown by Freedman (1988a) and later elaborated upon by Soszyński et al.. In that paper they made template light curves in the J , H , and K_s bands that could be scaled and phased using a

complete V or I band light curve and a single near-infrared observation.

The technique was used by the Araucaria project to derive higher accuracy mean-light magnitudes from single observations in the J and K bands. They successfully applied the method to Cepheids in IC 1613 (Pietrzyński et al. 2006), obtaining PL relations in each band from either one or two observations per star. Here we test this method to derive mean-light magnitudes for our near-infrared Cepheids.

The most fundamental piece of information in the template fitting technique is the period of the Cepheid. From this and the V or I light curve the time of maximum light in the reference (optical) band is predicted. The phase-lag between the reference band and the near-IR band is known and can be used to predict the time of maximum light in each of the near-IR bands. The amplitude of the light curve is scaled to the amplitude of the reference light curve. The template is then fit to each near-IR observation individually and the mean-light magnitude is calculated; a weighted mean of these values gives the best estimate of the mean-light magnitude of the Cepheid.

It is imperative to have highly precise periods for the Cepheids; if the time of maximum light is computed incorrectly then the relative phases of each data point will be erroneous and the mean-light magnitude will be incorrect. As the observations in Pietrzyński et al. (2006) were taken several years after the data for the reference light curves they took three more V band observations contemporaneously with the near-IR data. This allowed them to refine the periods and define the time of maximum light ($\phi = 0$) more accurately. The periods typically changed by 0.1 to 0.5%, but over 10 years this ΔP is sufficient to shift the time of maximum light by as much as $\phi = \pm 0.5$ compared to the original estimate.

We do not have contemporaneous optical data to further refine the periods. We adopted the periods from Pietrzyński et al. (2006) where they were available and Udalski et al. (2001) in all other cases. To account for the less accurate periods the fitting algorithm was altered to adopt a phase shift. The best-fit phase shift was calculated by stepping through the possible shifts with a step size of $\delta\phi = 0.001$ and minimizing the residuals of the points around the template light curve. Example light curves for two Cepheids are shown in Figure 8. The phase shifts were found to be anywhere in the range $-0.4 \leq \Delta\phi \leq 0.4$. To reiterate, shifts of this size could be induced by a period change of less than 1% over 10 years. At this point it is clear that we do not know the periods of the Cepheids well enough to predict the time of maximum light to the required degree of accuracy, and so cannot determine the phase of any of our data points with a high degree of confidence.

To confirm these thoughts the PL relations were plotted using the template mean-light magnitudes. The resulting apparent moduli showed marginal changes — at the level of 1σ — but no significant differences.

We conclude that although our knowledge of the periods of the Cepheids is good enough to derive a PL relation, it is not sufficient for determining the time of maximum light to the accuracy required for the template fitting technique. For the rest of this work we adopt the regular mean

values for J , H , and K_s as listed in Table 3.

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Date	Band	Exposure time	N_{dither}	N_{coadd}	Total exposure (s)
2011-09-09	J	20.38	9	2	367
	H	8.733	9	4	314
	K_s	14.56	9	2	262
2011-10-04	J	20.38	9	2	376
	H	8.733	9	4	314
	K_s	14.56	9	2	262
2011-11-03	J	20.38	9	2	367
	H	5.822	9	6	314
	K_s	5.822	9	6	314

Table 1: FourStar observations of IC 1613.

Date	Block	Epoch
2010-01-26	1	1
2010-02-06	1	2
2010-02-14	1	3
2010-02-25	1	4
2010-08-20	2	5
2010-08-28	2	6
2010-09-07	2	7
2010-09-17	2	8
2011-02-03	3	9
2011-02-13	3	10
2011-02-23	3	11
2011-03-05	3	12

Table 2: IRAC observations of IC 1613.

Table 3. Near- and mid-infrared mean magnitudes of Cepheids found in IC 1613.

OGLE ID	Period (days)	RA (hh:mm:ss)	Dec (dd:mm:ss)	J (mag)	σ_J (mag)	H (mag)	σ_H (mag)	K_s (mag)	σ_{K_s} (mag)	[3.6] (mag)	$\sigma_{[3.6]}$ (mag)	[4.5] (mag)	$\sigma_{[4.5]}$ (mag)	S71
V22 ^{a,b}	123.880	1:05:00.701	+02:10:48.60	15.878	0.004	15.644	0.003	15.454	0.002	15.369	0.009	15.457	0.017	V22
11446 ^a	41.630	1:04:59.740	+02:05:28.30	17.139	0.012	16.970	0.009	16.866	0.007	16.600	0.014	16.610	0.018	V20
736	23.45	1:04:32.130	+02:05:01.90	17.903	0.162	17.325	0.043	V2
7647 ^a	16.540	1:04:37.700	+02:09:08.40	18.052	0.020	17.916	0.016	17.854	0.014	17.656	0.032	17.694	0.031	V18
13738	16.37	1:05:02.810	+02:10:35.10	18.440	0.027	18.210	0.019	18.025	0.014	18.040	0.021	18.199	0.021	V16
7664 ^a	10.450	1:04:41.420	+02:08:24.20	19.038	0.042	18.825	0.032	18.726	0.027	18.585	0.025	18.637	0.039	V16
926 ^a	9.402	1:04:33.590	+02:07:45.60	19.016	0.041	18.838	0.029	18.784	0.025	18.597	0.029	18.541	0.045	V06
11589	8.409	1:04:51.510	+02:05:33.50	19.490	0.075	19.291	0.050	19.245	0.043	18.404	0.056	18.504	0.0585	V34
13808	7.557	1:04:59.740	+02:08:43.10	19.617	0.069	19.332	0.044	19.262	0.038
13759	7.333	1:04:52.510	+02:08:04.80	19.491	0.068	19.331	0.048	19.298	0.043	18.336	0.068	18.451	0.071	V7
18905	6.766	1:05:06.310	+02:12:33.90	19.734	0.050	19.481	0.038	19.441	0.033	18.989	0.079	19.310	0.104	V24
13943 ^a	6.751	1:04:51.670	+02:10:55.00	19.407	0.091	19.217	0.064	19.134	0.054	18.990	0.041	19.101	0.046	V24
3732	6.669	1:04:40.210	+02:01:24.80	19.250	0.041	19.117	0.080	V27
5037	6.31	1:04:49.140	+02:07:20.20	20.149	0.109	19.850	0.080	19.790	0.072	19.140	0.079	18.958	0.103	V26
3722	5.818	1:04:43.830	+02:01:04.70	19.734	0.074	19.317	0.082	V26
13911	5.717	1:04:51.600	+02:10:10.50	19.911	0.081	19.688	0.055	19.638	0.050	19.481	0.064	19.401	0.085	V17
13780	5.58	1:04:56.250	+02:08:21.60	19.965	0.087	19.717	0.060	19.684	0.053	19.501	0.058	19.407	0.082	V9
4875	5.138	1:04:48.980	+02:05:37.10	19.720	0.079	19.579	0.060	19.542	0.051	19.330	0.041	19.283	0.058	V14
15696	5.012	1:04:50.930	+02:14:30.60	19.426	0.072	19.645	0.078	V13
15670	4.849	1:04:53.290	+02:13:30.60	19.456	0.061	19.312	0.101	V13
14287	4.365	1:05:01.050	+02:09:11.80	20.484	0.104	20.252	0.079	20.202	0.070	19.582	0.091	19.692	0.134	V10
13784	4.045	1:04:59.848	+01:53:10.16	19.216	0.047	19.385	0.075	V10
6084	3.872	1:04:46.550	+02:07:28.10	19.667	0.087	19.329	0.095	V10

^aDetected by Freedman et al. (2009)

^bNo OGLE ID

Table 4. J band time series photometry of Cepheids in IC 1613.

OGLE ID	J_1^a (mag)	σ_{J_1} (mag)	J_2^b (mag)	σ_{J_2} (mag)	J_3^c (mag)	σ_{J_3} (mag)
V22	15.709	0.008	15.754	0.007	16.192	0.007
11446	17.079	0.014	17.205	0.057	17.397	0.028
736
7467	18.21	0.042	17.988	0.046	18.018	0.026
13738	18.67	0.079	18.454	0.038	18.36	0.044
7664	18.813	0.079	19.153	0.07	19.13	0.07
926	19.011	0.057	19.02	0.083	19.023	0.085
11589	19.354	0.14	19.457	0.126	19.654	0.126
13808	19.661	0.127	19.581	0.128	19.613	0.109
13759	19.497	0.13	19.459	0.113	19.519	0.111
18905	19.742	0.084	19.789	0.111	19.702	0.076
13943	19.407	0.091
3732
5037	20.056	0.18	19.975	0.16	21.408	0.268
3722
13911	20.017	0.148	19.807	0.137	19.935	0.137
13780	19.957	0.175	19.849	0.137	20.12	0.146
4875	19.718	0.134	19.743	0.144	19.703	0.135
15696
15670
14287
13784
6084

^aMJD₁ = 2455811

^bMJD₂ = 2455838

^cMJD₃ = 2455868

Table 5. H band time series photometry of Cepheids in IC 1613.

OGLE ID	H_1^a (mag)	σ_{H_1} (mag)	H_2^b (mag)	σ_{H_2} (mag)	H_3^c (mag)	σ_{H_3} (mag)
V22	15.305	0.007	15.357	0.009	15.652	0.006
11446	16.622	0.034	16.82	0.019	16.803	0.027
736
7467	17.844	0.049	17.687	0.041	17.681	0.043
13738	18.257	0.043	17.975	0.046	17.839	0.049
7664	18.495	0.073	18.712	0.093	18.617	0.098
926	18.67	0.077	18.737	0.071	18.641	0.069
11589	19.085	0.104	19.179	0.115	19.252	0.13
13808	19.232	0.101	19.162	0.102	19.127	0.097
13759	19.178	0.112	19.227	0.137	19.164	0.115
18905	19.328	0.133	19.177	0.091	19.227	0.087
13943	19.058	0.09
3732
5037	19.637	0.18	19.554	0.152
3722
13911	19.702	0.154	19.474	0.123	19.474	0.123
13780	19.511	0.122	19.448	0.156	19.67	0.176
4875	19.449	0.174	19.361	0.175	19.416	0.142
15696
15670
14287	20.023	0.283	20.073	0.215	19.958	0.168
13784
6084

^aMJD₁ = 2455811

^bMJD₂ = 2455838

^cMJD₃ = 2455868

Table 6. K_S band time series photometry of Cepheids in IC 1613.

OGLE ID	K_{S1}^a (mag)	$\sigma_{K_{S1}}$ (mag)	K_{S2}^b (mag)	$\sigma_{K_{S2}}$ (mag)	K_{S3}^c (mag)	$\sigma_{K_{S3}}$ (mag)
V22	15.159	0.005	15.189	0.006	15.532	0.007
11446	16.539	0.025	16.79	0.019	16.753	0.019
736
7467	17.772	0.06	17.603	0.047	17.661	0.048
13738	18.161	0.048	17.841	0.041	17.741	0.028
7664	18.427	0.072	18.628	0.093	18.615	0.091
926	18.635	0.07	18.732	0.078	18.58	0.104
11589	19.019	0.148	19.119	0.148	19.236	0.138
13808	19.147	0.218	19.082	0.112	19.064	0.116
13759	19.16	0.182	19.207	0.167	19.149	0.162
18905	19.262	0.129	19.337	0.127	19.317	0.122
13943	18.963	0.098
3732
5037	19.556	0.168
3722
13911	19.46	0.194	19.437	0.208	19.426	0.203
13780	19.566	0.183	19.637	0.187	19.56	0.183
4875	19.459	0.171	19.497	0.179	19.405	0.161
15696
15670
14287	19.996	0.219	20.063	0.216
13784
6084

^aMJD₁ = 2455811

^bMJD₂ = 2455838

^cMJD₃ = 2455868

Band	Zero-point (mag) ^a	Standard deviation	μ (mag)
<i>J</i>	19.07 ± 0.05	0.206	24.35 ± 0.05^b
<i>H</i>	18.69 ± 0.04	0.165	24.30 ± 0.04^b
<i>K_s</i>	18.64 ± 0.05	0.168	24.33 ± 0.05^b
[3.6]	18.51 ± 0.08	0.307	24.31 ± 0.09^c
[4.5]	18.50 ± 0.07	0.235	24.26 ± 0.08^c

Table 7: Mid-infrared period–luminosity relation zero–points for unblended Cepheids in IC 1613.

^aPL relations took the form $M = a(\log P - 1.0) + b$; the a coefficients are taken from Persson et al. (2004) (*J*, *H*, *K_S*) and Scowcroft et al. (2011) ([3.6], [4.5]).

^bDistance moduli were calculated using the LMC PL relation zero–points and uncertainties from Persson et al. (2004), and assuming $\mu_{0,LMC} = 18.48$. They have not been corrected for extinction.

^cDistance moduli were calculated using the MW PL relation zero–points and uncertainties from Monson et al. (2012). They have not been corrected for extinction.

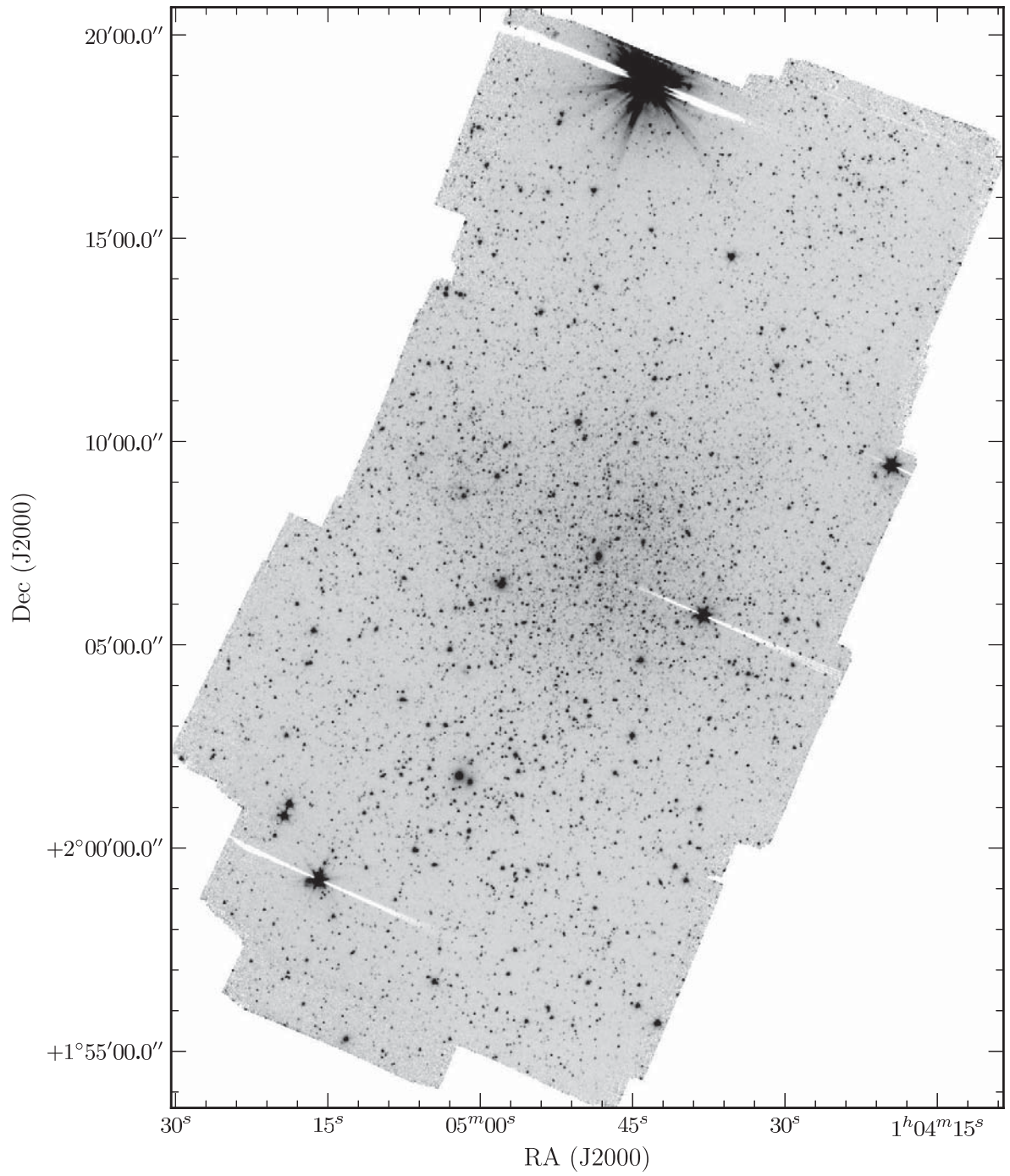


Fig. 1.— IRAC 3.6 μm science mosaic. The central third is covered in all epochs. Orientation: North is up, East is left.

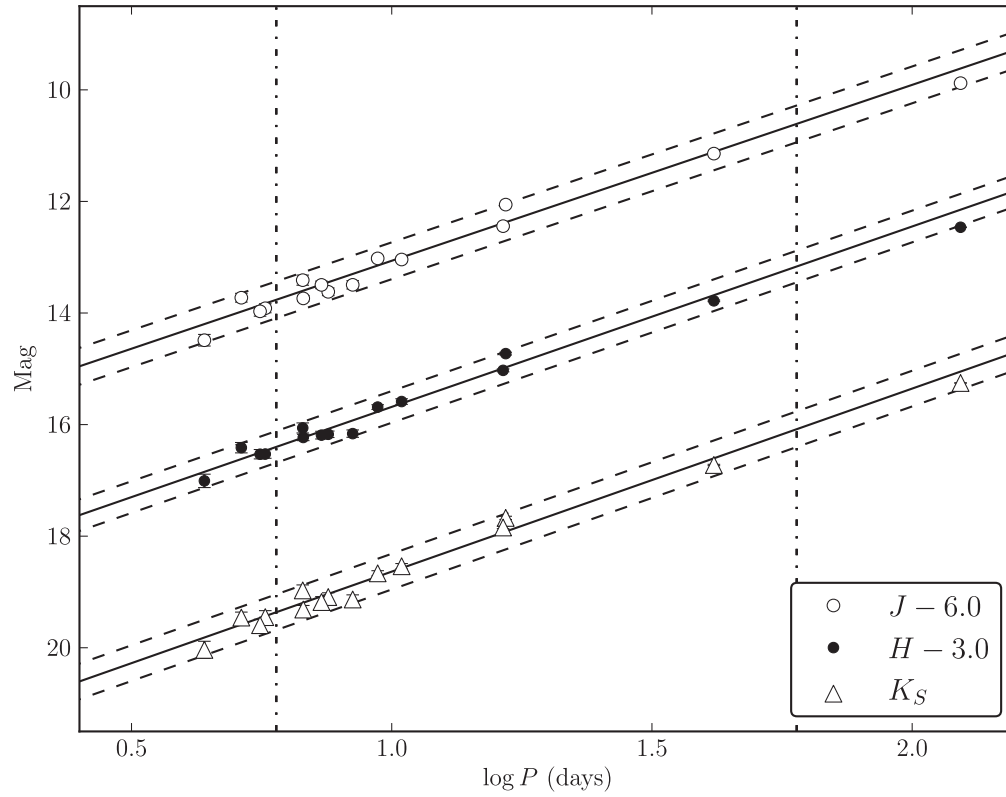


Fig. 2.— Period–luminosity relations in the J , H and K_s bands. The solid lines represent the fitted PL relations; the dashed lines show the $\pm 1\sigma$ changes in zero–point. The dot-dash lines show the period range used to fit the zero–point.

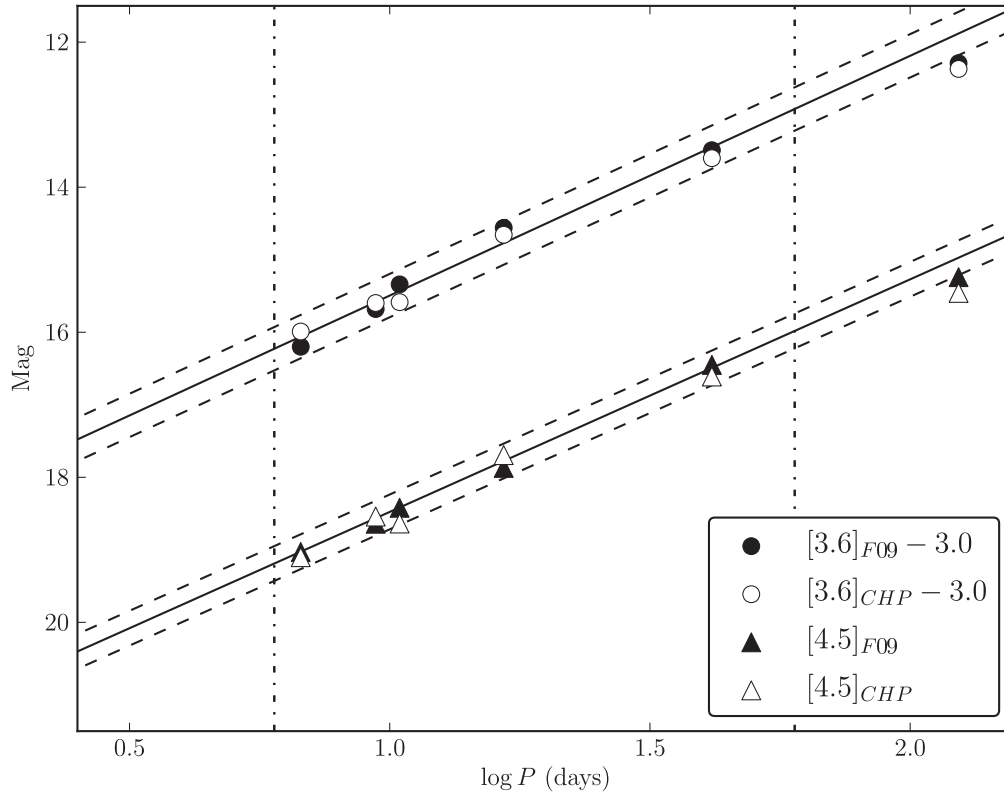


Fig. 3.— Comparison of photometry from this paper with that of Freedman et al. (2009) for the six Cepheids they detected. The magnitudes are not expected to be identical as theirs are single-epoch observations, while ours are averages over twelve phase points. The filled symbols are the cold mission data, the open symbols are the CHP data. Solid lines depict the PL fit to the CHP data (fixing the slopes to LMC values, using the sample with $6 \leq P \leq 60$ days) and are indistinguishable from the fits to the cold data. Dashed lines are $\pm 2\sigma$ around the fits.

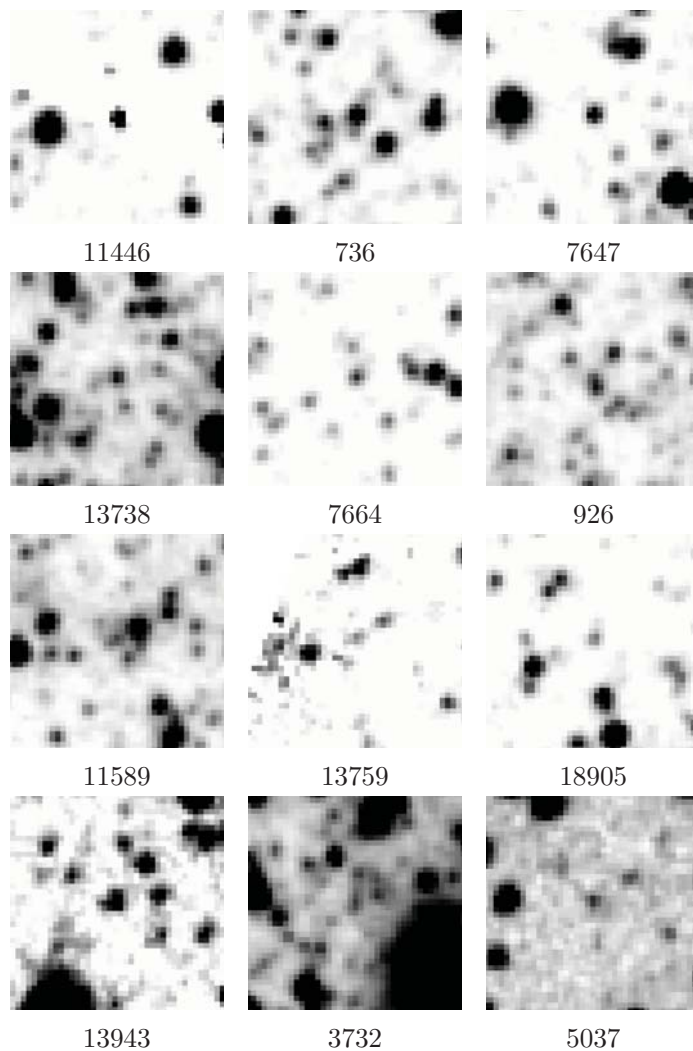


Fig. 4.— A montage of $3.6\mu\text{m}$ images of the Cepheids used in the mid-infrared PL relation. The Cepheids are the stars in the center of each image

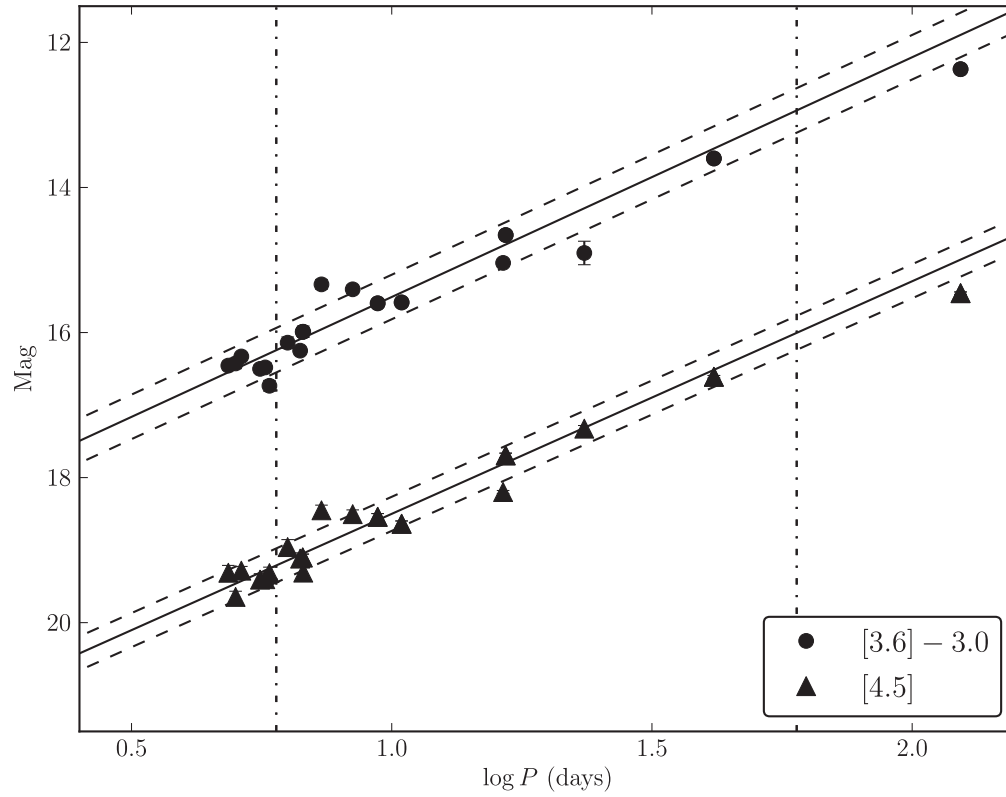


Fig. 5.— Period–luminosity relations in the 3.6 and 4.5 μm bands. The solid lines represent the fitted PL relations; the dashed lines show the $\pm 2\sigma$ changes in zero–point. The dot–dash lines show the period range used to fit the zero–point.

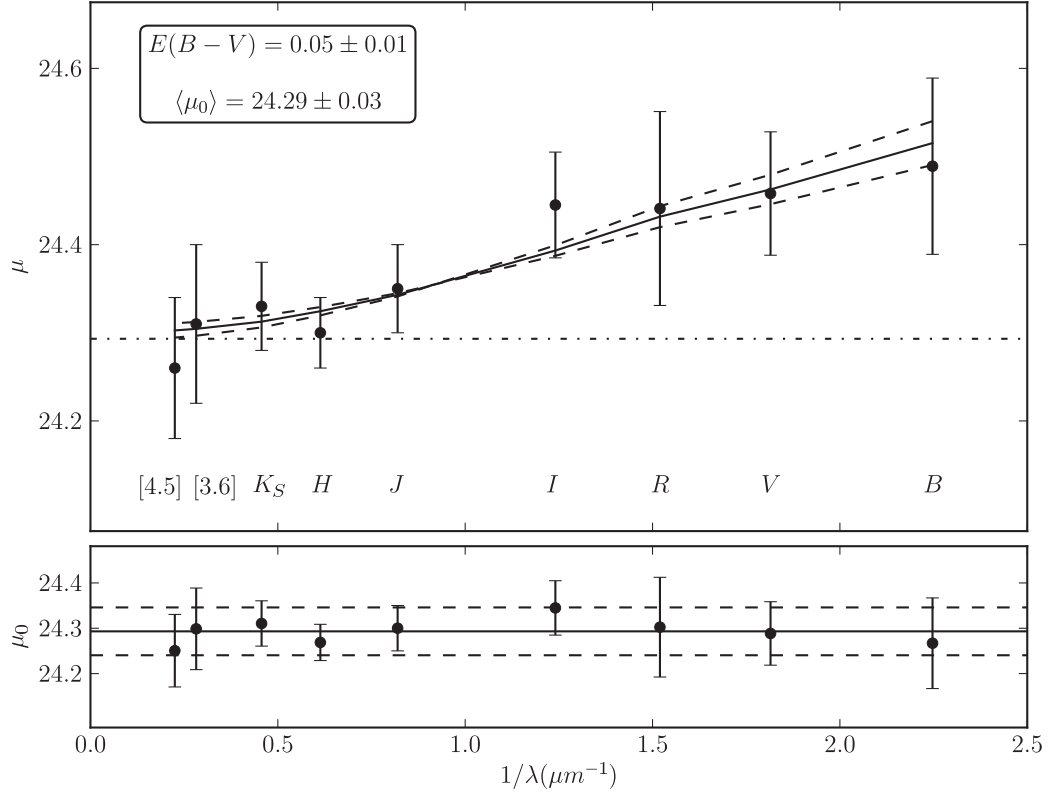


Fig. 6.— Fitting the IC 1613 distance moduli to the reddening laws of Cardelli et al. (1989) (B to K_s) and Indebetouw et al. (2005) (K_s to [4.5]). Points at B and R are taken from Freedman (1988a), V and I are from Udalski et al. (2001), J , H and K_s are FourStar data, and [3.6] and [4.5] are IRAC data. The distance moduli from B to R were refit using the LMC PL relations from Fouqué et al. (2007) and adopting $\mu_{LMC} = 18.48$. The solid line is the best-fit reddening law, the dashed lines are $\pm 1\sigma$ around the law, and the dot-dashed line is the resulting reddening-corrected distance modulus. The bottom panel shows the residuals of the extinction-corrected distance moduli around the mean value.

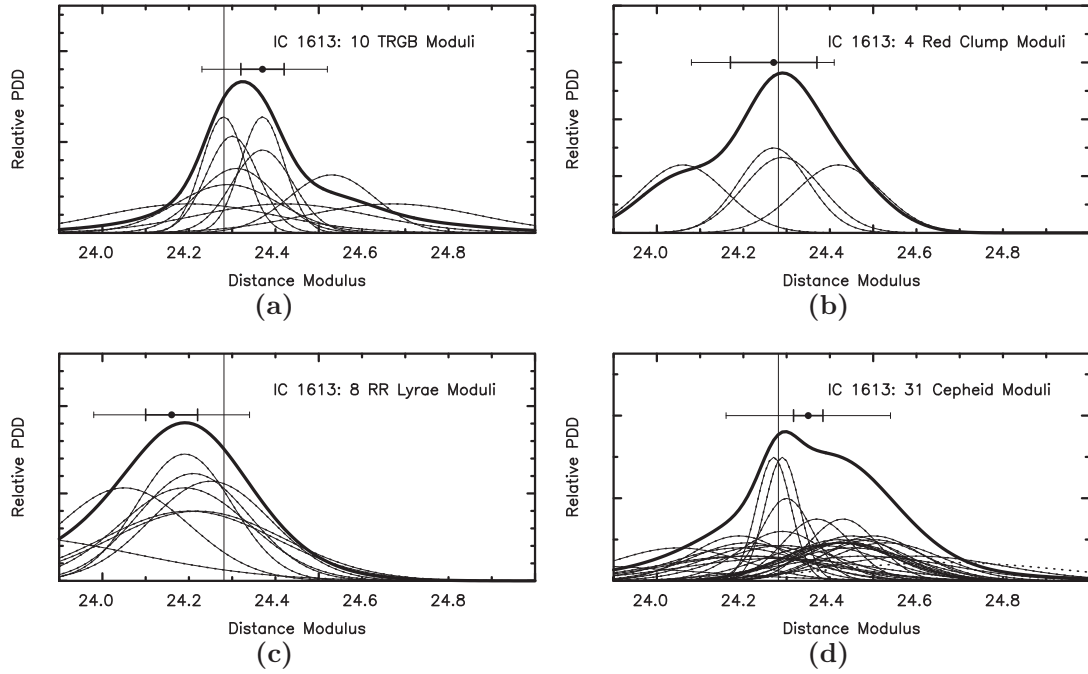


Fig. 7.— A montage of individual comparisons of distance moduli to IC 1613 as published over the years and broken down into four major methods: the TRGB (a), the Red Clump (b), the RR Lyraes (c) and the Classical Cepheids (d). Individual distance determinations are shown as unit-area gaussians. The cumulative distribution is shown as the thick solid line. The median value is shown as a solid point with error bars (see text for details). For consistent comparison, the vertical solid line marks the Cepheid distance determined in this paper.

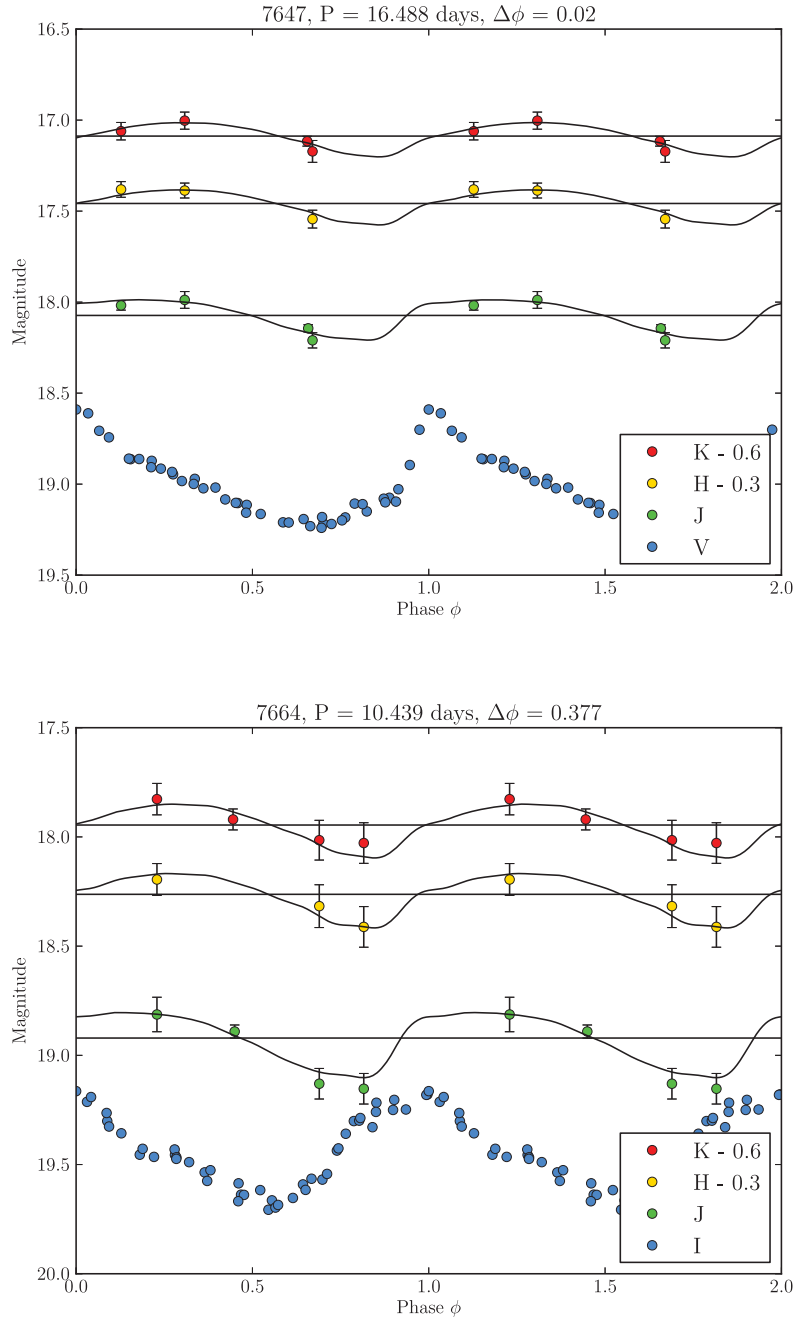


Fig. 8.— Light curves fit using the template method from Soszyński et al. (2005). A phase shift, $\Delta\phi$ was incorporated into the algorithm to account for possible period changes, and hence deviations from the predicted time of maximum light in the near-IR bands. The J and K_s light curves also contain points from Pietrzyński et al. (2006) where available. The Cepheid IDs use the OGLE numbering scheme.