

for the ATHENA X-ray Integral Field Unit

Presented by Stephen Smith NASA GSFC On behalf of the X-IFU array development team

Key contributors:

NASA GSFC: J.S Adams, S.R. Bandler, J.A. Chervenak, A.M. Datesman, M.E. Eckart, A.J. Ewin,
F.M. Finkbeiner, R. Hummatov, R.L. Kelley, C.A. Kilbourne, A.R. Miniussi, F.S. Porter, J.S. Sadleir,
K. Sakai, N. Wakeham, E.J. Wassell
SRON: H. Akamatsu, L. Gottardi, R.H. den Hartog, B.D. Jackson, J. van der Kuur.
NIST: D.A. Bennett, W.R. Doriese, G.C. Hilton, D. Swetz, J.N. Ullom.

Athena X-IFU Baseline Array Configuration



Large pixel array

- Currently assumed X-IFU baseline
 - Uniform Large Pixel Array LPA.
 - 3840 identical transition-edge sensor (TES) pixels, 0.25mm pitch.
 - $-\Delta E_{\text{FWHM}} = 2.5 \text{ eV} @ 7 \text{ keV}.$
 - Baseline changed in Nov 2016 to make use of defocusing optic (for high rate point source observations).
 - Reduces count-rate requirement to ~ few cps, enables reduced speed pixels (~x2-3) and more optimal use of available readout resources (greater engineering margin).
 - LPA-1 -> LPA-2





Transition-edge sensor arrays, 2016 state of the art



- Previously proposed "LPA-1" GSFC TES pixel.
 - Mo/Au Bilayer TES, target $T_C \sim 90$ mK, suspended on SiN (~ 1 µm).
 - Low impedance TES, $R_n = 10 \text{ m}\Omega$.
 - Composite Au/Bi electroplated absorbers (low heat-capacity and fast thermalization).
 - Thermal conductance $G_b = 200 \text{ pW/K}$, Fast ~ 400 µs decay times (originally for LPA-1, 10's cps / pixel).



3

Space Telescopes and Instrumentation 2018: Ultraviolet to Gamma Ray, SPIE, Austin, Texas, 2018

Transition-edge sensor arrays, 2016 state of the art



- GSFC devices were historically optimized for *Time Division Multiplexing* (TDM) approach under study as back-up.
 - See J. Ullom et al. 10699-60 @ 4:30pm
- However baseline readout approach uses *Frequency Division Multiplexing* (FDM).
 - See H. Akamatsu et al. 10699-58 @ 4:00pm
- In TDM TESs are DC biased (multiplexing via switching SQUIDs).
 - TES transition is independent of the muxing.
- In FDM the mux encoding via AC TES bias with different frequencies.
 - In Mo/Au TESs lead to frequency dependent variations in TES transition shape and degraded resolution performance.



Pixel optimization activities



• Technology development focused on 3 main areas:

(1) Pivot towards lower count-rate 'LPA-2'.

- x2-4 slower pixels, control of thermal conductance to the heat-sink via TES size / membrane thickness.

(2) Uniformity optimizations.

- Exploring better transition uniformity, less sensitive to environment (role of stripes, size effects).

5

(3) Reduce the impact of frequency dependent effects in the AC-TES.

Mixed arrays for transition and G_b studies

50 μm 75 μm 100 μm 120 μm 50 μm 75 μm 100 μm 120 μm





Space Telescopes and Instrumentation 2018: Ultraviolet to Gamma Ray, SPIE, Austin, Texas, 2018

shape.

DC transition studies – studying role of stripes



- LPA optimization activities has led to growing understand of geometry effects in TESs.
 - Exploring the role of stripes, historically used on larger TESs for noise and transition shape control (empirically).
 - Measurements in large arrays show undesirable variations in transition shape.
 - Presence of 'kinks' hard to predict and can impact array uniformity.

- Improving understand of how they come about / evolve See Wakeham et al., LTD-17, 2017.
- Maybe due to alignment difference between the different metal layers (stripes/stems).

Different transition shapes in different quadrants

of 32x32 array Q1 and Q3 have majority 'kinked' transition Q¹⁶ and Q4 had majority 'smooth'**1**40 nsition 120 Resistance (mΩ) $\alpha_{IV} = T/R dR/dT$ 100 $\overset{\geq}{\mathrm{08}}$ 80 60 40 'kink' 20 2 Ω 20 40 60 80 82 78 80 84 86 88 0 76 90 R/R_n Temperature (K)

Quadrant '



Quadrant 3



Space Telescopes and Instrumentation 2018: Ultraviolet to Gamma Ray, SPIE, Austin, Texas, 2018

100



DC transition studies – studying role of stripes





7

Space Telescopes and Instrumentation 2018: Ultraviolet to Gamma Ray, SPIE, Austin, Texas, 2018

Stephen Smith



First uniform kilo-pixel arrays of no-stripe pixels



- First uniform 32x32 arrays, no stripes, $50/75/100/120 \mu m$, 250 μm pitch.
- Example for 50 µm devices, show good DC transition shape and uniformity.
- Large transition phase space with no kinks, less sensitive to environment (Bfield)



Space Telescopes and Instrumentation 2018: Ultraviolet to 8 Gamma Ray, SPIE, Austin, Texas, 2018



DC transition studies – example 50 µm TES, no stripes



50 µm TES

- Good pulse shape and noise uniformity.
- DC performance is excellent, ΔE_{FWHM} consistently at 2.0 eV level.
- Smaller TES sizes. $G_b = 75 \text{ pW/K}$ at $T_c = 90 \text{ mK}$.

= slower speed $\tau = 2.2$ ms.

Very promising LPA2 DC pixel understudy for use with backup TDM readout.



Space Telescopes and Instrumentation 2018: Ultraviolet to 9 Gamma Ray, SPIE, Austin, Texas, 2018



AC Pixel optimization activities



- Large collaborative effort NASA/SRON/NIST to understand frequency effects in AC TES.
- Identified two main contributors:

1) AC, dissipative, losses (magnetic coupling to nearby metals) broadens transition (lower α). Limits access to most sensitive part of the transition used to achieve < 2.5 eV. [Sakai et al 2017].

2) **AC Josephson reactance**, periodic steps through the transition due to variation in Josephson inductance. Undesirable non-linearity and noise properties, hard to find good bias points. [Gottardi et al 2017].



Space Telescopes and Instrumentation 2018: Ultraviolet to 10 Gamma Ray, SPIE, Austin, Texas, 2018



AC loss and Josephson-effect mitigation strategies



- Comparison of many different geometries and improved theoretical understanding has led to optimization routes.
- Strategy is to explore higher resistance regime in GSFC Mo/Au TES.
 - AC loss independent of TES Z => higher Z bias point reduces impact of fixed AC loss.
 - High Z devices have small Josephson oscillations.
 - Seen for example in SRON Ti/Au very high-Z devices 200 mΩ [P. Khosropanah, 10699-57]
 - Challenge is to increase Z enough, without affecting other noise or uniformity properties.

1) 15 -> 50 m Ω/\Box bilayer sheet resistance. Thinner TES films. Now implemented and in testing.

2) Change aspect ratio (1:1 -> 1:0.25). Longer and thinner. Now implemented and in testing.



Space Telescopes and Instrumentation 2018: Ultraviolet to 11 Gamma Ray, SPIE, Austin, Texas, 2018



First high-Z devices with good DC performance





Space Telescopes and Instrumentation 2018: Ultraviolet to 12 Gamma Ray, SPIE, Austin, Texas, 2018

First results on high-Z TES under AC bias at SRON



- First spectral measurements on mixed arrays at SRON yielded AC-TES x-ray resolution comparable to DC.
- Best performance seen in 100/120 µm sizes. Smaller, slower 50 µm sizes not as good performance.
- Less structure + more access to lower bias points. => consistent access to < 2.5 eV 1-5 MHz range.
- Later presentation by H. Akamatsu will show more results and FDM testing from 1-5 MHz



Space Telescopes and Instrumentation 2018: Ultraviolet to 13 Gamma Ray, SPIE, Austin, Texas, 2018

First uniform high-Z arrays suitable for large scale FDM demonstrations



- Two 32x32 arrays screened at GSFC and sent to SRON for 40-pixel FDM testing.
- Good DC transition properties and uniformity, R vs T, pulse shape, ΔE .
- $<\Delta E_{FWHM} > = 1.95 \text{ eV}$ for 6 pixels tested.
- Larger TES size => Pixels are fast => for X-IFU may still need to slow these down.



Space Telescopes and Instrumentation 2018: Ultraviolet to 14 Gamma Ray, SPIE, Austin, Texas, 2018

First DC results from high aspect ratio 120 μm TESs





- First 120 µm high aspect ratio devices in DC testing.
- Preliminary measurements suggest no strong ΔE dependence on aspect ratio
- Added design flexibility: higher Z and lower G_{b.}
- First AC tests planned soon.

First large X-IFU array scale test parts yielded



- 90 mm diameter hexagonal chip and prototype detector array.
- 3540 sensors on 250 µm pitch.
- 960 pixels connected to bond pads.
- First DC tests planned for later this year.
- Later iteration will include coil-coupling for AC biased testing.





Space Telescopes and Instrumentation 2018: Ultraviolet to Gamma Ray, SPIE, Austin, Texas, 2018



Summary



- Improved DC and AC pixel designs:
 - Evolved original 3-stripe 140 μm TES design to smaller TESs without stripes.
 - Larger transition phase-space without 'kinks' -> Improves array uniformity.
 - Developed first high-Z Mo/Au TESs for reduced AC Loss and + Josephson effect in AC TES.
 - Improved access good transition regions.
 - Break through energy resolution < 2.5 eV results for AC TESs at both low and high frequency.
 - First high-Z uniform 32x32 arrays now delivered to SRON for 40-pixel FDM testing.
- Ongoing development activities:
 - First high aspect ratio devices in DC testing.
 - Designs offer addition parameter space for higher Z and low Gb.
 - AC testing planed for this fall.
 - -First large scale X-IFU testing planned for 1000-pixel testing later this year.

