Some New Results on Electron Transport in the Atmosphere

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Abstract. The penetration, diffusion and slowing down of electrons in a semi-infinite air medium has been studied by the Monte Carlo method. The results are applicable to the atmosphere at altitudes up to ~300 km. Most of the results pertain to monoenergetic electron beams injected into the atmosphere at a height of 300 km, either vertically downwards or with a pitch-angle distribution isotropic over the downward hemisphere. Some results were also obtained for various initial pitch angles between 0° and 90°. Information has been generated concerning the following topics: (a) the backscattering of electrons from the atmosphere, expressed in terms of backscattering coefficients, angular distributions and energy spectra of reflected electrons, for incident energies \( \langle T \rangle \) between 2 keV and 2 MeV; (b) energy deposition by electrons as a function of the altitude, down to ~80 km, for \( \langle T \rangle \) between 2 keV and 2 MeV; (c) the corresponding energy deposition by electron-produced bremsstrahlung, down to ~30 km; (d) the evolution of the electron flux spectrum as function of the atmospheric depth, for \( \langle T \rangle \) between 2 keV and 20 keV. As far as possible, the results have been expressed in a scaled form which reduces the explicit dependence on, and permits interpolation with respect to, \( \langle T \rangle \). Energy deposition results are given for incident electron beams with exponential and power-exponential spectra.

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1. INTRODUCTION

In the course of earlier Monte Carlo calculations of auroral luminosity patterns (BERGER, SELTZER and MAEDA, 1970; to be referred to as BSM) and bremsstrahlung flux spectra in the atmosphere (BERGER and SELTZER, 1972) we have obtained other electron and bremsstrahlung transport results which may be useful for the analysis and interpretation of observed atmospheric phenomena. These results are given here, and pertain to the following topics: (a) the backscattering of electrons from the atmosphere, as the result of multiple Coulomb scattering; (b) the altitude-dependence of energy deposition by electrons precipitated into the atmosphere; (c) the corresponding altitude-dependence of energy deposition by electron-produced bremsstrahlung; and (d) the modification of the energy spectrum of electron beams as they penetrate down into the atmosphere.

2. MONTE CARLO MODELS

The Monte Carlo method involves the simulation of all important physical processes by random sampling. One must sample the elastic scattering of electrons by atoms, the inelastic scattering of electrons by orbital electrons, the production of bremsstrahlung, and the Compton scattering and photoelectric absorption of the bremsstrahlung photons.

The number of elastic scatterings which an electron undergoes before it is slowed down to rest in the atmosphere is generally quite large. For example, with the cross sections used in BSM, one finds that the average number of collisions is 170 for an initial energy of 10 keV, 290 for 20 keV and 1060 for 100 keV. An alternative, less time-consuming, method has therefore been used in most of the calculations reported here, which is referred to as Monte Carlo Model A in BSM. This approach reduces the required computational effort through the combined use of random sampling and analytical multiple scattering theories. Each electron track to be sampled is
divided into a number of segments (usually not more than ~100) whose lengths are chosen so that the number of elastic scatterings per segment is large whereas the net multiple-scattering deflection and energy loss per segment are small. The angular deflections in successive segments are sampled from the distribution of Goudsmit and Saunders (1940), and the energy-losses from the distribution of Landau (1944) with the binding correction of Blunck and Leisegang (1950). These deflections and energy losses are then combined to construct the complete electron track. Monte Carlo Model A also includes the production and subsequent diffusion of secondary electrons, and the production, scattering and photo-electric absorption of secondary bremsstrahlung photons. The transport of bremsstrahlung photons is followed by conventional random sampling techniques (Fano, Spencer and Berger, 1959) involving the sampling of all successive Compton scatterings until photoelectric absorption occurs.

In order to obtain our backscattering results for initial electron energies ≤ 20 keV we have used another method, which is described in detail in BSM where it is referred to as Monte Carlo Model B. In this model all successive elastic scatterings of an electron are followed by random sampling, with use of the appropriate single-scattering cross section. The numerous inelastic collisions are lumped together and are treated in the continuous-slowing-down approximation, assuming a relation between electron energy and pathlength traveled given by stopping power theory. In this approximation, secondary knock-on electrons are not taken into account.
In a radiobiological calculation for a water medium (BERGER, 1972) a more elaborate method has been introduced which could be designated as Model C. In this model all hard inelastic collisions are sampled individually, whereas the soft inelastic collisions are treated in the continuous-slowing-down approximation with use of a "restricted stopping power" theory. Hard inelastic collisions are defined to be those in which knock-on electrons with an energy greater than a chosen threshold value, say 200 eV, are set in motion. This procedure makes allowance for energy-loss straggling and the subsequent transport of the knock-on electrons.

For each source energy and geometry of interest, a large set of electron tracks (typically 10,000) was simulated according to Model A or B or, in some cases, with both. The sampled tracks were then analyzed to obtain information about various transport phenomena including backscattering, energy deposition and electron flux spectra. Details of some of the sampling procedures and analysis are described in BERGER (1963) and BERGER and SELTZER (1968).

3. CROSS SECTIONS

In Monte Carlo Model A the cross section for the elastic scattering of electrons by atoms, needed for the evaluation of the Goudsmitt-Saunderson multiple-scattering distribution, is taken from the theory of MOTT (1929) which includes spin and relativistic effects. In Model B, applied at low energies, relativistic effects are unimportant and the Rutherford cross section can be used. In both models, the cross section at small angles is modified by a screening correction given by MOLIÈRE (1947, 1948). The electron stopping power values have been taken from the Bethe theory as formulated by ROHRlich and CARLSON (1954).

Because of the use of the Born approximation, and because of the incomplete treatment of atomic binding effects, the theoretical cross sections used in our Monte Carlo program to describe interactions between electrons and the atmospheric constituents are expected to be applicable only down to some cut-off energy, perhaps
comparable to the K-shell binding energies of the atoms involved. We estimate that the overall cross section uncertainty is still only 10% to 15% at 0.5 keV, but becomes greater at lower energies. We expect that the transport calculations of energy deposition will give reliable results under conditions such that the electrons have lost the predominant part of their initial energy, and have traveled a path-length equal to a large fraction of their initial range, when they reach an energy of 0.5 keV.

There is experimental evidence regarding the reliability of the calculations. It has been shown in BSM that the calculated spatial distribution of energy deposition in air is in good agreement with the results of a laboratory experiment by GRÜN (1957) with electron beams incident with energies as low as 5 keV. Comparisons with similar measurements of COHN and CALEDONIA (1970) in nitrogen indicate similarly good agreement down to 2 keV. In this paper it will be shown that the calculated backscattering coefficients for air are in agreement with experimental data down to 2 keV. There are at present no laboratory data which could be used to check the calculated electron flux spectra. It is possible that the spectral distributions are more sensitive to the cross section input than energy deposition distributions or backscattering coefficients. The electron flux spectra are therefore expected to be the most tentative of the results in this paper, and subject to revision in the light of better cross section information.*/

*New information on low-energy electron scattering and energy loss processes is gradually accumulating. We are beginning a new cycle of transport calculations according to Monte Carlo Model C, using data such as the semi-empirical energy loss functions of GREEN and PETERSON (1968), the inelastic electron scattering cross sections measured by OPAL, BEATY and PETERSON (1972) and calculated by OMIDVAR, KYLE and SULLIVAN (1970), and the measured elastic scattering cross sections of BROMBERG (1970) and SHYN, STOLARSKI and CARIGNAN (1972). Much work is yet needed to generate from these and similar data the comprehensive data base needed for the Monte Carlo calculations.
Electrons are assumed to be incident onto the boundary of a semi-infinite air medium with kinetic energy $T_0$ and obliquity $\theta_0$ (angle between the direction of incidence and the normal to the boundary plane). Electron transport is calculated as a function of the depth $z_m$ in the medium, i.e. the mass thickness (in g cm$^{-2}$) measured from the boundary plane. The boundary plane is treated as non-reentrant, so that electrons (or electron-produced bremsstrahlung photons), after escaping from the medium, cannot return to it. The incident electron beam current across the boundary plane is assumed to be either monodirectional (with a fixed obliquity angle $\theta_0$), or characterized by an angular distribution proportional to $\cos^2 \theta_0$.

In applications to the atmosphere, the obliquity angle $\theta_0$ (at sufficiently high magnetic latitudes) can be identified with the incident pitch angle; $z_m$ is the atmospheric depth; the cosine-law/distribution corresponds to the case of an incident electron flux isotropic over the downward hemisphere (to be denoted by the abbreviation IDH). As has been shown by REES (1964a), the relation between the atmospheric depth $z_m$ and the altitude $h$ is given by

$$ z_m = \int_0^\infty D(h) \left[ \frac{4 - 3x}{4 - 4x} \right]^{1/2} dh $$

(1)

where

$$ x = \cos^2 \lambda_m (1 + h/a), $$

(2)

$\lambda_m$ is the geomagnetic latitude, $a$ is the radius of the earth and $D(h)$ is the density of the atmosphere. In the present work we have assumed that $D(h)$ is given by the CIRA (1965) mean atmosphere for heights between 30 km and 300 km.

5. NEGLECT OF MAGNETIC MIRRORING

Electrons are deflected in the atmosphere not only as the result of multiple Coulomb scattering but also by magnetic mirroring. The combined action of both effects has been included in the calculations of WALT, McDONALD and FRANCIS (1967),
who solved the electron transport equation in the Fokker-Planck approximation, and in the Monte Carlo calculations of Wedde (1970) and McEntire (1972). The calculations reported here do not include the effects of magnetic mirroring. The results are nevertheless applicable to atmospheric problems, provided one limits the applications to altitude no greater than, say 300 km. We have therefore assumed that the atmospheric depth \( z_m = 0 \) (the boundary of the semi-infinite air medium where the characteristics of the incident electron beam are specified) corresponds to an altitude of 300 km.

This assumption has practically no effect on the interpretation of the calculated multiple Coulomb scattering results, because the mass thickness of the atmosphere above 300 km is negligibly small compared to the range of the electrons at the energies of interest. The magnetic deflection angles are generally small compared to Coulomb multiple-scattering deflections for electrons starting out at 300 km. They can be estimated assuming a magnetic dipole field and using the adiabatic invariance of the quantity \( \sin^2\theta / H \), where \( \theta \) is the pitch angle and \( H \) the magnetic field strength. Typical pitch-angle changes are given in Table 1 for electrons that start with various initial pitch angles at altitudes of 300 or 200 km and travel down to heights of 90 or 50 km. It can be seen that in most cases the magnetic deflection angles amount to only a few degrees and become significantly greater only when the initial pitch angle approaches 70°. Moreover, the fraction of the electrons with pitch angles greater than 70° is small for an incident electron flux isotropic over the downward hemisphere.

6. ELECTRON BACKSCATTERING

The backscattering of electrons from the ionosphere was observed in rocket experiments by McDiarmid, Rose and Budzinski (1961) and in the Injun III satellite experiments of O'Brien (1964). According to O'Brien, approximately 10% of the electrons reaching the atmosphere are backscattered into the magnetosphere, which
is consistent with the results of our calculations. More evidence on backscattering
has been obtained in an "electron echo" experiment by HENDRIKSON, McENTIRE and
WINKLER (1970) in which 40-keV electrons were injected into the atmosphere from
a rocket-borne accelerator at a height of 300 km at a geomagnetic latitude of
52° N, with emission pitch angles between 60° and 120°. The injected electrons
traveled back and forth between the mirror point near the rocket and the conjugate
southern mirror point, and several electron echos were detected for each emitted
pulse. The southern conjugate mirror point was almost at sea level, so that
Coulomb scattering rather than magnetic mirroring was primarily responsible for the
appearance of the echos. In an analysis of this experiment, McENTIRE (1972) has
made an elaborate Monte Carlo calculation which takes into account not only multiple
Coulomb scattering in the atmosphere (by a method similar to our Monte Carlo Model B)
but also the motion of the electrons under the action of the geomagnetic field.
These calculations are reported to account quite well for many features of the
observed electron echos.

Our calculations are aimed not so much toward the explanation of a particular
experiment as toward the production of background information of general interest.
First of all, we shall examine the validity of the multiple-scattering calculations
through the comparison with laboratory experiments. Such a comparison is possible
in respect to the number albedo $R_N$, i.e. the fraction of the incident electrons that
are backscattered from a semi-infinite medium. Several authors have measured the
number albedo for series of target materials with atomic numbers bracketing that of
air ($Z_{av} = 7.4$). We have interpolated graphically with respect to atomic number and
have thus obtained the experimental $R_N$-values shown in Fig. 1. Also shown in that
figure is the calculated curve of $R_N$ vs. the energy of incidence, $T_0$, for $T_0$
between 2 keV and 2 MeV. The calculated results have an estimated statistical uncertainty
(relative standard deviation) of 3%; including systematic error, the uncertainty
may be as large as 5%. The uncertainty of the experimental results is indicated in some cases, and can also be judged from the dispersion among the backscattering coefficients from different experiments. Within the combined limits of experimental and calculational error there seems to be good agreement between predicted and observed number-albedo values. The theoretical number albedo shows a systematic increase as $T^\circ_o$ is decreased. So do the experimental values, except that some experiments indicate a leveling off at an energy $T^\circ_o \sim 2$ keV.*

*Calculated backscattering coefficients for air have been given previously by MAEDA (1965). His results, for $T^\circ_o = 20$ keV, are in good agreement with the present results. At lower energies his backscattering coefficients are smaller and have a different energy dependence, showing a decrease rather than an increase with decreasing $T^\circ_o$. These discrepancies may be related to difficulties with Maeda's electron flux calculations discussed below in Section 9.

The calculation also provided results not available from experiments, namely, the number albedo $R_N$ and the energy albedo $R_E$ (fraction of incident energy back-scattered) as functions of the obliquity of the incident electrons. In Fig. 1, $R_E$ is shown together with $R_N$ for incident energies between 2 keV and 2000 keV, for perpendicular incidence ($\theta_o = 0^\circ$) and for a cosine-law source (IDH case in the atmosphere). Figs. 2a and 2b show $R_N$ and $R_E$ as functions of the incident obliquity angle $\theta_o$, for incident energies of 10, 100 and 1000 keV.

We have insufficient data to give a distribution of backscattered electrons differential in energy and direction. We can only show the angular distribution regardless of energy, and the energy spectrum regardless of the direction of emergence. The dependence of these single-variable distributions on the incident energy $T^\circ_o$ is greatly reduced if one normalizes them through division by the backscattering coefficient $R_N$. Normalized angular distributions $\vec{W}(\theta)$ are shown in Fig. 3 for $\theta_o = 0^\circ$ and for the IDH case, in the form of universal curves which are
applicable for all $T_o$-values between 2 keV and 2 MeV. Normalized cumulative energy spectra are shown in Figs. 4a and 4b for $\theta_o = 0^\circ$ and the IDH case. Again the results are insensitive to the value of $T_o$, particularly for $T_o < 100$ keV. The dependence of the normalized cumulative energy spectra of backscattered electrons on the incident obliquity angle $\theta_o$ is shown in Figs. 5a, b and c for incident energies of 10, 100 and 1000 keV.

7. ENERGY DEPOSITION BY ELECTRONS

The basic quantity of interest is the energy deposition function $A(z_m)$ which is defined as the energy deposited per unit mass thickness at a depth $z_m$ from the boundary of a semi-infinite medium. We shall assume that $A(z_m)$ is expressed in units of eV/(g cm$^{-2}$), and the depth $z_m$ in g cm$^{-2}$. The amount of energy deposited per cm$^3$ in the atmosphere at an altitude $h$, normalized to an incident beam of 1 electron cm$^{-2}$, is given by $D(h)A[h(z_m)]$, where $D(h)$ is the atmospheric density, in g cm$^{-3}$, and where $z_m$ and $h$ are related by Eq(1). We shall refer to the product $D(h)A[h(z_m)]$ as the altitude profile of energy deposition.*

*The notation is the same as in BSM where the symbol $\rho$ was used to indicate a radial spatial variable and $D$ to indicate the density. The energy deposition function $A$ was called penetration function in BSM.

The energy deposition function $A(z_m)$ has previously been computed in BSM with Monte Carlo Model B for electrons with incident energies between 2 keV and 20 keV. Using Monte Carlo Model A, we have now extended these calculations up to 2000 keV for the case of perpendicular incidence, and to 500 keV for the cosine-law (IDH) case. At 20 keV, the old and the new calculations are in close agreement. The accuracy of the calculated energy deposition functions is estimated to be 3% or better. As shown in BSM, there is good agreement between the calculations and experimental laboratory data in air obtained by GRUN (1957) at...
energies up to 54 keV. Our results are therefore also consistent with the empirical energy deposition function of REES (1963) at energies up to 300 keV, obtained by interpolation and extrapolation of Grüneis's data. Further experimental confirmation is available from the measurements of McLAUGHLIN and HUSSMANN (1969) at 100 and 400 keV, and the measurements of ROSENSTEIN, EISEN, and SILVERMAN (1971) at 2000 keV, in a medium (polystyrene) with an average atomic number close to that of air.

In order to minimize the dependence of the energy deposition function on the initial electron energy $T_o$, it is useful to tabulate it in scaled form, by expressing depths as fractions of the initial electron range. The simplest scaling parameter that can be used is the mean range $r_o$, calculated by integrating

$$
\text{mean range } r_o \text{ is numerically equal to the pathlength which the electron would travel if it lost energy at every point along its trajectory at a rate equal to the mean loss given by stopping-power theory (see, e.g. BERGER and SELTZER, 1964).}
$$

The mean range $r$ (sometimes also denoted as c.s.d.a. range because it is calculated in the continuous-slowing down approximation) is numerically equal to the pathlength which the electron would travel if it lost energy at every point along its trajectory at a rate equal to the mean loss given by stopping-power theory (see, e.g. BERGER and SELTZER, 1964).

For low values of $T_o$, say below 5 keV, the mean range $r_o$ is not very well defined because of lack of adequate information about low-energy cross sections. In BSM this difficulty was circumvented by using as a scaling parameter the so-called practical range $r_p$, which is better defined at low energies but must be obtained from a complete transport calculation. In the present work, for $T_o \geq 20$ keV, we have reverted to the use of the mean range $r_o$.

The scaled energy deposition function will be found tabulated in Table 2a for the case of perpendicular incidence, and in Table 2b for the IDH case. Each table
contains a high-energy part, consisting of the dimensionless quantity \( \frac{r_o}{T_o} A(z_m) \) as a function of \( z_m/r_o \), for \( T_o \geq 20 \text{ keV} \); and a low-energy part, excerpted from BSM, consisting of the dimensionless quantity \( \frac{r_p}{T_o} A(z_m) \) as a function of \( z_m/r_p \), for \( T_o \leq 20 \text{ keV} \). Together, the low and high-energy parts span the energy region from 2 keV to 2000 keV for perpendicular incidence, and from 2 keV to 500 keV for the IDH case. At the matching energy \( (T_o = 20 \text{ keV}) \) the two different scaled representations give the same values for \( A(z_m) \). The values of the ranges \( r_o \) and \( r_p \), needed to interpret the data in Tables 2a and 2b, are given in Table 3.
By interpolation with respect to $T_0$ one can readily obtain energy deposition functions and altitude profiles of energy deposition for any incident electron spectrum of interest, for example, exponential and power-exponential spectra of the form $T_0^\gamma \exp(-T_0/\alpha)$. The data in Tables 2a, 2b and 3 are given only for incident energies $T_0$ greater than 2 keV. However, the spectra of interest may contain electrons with lower energies that make a significant contribution to the energy deposition, particularly at very high altitudes. It is possible to include the contribution from these electrons by extrapolating the scaled energy deposition function, $(r_p/T_0) A(z_m)$, and the practical range, $r_p$, to lower energies. A simple approximation is to assume that the scaled energy deposition function retains the shape it has at 2 keV; exploratory Monte Carlo calculations indicate that the error incurred by this assumption is probably smaller than 20% even for $T_0$-values as low as 0.2 keV. Extrapolation of $r_p$ for the IDE case leads to estimated values of $4.7 \times 10^{-6}$, $1.8 \times 10^{-6}$ and $6.6 \times 10^{-7}$ g cm$^{-2}$ at 1, 0.5 and 0.2 keV, respectively.

When calculating the results given below, we have assumed that the exponential or power-exponential spectra extend down to $T_0 = 0$. The contributions of electrons with initial energies less than 2 keV to the energy deposition are shown in Table 4 for various atmospheric depths. In Fig. 6 we give the altitude profile of energy deposition for a purely exponential spectrum ($\gamma = 0$), assuming an incident flux isotropic over the downward hemisphere and e-folding energies $\alpha$ between 5 keV and 200 keV. We estimate that the uncertainty of these results due to the low-energy extrapolation is $\lesssim 20\%$ at 300 km, $\lesssim 10\%$ at 200 km, and $\lesssim 5\%$ at 150 km.
In Table 5 we give data for the case of a power-exponential spectrum $T_0^{-\gamma} \exp(-T/\alpha)$: the height $h$ at which the altitude profile peaks and the corresponding peak value $D(h)A(z_m(h))$, for various $\alpha-\gamma$ combinations. All of the above results are evaluated assuming in Eq.(1) a geomagnetic latitude of $90^\circ$. The dependence of the altitude profile on the geomagnetic latitude is slight, and is indicated in Table 6 for the example of a power-exponential spectrum with $\gamma = 0.5$ and $\alpha = 10$ keV.

8. ENERGY DEPOSITION BY ELECTRON-PRODUCED BREMSSTRAHLUNG

Electron beams deposit their energy in the atmosphere not only directly but also indirectly via bremsstrahlung. The amount of bremsstrahlung energy deposited is small but significant because the photons can penetrate down to altitudes of 30 km or lower, which the primarily electrons cannot reach.

The energy deposition function for bremsstrahlung, $A_{br}(z_m)$ has been obtained by evaluating the expression

$$A_{br}(z_m) = \int \overline{\psi}_o(k, z_m, T_0) k_{\mu_{en}}(k) dk$$

where $k$ is the bremsstrahlung photon energy, $\overline{\psi}_o(k, z_m, T_0)$ the bremsstrahlung flux spectrum at depth $z_m$ due to incident electrons of energy $T_0$, and $\mu_{en}(k)$ the photon energy absorption coefficient for air (HUBBELL, 1969). The bremsstrahlung flux spectra were taken from recent calculations (BERGER and SELTZER, 1972) for
the case of broad beams of electrons entering a semi-infinite medium with a cosine-law angular distribution (corresponding to the IDH case and wide-area precipitation into the atmosphere). In order to reduce the explicit dependence on $T_0$, we have found it useful to tabulate the quantity $A_{br}(z_m)/T_0$ as a function of the variable $z_m/T_0$. Such data are given in Table 7 and are sufficient to determine, by interpolation, the energy deposition by bremsstrahlung at atmospheric depths up to $20 \text{ g cm}^{-2}$, for incident electron energies between 20 keV and 2000 keV. By such an interpolation we have obtained altitude profiles of energy deposition by bremsstrahlung resulting from the uniform wide-area precipitation into the atmosphere of electron beams with exponential energy spectra. These results, for e-folding energies between 5 keV and 200 keV, are shown in Fig. 7.

9. ELECTRON FLUX SPECTRA

As an electron beam penetrates down into the atmosphere, the spectrum of primary electrons is shifted towards lower energies, and a buildup of low-energy secondary electrons takes place. Simple estimates of the change of the primary spectrum have been made by various authors, e.g. REES (1964b) and STOLARSKI (1968) who implicitly assumed that the spectrum at any depth $z_m$ is monoenergetic and concentrated at an energy

$$T(z_m) = T_0 - \int_{0}^{z_m} A(z_m')dz_m' ,$$

(4)

where $A(z_m)$ is the energy deposition function. This is equivalent to assuming that the electrons move along straight lines and lose energy at a rate per unit pathlength equal to $A(z_m)$.

The spectrum actually is not monoenergetic but rather broad because of the occurrence of two types of fluctuations.
First, the electrons do not really travel in straight lines, because of multiple-scattering angular deflections, so that the pathlength traveled down to depth $z_m$ is a stochastic quantity. Second, even for a given pathlength, the actual energy loss fluctuates around the mean given by the continuous-slowing-down approximation. Both types of fluctuations—pathlength and energy-loss straggling—have been taken into account in the present calculations which also include the buildup of secondary electrons.

We have calculated the spectrum only down to a cut-off energy $\Delta$, which was chosen to be 5% of the initial energy $T_0$, and have also computed the average.

*For the lowest source energy treated, 2 keV, the cut-off energy $\Delta$ was 0.1 $T_0 = 200$ eV.*

The number of electrons per unit depth, $n_{\Delta}(z_m)$, that reach an energy $\Delta$. The relation between the flux spectrum $F(T,z_m)$ and the energy deposition function $A(z_m)$ is given by

$$A(z_m) = \Delta n_{\Delta}(z_m) + \int_\Delta^T F(T,z_m) L_\Delta(T) dT,$$

where $F(T,z_m)$, in MeV$^{-1}$, is the spectrum of primary plus secondary electrons, and

**The electron flux differential in energy and direction is defined as the number of electrons per unit energy and solid angle that cross a unit area perpendicular to the direction of motion. The spectrum $F(T,z_m)$ is obtained from the doubly differential flux distribution by integrating over all directions of motion, and can thus be interpreted as the distribution with respect to energy of electrons crossing a small spherical probe at depth $z_m$.**

where $L_\Delta(T)$ is the restricted stopping power. $L_\Delta(T)$ includes energy losses from electron interactions resulting in atomic excitations or in ionizing events in which secondary electrons with energy smaller than $\Delta$ are released; it is less than or equal to the ordinary stopping power and can be computed with the use of Eqs(22-23) of BERGER and SELTZER (1964).
Comparison with a Previous Calculation. A Monte Carlo calculation of electron flux spectra, somewhat similar in scope to the present work, was done by MAEDA (1965) whose results were fitted to analytical expressions by MAEDA and AIKIN (1968) and used by REES (1969) for the analysis of satellite and rocket measurements of auroral electrons. Recently, SHEMANSKY, DONAHUE and ZIFF (1972) have studied the relation between the emission lines from excited nitrogen molecules and the energy flux spectrum of the auroral electron responsible for the excitation. Their analysis led them to the conclusion that Maeda's spectra are deficient in low-energy degraded primary electrons, and that the rate at which electrons from these spectra would ionize the atmosphere is up to \( \frac{1}{5} \) times smaller than the integrated ionization rate inherent in the incident electron flux. The present result supports the conclusions of Shemansky et al.

For a detailed comparison we have chosen the case of 20-keV electrons incident perpendicularly. Fig. 8 shows the electron spectrum \( F(T,z_m) \) calculated with our Monte Carlo Model A and the corresponding quantity, denoted as "differential energy spectrum" given earlier by MAEDA (1965). In Fig. 9 we show the distribution function \( n_{\Delta}(z_m) \) also obtained in our Monte Carlo calculation. When \( F(T,z_m) \) and \( n_{\Delta}(z_m) \) are inserted into the right-hand side of Eq(5), the energy deposition thus obtained is in agreement to within 5% or better with the value of the energy deposition function \( A(z_m) \) from Table 2a.* This shows that our energy deposition and flux spectrum calculations are indeed consistent in the sense that energy is conserved.

\* In the case under discussion, the contribution to energy deposition from electrons with energies \( \leq \Delta \) (not treated explicitly in the Monte Carlo calculation) is relatively minor, and can be approximated by the formula

\[
[\Delta \cdot n_{\Delta}(z_m)/A(z_m)] = 0.064 + 0.122 \left( \frac{z_m}{r_o} \right).
\]
It can be seen from Fig. 8 that Maeda's "differential spectrum" is lower than 
\(F(T, z_m)\) at all three depths shown, and that the discrepancy becomes more marked 
* at the lower spectral energies. When a weighted integral is taken over "differential 
spectrum" the resulting energy deposition values are smaller than \(A(z_m)\) from Table 2a 
by factors ranging from 2.4 to 5.5. This is consistent with the energy (or ionization)-content discrepancy 
noted by Shemansky et al for an incident electron spectrum with a \(T^{-0.25}\) power law.

The distribution \(n_A(z_m)\) is not available from Maeda's work. We have therefore 
* omitted the first term on the right-hand side of Eq(5) and have compensated for 
this by using as weighting factor the total stopping power \(L(T)\) rather than the 
restricted stopping power \(L_{\alpha}(T)\). The error incurred thereby is estimated to be 
small compared to the discrepancies by factors up to 5 which are being investigated.

We have attempted, but failed, to account completely for this discrepancy, 
but were at least able to reduce it considerably, by analyzing what quantity was 
actually computed in Maeda's work. The following points were considered:

(1) The Monte Carlo program for simulating electron tracks used by Maeda 
was based on the procedures of BERGER (1963) and was thus an ancestor of the Monte 
Carlo Model A used in the present work. The early version lacked certain refine- 
ments but was not essentially different from the current version.

(2) The early version of the program included only primary electrons. How- 
ever, it can be seen from Fig. 8 that for the 20-keV source the contribution of 
secondary electrons becomes significant only at spectral energies \(\lesssim 4\) keV.

(3) The early version of the program provided as output not the electron 
flux spectrum in a semi-infinite medium, but the electron current ** transmitted

**In the Monte Carlo calculation the transmitted current is estimated by counting 
the electron tracks that cross a unit area of the exit surface of the slab tar- 
get. Once the electrons have left the slab they are not allowed to return to it, so 
that the exit surface is crossed at most once. The flux in a semi-infinite medium is estimating by computing the lengths of electron tracks in many thin 
layers of the medium parallel to the boundary plane. Each layer of the medium 
is allowed to be crossed repeatedly by the same electron, in the direction away 
from or toward the boundary plane.

through a slab target. This current, denoted by Maeda as "differential energy 
spectrum" differs both conceptually and numerically from the electron flux 
spectrum as defined above.
In order to determine whether the discrepancy could be accounted for by the substitution of current for flux in Maeda's work, we have made model studies in which both quantities were calculated from the same set of sampled electron Monte Carlo histories. A typical set of results is shown in Fig. 10 which gives the total flux (denoted in this instance as $F_{4\pi}(T,z_m)$), the contribution to the flux from electrons moving in the direction of increasing depth, $F_{2\pi}(T,z_m)$, and the current of electrons transmitted through a slab target of thickness $z_m$, $J(T,z_m)$. Of these three distributions, $J(T,z_m)$ is closest to Maeda's "differential energy spectrum", but is still higher by a factor of 2 or more.

Some New Results. We would like to repeat that the spectra given here are preliminary, because they are calculated with oversimplified cross sections and by applying Monte Carlo Model A at lower energies than is, strictly speaking, permissible. Furthermore it would be desirable to extend the spectra below the cut-off $\Delta$ of the present work ($\gamma$ or $T_0$). In any case, we have verified that the energy content of the spectra is correct, in the sense that the relation between energy deposition and integrated spectra according to Eq(5) is satisfied (to within 5% at 20 keV, 5-10% at 10 and 5 keV, and 10-20% at 2 keV).*

*The decreasing accuracy at low source energies results from numerical approximations of the Monte Carlo scheme. At 2 keV we have found it necessary to use the total rather than the restricted stopping power in Eq(5) to get the desired energy balance, corresponding to the fact that the Monte Carlo treatment at extremely low energies approaches the continuous-slowing-down approximation.

The results to be shown indicate that the dependence of the electron flux spectra on the initial energy $T_0$ and the depth $z_m$ can be simplified considerably by certain scaling procedures. The dependence on $T_0$ can be minimized by expressing the spectra as functions of the ratio $T/T_0$, where $T$ is the spectral energy. The dependence on the depth $z_m$ can be reduced by plotting instead of $F(T,z_m)$ the
dimensionless quantity $\left[ \frac{T_o}{r_p A(z_m)} \right] F(T_o z_m)$ vs $T/T_o$, where $r_p$ is the practical range. The results in the following figures are normalized to one incident electron.

The residual dependence of the scaled flux spectra on $T_o$ is illustrated in Fig. 11, for $T_o$ between 2 and 20 keV, and can be seen to be rather small. The flux spectra shown in the remaining figures all pertain to electrons incident with energy $T_o = 10$ keV. The change of the spectral shape with increasing depth is illustrated by Fig. 12. As expected, there is a spectral peak at or near $T_o$, due to primary electrons which have lost only a small part of their energy. This peak disappears with increasing depth. There is also a low-energy component in the spectra, contributed (a) by primary electrons that have lost most of their energy in the course of traveling deep into the medium and were eventually turned around; (b) by low-energy secondary electrons resulting from inelastic collisions. Fig. 13 separates the flux spectra into the contributions from electrons traveling in the direction of increasing depth (DOWN) or decreasing depth (UP). At shallow depths one can distinguish three spectral regions. Near $T_o$, downward-directed primary electrons predominate; at somewhat lower energies, upward-directed electrons predominate, and at low energies the number of upward and downward directed electrons is approximately equal, indicating that the low-energy flux component is more or less isotropic. At great depths the downward-directed flux exceeds the upward-directed flux, except at very low energies where the two components remain approximately equal. Finally, Fig. 14 shows the dependence of the flux spectra on the direction of incidence, $\theta_o$. This dependence is more pronounced for spectral energies near $T_o$ than at low energies, because the low-energy electrons have undergone so many interactions that they have "forgotten" their initial direction.
Acknowledgement. The authors are indebted to Dr. M. Rees and Dr. L. V. Spencer for many helpful comments.


BERGER M.J. 1970 Z. Naturf. 12a, 89.


BERGER M.J. 1970 Z. Naturf. 2a, 133.

BERGER M.J. 1970 Z. Naturf. 3a, 78.


Reference is also made to the following unpublished material:

MCENTIRE R.W. 1972 The electron echo experiment - A comparison between the observed and predicted trajectories for electrons artificially injected into the magnetosphere. University of Minnesota School of Physics and Astronomy, Cosmic Physics Tech. Rep. No. 159. Also, Ph.D. Dissertation, Univ. of Minnesota.
Table 1. Change of pitch angle, $\Delta \theta$, due to the action of the geomagnetic field, for electrons with initial pitch angle $\theta_0$ traveling from height $h_0$ to height $h_1$. $P(\theta_0)$ is the percentage of electrons in a cosine-law distribution (IDH case) that have initial pitch-angles greater than $\theta_0$.  

<table>
<thead>
<tr>
<th>$\theta_0$ (deg)</th>
<th>$h_0 = 300$ km</th>
<th>$h_0 = 200$ km</th>
<th>$h_0 = 300$ km</th>
<th>$h_0 = 200$ km</th>
<th>$P(\theta_0)$ (%)</th>
</tr>
</thead>
<tbody>
<tr>
<td>0</td>
<td>0.0</td>
<td>0.0</td>
<td>0.0</td>
<td>0.0</td>
<td>100.0</td>
</tr>
<tr>
<td>10</td>
<td>0.6</td>
<td>0.5</td>
<td>0.5</td>
<td>0.3</td>
<td>97.0</td>
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<td>1.0</td>
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<td>1.6</td>
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<td>75.0</td>
</tr>
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<td>2.4</td>
<td>1.3</td>
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<td>10.4</td>
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<td>12.9</td>
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<td>5.3</td>
<td>9.5</td>
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<tr>
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<td>--</td>
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<td>7.2</td>
<td>6.7</td>
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Table 2a. Scaled energy deposition function, for the case of a perpendicularly incident electron flux.
<table>
<thead>
<tr>
<th>$z_m/r_p$</th>
<th>$T_o$, keV</th>
<th>$(r_p/T_o) A(z_m)$</th>
<th>$z_m/r_o$</th>
<th>$T_o$, keV</th>
<th>$(r_o/T_o) A(z_m)$</th>
</tr>
</thead>
<tbody>
<tr>
<td></td>
<td>2</td>
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<td>1.27</td>
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<td>1.25</td>
<td>1.24</td>
<td>1.25</td>
<td>1.56</td>
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<tr>
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<td>1.21</td>
<td>1.21</td>
<td>1.21</td>
<td>1.50</td>
</tr>
<tr>
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<td>1.13</td>
<td>1.14</td>
<td>1.15</td>
<td>1.38</td>
</tr>
<tr>
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<td>1.03</td>
<td>1.05</td>
<td>1.07</td>
<td>1.17</td>
</tr>
<tr>
<td>0.5</td>
<td>0.89</td>
<td>0.91</td>
<td>0.93</td>
<td>0.94</td>
<td>0.90</td>
</tr>
<tr>
<td>0.6</td>
<td>0.72</td>
<td>0.76</td>
<td>0.78</td>
<td>0.77</td>
<td>0.60</td>
</tr>
<tr>
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<td>0.57</td>
<td>0.60</td>
<td>0.57</td>
<td>0.32</td>
</tr>
<tr>
<td>0.8</td>
<td>0.37</td>
<td>0.40</td>
<td>0.40</td>
<td>0.39</td>
<td>0.13</td>
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<tr>
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<td>0.23</td>
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</tr>
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<td>0.13</td>
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<td>0.11</td>
<td>0.12</td>
<td>0.01</td>
</tr>
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</table>

Table 2b. Scaled energy deposition function, for the case of an incident electron flux isotropic over the downward hemisphere.
Table 3. Mean range, $r_o$, and practical range, $r_p$, for electrons in air.

<table>
<thead>
<tr>
<th>$T_o$ (keV)</th>
<th>$r_o$ (g cm$^{-2}$)</th>
<th>$r_p$ (g cm$^{-2}$)</th>
<th>$r_p$ (g cm$^{-2}$)</th>
</tr>
</thead>
<tbody>
<tr>
<td>2</td>
<td>1.86(-5)</td>
<td>1.52(-5)</td>
<td>1.40(-5)</td>
</tr>
<tr>
<td>5</td>
<td>8.73(-5)</td>
<td>7.19(-5)</td>
<td>6.65(-5)</td>
</tr>
<tr>
<td>10</td>
<td>2.91(-4)</td>
<td>2.48(-4)</td>
<td>2.30(-4)</td>
</tr>
<tr>
<td>20</td>
<td>9.82(-4)</td>
<td>8.45(-4)</td>
<td>7.83(-4)</td>
</tr>
<tr>
<td>50</td>
<td>4.92(-3)</td>
<td></td>
<td></td>
</tr>
<tr>
<td>100</td>
<td>1.63(-2)</td>
<td></td>
<td></td>
</tr>
<tr>
<td>200</td>
<td>5.09(-2)</td>
<td></td>
<td></td>
</tr>
<tr>
<td>500</td>
<td>2.00(-1)</td>
<td></td>
<td></td>
</tr>
<tr>
<td>1000</td>
<td>4.91(-1)</td>
<td></td>
<td></td>
</tr>
<tr>
<td>2000</td>
<td>1.08</td>
<td></td>
<td></td>
</tr>
</tbody>
</table>

*a/* For perpendicular incidence

*b/* For the IDH case
Table 4. Estimated contribution of electrons with initial energies $T_o \leq 2$ keV to the energy deposition at various atmospheric depths. The incident electron flux is assumed to be isotropic over the downward hemisphere with an exponential energy spectrum proportional to $\exp\left(-T_o/a\right)$.

<table>
<thead>
<tr>
<th>$z_m$ (g cm$^{-2}$)</th>
<th>$h$ (km)</th>
<th>Percent Contribution</th>
</tr>
</thead>
<tbody>
<tr>
<td></td>
<td></td>
<td>$\alpha = 5$ keV</td>
</tr>
<tr>
<td>0</td>
<td>300</td>
<td>63</td>
</tr>
<tr>
<td>$10^{-6}$</td>
<td>194.6</td>
<td>49</td>
</tr>
<tr>
<td>$2 \times 10^{-6}$</td>
<td>173.3</td>
<td>40</td>
</tr>
<tr>
<td>$5 \times 10^{-6}$</td>
<td>149.8</td>
<td>21</td>
</tr>
<tr>
<td>$10^{-5}$</td>
<td>135.7</td>
<td>7</td>
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</table>
Table 5. Height $\hat{h}$ at which the altitude profile of energy deposition peaks, and the corresponding value $D(\hat{h})A(\hat{h})$ of the altitude profile, for the case of an incident electron flux isotropic over the downward hemisphere with a power-exponential energy spectrum proportional to $T_0^\gamma \exp(-T_0/a)$. Assumed geomagnetic latitude is $\lambda_m = 90^\circ$.

<table>
<thead>
<tr>
<th>$\alpha$ (keV)</th>
<th>$\gamma = 0$</th>
<th>2</th>
<th>5</th>
<th>10</th>
<th>20 keV</th>
</tr>
</thead>
<tbody>
<tr>
<td>5</td>
<td>108</td>
<td>106</td>
<td>104</td>
<td>100</td>
<td>96</td>
</tr>
<tr>
<td>10</td>
<td>100</td>
<td>100</td>
<td>98</td>
<td>96</td>
<td>93</td>
</tr>
<tr>
<td>20</td>
<td>93</td>
<td>93</td>
<td>92</td>
<td>91</td>
<td>90</td>
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</tr>
<tr>
<td>200</td>
<td>73</td>
<td>73</td>
<td>73</td>
<td>73</td>
<td>73</td>
</tr>
</tbody>
</table>

Altitude profile of energy deposition, $D(\hat{h})A(\hat{h})$, eV/cm

<table>
<thead>
<tr>
<th>$\alpha$ (keV)</th>
<th>$\gamma = 0$</th>
<th>2</th>
<th>5</th>
<th>10</th>
<th>20 keV</th>
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</thead>
<tbody>
<tr>
<td>5</td>
<td>1.29(-3)</td>
<td>2.05(-3)</td>
<td>3.37(-3)</td>
<td>6.07(-3)</td>
<td>1.22(-2)</td>
</tr>
<tr>
<td>10</td>
<td>3.14</td>
<td>3.98</td>
<td>5.33</td>
<td>7.82</td>
<td>1.36</td>
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<tr>
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<tr>
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<td>2.00(-2)</td>
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<td>2.89</td>
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<tr>
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<td>3.93</td>
<td>4.02</td>
<td>4.15</td>
<td>4.38</td>
<td>4.84</td>
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<td>7.64</td>
<td>7.77</td>
<td>7.99</td>
<td>8.43</td>
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</table>
Table 6. Dependence of the altitude profile of energy deposition, $D(h)A(h)$, on the geomagnetic latitude $\lambda_m$, for the case of an incident electron flux isotropic over the downward hemisphere with a power-exponential energy spectrum proportional to $T_0^\gamma \exp(-T_0/a)$, with $\gamma = 0.5$ and $a = 10$ keV. $D(h)A(h)$ is given in units of eV/cm$^3$. Also shown is the height $h$ at which the altitude profile peaks.

<table>
<thead>
<tr>
<th>$\lambda_m$ (deg)</th>
<th>90°</th>
<th>75°</th>
<th>60°</th>
<th>45°</th>
<th>30°</th>
<th>15°</th>
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<tbody>
<tr>
<td>200</td>
<td>1.73(-5)</td>
<td>1.73(-5)</td>
<td>1.73(-5)</td>
<td>1.72(-5)</td>
<td>1.70(-5)</td>
<td>1.60(-5)</td>
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<tr>
<td>180</td>
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<td>3.43</td>
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<td>3.35</td>
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<td>7.75</td>
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<td>1.25(-4)</td>
<td>1.24(-4)</td>
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<td>2.89(-4)</td>
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<tr>
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<td>1.63(-5)</td>
<td>4.61(-6)</td>
<td>-</td>
</tr>
</tbody>
</table>

$h$ (km) 98 98 98 99 100 103
Table 7. Energy deposition function for bremsstrahlung for the case of uniform wide-area precipitation of an electron flux isotropic over the downward hemisphere. The quantity given is $A_{br}(z_m)/T_0$, in units of $\text{cm}^2 \text{g}^{-1}$. Numbers in parentheses indicate powers of ten.

<table>
<thead>
<tr>
<th>$z/T_0$</th>
<th>20</th>
<th>50</th>
<th>100</th>
<th>200</th>
<th>500</th>
<th>1000</th>
<th>2000</th>
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<tbody>
<tr>
<td>$g \text{cm}^{-2} \text{keV}^{-1}$</td>
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<td></td>
<td></td>
</tr>
<tr>
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<td>7.0(-4)</td>
<td>5.7(-4)</td>
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<td>4.1(-4)</td>
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<td>7.0</td>
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<td>4.9</td>
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<td>2.6</td>
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<td>2.4</td>
<td>2.8</td>
<td>4.0</td>
</tr>
<tr>
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<td>1.7</td>
<td>1.5</td>
<td>1.5</td>
<td>2.0</td>
<td>2.9</td>
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Fig. 1. Backscattering coefficients $R_N$ (number albedo) and $R_E$ (energy albedo) for monoenergetic electrons incident on a semi-infinite air medium. The calculated coefficients are for perpendicular incidence ($\theta_o = 0^\circ$) or for an incident flux isotropic over the downward hemisphere (IDH). The experimental values of $R_N$ are for perpendicular incidence and have been obtained by interpolating to atomic number $Z = 7.4$ the data of the following authors:

B: BISHOP (1966); H: HEINRICH (1966); WP: WEINRYB and PHILIBERT (1964); TV: TRUMP and VAN DE GRAAFF (1949); SA: SALDICK and ALLEN (1954); WT: WRIGHT and TRUMP (1962); GG: GLAZUNOV and GUGLYA (1964); S: STERNGLASS (1954); P: PALLUEL (1947); VA: VERDIER and ARNAL (1968).

Fig. 2. Backscattering of monoenergetic electrons from a semi-infinite air medium, as a function of the incident obliquity angle $\theta_o$.

a. Number albedo, $R_N$

b. Energy albedo, $R_E$

Fig. 3. Angular distribution $W(\theta)$ of the current of backscattered electrons emerging from a semi-infinite air medium. The distributions shown are normalized to unity over a $2\pi$-solid angle.

Fig. 4. Cumulative energy spectra of electrons backscattered from a semi-infinite air medium. The curves show, as a function of the incident energy $T_o$, the fraction of the backscattered electrons that have energy $T$ greater than $0.2 T_o$, $0.4 T_o$, $0.6 T_o$ or $0.8 T_o$.

a. Perpendicular incidence, $\theta_o = 0^\circ$

b. Incident flux isotropic over downward hemisphere, $\theta_o$: IDH.

Fig. 5. Cumulative energy spectra of electrons backscattered from a semi-infinite air medium. The curves show, as a function of the incident obliquity angle $\theta_o$, the fraction of the backscattered electrons with energy $T$ greater than $0.2 T_o$, $0.4 T_o$, $0.6 T_o$ or $0.8 T_o$, where $T_o$ is the incident energy.

a. $T_o = 10$ keV   b. $T_o = 100$ keV   c. $T_o = 1000$ keV
Fig. 6. Altitude profile of energy deposition in the atmosphere, for an incident electron flux isotropic over the downward hemisphere with an exponential energy spectrum $\frac{1}{a} \exp(-T_o/a)$. The relation between atmospheric depth and altitude is assumed to be that given by the CIRA (1965) Mean Atmosphere at geomagnetic latitude $\lambda_m = 90^\circ$. Results are normalized to one incident electron/cm$^2$.

Fig. 7. Altitude profile of energy deposition by electron-produced bremsstrahlung in the atmosphere, for the case of wide-area precipitation of an electron flux incident isotropically over the downward hemisphere with an exponential spectrum $\frac{1}{a} \exp(-T_o/a)$. The relation between atmospheric depth and altitude is assumed to be that given by the CIRA (1965) Mean Atmosphere at geomagnetic latitude $\lambda_m = 90^\circ$. Results are normalized to one incident electron/cm$^2$. Dashed portions of curves are extrapolations.

Fig. 8. Comparison of the energy flux spectrum $F(T,z_m)$ with Maeda's "differential energy distribution", for 20-keV electrons incident perpendicularly on an air medium. The dotted curves indicate the contribution to $F(T,z_m)$ from degraded primary electrons.

Fig. 9. Depth distribution of electrons at cut-off energy $\Delta$. The ordinate is the dimensionless quantity $r_o n_{\Delta}(z_m)$, where $n_{\Delta}(z_m)$ is the number of electrons per unit depth whose energy falls below $\Delta$. P and S indicate primary and secondary electrons, respectively. The normalization corresponds to one incident primary electron.

Fig. 10. Comparison of various spectral distributions, at a depth $z_m = 0.5 r_o$, for 20-keV electrons incident perpendicularly on an air medium.

- $F_{4\pi}(T,z_m)$: total flux spectrum
- $F_{2\pi}(T,z_m)$: contribution to total flux spectrum from electrons moving in the direction of increasing depth
- $J(T,z_m)$: spectrum of the current of electrons transmitted through a slab target with a thickness equal to $z_m$.

Fig. 11. Dependence of the electron flux spectrum $F(T,z_m)$ on the incident energy $T_o$. 
Fig. 12. Dependence of the electron flux spectrum on the depth $z_m$, for electrons incident with an energy of 10 keV.

Perpendicular incidence; $r_p = 2.48 \times 10^{-4}$ g cm$^{-2}$

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Incident flux isotropic over the downward hemisphere;

$r_p = 2.30 \times 10^{-4}$ g cm$^{-2}$

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<tr>
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<td>0.40</td>
<td>0.58</td>
<td>0.76</td>
<td>0.94</td>
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Fig. 13. Directional characteristics of the electron flux spectrum, for electrons incident with an energy of 10 keV. The curves labelled DOWN and UP represent the contributions from electrons moving in the direction of increasing and decreasing depth, respectively. The curve marked TOTAL is the sum of DOWN and UP.

Fig. 14. Dependence of the flux spectrum $F(T, z_m)$ on the obliquity (pitch angle) $\theta_o$ of the incident electrons, for an incident electron energy of 10 keV; the scaling parameter $r_p$ in this figure is the practical range for the case of perpendicular incidence.
Fig. 1

- Backscattering coefficient

- $T_0$, keV

- $\theta_0$, IDH

- $R_E$ and $R_N$
Fig. 2a
$T_0 = 10$ keV

$100$ keV

$1000$ keV
Fig. 3

$\theta_o = 0^o$

$2 \leq T_o \leq 2000$ keV

$\theta_o$ : IDH

$W(\theta)$, sr$^{-1}$

$\cos \theta$
Fig. 4b
1.0
0.8
0.6
0.4
0.2
0.0

FRACTION

$\cos \theta_0$

$T > 0.2T_0$

0.4
0.6
0.8

$T_0 = 10$ keV

Fig. 5a
Fig. 5b
Fig. 5c

T > 0.2T₀

FRACTION

0.8

0.6

0.4

0.2

COS θ₀

T₀ = 1000 keV
Energy deposition by electrons

\[ S(T_0) = \frac{1}{\alpha} e^{-T_0/\alpha} \]

\( \theta_0 \) : IDH

Fig. 6
ENERGY DEPOSITION BY BREMSSTRAHLUNG

\[ S(T_0) = \frac{1}{\alpha} e^{-T_0/\alpha} \]

\[ \theta_0: \text{IDH} \]

\[ \alpha = 5 \text{ keV} \]
$T_o = 20 \text{ keV}$
$\theta_o = 0^\circ$

- **TOTAL FLUX, $F(T,z_m)$**
- **MAEDA'S "DIFFERENTIAL ENERGY DISTRIBUTION"**

$z_m = 0.1r_o$

**Fig. 8**
Fig. 9

$T_0 = 20\text{ keV}$
$
\Delta = 1\text{ keV}$

$\theta_0 = 0^\circ$
Fig. 10

SPECTRUM, keV$^{-1}$

$T_0 = 20$ keV

$\theta_0 = 0^\circ$

$z_m = 0.5 r_0$
\[ T_0 = 2 \text{keV} \]

\[ \theta_0 = 0^\circ \]

\[ z_m \approx 0.57 r_p \]

Fig. 11
Fig. 12

$T_0 = 10$ keV
$\theta_0 = 0^\circ$

$T_0 = 10$ keV
$\theta_0$ : IDH
$T_0 = 10$ keV
$\theta_0 = 0^\circ$
$z_m \approx 0.025 r_p$

$T_0 = 10$ keV
$\theta_0 :$ IDH
$z_m \approx 0.025 r_p$

$T_0 = 10$ keV
$\theta_0 = 0^\circ$
$z_m \approx 0.57 r_p$

$T_0 = 10$ keV
$\theta_0 :$ IDH
$z_m \approx 0.57 r_p$

Fig. 13
To = 10 keV
zm ≈ 0.025 r_p

\[ T_0 = 10 \text{ keV} \\
zm \approx 0.025 \ r_p \]

\[ \theta_0 = 89^\circ \]

\[ \theta_0 = 0^\circ, 30^\circ, 45^\circ, 60^\circ, 75^\circ, 89^\circ \]

\[ T_0 = 10 \text{ keV} \\
zm \approx 0.57 \ r_p \]

Fig. 16