ROTATIONAL MODULATION OF HYDROGEN LYMAN ALPHA FLUX FROM 44i BOOTIS*

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* Based on observations with the International Ultraviolet Explorer (IUE) obtained at the Villafranca satellite tracking station of the European Space Agency and at NASA/Goddard Space Flight Center.

ABSTRACT

We present IUE observations that cover the entire 6.4 hour orbital cycle of the late-type contact binary 44i Bootis. The intrinsic stellar hydrogen Lyman alpha emission flux was determined from low-resolution IUE spectra, compensating for geocoronal emission and for interstellar absorption. The variation of the stellar Lyman alpha emission flux correlates well with the variation of the CIII and CIV emission fluxes, and it shows orbital modulation in phase with the visual light curve. The ratio of Lyman alpha to CIII flux (15 to 20) is similar to that observed in solar active regions. Hydrogen Lyman alpha emission is thus one of the most important cooling channels in the outer atmosphere of 44i Boo. We obtained a high-resolution spectrum of the Lyman alpha line between orbital phases 0.0 and 0.6. The integrated flux in the observed high-resolution Lyman alpha profile is consistent with the fluxes determined using low-resolution spectra, and the composite profile indicates that both components of this binary have equally active chromospheres and transition regions. The uncertainty in the interstellar hydrogen column density cannot mimic the observed variation in the integrated Lyman alpha flux, because the stellar line is very much broader than the interstellar absorption.

Key Words: magnetic activity-chromospheres-Lyman alpha emission-ultraviolet spectroscopy - contact binaries

1. INTRODUCTION

The atmospheres of late-type stars consist of relatively cool photospheres, warm chromospheres, and hot coronae. The chromosphere and corona are separated by a thin transition region, in which the temperature changes abruptly from about 7000 K to several million K.

The Lyman alpha line at 1216 Angstroms plays a crucial role in the energetic relationship between the chromospheres and coronae of cool stars. Unfortunately, the possible correlation between the the Lyman alpha line flux and other chromospheric and transition region lines has not yet been investigated, mainly because the very strong Lyman alpha line is generally overexposed on all IUE SWP spectra taken with exposure times long enough to make other interesting spectral features, such as CIV 1550 and CIII 1335, visible. In addition, the stellar emission in low-resolution spectra is severely contaminated by scattered solar Lyman alpha emission from the geocoronal environment of the IUE satellite. Further, interstellar absorption can remove a significant fraction of the intrinsic stellar flux from the line.

Nearby contact binaries are good targets to study the Lyman alpha emission, because they are expected to have strong emission and because their emission lines are rotationally broaded so that interstellar absorption does not remove a large fraction of the total flux. In addition, the short periods of contact binaries allow an orbital modulation study to be made. It is possible to perform a simultaneous study of the Lyman alpha line and many chromospheric and transition region emission lines from the contact binary system 44i Bootis with short SWP exposures (which would not be saturated at Lyman alpha). In addition, 44i Boo is bright enough that a high resolution Lyman alpha spectrum can be obtained (although over a large spread in orbital phases), allowing an important check on the techniques used to remove the effects of geocoronal emission and interstellar absorption from the low-resolution spectra.

2. OBSERVATIONS AND CORRECTION FOR THE GEOCORONAL EMISSION AND INTERSTELLAR ABSORPTION

44i Bootis (HD 133640; SAO 45537) is a contact binary at a distance of 12 pc. It has been well studied with the IUE (Rucinski and Vilhu 1983) and with the EINSTEIN (Crudace and Dupree 1984) and EXOSAT (Vilhu and Heise 1986) observatories.

The amplitudes of the radial velocity variations are $K_1=115$ km s$^{-1}$ and $K_2=231$ km s$^{-1}$ (Batten, Fletcher, and Mann 1978). The mass ratio is $M_2/M_1=0.50$, and the orbital inclination is $77^\circ$ (Rucinski 1973). The ephemeris of light minimum given in Vilhu and Heise (1986) needs to be corrected by 0.015 days to be consistent with the July 1986 photometry of Al-Naimy et al. (1986). We used the orbital ephemeris JD(min)=2439852.5053 and an orbital period of 0.2678159 days.

The low-resolution SWP observations were performed on 28 June 1986 during a contiguous US2+ESA double shift (16 hours), covering roughly one orbital cycle. A high-resolution spectrum was obtained on 6 July 1986, with an exposure begun at about primary minimum lasting until secondary minimum (one half of the orbital cycle).
Fig. 1 Model profiles for the Lyman alpha emission line for 44i Boo. The models include the rotational and intrinsic broadening of both components and the absorption due to interstellar hydrogen. Models were computed for three hydrogen column densities ($10^{17}$, $10^{18}$, and $10^{19}$ atoms cm$^{-2}$) and for three orbital phases. The intensity scale is arbitrary, and wavelength (horizontal) scale (km s$^{-1}$) has its origin at the binary's center of mass.

Fig. 2 The fraction of stellar Lyman alpha emission that is absorbed as a function of the hydrogen column density and orbital phase.

Fig. 3 The stellar Lyman alpha, CII 1335, and CIV 1550 emission line light curves of 44i Boo. The visual points from the IUE Fine Error Sensor (FES) and the computed synthetic visual light curve (SYNTH) are marked.
All reduction and analysis of the IUE observations was performed using software at the Colorado Regional Data Analysis Facility (RDAF). After applying the standard flux calibration to the one-dimensional extracted spectra, the Lyman alpha, C II (1335 Angstroms), and C IV (1550 Angstroms) were determined by fitting Gaussian emission profile to the spectral line and a quadratic function to the local background.

Solar Lyman alpha emission is resonantly scattered into the line of sight by hydrogen in the "geocorona". The major axis of the elliptical large aperture is oriented roughly perpendicular to the dispersion in low-resolution spectra and roughly parallel to it in high-resolution spectra. From the geosynchronous orbital position of the IUE spacecraft, geocoronal emission is seen from all directions, therefore uniformly illuminating the aperture. The intensity of the two-dimensional (spatial and spectral dimensions) geocoronal profile varies with the angle between the line of sight and the sun (the Beta angle), the angle between the line of sight and the Earth, and the intrinsic variability of the solar Lyman alpha (Clarke 1982), but the shape of the profile is due solely to instrumental factors. This profile shows light loss near the edge of the aperture due to the point-spread function of the telescope and is roughly flat across the center of the aperture.

We subtracted the geocoronal profile from the low-resolution SWP images using the method developed by Neff et al (1986). The correction procedure operates on the standard line-by-line (ELBL) files, which include a series of spatially resolved spectra. These spatially resolved spectra are treated as an "image" of the aperture. Geocoronal emission is present throughout the aperture, while the stellar emission is confined to the center of the aperture. We fitted a model profile perpendicular to the dispersion (i.e. in the spatial dimension) at each wavelength within the aperture. This profile was subtracted from each spatial slice to produce a corrected profile shows light loss near the edge of the aperture due to the point-spread function of the telescope and is roughly flat across the center of the aperture.

The geocoronal model profile consists of gaussian wings fit to the outer portion of the aperture (where light loss is significant) and a linear interpolation across the center of the aperture (the central 11 lines in the line-by-line spectrum). We therefore use the information within the aperture but outside the region that contains stellar emission to determine the intensity level of the geocoronal emission. Thus, this procedure will work with any spectrum in the IUE archives. This profile has proven to give a satisfactory fit in cases when only the geocorona was observed (sky background or quasar spectra, Neff et al. 1986). For this particular set of observations the stellar emission was strong and exposure times short, so the geocoronal contribution was relatively small.

Because the Lyman alpha line profile is not resolved in low-resolution observations, the only way to estimate how much of the stellar flux is removed by interstellar absorption is to model both the stellar emission and interstellar absorption profiles. The only observational constraint available in the low-resolution spectra is the total integrated flux in the earth-received profile. Other parameters of the model must be constrained with prior knowledge or else varied between their reasonable limits. We used the method described by Rucinski, Vilhu, and Whelan (1985), who argued that since the stellar profiles of contact binaries are very broad (rapid rotators), the interstellar absorption is not extremely large unless the interstellar hydrogen column density is greater than 10^{18} cm^{-2}. For narrow-lined stars the situation is quite different, as even very small column densities cause the interstellar absorption to affect the entire stellar emission profile.
Fig. 4 The high-resolution Lyman alpha spectrum of 44i Boo obtained 6 July 1986, between orbital phases 0.92-0.57. The theoretical profile is shown by the dashed line. The geocoronal emission is clearly visible inside the broad stellar emission wings. The positions of the interstellar neutral hydrogen and deuterium are marked.

4. DISCUSSION AND SUMMARY

The symmetry of the high-resolution Lyman alpha profile and the lack of any clear rotational modulation of the line fluxes imply that both components of the 44i Boötis system must have similar chromospheres and transition regions. This should be expected, because both components have almost equal effective temperatures and apparently equal angular velocities. To a zeroth approximation, each component would therefore have equal "dynamo numbers".

The classical theory of contact binaries assumes "a common convective envelope" where both stars are embedded. However, more developed theories deviate from thermal equilibrium, producing quite unequal convective zones for the component stars (see for example Rahnen 1981 and 1982; Rahnen and Vilhu 1982). If the chromospheric and transition region line fluxes are indicative of the magnetic activity, then the dynamo power of the components should be equal, and the physical properties of the dynamo layers (such as temperature, density, rotation, differential rotation, turbulent velocities, etc.) should be equal as a consequence. However, as pointed by Vilhu and Heise (1986), because contact binaries are close to the observed saturation limits of all chromospheric and transition region indicators, this test is very insensitive to the dynamo-related parameters.

The total CII and Lyman alpha intensities of solar active regions are correlated, and a linear relation $F(\text{Lyman alpha}) = 25F(\text{CII})$ for a sample of active regions follows from Schrijver et al. (1986). We find for 44i Boo $F(\text{Lyman alpha}) = (18 \pm 6)F(\text{CII})$, quite consistent with the solar active region data. We note, however, that the absolute values of the surface fluxes in 44i Boo are much larger, indicating higher surface coverage of magnetic field than in solar plages. In the most intense solar active region studied by Schrijver et al. $F(\text{Lyman alpha}) = 2.5E6 ~\text{erg s}^{-1} \text{cm}^{-2}$, compared with the mean value $6.5E6 ~\text{erg s}^{-1} \text{cm}^{-2}$ for 44i Boo.

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