Final Technical Report

Grant: NAG5-1858

"Guaranteed Time Observations Support for Goddard High Resolution Spectrograph (GHRS) on HST"

Edward Beaver, Principal Investigator


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Final Technical Report

Guaranteed Time Observations Support for
Goddard High Resolution Spectrograph (GHRS) on HST

NAG5-1858


Ed Beaver, P.I.

University of California, San Diego
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1. Introduction

We assemble this final grant report by combining our previously submitted progress reports with the last year's progress report. Section 2 is the progress report for the June 1, 1991 to Nov. 14, 1995 period. Section 4 is the progress report for the Nov. 14, 1996 to Dec. 31, 1996 period. Section 5 is the progress report for the Nov. 14 to Aug. 31, 1997 period. Section 5 is the new progress report for the Sept. 15, 1997 to Nov. 14, 1998 final period. Section 3 is a summary of our spare detector high voltage transient tests activity in 1992 in support of the renewed safe operation of the GHRS HST D1 detector.

Note that we have left the format of each progress report the same as originally sent out. The slight differences in format presentation are thus intended.

2. June 1, 1991 to November 14, 1995 Submitted Progress Report

Technical Progress Report

Guaranteed Time Observations Support for
Goddard High Resolution Spectrograph (GHRS) on HST

NAG5-1858

June 1, 1991 — November 14, 1995

Ed Beaver, P.I.

University of California, San Diego
I. Overall Progress

A. Results Obtained During the Reporting Period

1. GTO Observations
The goals of the GTO effort are for investigations defined in previous years by the IDT to be carried out as HST observations and for the results to be communicated to the scientific community and to the public. The following is a listing of GHRS GTO observations in the report period.

Scheduled and completed, with acceptable data:

<table>
<thead>
<tr>
<th>Yr.day</th>
<th>Prop</th>
<th>Cy</th>
<th>Old #</th>
<th>Title</th>
<th>Target</th>
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<td>1190</td>
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<td>FUV Em Line profiles, W Serpentis Binaries</td>
<td>W Ser (1/1)</td>
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<td>1191</td>
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<td>1191</td>
<td>Phys Cond, Low z Abs Line Sys in QSOs</td>
<td>PKS2135-14 (1/3)*</td>
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<tr>
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<td>1191</td>
<td>Phys Cond, Low z Abs Line Sys in QSOs</td>
<td>PG1630+337 (2/3)</td>
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<td>4</td>
<td>1191</td>
<td>Phys Cond, Low z Abs Line Sys in QSOs</td>
<td>PG1630+337 (2/3)</td>
</tr>
<tr>
<td>95.071</td>
<td>6237</td>
<td>5</td>
<td>1191</td>
<td>Phys Cond, Low z Abs Line Sys in QSOs</td>
<td>PG0935 (3/3)</td>
</tr>
<tr>
<td>92.007</td>
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<td>UV Emission Line Profiles, QSOs &amp; AGN</td>
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<td>3C286 (1/3)</td>
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<td>Ly α Region of QSOs w Strong Absorp Lines</td>
<td>PKS1229-02 (2/3)</td>
</tr>
<tr>
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<td>3939</td>
<td>2</td>
<td>1193</td>
<td>Ly α Region of QSOs w Strong Absorp Lines</td>
<td>PKS1229-02 (2/3)</td>
</tr>
<tr>
<td>94.094</td>
<td>5176</td>
<td>4</td>
<td>1193</td>
<td>Ly α Region of QSOs w Strong Absorp Lines</td>
<td>PKS1229-02 Im (2/3)</td>
</tr>
<tr>
<td>91.339</td>
<td>1193</td>
<td>1</td>
<td>1193</td>
<td>Ly α Region of QSOs w Strong Absorp Lines</td>
<td>3C196 (3/3)</td>
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<tr>
<td>92.270</td>
<td>3939</td>
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<td>1193</td>
<td>Ly α Region of QSOs w Strong Absorp Lines</td>
<td>3C196 (3/3)</td>
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<tr>
<td>94.083</td>
<td>5176</td>
<td>4</td>
<td>1193</td>
<td>Ly α Region of QSOs w Strong Absorp Lines</td>
<td>3C196 Image (3/3)</td>
</tr>
</tbody>
</table>

*Note: the (1/3) indicates this is the first of 3 targets in this program.

2. Publications (including partial support)

a) Papers Published or In Press


c) Abstracts and Invited Talks

Abstracts for 179th AAS Meeting, Atlanta, Georgia, January 1992


Abstracts for 180th AAS Meeting, Columbus, Ohio, June 1992


Abstracts for 182nd AAS Meeting, Berkeley, California, June 1993


Abstracts for the 183rd AAS Meeting, Washington, D.C., January 1994


Abstracts for the 184th AAS Meeting, Minneapolis, MN, May-June 1994


Abstracts for the 185th AAS Meeting, Tucson, AZ, January 1995


3. Public Outreach

Informing the public is also a part of our mission. GHRS GTO team members regularly give public information lectures. W. Baity is the host of and science consultant to a semi-monthly UCSD-TV series on astronomy and space science. In the area of public information, we include the following:

a) GHRS GTO team members at UCSD gave public information lectures and furnished a display for the ASP “Universe ’93” meeting in San Diego in July 1993.

b) GHRS team members at UCSD have maintained informative pages on the World Wide Web since June, 1994. to communicate our results to scientific colleagues and to the public.
None per se, but in order to complete detailed analysis of GHRS GTO data, we will likely need a no-cost extension beyond the end of the GTO support period in December 1996.
The GHRS GTO at UCSD team has successfully completed the observation phase of their GTO investigations and for the next year will perform analyses of the observations listed in the first section above and complete work on previously obtained data. The HST has been performing up to its original design specifications, which is a testimonial to all those who worked hard on this project and its maintenance and repair.

3. GHRS Spare Detector High Voltage Transient Confidence Tests at UCSD

In 1992 UCSD was asked by NASA to perform high voltage transient tests on the spare GHRS D2 detector. These tests were necessitated by the intermittent failure of the flight GHRS side 1 low voltage power supply, which in turn caused high voltage transients on the D1 detector. Around July 1991, the high voltage operation of the D1 Detector was discontinued until a plan for safe operation was developed.

During the GHRS Detector phase, rapid high voltage changes caused damage to the detectors. The detector failure mode was an elevated (100 times higher than spec) detector background level due to light discharges in the Digicon high voltage front end ceramic spacer cylinder. Even though design safeguards were built into later flight designs (a chrome oxide surface was added to the ceramic spacer), NASA needed this additional ground test simulation data in order to characterize the performance of the flight detector after being subjected to a series of high voltage transients on a GHRS flight detector.

After the D2 spare detector was transferred to non flight status, UCSD personal transported the D2 detector from Ball Aerospace, Bolder by ground transportation (i.e. rental car) to our UCSD test lab. UCSD then converted the FOS vacuum chamber and data acquisition test stand to acquire data from the GHRS D2 detector head assembly. In addition a 386 IBM computer with new software was added to be a real time interface to the GHRS high voltage power supply so that the high voltage power supply could be sequenced on and off in a preset pattern.

After each high voltage pattern, three measurements were made: 1. a dark count to measure any background changes, 2. a detector faceplate mask edge response scan to measure any resolution changes, and 3. a count rate measurement on a stable ultraviolet light source to see if the detector quantum efficiency changed. These measurements are summarized in Table 1. We started out on 10/2/92 with the potentially least damaging high voltage transient sequence of a step high voltage off (from ~22 kilovolt to 0 volts) for 5 seconds then a step turn on (0 volts to 22 kilovolts). Since this was only 1 cycle, the background, edge response, and standard lamp were measured immediately after this sequence. The high voltage transient tests were finished on 10/28/92 with the sequence of 100 cycles of a 30 second on/off high voltage pattern. Note that nominal D2 detector performance before these transient tests were the following: 1. background = 0.001 cts/sec/diode, 2. Edge full width half max = 50 microns, and 3. UV response = 1.00.

A statistical analysis of the data from the sequence of high voltage transient test patterns shown in Table 1 shows no significant changes in the D2 detector performance. Thus it has been demonstrated that the flight GHRS Detector can safely survive a reasonable series of high voltage transients. GHRS NASA Goddard engineer Frank Rebar requested UCSD to write up this high voltage test simulation into a report, so that he could use it to support the high voltage operation of the D1 detector on the HST. We delivered this written report to Mr. Rebar around December 1992. Shortly thereafter the D1 detector was successfully brought back to high voltage operation for the rest of the GHRS mission lifetime on the HST.
TABLE 1: HIGH VOLTAGE TRANSIENT DATA FOR SPARE D2 DETECTOR TESTS

<table>
<thead>
<tr>
<th>DATE</th>
<th>TIME OFF (SEC)</th>
<th>CYCLE (#)</th>
<th>BACKGROUND (10^-3 C/D/S)</th>
<th>EDGE FWHM (MICRONS)</th>
<th>UV RESPONSE (NORMALIZED)</th>
</tr>
</thead>
<tbody>
<tr>
<td></td>
<td></td>
<td></td>
<td>BEFORE</td>
<td>AFTER</td>
<td>BEFORE</td>
</tr>
<tr>
<td>10/2/92</td>
<td>5</td>
<td>1</td>
<td>0.55</td>
<td>0.91</td>
<td>49.7</td>
</tr>
<tr>
<td>10/7/92</td>
<td>10</td>
<td>1</td>
<td>1.27</td>
<td>0.91</td>
<td>48.8</td>
</tr>
<tr>
<td>10/12/92</td>
<td>30</td>
<td>1</td>
<td>1.82</td>
<td>1.82</td>
<td>48</td>
</tr>
<tr>
<td>10/16/92</td>
<td>300</td>
<td>1</td>
<td>0.36</td>
<td>1.45</td>
<td>44.3</td>
</tr>
<tr>
<td>10/19/92</td>
<td>10</td>
<td>10</td>
<td>0.91</td>
<td>0.91</td>
<td>48.8</td>
</tr>
<tr>
<td>10/19/92</td>
<td>30</td>
<td>10</td>
<td>0.55</td>
<td>0.91</td>
<td>47.4</td>
</tr>
<tr>
<td>10/23/92</td>
<td>300</td>
<td>10</td>
<td>0.73</td>
<td>0.91</td>
<td>44.8</td>
</tr>
<tr>
<td>10/26/92</td>
<td>30</td>
<td>100</td>
<td>1.09</td>
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<td>48.6</td>
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<td>10/28/92</td>
<td>30</td>
<td>100</td>
<td>0.91</td>
<td>0.55</td>
<td>50.5</td>
</tr>
</tbody>
</table>

AVERAGE OF COLUMN ENTRIES

| BACKGROUND (10^-3 C/D/S) | 1.01 | 1.09 | 47.93 | 48.45 | 1.037 | 0.963 |
| EDGE FWHM (MICRONS)     | 1.09 | 0.73 | 48    | 52    | 1.050 | 1.035 |

3 SIGMA ERROR OF ENTRIES

| 1.29 | 1.29 | TBD   | TBD   |


Technical Progress Report

Guaranteed Time Observations Support for Goddard High Resolution Spectrograph (GHRS) on HST

NAG5-1858
I. Overall Progress

A. Results Obtained During the Reporting Period

1. GTO Observations

The goals of the GTO effort are for investigations defined in previous years by the IDT to be carried out as HST observations and for the results to be communicated to the scientific community and to the public.

2. Publications (including partial support)

a) Papers Published or In Press


b) Conference Proceedings

None.

c) Abstracts and Invited Talks


3. Public Outreach

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a) GHRS GTO team members at UCSD gave public information lectures at the Fleet Space Theater, San Diego’s planetarium and at various public and school events.
b) GHRS team members at UCSD have maintained informative pages on the World Wide Web since June, 1994.

B. Scientific Significance of Selected GTO Results

PKS 1229-021: We completed the model fits to the UV spectrum of this 21 cm absorber. The best estimate the column density places this system clearly within the damped Lyα sample. We performed a careful extraction of the optical spectrum of the nearby galaxy in the spatial region which corresponds to the location of the galaxy detected in the WFPC2 image. The only detected feature is [O II] in emission at the redshift of the absorption system. The deduced star formation rate based on the [O II] line is 2 solar masses per year. All of the measurements have been completed, as has most of the paper, which will be submitted shortly.

PG1630+377: We made preliminary measurements of all lines in the spectrum of the bright QSO PG1630+377 to measure and set limits on the He I/H I ratio. We found one possible line of He I λ584 associated with a Lyman limit, and we have set limits on the He I/H I ratio in the forest. In the Lyman limit system, we derive N(H I)/N(He I)=100, assuming that the absorption line is a single component. We used photoionization models (produced using CLOUDY) and the best published estimate of the metagalactic UV background to derive the properties of the cloud producing the He I absorption. In order to match the derived abundance ratio, a very high ionization parameter is required, with a correspondingly low density and H I fraction, leading to an unrealistically large cloud. More realistic cloud sizes require that the He I column could be considerably larger than measured (as would be the case if the He I absorption line is a blend of a number of unresolved lines). Alternatively, a harder power law continuum could require a lower ionization parameter and clouds of only 30 kpc diameter.

PG0935+417: Our analysis indicates that measurement of lines of molecular H in this spectrum is not possible, due to contamination by other lines in that system, the Lyα forest, and lines in an absorption system at the emission redshift which we have detected in the UV spectrum. We have acquired and reduced high resolution optical data from the Keck telescope that we can use with absorption lines in the UV spectrum to study both the damped Lyα system in which we had hoped to detect molecular H and the system at the emission redshift. The signal-to-noise of these optical data is too low, and we require additional integration time, for which we have been allocated telescope time in January, 1997.

II. Current Problems

None per se, but in order to complete detailed analysis of GHRS GTO data, we will need a no-cost extension beyond the end of the GTO support period in December 1996.

III. Plans for the Next Year of Support

The GHRS GTO at UCSD team has successfully completed the observation phase of their GTO investigations and for the next year will perform analyses of the observations listed in the first section above and complete work on previously obtained data. The HST has been performing up to its original design specifications, which is a testimonial to all those who worked hard on this project and its maintenance and repair.


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I. Overall Progress

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B. Scientific Significance of Selected GTO Results

PKS2135-147: We use new UV and optical spectra and an archival HST-WFPC2 image to study the z_a-z_e absorber in the z_e = 0.20 QSO PKS 2135-147. The UV spectra show strong z_a-z_e absorption lines of C IV, N V, O VI, Ly-alpha and Ly-beta. The z_a-z_e line profiles are resolved, with deconvolved FWHM of 270 to 450 km/s. Lower limits on the total column densities are of order 10^{15} cm^{-2} for all ions. If the absorber is photoionized by the QSO on the derived relative columns in C IV and H I are roughly correct, then the metallicity must be at least solar. The location of the z_a-z_e absorber remains uncertain. Two L* galaxies in a small cluster centered on PKS 2135-147 lie within 36h^{-1} kpc projected distance and have redshifts consistent with causing or contributing to the z_a-z_e lines. The extensive halo of the QSO's host galaxy could also contribute. Calculations show that the QSO is bright enough to photoionize gas up to O VI in the low-density halos of the host and nearby cluster galaxies. Nonetheless, there is indirect evidence for absorption much nearer the QSO, namely (1) the derived high (albeit uncertain) metallicity. (2) the relatively strong N V absorption lines, which might be caused by a higher nitrogen abundance in the metal-rich gas, and (3) strong, lobe-dominated steep-spectrum radio emission, which is known to correlate with a much higher incidence of (probably intrinsic) z_a-z_e lines. We propose that the CIV/NV/OVI line ratios can be used as a general diagnostic of intrinsic versus intervening absorption, as long as the line saturation effects are understood.

A postscript file version of this paper on our GRS GTO observation of PKS2135-147 is located on the Wold Wide Web at http://xxx.lanl.gov/ps/astro-ph/9705240.

Damped Lyman Alpha Systems: We have continued our studies of low redshift damped Lyman alpha systems. The purpose of this program is to study the evolution galaxies from high redshifts to the present day. While ground based observations are sufficient at high redshifts, HST observations are required to measure features in the UV at times closer to the present. For this program we obtained GTO imaging and spectroscopic observations of four quasars. Observations of two have already been published. In this period, we made use of the unique capabilities of the Keck telescope to acquire high resolution, high signal to noise observations of metal lines in the optical region of our target quasars.

We wish to understand the evolution of the abundance of the elements at low redshift. This is complicated by the absorption of some elements onto dust grains. However, the combination of HST and Keck observations allows us to both measure absolute abundances of the elements and to determine whether the abundances show the signature of metallicity variations of dust absorption. We have obtained complete or partial high resolution data sets on PG 0935+417, PKS 1229-021, and 3CR 286, including lines of Zn II, Cr II, Mn II, Ti II, and Ca II and begun detailed fits of the line profiles to measure the abundances.

PG 0935+417 also has an absorption system at the emission redshift, and a "mini" broad absorption line system more then 50,000 km/s blueward of the emission redshift. Fortuitously, lines of these systems fall in both our HST and ground based spectra. Modeling of the physical conditions in both these systems using both data sets has begun.

PKS 1229-021: During this period of performance we successfully acquired Keck hi-res data (6 km/s) on this object and modeled the absorption line profiles of the optical spectrum in order to determine abundances. All of the measurements when combined with our HST data indicate a Mn abundance down by a factor of 10 with respect to solar, despite a look back time of only 3.4x10^{10} years. The Ti/Mn ratio indicates that this is not due to depletion of metals onto dust grains. Our paper on PKS 1229-021 entitled HST and Keck Observations of PKS 1229-021 - The Star Formation Rate at z=.395 will be submitted to ApJ shortly.

PG1630+377: We made preliminary measurements of all lines in the spectrum of the bright QSO PG1630+377 to measure and set limits on the He I/H I ratio. We found one possible line of He I λ584 associated with a Lyman limit, and we have set limits on the He I/H I ratio in the forest. In the Lyman limit system, we derive N(H I)/N(He I)=100, assuming that the absorption line is a single component. We used photoionization models (produced using CLOUDY) and the best published estimate of the metagalactic UV background to derive the properties of the cloud producing the He I absorption. In order to match the derived...
abundance ratio, a very high ionization parameter is required, with a correspondingly low density and H I fraction, leading to an unrealistically large cloud. More realistic cloud sizes require that the He I column could be considerably larger than measured (as would be the case if the He I absorption line is a blend of a number of unresolved lines). Alternatively, a harder power law continuum could require a lower ionization parameter and clouds of only 30 kpc diameter.

PG0935+417: During this period of performance we acquired and reduced high resolution optical data from the Keck telescope that we can use with absorption lines in the UV spectrum to study both the damped Lyα system in which we had hoped to detect molecular H and the system at the emission redshift. Unfortunately the signal-to-noise of some of the regions of the optical data is too low due to cloudy weather problems during the observation, and we require additional integration time. Fortunately we have been allocated Keck telescope time for PG0935+417 on December 24, 1997.

II. Current Problems
None per se, but in order to complete detailed analysis of GHRS GTO data, we will need a no-cost extension beyond the end of the GTO support period of November 15, 1997.

III. Plans for the Next Year of Support
The GHRS GTO at UCSD team has successfully completed the observation phase of their GTO investigations and for the next year will complete work on previously obtained data.

1. GTO Observations

No Observations this period.

2. Publications

a) Papers Published, In Press, in draft, or in preparation

Cohen, R. D., Beaver, E. A., Diplas, Athanassios, Junkkarinen, V., and Lyons, R. W., "HST and Keck Observations of PKS 1229-021 - The Star Formation Rate at z=0.395", in draft form to be submitted to Ap. J.


Hamann, F., Barlow, T., Beaver, E., Cohen, R. D., Junkkarinen, V., and Burbidge, E. M., "High-Velocity Narrow-Line Absorbers in QSO's", in preparation form, to be submitted to ApJ.

Junkkarinen, Beaver, Cohen and Smith, "Metal abundances in the Low-Redshift Damped Lyman alpha Absorption System in 3CR 286", in preparation form, to be submitted to ApJ.

b) Conference Proceedings


B. Scientific Significance of Selected GTO Results

Damped Lyman Alpha Systems: Our analysis of spectra of low redshift damped Lyman alpha systems continued. We are combining high resolution optical observations of metal absorption lines obtained at Keck Observatory with observations of Lyman alpha obtained with HST. This combination allows us to determine absolute metal abundances of important ionic species, including some with very low abundances and weak lines.

3C286: Our Keck optical spectra of 3C286 are completely analyzed. The results are particularly accurate because of the simple velocity structure in this system. The simple velocity structure also allows a measurement of the diagnostically important Mg abundance. Results are as follows for logarithmic abundances relative to solar abundances:

\[
\begin{align*}
[Mg II/H] &= -1.12; [Ti II/H] = -1.54; [Cr II/H] = -1.440; [Mn II/H] = -1.8; \\
[Fe II/H] &= -1.70; \text{ and } [Zn II/H] = -1.39.
\end{align*}
\]

The debate over the abundances of the damped Lyman alpha systems centers on whether the relative abundances are indicative of absorption onto dust grains, or whether they show an abundance pattern indicative of type II supernova abundances. At first glance, the abundances of the Fe group elements Cr, Fe, and Zn do not appear compatible with either interpretation. The Cr/Zn ratio suggests that little or no dust is present, but in that case the low Fe abundance appears inexplicable. Fe, Zn, and Cr should have the same intrinsic abundance because they should have the same nucleosynthetic history, yet in the absence of dust, Cr appears 80% overabundant with respect to Fe. However, if we use the formalism of Vladilo (1998, ApJ, 493, 583), who proposes that the composition of dust grains is the same regardless of the amount of dust, we find that the measured abundances are compatible with a type II abundance pattern on which has been imposed the signature of a small amount of dust depletion. The only exception appears to be Cr, whose abundance is 20% higher relative than expected relative to Fe. This is difficult to explain, but within the scatter measured for
Galactic stars. The derived relative abundances of Mg and Ti abundances are high relative to solar, while the abundance of Mn is low. These results are consistent with predictions of abundances dominated by the output of type II supernovae. This new 3CR 286 work will likely be submitted for publication as "Metal abundances in the Low-Redshift Damped Lyman alpha Absorption System in 3CR 286" by Junkkarinen, Beaver, Cohen and Smith.

PKS 1229-021: At the 192nd meeting of the American Astronomical Society we presented a paper entitled "High Resolution Spectra of Low Redshift Damped Lyα Absorption Systems". This paper discussed our most recent scientific interpretations of our PKS1229-021 and 3C286 optical and HST spectra. The abstract for this paper follows-

"We have been able to form a fairly complete picture of the galaxy responsible for the z=0.395 absorption line system in PKS 1229-021 by combining Keck HIRES and LRIS spectroscopy with observations taken with the Hubble Space Telescope. The image of the absorber is consistent with the inclined disk of a moderately luminous spiral galaxy. We have not been able to detect the continuum from this galaxy spectroscopically, but our LRIS spectra show emission from (0 II) 3372 which can be interpreted to be indicative of star formation at the rate of a few solar masses per year. The HIRES spectra clearly show an "edge-leading" absorption profile. Prochaska and Wolfe have predicted that the velocity of the center of mass of the absorbing galaxy should fall near one edge of the absorption profile if the damped Lyα systems are due to the rotating disks of spiral galaxies. The [0 II] emission velocity is consistent with this, but there is some ambiguity due to the doublet nature of the [0 II] emission. Although the absorption lines of the abundant elements are saturated in the components which correspond to the H I absorption, we have been able to measure accurate column densities for Ca II, Ti II and Mn II for comparison with the H I column density determined from low resolution HST/FOS spectra. The abundances are compatible with approximately 0.1 of solar, with little or no dust, but they are also consistent with lines of sight toward ξ 0ph through warm interstellar clouds."

We have added additional new material to our GHRS PKS 1229-021 paper (in draft) concerning the character of the intervening galaxy, based on discussions with Dr. A. Wolf and others. In addition some issues were brought up by readers of our AAS paper that required changes to our draft paper.

PG 0935+417: Our analysis of the spectrum of PG 0935+417 is less advanced. However, we have summed all out high resolution spectra and begun analysis of the absorption features.

PG1630+377: For a more accurate measurement of line positions, we used the SPECFIT code to deblend all lines in our G140L and H27 HST spectra. Interstellar lines were then used to find to find absolute line position wavelengths. This additional work was required because we want to determine how many lines in the g160L spectrum are coincident with the HeI (584.334 A) wavelengths positions associated with the Lyman alpha cloud system lines in the H27 spectrum. Initially using a spreadsheet type of analysis, we initially found out of the 101 lines in the HeI spectrum and 64 lines the G160L spectrum, eleven G160L lines were found coincident with the HeI wavelengths associated with the Lyman alpha system lines in the H27 spectrum. The issue in these type of measurements is generally whether this positive detection of eleven line coincidences is statistically significant or simply due to random overlaps with Lyα cloud lines in the G160L wavelength range.

In order to look at the statistical significance of this positive detection result, we developed a Monte Carlo code that simulates our line measurements technique. This IDL coded software reads in the H27 Lyα line list and then calculates the HeI line positions associated with the redshift of the H27 lines. The code then generates a simulated G160L random line spectrum with random wavelength positions, but the number of lines equal to that found in the actual G160L spectrum. Next the code finds the HeI H27 Lyα system lines coincident with those in the random G160L lines. This procedure is then looped a thousand times so that a distribution of coincidences is generated.

The result from the Monte Carlo simulation is that a coincidence of eleven lines occurs with a 5.04 percent of the time probability. Generally for these type of analysis, a positive, nonrandom detection is claimed if the process occurs with less than a 3 percent frequency (preferably much less than 3%). Thus the eleven lines that we detect in our G140L spectrum at HeI wavelengths associated with H27 Lyα cloud systems is not quite statistically significant.
We are continuing to try to improve our result by constraining the line selection process by, for example, only looking for coincidences for large equivalent width H27 Lyo cloud system lines only. If we can achieve the 3 percent or less criterion with these additional line constraints, then this work is probably worth submitting a paper to a refereed journal. Otherwise this PG1630 work will make a nice non-referried workshop article.

7. Conclusion

We conclude this final report with a complete list of our GHRS papers for this Grant. The papers in draft or preparation should be published over the next several years. We will notify the GHRS Principal Investigator as these papers are published. Although the Principal Investigator of this Grant is now retired from UCSD, he will continue to work with the lead authors of the “to be published papers” until they are published.

PUBLISHED WORK (Chronological order)


Papers in Draft or Preparation


Cohen, R. D., Beaver, E. A., Hamann, F. W., Junkkarinen, V. T., Lyons, R. W., "A Search for He I Absorption in Lyman alpha Forest Clouds in Spectra of QSOPG 1630+377", in Preparation, likely to be submitted as a workshop publication.