Formation Flying in Earth, Libration, and Distant Retrograde Orbits

David Foltz
NASA - Goddard Space Flight Center

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Agenda

I. Formation flying – current and future

II. LEO Formations
   Background on perturbation theory / accelerations
   - Two body motion
   - Perturbations and accelerations

   LEO formation flying
   - Rotating frames
   - Review of CW equations,
   - Lambert problems, Shuttle
   - The EO-1 mission
   - Realities of operations

III. Control strategies for formation flight in the vicinity of the libration points
   - Libration missions
   - Natural and controlled libration orbit formations
     - Natural motion
     - Non-Natural motion

IV. Distant Retrograde Orbit Formations

V. References
   - All references are textbooks and published papers
   - Reference(s) used listed on each slide, lower left, as ref#
NASA Themes and Libration Orbits

NASA Enterprises of Space Sciences (SSE) and Earth Sciences (ESE) are a combination of several programs and themes.

- Recent SEC missions include ACE, SOHO, and the L1/L2 WIND mission. The Living With a Star (LWS) portion of SEC may require libration orbits at the L1 and L3 Sun-Earth libration points.

- Structure and Evolution of the Universe (SEU) currently has MAP and the future Micro Arc-second X-ray Imaging Mission (MAXIM) and Constellation-X missions.

- Space Sciences’ Origins libration missions are the James Webb Space Telescope (JWST) and The Terrestrial Planet Finder (TPF).

- The Triana mission is the lone ESE mission not orbiting the Earth.

- A major challenge is formation flying components of Constellation-X, MAXIM, TPF, and Stellar Imager.
Earth Science Launches
Low Earth Orbit Formations

The ‘a.m.’ train
~ 705km, 98° inclination,
10:30 .pm. Descending node sun-sync
- Terra (99): Earth Observatory
- Landsat-7(99): Advanced land imager
- SAC-C(00): Argentina s/c
- EO-1(00) : Hyperspectral inst.

The ‘p.m.’ train
~ 705km, 98° inclination,
1:30 .pm. Ascending node sun-sync
- Aqua (02)
- Aura (04)
- Calipso (05)
- CloudSat (05)
- Parasol (04)
- OCO (tbd)
Space Science Launches
Possible Libration Orbit missions

- FKSI (Fourier Kelvin Stellar Interferometer): near IR interferometer
- JWST (James Webb Space Telescope): deployable, ~6.6 m, L2
- Constellation – X: formation flying in librations orbit
- SAFIR (Single Aperture Far IR): 10 m deployable at L2,
- Deep space robotic or human-assisted servicing
- Membrane telescopes
- Very Large Space Telescope (UV-OIR): 10 m deployable or assembled in LEO, GEO or libration orbit
- MAXIM: Multiple X ray s/c
- Stellar Imager: multiple s/c form a fizeau interferometer
- TPF (Terrestrial Planet Finder): Interferometer at L2
- 30 m single dish telescopes
- SPECS (Submillimeter Probe of the Evolution of Cosmic Structure): Interferometer 1 km at L2
Future Mission Challenges
Considering science and operations

**Orbit Challenges**
- Biased orbits when using large sun shades
- Shadow restrictions
- Very small amplitudes
- Reorientation and Lissajous classes
- Rendezvous and formation flying
- Low thrust transfers
- Quasi-stationary orbits
- Earth-moon libration orbits
- Equilateral libration orbits: $L_4$ & $L_5$

**Operational Challenges**
- Servicing of resources in libration orbits
- Minimal fuel
- Constrained communications
- Limited $\Delta V$ directions
- Solar sail applications
- Continuous control to reference trajectories
- Tethered missions
- Human exploration
Background on perturbation theory / accelerations

- Two Body Motion
- Atmospheric Drag
- Potential Models Forces
- Solar Radiation Pressure
Two - Body Motion

Newton's law of gravity of force is inversely proportional to distance.

Vector direction from $r_2$ to $r_1$ and $r_1$ to $r_2$.

Subtract one from other and define vector $r$, and gravitational constants.

Fundamental Equation of Motion

$$\vec{r} = \vec{\eta} - \vec{\eta}_2$$

$$\vec{r}'' = \left( \frac{-Gm_1m_2}{|\vec{r}|^3} \right) \vec{r}$$

$$\mu = \frac{Gm_1 + m_2}{r^2}$$
FORCES ON
PROPAGATED ORBIT

- Equation Of Motion Propagated.
  \[ \frac{m \, d^2 \vec{r}}{dt^2} = -\frac{\mu \vec{r}}{r^3} + \text{accelerations} \]

- External Accelerations Caused By Perturbations
  \[ \vec{a} = \vec{a}_{\text{nonspherical}} + \vec{a}_{\text{drag}} + \vec{a}_{\text{3body}} + \vec{a}_{\text{srp}} + \vec{a}_{\text{tides}} + \vec{a}_{\text{other}} \]
Gaussian Lagrange Planetary Equations

- Changes in Keplerian motion due to perturbations in terms of the applied force. These are a set of differential equations in orbital elements that provide analytic solutions to problems involving perturbations from Keplerian orbits. For a given disturbing function, $R$, they are given by

$$\dot{a} = \frac{2}{na} \frac{\partial R}{\partial M}$$

$$\dot{e} = \frac{1 - e^2}{na^2e} \frac{\partial R}{\partial M} - \frac{\sqrt{1 - e^2}}{na^2e} \frac{\partial R}{\partial \omega}$$

$$\dot{i} = \frac{\cos I}{na^2 \sqrt{1 - e^2} \sin I} \frac{\partial R}{\partial \omega} - \frac{1}{na^2 \sqrt{1 - e^2} \sin I} \frac{\partial R}{\partial \Omega}$$

$$\dot{\omega} = -\frac{\cos I}{na^2 \sqrt{1 - e^2} \sin I} \frac{\partial R}{\partial I} + \frac{\sqrt{1 - e^2}}{na^2 e} \frac{\partial R}{\partial e}$$

$$\dot{\Omega} = \frac{1}{na^2 \sqrt{1 - e^2} \sin I} \frac{\partial R}{\partial I}$$

$$\dot{M} = n - \frac{1 - e^2}{na^2 e} \frac{\partial R}{\partial e} - \frac{2}{na} \frac{\partial R}{\partial a}$$
Geopotential

- Spherical Harmonics break down into three types of terms
  Zonal – symmetrical about the polar axis
  Sectorial – longitude variations
  Tesseral – combinations of the two to model specific regions
- J2 accounts for most of non-spherical mass
- Shading in figures indicates additional mass
Potential Accelerations

\[ \Phi = \frac{\mu}{r} \left[ 1 - \sum_{l=2}^{\infty} J_1 \left( \frac{R}{r} \right)^l P_l \left( \sin \phi \right) + \sum_{l=1}^{\infty} \sum_{m=1}^{l} \frac{1}{l+1} P_{lm} \left( \sin \phi \right) \left( C_{lm} \cos m \lambda + S_{lm} \sin m \lambda \right) \right] \]

The coordinates of P are now expressed in spherical coordinates \((r, \phi, \lambda)\), where \(\phi\) is the geocentric latitude and \(\lambda\) is the longitude. \(R\) is the equatorial radius of the primary body and \(P_{lm}(\sin \phi)\) is the Legendre’s Associated Functions of degree and \(\lambda\) order \(m\).

The coefficients \(C_{lm}\) and \(S_{lm}\) are referred to as spherical harmonic coefficients. If \(m = 0\) the coefficients are referred to as zonal harmonics. If \(1 \neq m \neq 0\) they are referred to as tesseral harmonics, and if \(1 = m \neq 0\), they are called sectoral harmonics.

Simplified J2 acceleration model for analysis with acceleration in inertial coordinates

\[ a = \nabla \phi = \frac{\delta \phi}{\delta x} \hat{i} + \frac{\delta \phi}{\delta y} \hat{j} + \frac{\delta \phi}{\delta z} \hat{k} \]

\[ a_i = \frac{-3J_2\mu R_e^2 r_i}{2r^5} \left( 1 - \frac{5r_k^2}{r^2} \right) \]

\[ a_j = \frac{-3J_2\mu R_e^2 r_j}{2r^5} \left( 1 - \frac{5r_k^2}{r^2} \right) \]

\[ a_k = \frac{-3J_2\mu R_e^2 r_k}{2r^5} \left( 3 - \frac{5r_k^2}{r^2} \right) \]
Atmospheric Drag

- Atmospheric Drag Force On The Spacecraft Is A Result Of Solar Effects On The Earth's Atmosphere
- The Two Solar Effects:
  - Direct Heating of the Atmosphere
  - Interaction of Solar Particles (Solar Wind) with the Earth's Magnetic Field
- NASA / GSFC Flight Dynamics Analysis Branch Uses Several models:
  - Harris-Priester
    - Models direct heating only
    - Converts flux value to density
  - Jacchia-Roberts or MSIS
    - Models both effects
    - Converts to exospheric temp. And then to atmospheric density
    - Contains lag heating terms
Solar Flux Prediction

Historical Solar Flux, F10.7 cm values
Observed and Predicted (+2s) 1945-2002
Drag Acceleration

- Acceleration defined as

\[ a = \frac{1}{2} \frac{C_d A \rho v_a^2 \hat{v}}{m} \]

\[ A = \text{Spacecraft cross sectional area, (m}^2) \]
\[ C_d = \text{Spacecraft Coefficient of Drag, unitless} \]
\[ m = \text{mass, (kg)} \]
\[ \rho = \text{atmospheric density, (kg/m}^3) \]
\[ v_a = \text{s/c velocity wrt to atmosphere, (km/s)} \]
\[ \hat{v} = \text{inertial spacecraft velocity unit vector} \]
\[ \frac{C_d A}{m} = \text{Spacecraft ballistic property} \]

- Planetary Equation for semi-major axis decay rate of circular orbit
  (Wertz/Vallado p629), small effect in e

\[ \Delta a = - \frac{2\pi C_d A \rho \alpha^2}{m} \]
Solar Radiation Pressure Acceleration

\[ \vec{F} = -\frac{1}{c} \vec{G} A_{s/c} \hat{s}, \text{ but} \]

\[ G = \frac{1350 \text{ watts/m}^2}{c} = \frac{4.5 \times 10^{-6} \text{ watt sec/m}^3}{3 \times 10^8 \text{ m/s}} \]

\[ = 4.5 \times 10^{-6} \text{ N/m}^2 \equiv P_{SR} \]

Therefore,

\[ \vec{F} = -P_{SR} A_{s/c} \hat{s} \quad (3) \]

\[ \vec{F} = \vec{F}_s + \vec{F}_n \]

Where \( \vec{F}_s \) is the force in the solar direction and \( \vec{F}_n \) is the force normal to the surface.

\[ \vec{F} = -P_{SP} A_{s/c} [\alpha \hat{s} + 2 \gamma \hat{n}] \]

where

\[ \alpha \equiv \text{absorptivity coefficient, } 0 \leq \alpha \leq 1, \quad \alpha = 1 - \gamma \]

\[ \gamma \equiv \text{reflectivity coefficient of specular reflection, } 0 \leq \gamma \leq 1 \]

\( \hat{n} \) is a unit vector normal to the surface, \( A_{s/c} \)

\( A_{s/c} \) is the area normal to the sun direction.

From Lagrange's planetary equations

\[ \frac{da}{dt} = \frac{2}{\mu \sqrt{1-e^2}} \left\{ \sin \nu \vec{F}_R + \frac{a(1-e^2)}{r} F_i \right\} \]

Where \( \vec{F}_R \) and \( F_i \) are the radial and in-track solar pressure forces.
Other Perturbations

- **Third Body**
  \[ a_{3b} = -\mu/r^3 \hat{r} = \mu(r_j/r_j^3 - r_k/r_k^3) \]
  - Where \( r_j \) is distance from s/c to body and \( r_k \) is distance from body to Earth

- **Thrust** — from maneuvers and out gassing from instrumentation and materials
  Inertial acceleration: \( \ddot{x} = T_x/m, \ddot{y} = T_y/m, \ddot{z} = T_z/m \)

- **Tides, others**
Ballistic Coefficient

- Area (A) is calculated based on spacecraft model.
  - Typically held constant over the entire orbit
  - Variable is possible, but more complicated to model
  - Effects of fixed vs. articulated solar array

- Coefficient of Drag (C_d) is defined based on the shape of an object.
  - The spacecraft is typically made up of many objects of different shapes.
  - We typically use 2.0 to 2.2 (C_d for a sphere or flat plate) held constant over the entire orbit because it represents an average

- For 3 axis, 1 rev per orbit, earth pointing s/c: A and C_d do not change drastically over an orbit wrt velocity vector
  - Geometry of solar panel, antenna pointing, rotating instruments

- Inertial pointing spacecraft could have drastic changes in B_c over an orbit
Ballistic Effects

Varying the mass to area yields different decay rates
Sample: 100kg with area of 1, 10, 25, and 50m², C_d=2.2

FreeFlyer Plot Window
2/11/2003
Numerical Integration

- Solutions to ordinary differential equations (ODEs) to solve the equations of motion.
- Includes a numerical integration of all accelerations to solve the equations of motion
- Typical integrators are based on
  - Runge-Kutta
    formula for $y_{n+1}$, namely:
    $$ y_{n+1} = y_n + (1/6)(k_1 + 2k_2 + 2k_3 + k_4) $$
    is simply the $y$-value of the current point plus a weighted average of four different $y$-jump estimates for the interval, with the estimates based on the slope at the midpoint being weighted twice as heavily as the those using the slope at the end-points.
  - Cowell-Moulton
  - Multi-Conic (patched)
  - Matlab ODE 4/5 is a variable step RK
Coordinate Systems

• Origin of reference frames:
  Planet
  Barycenter
  Topographic
• Reference planes:
  Equator – equinox
  Ecliptic – equinox
  Equator – local meridian
  Horizon – local meridian

Most used systems
  GCI – Integration of EOM
  ECEF – Navigation
  Topographic – Ground station
Describing Motion Near a Known Orbit

A local system can be established by selection of a central s/c or center point and using the Cartesian elements to construct the local system that rotates with respect to a fixed point (spacecraft)

\[
\bar{r}^* = \bar{r}^* (t) \\
\bar{v}^* = \bar{v}^* (t) \\
\bar{r} = \bar{r} (t) \quad \text{Known (reference Orbit)} \\
\bar{v} = \bar{v} (t)
\]

\[
\delta \bar{r} (t) = \bar{r} (t) - \bar{r}^* (t) \quad \text{Relative Motion}
\]

\[
\delta \bar{v} (t) = \bar{v} (t) - \bar{v}^* (t)
\]

\[
\frac{d^2}{dt^2} \bar{r} (t) = 1 / m \bar{F} (\bar{r}, \bar{v}) = \bar{f} (\bar{r}, \bar{v})
\]

\[
\frac{d^2}{dt^2} \delta \bar{r} = \frac{d^2}{dt^2} (\bar{r} (t) - \bar{r}^* (t)) = \frac{d^2}{dt^2} \bar{r} (t) - \frac{d^2}{dt^2} \bar{r}^* (t) = \bar{f} (\bar{r}, \bar{v}) - \bar{f} (\bar{r}^*, \bar{v}^*)
\]

\[
\frac{d^2}{dt^2} \delta \bar{r} (t) = \bar{f} (\bar{r}, \bar{v}) - \bar{f} (\bar{r}^*, \bar{v}^*)
\]
As it stands (1) is exact. However if $\delta \mathbf{r}$ is sufficiently close, the term $\mathbf{f}(\mathbf{r})$ can be expanded via Taylor’s series ...

$$\mathbf{f}(\mathbf{r}) = \mathbf{f}(\mathbf{r}^* + \delta \mathbf{r}) = \mathbf{f}(\mathbf{r}^*) + \left. \frac{\partial \mathbf{f}}{\partial \mathbf{r}} \right|_{\mathbf{r}=\mathbf{r}^*} \delta \mathbf{r} + ...$$

Substituting in yields a linear set of ODEs

$$\frac{d^2}{dt^2} \delta \mathbf{r}(t) = \mathbf{f}(\mathbf{r}) - \mathbf{f}(\mathbf{r}^*) = \mathbf{f}(\mathbf{r}^*) + \left. \frac{\partial \mathbf{f}}{\partial \mathbf{r}} \right|_{\mathbf{r}=\mathbf{r}^*} \delta \mathbf{r} - \mathbf{f}(\mathbf{r}^*)$$

This is important since it will be our starting point for everything that follows
Describing Motion Near a Known Orbit

- Describe motion taking place near a circular orbit
- A natural coordinate frame is one that rotates with the circular orbit

\[
\hat{R} = \frac{\vec{r}^*}{|\vec{r}^*|} \quad \hat{N} = \frac{\vec{r}^* \times \vec{v}^*}{|\vec{r}^* \times \vec{v}^*|} \quad \hat{T} = \hat{N} \times \hat{R}
\]

The frame described is known as
- Hill’s
- Clohessy-Wiltshire
- LVLH
- RTN
- RAC
- RIC
Describing Motion Near a Known Orbit

Any vector will now be given by: \[ \vec{A} = x\hat{R} + y\hat{T} + z\hat{N} = \begin{bmatrix} x \\ y \\ z \end{bmatrix} \]

Now we can evaluate \[ \vec{f} = \left( -\frac{\mu}{r^3} \vec{r} \right) \Rightarrow f_i = \left( -\frac{\mu}{r^3} x_i \right) \]

\[ \frac{\partial \vec{f}}{\partial \vec{r}} = \frac{\partial f_i}{\partial x_j} = \frac{\partial}{\partial x_j} \left( -\frac{\mu}{r^3} x_i \right) = \frac{3\mu}{r^5} x_i x_j - \frac{\mu}{r^3} \delta_{ij} \]

\[ \frac{\partial \vec{f}}{\partial \vec{r}} = \frac{\mu}{r^5} (x_i x_j - r^2 \delta_{ij}) \]

In RTN =RR

\[ \vec{r} = x=r=r \quad \vec{r}^* = \frac{\mu}{r^5} \left( \begin{array}{ccc} 2x^2 - y^2 - z^2 & 3xy & 3xz \\ 3xy & 2y^2 - x^2 - z^2 & 3yz \\ 3xz & 3yz & 2z^2 - y^2 - x^2 \end{array} \right) \]

\[ \frac{\partial \vec{f}}{\partial r} \bigg|_{r^*} = \frac{\mu}{r^3} \left( \begin{array}{ccc} 2 & 0 & 0 \\ 0 & -1 & 0 \\ 0 & 0 & -1 \end{array} \right) n^2 = \left( \begin{array}{ccc} 2 & 0 & 0 \\ 0 & -1 & 0 \\ 0 & 0 & -1 \end{array} \right) \]

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Transforming the EOM

Now convert Newton's 2nd law to RTN frame

\[
\frac{d}{dt} \left( \frac{d}{dt} \right)_{\text{moving}} = \left( \frac{d}{dt} \right)_{\text{fixed}} - \omega \times \delta r
\]

Newton's law involves 2nd derivatives:

\[
\frac{D^2}{Dt^2} \delta \vec{r} = \frac{d^2}{dt^2} \delta \vec{r} - \frac{d \omega^2}{dt^2} \times \delta \vec{r} - 2 \omega \times \frac{d}{dt} \delta \vec{r} + \omega \times (\omega \times \delta \vec{r})
\]

\[
\frac{D^2}{Dt^2} \delta \vec{r} = \frac{d^2}{dt^2} \delta \vec{r} - \frac{d \omega^2}{dt^2} \times \delta \vec{r} - 2 \omega \times \frac{D}{Dt} \delta \vec{r} + \omega \times (\omega \times \delta \vec{r})
\]

\[
\frac{d \omega}{dt} = \frac{D \omega}{Dt} = 0
\]

\[
\frac{D}{Dt} \delta \vec{r} = \left( \begin{array}{c} \dot{x} \\ \dot{y} \\ \dot{z} \end{array} \right)
\]

\[
\omega \times \delta \vec{r} = \left( \begin{array}{c} -ny \\ nx \\ 0 \end{array} \right)
\]

\[
\omega \times (\omega \times \delta \vec{r}) = \left( \begin{array}{c} -n^2x \\ -n^2y \\ 0 \end{array} \right)
\]

\[
\frac{d^2}{dt^2} \delta \vec{r} = \frac{\partial^2 f}{\partial r^2} \bigg|_{\vec{r}} \cdot \delta \vec{r} = \left( \begin{array}{c} -2n^2x \\ -n^2y \\ -n^2z \end{array} \right)
\]
Transforming the EOM yields Clohessy-Wiltshire Equations

\[
\frac{\ddot{r}}{Dt^2} = \begin{pmatrix} -2n^2x \\ -n^2y \\ -n^2z \end{pmatrix} - 2 \begin{pmatrix} -n^2\dot{x} \\ -n^2\dot{y} \end{pmatrix} - \begin{pmatrix} -n^2x \\ -n^2y \end{pmatrix}
\]

\[
\begin{pmatrix} \dot{x} \\ \dot{y} \\ \dot{z} \end{pmatrix} = \begin{pmatrix} 3n^2x + 2n^2\dot{y} \\ -2n\dot{x} \\ -nz \end{pmatrix}
\]

\[
\dot{y} = -2n\dot{x} \Rightarrow \dot{y} = -2nx + k_1
\]

\[
\dot{x} = -n^2x + 2nk_1
\]

\[
x = x_0 \cos(nt) + \frac{v_0}{n} \sin(nt) + \frac{2k_1}{n}
\]

\[
y = -2x_0 \sin(nt) + \frac{2v_0}{n} \cos(nt) + \frac{2k_1}{n} t + y_0
\]

A “balance” form will have no secular growth, \(k_1=0\)

\[
z = z_0 \cos(nt) + \frac{\omega_0}{n} \sin(nt)
\]

Note that the y-motion (associated with T) has twice the amplitude of the x motion (R)
Relative Motion

A numerical simulation using RK8/9 and point mass
Effect of Velocity (1 m/s) or Position (1 m) Difference

- An Along-track separation remain constant
- A 1 m radial position difference yields an along-track motion
- A 1 m/s along-track velocity yields an along-track motion
- A 1 m/s radial velocity yields a shifted circular motion
Shuttle Vbar / Rbar

- Shuttle approach strategies
  - Vbar – Velocity vector direction in an LVLH (CW) coordinate system
  - Rbar – Radial vector direction in an LVLH (CW) coordinate system

- Passively safe trajectories – Planned trajectories that make use of predictable CW motion if a maneuver is not performed.

- Consideration of ballistic differences – Relative CW motion considering the difference in the drag profiles.

Graphics Ref: Collins, Meissinger, and Bell, Small Orbit Transfer Vehicle (OTV) for On-Orbit Satellite Servicing and Resupply, 15th USU Small Satellite Conference, 2001
What Goes Wrong with an Ellipse

In stable – space notation, linearization is written as

\[
\frac{d}{dt} \delta \mathbf{S} = A \delta \mathbf{S}
\]

\[
A = \begin{pmatrix}
0 & I \\
\frac{\partial f}{\partial \mathbf{x}} & 0
\end{pmatrix}
\]

Since the equation is linear

\[
\delta \mathbf{S} = \Phi \delta \mathbf{S}_0 \implies \frac{d}{dt} \Phi = A \Phi
\]

\[
\Phi = I + \int_{t_0}^{t} dt' A(\mathbf{r}^*(t')) \Phi(t',t_0)
\]

Which has no closed form solution if

\[
[A(t_1), A(t_2)] = 0
\]
Lambert Problem

- Consider two trajectories \( r(t) \) and \( R(t) \).
- Transfer from \( r(t) \) to \( R(t) \) is affected by two \( \Delta V \)s
  - First \( \Delta V_1 \) is designed to match the velocity of a transfer trajectory \( \mathcal{R}(t) \) at time \( t_2 \)
  - Second \( \Delta V_1 \) is designed to match the velocity of \( R(t) \) where the transfer intersects at time \( t_3 \)
- Lambert problem:
  Determine the two \( \Delta V \)s
Lambert Problem

- The most general way to solve the problem is to use to numerically integrate \( r(t) \), \( R(t) \), and \( \Phi(t) \) using a shooting method to determine \( \delta V_i \) and then simply subtracting to determine \( \delta V_f \).
- However, this is relatively expensive (prohibitively onboard) and is not necessary when \( r(t) \) and \( R(t) \) are close.
- For the case when \( r(t) \) and \( R(t) \) are nearby, say in a stationkeeping situation, then linearization can be used.
- Taking \( r(t) \), \( R(t) \), and \( \Phi(t_3,t_2) \) as known, we can determine \( \delta V_i \) and \( \delta V_f \) using simple matrix methods to compute a 'single pass'.
The Hohmann Transfer

\[ v_p^2 = \frac{\mu}{r_p} \left( \frac{2 - \frac{1}{a}}{r_a + r_p} \right) \]

\[ a = \frac{r_a + r_p}{2} \]

\[ v_p = \frac{\mu}{\sqrt{r_a r_p}} \]

\[ v_o = \sqrt{\frac{2r_p}{r_a + r_p}} \]

\[ V_e = \sqrt{\frac{\mu}{r_p}} \]

\[ \Delta V_T = \Delta V_1 + \Delta V_2 \]

\[ \Delta V_1 = \sqrt{\frac{\mu}{r_1}} \left( \sqrt{\frac{2r_2}{r_1 + r_2}} - 1 \right) \]

\[ \Delta V_2 = \sqrt{\frac{\mu}{r_2}} \left( 1 - \sqrt{\frac{2r_1}{r_1 + r_2}} \right) \]

where:

- \( r_1 \) and \( r_2 \) are the radii of the elliptical orbit
- \( \mu \) is the gravitational constant
- \( a \) is the semi-major axis of the ellipse
EO-1 GSFC Formation Flying
New Millennium Requirements

- **Enhanced Formation Flying (EFF)**
  
The Enhanced Formation Flying (EFF) technology shall provide the autonomous capability of flying over the same ground track of another spacecraft at a fixed separation in time.

- **Ground track Control**
  
  EO-1 shall fly over the same ground track as Landsat-7. EFF shall predict and plan formation control maneuvers or Da maneuvers to maintain the ground track if necessary.

- **Formation Control**
  
  Predict and plan formation flying maneuvers to meet a nominal 1 minute spacecraft separation with a +/- 6 seconds tolerance. Plan maneuver in 12 hours with a 2 day notification to ground.

- **Autonomy**
  
  The onboard flight software, called the EFF, shall provide the interface between the ACS / C&DH and the AutoCon™ system for Autonomy for transfer of all data and tables.
Formation Flying Maintenance Description
Landsat-7 and EO-1

Different Ballistic Coefficients and Relative Motion
EO-1 Formation Flying Algorithm

- Determine \((r_1, v_1)\) at \(t_0\) (where you are at time \(t_0\)).
- Determine \((R_1, V_1)\) at \(t_1\) (where you want to be at time \(t_1\)).
- Project \((R_1, V_1)\) through \(-\Delta t\) to determine \((r_0, v_0)\) (where you should be at time \(t_0\)).
- Compute \((\delta r_0, \delta v_0)\) (difference between where you are and where you want to be at \(t_0\)).

Keplerian State \((t_1)\) \quad \downarrow \quad \delta r(t_0) \quad \delta v(t_0) \quad \Delta t \quad \rightarrow \quad \text{Keplerian State } K(t_f)
State Transition Matrix

A state transition matrix, $F(t_1, t_0)$, can be constructed that will be a function of both $t_1$ and $t_0$ while satisfying matrix differential equation relationships. The initial conditions of $F(t_1, t_0)$ are the identity matrix. Having partitioned the state transition matrix, $F(t_1, t_0)$ for time $t_0 < t_1$:

$$
\Phi(t_1, t_0) = \begin{bmatrix}
\Phi_1(t_1, t_0) & \Phi_2(t_1, t_0) \\
\Phi_3(t_1, t_0) & \Phi_4(t_1, t_0)
\end{bmatrix}
$$

We find the inverse may be directly obtained by employing symplectic properties:

$$
\Phi^{-1}(t_1, t_0) = \begin{bmatrix}
(\Phi_4(t_1, t_0))^T & (\Phi_2(t_1, t_0))^T \\
(\Phi_3(t_1, t_0))^T & (\Phi_1(t_1, t_0))^T
\end{bmatrix}
\quad \Phi^{-1}(t_1, t_0) = \Phi(t_0, t_1) = \begin{bmatrix}
\Phi_1(t_0, t_1) & \Phi_2(t_0, t_1) \\
\Phi_3(t_0, t_1) & \Phi_4(t_0, t_1)
\end{bmatrix}
$$

$F(t_0, t_1)$ is based on a propagation forward in time from $t_0$ to $t_1$ (the navigation matrix) $F(t_1, t_0)$ is based on a propagation backward in time from $t_1$ to $t_0$, (the guidance matrix).

We can further define the elements of the transition matrices as follows:

$$
\tilde{R}(t_1) \equiv \Phi_1(t_1, t_0) \quad \tilde{R}^*(t_0) \equiv \Phi_1(t_0, t_1)
$$

$$
R(t_1) \equiv \Phi_2(t_1, t_0) \quad R^*(t_0) \equiv \Phi_2(t_0, t_1)
$$

$$
\tilde{V}(t_1) \equiv \Phi_3(t_1, t_0) \quad \tilde{V}^*(t_0) \equiv \Phi_3(t_0, t_1)
$$

$$
V(t_1) \equiv \Phi_4(t_1, t_0) \quad V^*(t_0) \equiv \Phi_4(t_0, t_1)
$$

$$
\begin{bmatrix}
\tilde{R}^*(t_0) & R^*(t_0) \\
\tilde{V}^*(t_0) & V^*(t_0)
\end{bmatrix} = 
\begin{bmatrix}
V^T(t_1) & -R(t_1) \\
-\tilde{V}^T(t_1) & \tilde{R}(t_1)
\end{bmatrix}
$$

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Enhanced Formation Flying Algorithm

The Algorithm is found from the STM and is based on the simplectic nature (navigation and guidance matrices) of the STM)

\[ R(t_1) \]
\[ R(t_1) \]

\[ \delta r_0 = (r_1 - r_0) \quad \delta v_0 = (v_1 - v_0) \]

\[ [R(t_1)] = \frac{r_0}{\mu} (1 - F) \left[ (R_1 - r_0) v_0^T - (V_1 - v_0) r_0^T \right] + \frac{C}{\mu} [V_1 v_0^T] + G[I] \]

\[ [\bar{R}(t_1)] = \frac{R_1}{\mu} \left[ (V_1 - v_0)(V_1 - v_0)^T \right] + \frac{1}{r_0^3} [r_0 (1 - F) R_1 r_0^T + CV_1 r_0^T] + F[I] \]

\[ \Delta v_0 = \left\{ \left[ \bar{R}^T(t_1) \right] [-R^T(t_1)] \right\} \delta r_0 - \delta v_0 \]

where F and G are found from Gauss problem and the f & g series and C found through universal variable formulation
EO-1 AutoCon™ Functional Description

- **Ground Uplink**
  - Landsat-7 Data

- **Relative Navigation**
  - EO-1/LS-7

- **On-Board Navigation**
  - Orbit Determination
  - Attitude Determination

- **C&DH / ACS**
  - Generate Commands
  - Execute Burn

- **Thrusters**

- **Decision Rules**
  - Performance Objectives
  - Flight Constraints

- **How do I get there?**
EO-1 Formation Flying Subsystem Interfaces

- EFF Subsystem
- AutoCon-F
  - GSFC
  - JPL
  - GPS Data Smoother
- Stored Command Processor
- Cmd Load
Difference in EO-1 Onboard and Ground Maneuver Quantized DVs

Quantized - EO-1 rounded maneuver durations to nearest second

<table>
<thead>
<tr>
<th></th>
<th></th>
<th></th>
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<th></th>
<th></th>
<th></th>
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</thead>
<tbody>
<tr>
<td>Δt (sec)</td>
<td>8.33</td>
<td>3.55</td>
<td>4.98</td>
<td>2.43</td>
<td>1.08</td>
<td>2.38</td>
<td>5.29</td>
<td>2.19</td>
<td>3.55</td>
</tr>
<tr>
<td>ΔV (m/s)</td>
<td>4.33</td>
<td>7.93</td>
<td>0.00</td>
<td>3.79</td>
<td>1.62</td>
<td>0.26</td>
<td>1.85</td>
<td>5.20</td>
<td>7.93</td>
</tr>
<tr>
<td>ΔV (m/s²)</td>
<td>2e-7</td>
<td>3e-7</td>
<td>1e-7</td>
<td>3e-7</td>
<td>6e-6</td>
<td>1e-7</td>
<td>8e-4</td>
<td>4e-7</td>
<td>3e-7</td>
</tr>
<tr>
<td>Δt (sec)</td>
<td>2e-1</td>
<td>2e-7</td>
<td>0.0</td>
<td>2e-7</td>
<td>3e-3</td>
<td>1e-7</td>
<td>3e-4</td>
<td>3e-3</td>
<td>3e-3</td>
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<td>ΔV (m/s)</td>
<td>.0008</td>
<td>.0008</td>
<td>.0001</td>
<td>.0001</td>
<td>.0588</td>
<td>.0001</td>
<td>1.569</td>
<td>.0016</td>
<td>.0008</td>
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<td>ΔV (m/s²)</td>
<td>.0030</td>
<td>.0030</td>
<td>.0005</td>
<td>.0005</td>
<td>14.23</td>
<td>.0007</td>
<td>1.572</td>
<td>.0002</td>
<td>3.57</td>
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</tbody>
</table>

**Inclination Maneuver Validation**: Computed ΔV at node crossing, of ~ 24 cm/s (114 sec duration), Ground validation gave same results.
Difference in EO-1 Onboard and Ground Maneuver Three-Axis $\Delta$Vs

EO-1 maneuver computations in all three axis

<table>
<thead>
<tr>
<th></th>
<th>Altitude</th>
<th>$\Delta V_{yz}$</th>
<th>$\Delta V_{zx}$</th>
<th>$\Delta V_{xy}$</th>
</tr>
</thead>
<tbody>
<tr>
<td>Auto</td>
<td>3.15</td>
<td>0.122</td>
<td>0.196</td>
<td>0.032</td>
</tr>
<tr>
<td>Auto</td>
<td>8.45</td>
<td>0.683</td>
<td>0.308</td>
<td>0.008</td>
</tr>
<tr>
<td>Auto</td>
<td>10.85</td>
<td>-5e-4</td>
<td>0.000</td>
<td>0.000</td>
</tr>
<tr>
<td>Auto</td>
<td>11.86</td>
<td>0.178</td>
<td>0.015</td>
<td>-0.0102</td>
</tr>
<tr>
<td>Semi-auto</td>
<td>12.64</td>
<td>0.312</td>
<td>0.024</td>
<td>0.00091</td>
</tr>
<tr>
<td>Semi-auto</td>
<td>14.76</td>
<td>0.188</td>
<td>0.013</td>
<td>0.000</td>
</tr>
<tr>
<td>Semi-auto</td>
<td>15.38</td>
<td>-0.256</td>
<td>-0.0016</td>
<td>-0.0633</td>
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<tr>
<td>Manual</td>
<td>15.58</td>
<td>10.41</td>
<td>.6682</td>
<td>-.0117</td>
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<tr>
<td>Manual</td>
<td>15.47</td>
<td>0.002</td>
<td>.0001</td>
<td>-.0307</td>
</tr>
</tbody>
</table>

$\Delta V_{yz}$, $\Delta V_{zx}$, $\Delta V_{xy}$ in meters per second (m/s).
Formation Data from Definitive Navigation Solutions

Radial vs. along-track separation over all formation maneuvers (range of 425-490km)

Ground-track separation over all formation maneuvers maintained to 3km
Formation Data from Definitive Navigation Solutions

Along-track separation vs. Time over all formation maneuvers (range of 425-490 km)

Semi-major axis of EO-1 and LS-7 over all formation maneuvers
Formation Data from Definitive Navigation Solutions

Frozen Orbit eccentricity over all formation maneuvers (range of 0.001125 - 0.001250)

Frozen Orbit $\omega$ vs. ecc. over all formation maneuvers. $\omega$ range of 90+/- 5 deg.
EO-1 Summary / Conclusions

- A demonstrated, validated fully non-linear autonomous system

- A formation flying algorithm that incorporates
  - Intrack velocity changes for semi-major axis ground-track control
  - Radial changes for formation maintenance and eccentricity control
  - Crosstrack changes for inclination control or node changes
  - Any combination of the above for maintenance maneuvers
Summary / Conclusions

- Proven executive flight code
- Scripting language alters behavior w/o flight software changes
- I/F for Tlm and Cmds
- Incorporates fuzzy logic for multiple constraint checking for maneuver planning and control
- Single or multiple maneuver computations.
- Multiple or generalized navigation inputs (GPS, Uplinks).
- Attitude (quaternion) required of the spacecraft to meet the ΔV components
- Maintenance of combinations of Keperlian orbit requirements sma, inclination, eccentricity, etc.

Enables Autonomous StationKeeping, Formation Flying and Multiple Spacecraft Missions
CONTROL STRATEGIES FOR FORMATION FLIGHT IN THE VICINITY OF THE LIBRATION POINTS
# NASA Libration Missions

## L1 Missions

<table>
<thead>
<tr>
<th>Mission</th>
<th>Description</th>
</tr>
</thead>
<tbody>
<tr>
<td>ISEE-3/ICE</td>
<td>L1 Halo Orbit, Direct Transfer, L2 Pseudo Orbit, Comet Mission</td>
</tr>
<tr>
<td>WIND (94-04)</td>
<td>Multiple Lunar Gravity Assist - Pseudo-L1/2 Orbit</td>
</tr>
<tr>
<td>SOHO (95-04)</td>
<td>Large Halo, Direct Transfer</td>
</tr>
<tr>
<td>ACE (97-04)</td>
<td>Small Amplitude Lissajous, Direct Transfer</td>
</tr>
<tr>
<td>GENESIS (01-04)</td>
<td>Lissajous Orbit, Direct Transfer, Return Via L2 Transfer</td>
</tr>
<tr>
<td>TRIANA</td>
<td>L1 Lissajous Constrained, Direct Transfer</td>
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</table>

## L2 Missions

<table>
<thead>
<tr>
<th>Mission</th>
<th>Description</th>
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</thead>
<tbody>
<tr>
<td>GEOTAIL (1992)</td>
<td>L2 Pseudo Orbit, Gravity Assist</td>
</tr>
<tr>
<td>MAP (2001-04)</td>
<td>Orbit, Lissajous Constrained, Gravity Assist</td>
</tr>
<tr>
<td>JWST (~2012)</td>
<td>Large Lissajous, Direct Transfer</td>
</tr>
<tr>
<td>CONSTELLATION-X</td>
<td>Lissajous Constellation, Direct Transfer?, Multiple S/C</td>
</tr>
<tr>
<td>SPECS</td>
<td>Lissajous, Direct Transfer?, Tethered S/C</td>
</tr>
<tr>
<td>MAXIM</td>
<td>Lissajous, Formation Flying of Multiple S/C</td>
</tr>
<tr>
<td>TPF</td>
<td>Lissajous, Formation Flying of Multiple S/C</td>
</tr>
</tbody>
</table>

(Previous missions marked in blue)
Mission: Investigate Solar-Terrestrial relationships, Solar Wind, Magnetosphere, and Cosmic Rays
Lissajous Orbit: L1 Libration Halo Orbit, Ax=\(~175,000\)km, Ay = 660,000km, Az=120,000km, Class I
Spacecraft: Mass=480Kg, Spin stabilized,
Notable: First Ever Libration Orbiter, First Ever Comet Encounter
**Mission:** Investigate Solar-Terrestrial Relationships, Solar Wind, Magnetosphere

**Launch:** Nov., 1994, Multiple Lunar Gravity Assist

**Lissajous Orbit:** Originally an L1 Lissajous Constrained Orbit, Ax~10,000km, Ay~350,000km, Az~250,000km, Class I

**Spacecraft:** Mass=1254kg, Spin Stabilized,

**Notable:** First Ever Multiple Gravity Assist Towards L1
Mission: Produce an Accurate Full-sky Map of the Cosmic Microwave Background Temperature Fluctuations (Anisotropy)
Launch: Summer 2001, Gravity Assist Transfer
Lissajous Orbit: L2 Lissajous Constrained Orbit Ay~ 264,000km, Ax~tbd, Ay~ 264,000km, Class II
Spacecraft: Mass=818kg, Three Axis Stabilized,
Notable: First Gravity Assisted Constrained L2 Lissajous Orbit; Map-earth Vector Remains Between 0.5° and 10° off the Sun-earth Vector to Satisfy Communications Requirements While Avoiding Eclipses
Launch: JWST~2010, Direct Transfer
Lissajous Orbit: L2 large lissajous, Ay~ 294,000km, Ax~800,000km, Az~ 131,000km, Class I or II
Spacecraft: Mass~6000kg, Three Axis Stabilized, ‘Star’ Pointing
Notable: Observations in the Infrared Part of the Spectrum. Important That the Telescope Be Kept at Low Temperatures, ~30K. Large Solar Shade/Solar Sail
The linearized equations of motion for a S/C close to the libration point are calculated at the respective libration point.

- Linearized Eq. Of Motion Based on Inertial X, Y, Z Using
  
  \[
  \begin{align*}
  x &= x_0 + x, \\
  y &= y_0 + y, \\
  z &= z_0 + z
  \end{align*}
  \]

  \[
  \ddot{x} - 2n\dot{y} = U_{xx} x, \quad \dot{y} + 2n\dot{x} = U_{yy} y, \quad \ddot{z} = U_{zz} z
  \]

- Pseudopotential:

  \[
  U_{xx} = \frac{\partial^2 U}{\partial x^2}, \quad U_{yy} = \frac{\partial^2 U}{\partial y^2}, \quad U_{zz} = \frac{\partial^2 U}{\partial z^2}.
  \]
A State Space Model

\[ \dot{x}^j = A^j x^j \]

where

\[ x^j = \begin{bmatrix} x^j \\ y^j \\ z^j \\ x^j \\ y^j \\ z^j \end{bmatrix}, \quad A^j = \begin{bmatrix} 0 & 0 & 0 & 1 & 0 & 0 \\ 0 & 0 & 0 & 0 & 1 & 0 \\ 0 & 0 & 0 & 0 & 0 & 1 \\ U_{XX} & 0 & 0 & 0 & 2n & 0 \\ 0 & U_{YY} & 0 & -2n & 0 & 0 \\ 0 & 0 & U_{ZZ} & 0 & 0 & 0 \end{bmatrix} \]

\[ n = \sqrt{\frac{G(M_1 + M_2)}{D}} \]
Reference Motions

- **Natural Formations**
  - String of Pearls
  - Others: Identify via Floquet controller (CR3BP)
    - Quasi-Periodic Relative Orbits (2D-Torus)
    - Nearly Periodic Relative Orbits
    - Slowly Expanding Nearly Vertical Orbits + Stable Manifolds

- **Non-Natural Formations**
  - Fixed Relative Distance and Orientation RLP
    - Inertial
  - Fixed Relative Distance, Free Orientation
  - Fixed Relative Distance & Rotation Rate
  - Aspherical Configurations (Position & Rates)
Natural Formations.
Natural Formations: String of Pearls
CR3BP Analysis of Phase Space Eigenstructure
Near Halo Orbit

Floquet Decomposition of $\Phi(t, 0)$:

$$\Phi(t, 0) = P(t) e^{\lambda t} P(0) = \left\{ P(t) S \right\} e^{\lambda t} \left\{ P(0) S \right\}^{-1}$$

Floquet Modal Matrix:

$$E(t) = P(t) S = \Phi(t, 0) E(0) e^{-\lambda t}$$

Solution to Variational Eqn. in terms of Floquet Modes:

$$\delta \vec{x}(t) = \sum_{j=1}^{6} \delta \vec{x}_j(t) = \sum_{j=1}^{6} c_j(t) \vec{e}_j(t) = E(t) \vec{c}$$
Natural Formations:
Quasi-Periodic Relative Orbits → 2-D Torus
Floquet Controller
(Remove Unstable + 2 of the 4 Center Modes)

Find $\Delta \vec{v}$ that removes undesired response modes:

$$\sum_{j=1}^{6} \delta \vec{x}_j + \begin{bmatrix} 0_3 \\ I_3 \end{bmatrix} \Delta \vec{v} = \sum_{j=2,3,4} \left( 1 + \alpha_j \right) \delta \vec{x}_j$$

Remove Modes 1, 3, and 4:

$$\begin{bmatrix} \bar{\alpha} \\ \Delta \vec{v} \end{bmatrix} = \begin{bmatrix} \delta \vec{x}_{2,2} & \delta \vec{x}_{5,2} & \delta \vec{x}_{6,2} & 0_3 \\ \delta \vec{x}_{2,\bar{v}} & \delta \vec{x}_{5,\bar{v}} & \delta \vec{x}_{6,\bar{v}} & -I_3 \end{bmatrix}^{-1} \left( \delta \vec{x}_1 + \delta \vec{x}_3 + \delta \vec{x}_4 \right)$$

Remove Modes 1, 5, and 6:

$$\begin{bmatrix} \bar{\alpha} \\ \Delta \vec{v} \end{bmatrix} = \begin{bmatrix} \delta \vec{x}_{2,2} & \delta \vec{x}_{3,2} & \delta \vec{x}_{4,2} & 0_3 \\ \delta \vec{x}_{2,\bar{v}} & \delta \vec{x}_{3,\bar{v}} & \delta \vec{x}_{4,\bar{v}} & -I_3 \end{bmatrix}^{-1} \left( \delta \vec{x}_1 + \delta \vec{x}_5 + \delta \vec{x}_6 \right)$$
Deployment into Torus
(Remove Modes 1, 5, and 6)

\[ \vec{r}(0) = \begin{bmatrix} 50 & 0 & 0 \end{bmatrix} \text{ m} \]

\[ \dot{\vec{r}}(0) = \begin{bmatrix} 1 & -1 & 1 \end{bmatrix} \text{ m/sec} \]
Deployment into Natural Orbits
(Remove Modes 1, 3, and 4)

\[ \vec{r}(0) = [r_0 \ 0 \ 0] \text{ m} \]

\[ \vec{v}(0) = [1 \ -1 \ 1] \text{ m/sec} \]
Natural Formations:
Nearly Periodic Relative Motion

10 Revolutions = 1,800 days
Evolution of Nearly Vertical Orbits Along the $yz$-Plane
Natural Formations: Slowly Expanding Vertical Orbits

100 Revolutions = 18,000 days

\[ \vec{r}(0) \quad \vec{r}(t_f) \]

Origin = Chief S/C
Non-Natural Formations

- Fixed Relative Distance and Orientation
  - Fixed in Inertial Frame
  - Fixed in Rotating Frame

- Spherical Configurations (Inertial or RLP)
  - Fixed Relative Distance, Free Orientation
  - Fixed Relative Distance & Rotation Rate

- Aspherical Configurations (Position & Rates)
  - Parabolic
  - Others
Formations Fixed in the Inertial Frame

\[ \overline{\rho} = \rho \hat{Y} \]
Formations Fixed in the Rotating Frame

\[ \vec{\rho} = \rho \hat{y} \]
2-S/C Formation Model in the Sun-Earth-Moon System

Relative EOMs:
\[ \ddot{\mathbf{r}}(t) = \Delta f(\mathbf{r}(t)) + \mathbf{u}(t) \]

Ephemeris System = Sun+Earth+Moon Ephemeris + SRP
Discrete & Continuous Control
Linear Targeter

\[
\begin{bmatrix}
\delta r_{k+1}^- \\
\delta \nu_{k+1}^-
\end{bmatrix}
= \Phi(t_{k+1}, t_k)
\begin{bmatrix}
\delta r_{k+1}^+ \\
\delta \nu_{k+1}^+
\end{bmatrix}
= \begin{bmatrix} A_k & B_k \\ C_k & D_k \end{bmatrix}
\begin{bmatrix}
\delta r_{k+1}^- \\
\delta \nu_{k+1}^- + \Delta \nu_k
\end{bmatrix}
\]

\[
\Delta \nu_k = B_k^{-1} \left( \delta r_{k+1}^- - A_k \delta r_k \right) - \delta \nu_k^-
\]
Discrete Control: Linear Targeter

\[ \ddot{\rho} = (10 \text{ m}) \dot{Y} \]
\[ \dot{\rho} = 0 \]
Achievable Accuracy via Targeter Scheme

\[
\overline{r}^\circ = r\hat{Y}
\]
Continuous Control: LQR vs. Input Feedback Linearization

- LQR for **Time-Varying Nominal Motions**

\[
\begin{align*}
\dot{x}(t) &= \begin{bmatrix} \dot{r} \\ \dot{\bar{r}} \end{bmatrix}^T = f(t, x(t), \bar{u}(t)) \\
\dot{P} &= -A^T(t)P(t) - P(t)A(t) + P(t)B(t)R^{-1}B^T(t)P(t) - Q \rightarrow P(t_f) = 0
\end{align*}
\]

\[
\begin{align*}
\bar{u}(t) &= \bar{u}^\circ(t) + \left\{ -R^{-1}B^T(t)(\bar{x}(t) - \bar{x}^\circ(t)) \right\} \\
&= \text{Optimal Control, Relative to Nominal, from LQR}
\end{align*}
\]

- Input Feedback Linearization (IFL)

\[
\begin{align*}
\ddot{r}(t) &= \bar{F}(\bar{r}(t)) + \bar{u}(t) \\
\bar{u}(t) &= \bar{F}(\bar{r}(t)) + \bar{g}(\bar{r}(t), \dot{\bar{r}}(t)) \\
&= \text{Anihilate Natural Dynamics}
\end{align*}
\]
LQR Goals

\[
\min J = \frac{1}{2} \int_{t_0}^{t_f} \begin{bmatrix} \delta x(t) & Q \delta x(t) + \delta u_d(t) \end{bmatrix}^T \begin{bmatrix} 0 & R \delta u_d(t) \end{bmatrix} dt
\]

\[
\delta x(t) = x(t) - \bar{x}(t)
\]
\[
\delta u_d(t) = u_d(t) - \bar{u}_d(t)
\]

\[
Q = \text{diag}(10^{12}, 10^{12}, 10^{12}, 10^5, 10^5)
\]

\[
R = \text{diag}(1, 1, 1)
\]
LQR Process

Step 1: Evaluate the Jacobian matrix, at time $t_i$, associated with nominal path

$\rightarrow A(t_i)$ evaluated on $\bar{x}^o(t_i)$

Step 2: Numerically integrate the differential Riccatti Equation backwards in time

from $t_i \rightarrow t_{i-1}$, subject to $P(t_N) = 0$.

$$\dot{P}(t_i) = -A^T(t_i)P(t_i) - P(t_i)A(t_i) + P(t_i)BR^{-1}B^TP(t_i) - Q$$

Step 3: Compute and store the controller gain matrix

$$K(t_i) = R^{-1}B^TP(t_i)$$

Step 4: Repeat steps 1-3 until $t_{i-1} = t_0$

Step 5: Numerically integrate the perturbed trajectory forward in time

from $t_0 \rightarrow t_N$, subject to $\bar{x}(t_0) = \bar{x}_0$.

$\rightarrow$ Recall $K(t)$ from stored data

$\rightarrow$ The new integration step size is defined by the sampled gain data

$\rightarrow$ Substitute $\bar{x}^o(t)$ into EOMs and solve for $\bar{u}_d^o(t)$

$\rightarrow \bar{u}_d(t) = \bar{u}_d^o(t) - K(t)(\bar{x}(t) - \bar{x}^o(t))$

$\rightarrow$ Apply the computed control input to the perturbed EOMs
Step 1: Define, analytically, the desired response characteristics

\[ \ddot{\bar{r}} = \ddot{\bar{r}}^o - 2\omega_n (\dot{\bar{r}} - \dot{\bar{r}}^o) - \omega_n^2 (\bar{r} - \bar{r}^o) = g(\bar{r}, \dot{\bar{r}}) \]

Step 2: Begin numerical integration of perturbed path

Step 3: At each point in time, compute and apply the control input necessary to achieve the desired response characteristics:

\[ \bar{u}_d(t) = -f(\bar{r}^{P_2D}, \dot{\bar{r}}^{P_2D}) + g(\bar{r}, \dot{\bar{r}}) \]

\[ \text{Annihilate Natural Dynamics} \]
\[ \text{Reflects desired response} \]
LQR vs. IFL Comparison

\[ \rho = 5000 \text{ km}, \ \xi = 90^\circ, \ \beta = 0^\circ \]

\[ \delta x(0) = \begin{bmatrix} 7 \text{ km} & -5 \text{ km} & 3.5 \text{ km} & 1 \text{ mps} & -1 \text{ mps} & 1 \text{ mps} \end{bmatrix}^T \]

Dynamic Response

Control Acceleration History

Dynamic Response Modeled in the CR3BP
Nominal State Fixed in the Rotating Frame
Output Feedback Linearization
(Radial Distance Control)

Formation Dynamics
\[ \ddot{r} = \Delta f (\vec{r}) + \vec{u}(t) \quad \rightarrow \quad \text{Generalized Relative EOMs} \]
\[ y = l(\vec{r}) \quad \rightarrow \quad \text{Measured Output} \]

Measured Output Response (Radial Distance)
\[
\ddot{y} = \frac{d^2 l}{dt^2} = p(\vec{r}, \dot{\vec{r}}) + q(\vec{r}, \dot{\vec{r}}) \dot{\vec{r}}^T \vec{r} = g(\vec{r}, \dot{\vec{r}})
\]

Scalar Nonlinear Constraint on Control Inputs
\[
h(\vec{r}(t), \dot{\vec{r}}(t)) - \vec{u}(t)^T \vec{r}(t) = 0
\]
Output Feedback Linearization (OFL)
(Radial Distance Control in the Ephemeris Model)

\[ y = l(\vec{r}, \dot{\vec{r}}) \]

<table>
<thead>
<tr>
<th>Control Law</th>
</tr>
</thead>
<tbody>
<tr>
<td>[ \vec{u}(t) = \frac{h(\vec{r}, \dot{\vec{r}})}{r} \hat{r} ] Geometric Approach: Radial inputs only</td>
</tr>
<tr>
<td>[ \vec{u}(t) = \left{ \frac{g(\vec{r}, \dot{\vec{r}})}{r} - \frac{\dot{\vec{r}}^T \dot{\vec{r}}}{r^2} \right} \vec{r} + \left( \frac{\dot{r}}{r} \right) \dot{r} - \Delta \vec{f}(\vec{r}) ]</td>
</tr>
<tr>
<td>[ \vec{u}(t) = \left{ \frac{1}{2} \frac{g(\vec{r}, \dot{\vec{r}})}{r^2} - \frac{\dot{\vec{r}}^T \dot{\vec{r}}}{r^2} \right} \vec{r} - \Delta \vec{f}(\vec{r}) ]</td>
</tr>
<tr>
<td>[ \vec{u}(t) = \left{ -rg(\vec{r}, \dot{\vec{r}}) - \frac{\dot{\vec{r}}^T \dot{\vec{r}}}{r^2} \right} \vec{r} + 3 \left( \frac{\dot{r}}{r} \right) \dot{r} - \Delta \vec{f}(\vec{r}) ]</td>
</tr>
</tbody>
</table>

- Critically damped output response achieved in all cases
- Total ΔV can vary significantly for these four controllers
OFL Control of Spherical Formations in the Ephemeris Model

\[ \vec{r}(0) = \begin{bmatrix} 12 & -5 & 3 \end{bmatrix} \text{ km} \quad \dot{\vec{r}}(0) = \begin{bmatrix} 1 & -1 & 1 \end{bmatrix} \text{ m/sec} \]

Relative Dynamics as Observed in the Inertial Frame
OFL Controlled Response of Deputy S/C  
Radial Distance + Rotation Rate Tracking

Radial Error Response (Critically Damped):
\[ \delta\ddot{r} = -2\omega_n \delta\dot{r} - \omega_n^2 \delta r \]

Rotation Rate Error Response (Exponential Decay):
\[ \delta\dot{\theta} = \delta\dot{\theta}_0 e^{-t/T} \rightarrow \delta\ddot{\theta} = -\left( \frac{\delta\dot{\theta}_0}{T} \right) e^{-t/T} = -\delta\dot{\theta}/T \]

\[ \ddot{\theta} = g_{\theta}(t) = \ddot{\theta}_n - \frac{(\dot{\theta} - \dot{\theta}_n)}{T} = \ddot{\theta}_n - k \omega_n (\dot{\theta} - \dot{\theta}_n) \]
Equations of Motion in the Relative Rotating Frame

\[ \ddot{r} - r \dot{\theta}^2 = f_r + u_r \]
\[ f_r = \Delta \vec{f} \cdot \hat{r}, \quad f_\theta = \Delta \vec{f} \cdot \hat{\theta}, \quad f_h = \Delta \vec{f} \cdot \hat{h} \]

\[ r \ddot{\theta} + 2r \dot{\theta} = f_\theta + u_\theta \]
\[ u_r = \vec{u} \cdot \hat{r}, \quad u_\theta = \vec{u} \cdot \hat{\theta}, \quad u_h = \vec{u} \cdot \hat{h} \]

Rearrange to isolate the radial and rotational accelerations:

\[ \ddot{r} = f_r + u_r + r \dot{\theta}^2 = g_r(t) \]

\[ r \ddot{\theta} = f_\theta + u_\theta - 2r \dot{\theta} = rg_\theta(t) \]

Solve for the Control Inputs:

\[ u_r(t) = g_r(t) - f_r - r \dot{\theta}^2 \]
\[ u_\theta(t) = rg_\theta(t) - f_\theta + 2r \dot{\theta} \]
\[ u_h(t) = -f_h \text{ (constraint) } \]
OFL Control of Spherical Formations
Radial Dist. + Rotation Rate

Quadratic Growth in Cost w/ Rotation Rate

Linear Growth in Cost w/ Radial Distance
Inertially Fixed Formations in the Ephemeris Model

\[ \mathbf{e} \quad \text{inertially fixed formation pointing vector (focal line)} \]

- 500 m
- 100 km
- Chief S/C
Nominal Formation Keeping Cost
(Configurations Fixed in the RLP Frame)

\[ \Delta V = \int_{0}^{180 \text{ days}} \sqrt{\mathbf{u}^\circ(t) \cdot \mathbf{u}^\circ(t)} \, dt \]

\( \dot{r}^\circ = 5000 \text{ km} \)

\( A_z = 0.2 \times 10^6 \text{ km} \)

\( A_z = 0.7 \times 10^6 \text{ km} \)

\( A_z = 1.2 \times 10^6 \text{ km} \)
Max./Min. Cost Formations
(Configurations Fixed in the RLP Frame)

Maximum Cost Formation

Minimum Cost Formations

Nominal Relative Dynamics in the Synodic Rotating Frame
Formation Keeping Cost Variation
Along the SEM $L_1$ and $L_2$ Halo Families
(Configurations Fixed in the RLP Frame)
Conclusions

- Continuous Control in the Ephemeris Model:
  - Non-Natural Formations
    - LQR/IFL $\rightarrow$ essentially identical responses & control inputs
    - IFL appears to have some advantages over LQR in this case
    - OFL $\rightarrow$ spherical configurations + unnatural rates
    - Low acceleration levels $\rightarrow$ Implementation Issues

- Discrete Control of Non-Natural Formations
  - Targeter Approach
    - Small relative separations $\rightarrow$ Good accuracy
    - Large relative separations $\rightarrow$ Require nearly continuous control
    - Extremely Small $\Delta V$'s ($10^{-5}$ m/sec)

- Natural Formations
  - Nearly periodic & quasi-periodic formations in the RLP frame
  - Floquet controller: numerically ID solutions + stable manifolds
Some Examples from Simulations

- A simple formation about the Sun-Earth L1
  - Using CRTB based on L1 dynamics
  - Errors associated with perturbations

- A more complex Fizeau-type interferometer fizeau interferometer.
  - Composed of 30 small spacecraft at L2
  - Formation maintenance, rotation, and slewing
A State Space Model

• A common approximation in research of this type of orbit models the dynamics using CRTB approximations.

• The Linearized Equations of Motion for a S/C Close to the Libration Point Are Calculated at the Respective Libration Point.

  • Linearized Eq. Of Motion Based on Inertial X, Y, Z Using

    \[ \begin{align*}
    \dot{x} &= 2n \dot{y} = U_{xx} x, \\
    \dot{y} + 2n \dot{x} &= U_{yy} y, \\
    \dot{z} &= U_{zz} z
    \end{align*} \]

• Pseudopotential:

    \[ U_{xx} = \frac{\partial^2 U}{\partial X^2}, \quad U_{yy} = \frac{\partial^2 U}{\partial Y^2}, \quad U_{zz} = \frac{\partial^2 U}{\partial Z^2}. \]

\[
\begin{align*}
\begin{bmatrix}
x' \\
y' \\
z'
\end{bmatrix} &= A^i \begin{bmatrix}
x \\
y \\
z
\end{bmatrix}, \\
A^i &= \begin{bmatrix}
0 & 0 & 0 & 1 & 0 & 0 \\
0 & 0 & 0 & 0 & 1 & 0 \\
0 & 0 & 0 & 0 & 0 & 1 \\
U_{xx} & 0 & 0 & 0 & 2n & 0 \\
0 & U_{yy} & 0 & -2n & 0 & 0 \\
0 & 0 & U_{zz} & 0 & 0 & 0
\end{bmatrix}
\end{align*}
\]
Periodic Reference Orbit

\[
x = -A_x \sin(\omega_{xy} t + \varphi_{xy}) \quad \dot{x} = -A_x \omega_{xy} \cos(\omega_{xy} t + \varphi_{xy}) \\
y = -A_y \cos(\omega_{xy} t + \varphi_{xy}) \quad \dot{y} = A_y \omega_{xy} \sin(\omega_{xy} t + \varphi_{xy}) \\
z = A_z \sin(\omega_z t + \varphi_z) \quad \dot{z} = A_z \omega_z \cos(\omega_z t + \varphi_z)
\]

\(A = \text{Amplitude}\)
\(\omega = \text{frequency}\)
\(\phi = \text{Phase angle}\)
Centralized LQR Design

State Dynamics \[ \dot{x}^j = A^j x^j + B^j u^j + a_{solpres} + a_{FB} \]

Performance Index to Minimize \[ J = \frac{1}{2} \int_0^\infty \{(x - x^R)^T Q (x - x^R) + u^T R u\} dt \]

Control \[ u^j = - [R^j]^{-1} [B^j]^T S x \]

Algebraic Riccati Eq. \[ S (B R^{-1} B^T) S - SA - A^T S - Q = 0 \]

B Maps Control Input From Control Space to State Space
\[ B^j = \begin{bmatrix} O_{3\times3} \\ I_3 \end{bmatrix} \]

Q Is Weight of State Error
\[ Q^j = \begin{bmatrix} \frac{1}{p} I_3 & O_{3\times3} \\ O_{3\times3} & \frac{1}{q} I_3 \end{bmatrix} \]

R Is Weight of Control
\[ R^j = \frac{1}{r} I_3 \]
Centralized LQR Design
Disturbance Accommodation Model

- The $A$ Matrix Does Not Include the Perturbation Disturbances nor Exactly Equal the Reference

- Disturbance Accommodation Model Allows the States to Have Non-zero Variations From the Reference in Response to the Perturbations Without Inducing Additional Control Effort

- The Periodic Disturbances Are Determined by Calculating the Power Spectral Density of the Optimal Control [Hoff93] To Find a Suitable Set of Frequencies.

**Unperturbed**

$\omega_{x,y,z} = 4.26106e-7 \text{ rad/s}$

**Perturbed**

$\omega_x = 1.5979e-7 \text{ rad/s}$

$2.6632e-6 \text{ rad/s}$

$\omega_y = 2.6632e-6 \text{ rad/s}$

$\omega_z = 2.6632e-6 \text{ rad/s}$
Disturbance Accommodation Model

Without Disturbance Accommodation

With Disturbance Accommodation

Disturbance accommodation state is out of phase with state error, absorbing unnecessary control effort.
Motion of Formation Flyer With Respect to Reference Spacecraft, in Local (S/C-1) Coordinates
ΔV Maintenance in Libration Orbit Formation
Stellar Imager Concept

(Using conceptual distances and control requirements to analyze formation possibilities)

- Stellar Imager (SI) concept for a space-based, UV-optical interferometer, proposed by Carpenter and Schrijver at NASA / GSFC (Magnetic fields, Stellar structures and dynamics)

- 500-meter diameter Fizeau-type interferometer composed of 30 small drone satellites

- Hub satellite lies halfway between the surface of a sphere containing the drones and the sphere origin.

- Focal lengths of both 0.5 km and 4 km, with radius of the sphere either 1 km or 8 km.

- L2 Libration orbit to meet science, spacecraft, and environmental conditions
• Three different scenarios make up the SI formation control problem; maintaining the Lissajous orbit, slewing the formation, and reconfiguring.
• Using a LQR with position updates, the hub maintains an orbit while drones maintain a geometric formation.

The magenta circles represent drones at the beginning of the simulation, and the red circles represent drones at the end of the simulation. The hub is the black asterisk at the origin.

**SI Slewing Geometry**

**Formation ΔV Cost per slewing maneuver**

<table>
<thead>
<tr>
<th>Focal Length (km)</th>
<th>Slew Angle (deg)</th>
<th>Hub (m/s)</th>
<th>Drone 2 (m/s)</th>
<th>Drone 31 (m/s)</th>
</tr>
</thead>
<tbody>
<tr>
<td>0.5</td>
<td>30</td>
<td>1.0705</td>
<td>0.8571</td>
<td>0.8507</td>
</tr>
<tr>
<td>0.5</td>
<td>90</td>
<td>1.3555</td>
<td>0.9395</td>
<td>0.9587</td>
</tr>
<tr>
<td>1</td>
<td>30</td>
<td>1.2688</td>
<td>1.1389</td>
<td>1.1315</td>
</tr>
<tr>
<td>1</td>
<td>90</td>
<td>1.3870</td>
<td>2.1907</td>
<td>2.1932</td>
</tr>
</tbody>
</table>
Stellar Imager Mission Study Example Requirements

- Maintain an orbit about the Sun-Earth L2 co-linear point
- Slew and rotate the Fizeau system about the sky, movement of few km and attitude adjustments of up to 180deg
- While imaging ‘drones’ must maintain position
  - ~3 nanometers radially from Hub
  - ~0.5 millimeters along the spherical surface
- Accuracy of pointing is 5 milli-arcseconds, rotation about axis < 10 deg

3-Tiered Formation Control Effort:

Coarse - RF ranging, star trackers, and thrusters ~ centimeters
Intermediate - Laser ranging and micro-N thrusters control ~ 50 microns
Precision - Optics adjusted, phase diversity, wave front. ~ nanometers
This analysis uses high fidelity dynamics based on a software named Generator that Purdue University has developed along with GSFC.

- Creates realistic lissajous orbits as compared to CRTB motion.
- Uses sun, Earth, lunar, planetary ephemeris data.
- Generator accounts for eccentricity and solar radiation pressure.
- Lissajous orbit is more an accurate reference orbit.
- Numerically computes and outputs the linearized dynamics matrix, $A$, for a single satellite at each epoch.
- Data used onboard for autonomous
  - computation by simple uploads or
  - onboard computation as a background task of the 36 matrix elements and
  - the state vector.

- Origin in figure is Earth
- Solar rotating coordinates
State Space Controller Development, LQR Design

- Rotating Coordinates of SI: \( X = X_0 + x, \ Y = Y_0 + y, \ Z = Z_0 + z \)
  where the open-loop linearized EOM about L2 can be expressed as \( \dot{x} = Ax \)
  and the \( A \) matrix is the Generator Output \( x = [x \ y \ z \ \dot{x} \ \dot{y} \ \dot{z}]^T \)

- The STM is created from the dynamics partials output from the Generator and assumes to be constant over an analysis time period
  \( \Phi(t - t_0) = e^{A(t-t_0)} = \sum_{k=0}^{\infty} \frac{A^k(t-t_0)^k}{k!} \)

State Dynamics and Error \( \Rightarrow \dot{x} = Ax + Bu \)
  \( \tilde{x}(t) = x(t) - x_{ref}(t) \)

Performance Index to Minimize \( \Rightarrow J = \int_{t_0}^{t_f} \{ \tilde{x}^T(\tau)W\tilde{x}(\tau) + u^T(\tau)Vu(\tau) \} d\tau \)

Control and Closed Loop Dynamics \( \Rightarrow u = -K(t)\tilde{x} \)
  \( \dot{\tilde{x}} = (A - BK(t))\tilde{x} \)

Algebraic Riccati Eq. \( \Rightarrow S(BR^{-1}B^T)S - SA - A^T S - Q = 0 \)

time invariant system
• Expanding for the SI collector (hub) and mirrors (drones) yields a controller
\[
\dot{\mathbf{x}}_2 = A\mathbf{x}_2 + Bu_2 - Bu_1
\]

• Redefine A and B such that

\[
A = \begin{bmatrix}
A_1 & A_2 & \cdots & A_j
\end{bmatrix},
\quad
B = \begin{bmatrix}
B_1 & B_i & \cdots & B_i \\
-B_1 & B_i & \cdots & B_i \\
\vdots & \ddots & \ddots & \vdots \\
-B_1 & \cdots & -B_1 & B_j
\end{bmatrix}
\]

- \(B^j\) Maps Control Input From Control Space to State Space

- \(W\) is Weight of State Error

\[
W = \begin{bmatrix}
\frac{1}{p}I_3 & O_{3\times3} \\
O_{3\times3} & \frac{1}{q}I_3
\end{bmatrix}
\]

- \(V\) is Weight of Control

\[
V^j = \frac{1}{r}I_3
\]
Simplified extended Kalman Filter

- Using dynamics described and linear measurements augmented by zero-mean white Gaussian process and measurement noise

- Discretized state dynamics for the filter are: \( \tilde{x}_{k+1} = A_d \tilde{x}_k + B_d u_k + w \)
  where \( w \) is the random process noise

- The non-linear measurement model is \( y_k = m(\tilde{x}_k) + \nu \)
  and the covariances of process and measurement noise are
  \[ E[ww^T] = Q \]
  \[ E[\nu\nu^T] = R \]

- Hub measurements are range(r) and azimuth(az) / elevation(el) angles from Earth Drone measurements are r, az, el from drone to hub

\[
\begin{align*}
r &= \sqrt{x^2 + y^2 + z^2} \\
el &= \sin^{-1}\left(\frac{-z}{r}\right), \text{ and } az = \sin^{-1}\left(\frac{x}{r \cos(el)}\right)
\end{align*}
\]
Simulation Matrix Initial Values

- Continuous state weighting and control chosen as

\[
V = \begin{bmatrix} 1 & 1 \end{bmatrix}, \\
W = \begin{bmatrix} \text{1e6} & \text{1e6} \end{bmatrix}
\]

- The process and measurement noise Covariance (hub and drone) are

\[
Q_c = \begin{bmatrix} 0 & 0 \\ 0 & 1e-6 \end{bmatrix}, \\
R = \begin{bmatrix} 0.1^2 \\ \left(\frac{0.3}{1500000}\right)^2 \end{bmatrix}, \\
R = \begin{bmatrix} 0.0001^2 \\ \left(\frac{0.0003}{0.5}\right)^2 \end{bmatrix}
\]

- Initial covariances

\[
P_1(+) = \begin{bmatrix} 1 & 1 \\ 86.4^2 & 86.4^2 \end{bmatrix}, \\
P_1(+) = \begin{bmatrix} .001^2 \\ .0864^2 \end{bmatrix}
\]

\[
P_1(+) = \begin{bmatrix} .001^2 \\ .0864^2 \\ 107 \end{bmatrix}
\]
Results – Libration Orbit Maintenance

- Only Hub spacecraft was simulated for maintenance
- Tracking errors for 1 year: Position and Velocity
- Steady State errors of 250 meters and .075 cm/s
Results – Libration Orbit Maintenance

- Estimation errors for 12 simulations for 1 year: Position and Velocity
- Estimation errors of 250 meters and $2 \times 10^{-4}$ m/s in each component
Results – Formation Slewing

- Length of simulation is 24 hours
- Maneuver frequency is 1 per minute
- Using a constant $A$ from day-2 of the previous simulation
- Tuning parameters are same but strength of process noise is

$$Q_r = \begin{bmatrix} 0 & 0 & 0 \\ 0 & 0 & 1e^{-24} \\ 0 & 1e^{-24} & 1e^{-24} \end{bmatrix}$$

Formation Slewing:
90° simulation shown

Purple - Begin
Red - End
Results – Formation Slewing

- Tracking errors for 24 hours: Position and Velocity
- Steady State errors of 50 meters - hub, 3 meters - drone and 5 millimeters/sec – hub, and 1 millimeter/s - drone

HUB

DRONE

Represents both 0.5 and 4 km focal lengths
Results – Formation Slewing

- Estimation errors; 12 simulations for 24 hours: position and velocity
- Estimation 3σ errors of ~50 km and ~1 millimeter/s for all scenarios

Hub estimation 0.5 km separation / 90 deg slew

Drone estimation 0.5 km separation / 90 deg slew
# Results – Formation Slewing

## Formation Slewing Average ΔVs (12 simulations)

<table>
<thead>
<tr>
<th>Focal Length (km)</th>
<th>Slew Angle (deg)</th>
<th>Hub (m/s)</th>
<th>Drone 2 (m/s)</th>
<th>Drone 31 (m/s)</th>
</tr>
</thead>
<tbody>
<tr>
<td>0.5</td>
<td>30</td>
<td>1.0705</td>
<td>0.8271</td>
<td>0.8307</td>
</tr>
<tr>
<td>0.5</td>
<td>90</td>
<td>1.1355</td>
<td>0.9395</td>
<td>0.9587</td>
</tr>
<tr>
<td>4</td>
<td>30</td>
<td>1.2688</td>
<td>1.1189</td>
<td>1.1315</td>
</tr>
<tr>
<td>4</td>
<td>90</td>
<td>1.8570</td>
<td>2.1907</td>
<td>2.1932</td>
</tr>
</tbody>
</table>

## Formation Slewing Average Propellant Mass

<table>
<thead>
<tr>
<th>Focal Length (km)</th>
<th>Slew Angle (deg)</th>
<th>Hub mass-prop (g)</th>
<th>Drone 2 mass-prop (g)</th>
<th>Drone 31 mass-prop (g)</th>
</tr>
</thead>
<tbody>
<tr>
<td>0.5</td>
<td>30</td>
<td>6.0018</td>
<td>0.8431</td>
<td>0.8468</td>
</tr>
<tr>
<td>0.5</td>
<td>90</td>
<td>6.3662</td>
<td>0.9577</td>
<td>0.9773</td>
</tr>
<tr>
<td>4</td>
<td>30</td>
<td>7.1135</td>
<td>1.1406</td>
<td>1.1534</td>
</tr>
<tr>
<td>4</td>
<td>90</td>
<td>10.4112</td>
<td>2.2331</td>
<td>2.2357</td>
</tr>
</tbody>
</table>
# Results – Formation Slewing

## Formation Slewing Average ΔVs (without noise)

<table>
<thead>
<tr>
<th>Focal Length (km)</th>
<th>Slew Angle (deg)</th>
<th>Hub (m/s)</th>
<th>Drone 2 (m/s)</th>
<th>Drone 31 (m/s)</th>
</tr>
</thead>
<tbody>
<tr>
<td>0.5</td>
<td>30</td>
<td>0.0504</td>
<td>0.0853</td>
<td>0.0998</td>
</tr>
<tr>
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<td>90</td>
<td>0.1581</td>
<td>0.2150</td>
<td>0.2315</td>
</tr>
<tr>
<td>4</td>
<td>30</td>
<td>0.4420</td>
<td>0.5896</td>
<td>0.6441</td>
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<tr>
<td>4</td>
<td>90</td>
<td>1.3945</td>
<td>1.9446</td>
<td>1.9469</td>
</tr>
</tbody>
</table>

## Formation Slewing Average Propellant Mass (without noise)

<table>
<thead>
<tr>
<th>Focal Length (km)</th>
<th>Slew Angle (deg)</th>
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<th>Drone 2 mass-prop (g)</th>
<th>Drone 31 mass-prop (g)</th>
</tr>
</thead>
<tbody>
<tr>
<td>0.5</td>
<td>30</td>
<td>0.2826</td>
<td>0.0870</td>
<td>0.1017</td>
</tr>
<tr>
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</tr>
<tr>
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<td>0.6010</td>
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<tr>
<td>4</td>
<td>90</td>
<td>7.8182</td>
<td>1.9822</td>
<td>1.9846</td>
</tr>
</tbody>
</table>
Results – Formation Reorientation

- Rotation about the line of sight
- Length of simulation is 24 hours
- Maneuver frequency is 1 per minute
- Using a constant $A$ from day-2 of the previous simulation
- Tuning parameters are same as slewing
- Reorientation of 4 drones $90^\circ$
Results – Formation Reorientation

- Tracking errors: Position and Velocity
- Steady State errors of 40 meters - hub, 4 meters - drone and 8 millimeters/sec – hub, and 1.5 millimeter/s - drone

**HUB**

**DRONE**
Results – Formation Reorientation

- The steady-state estimation $3\sigma$ values are: $x \sim 30$ meters, $y$ and $z \sim 50$ meters
- The steady-state estimation $3\sigma$ velocity values are about 1 millimeter per second.
- For any drone and either focal length, the steady-state $3\sigma$ position values are less than 0.1 meters, and the steady-state velocity $3\sigma$ values are less than $1e-6$ meters per second.

### Formation Reorientation Average $\Delta V$s

<table>
<thead>
<tr>
<th>Focal Length (km)</th>
<th>Slew Angle (deg)</th>
<th>Hub (m/s)</th>
<th>Drone 2 (m/s)</th>
<th>Drone 31 (m/s)</th>
</tr>
</thead>
<tbody>
<tr>
<td>0.5</td>
<td>90</td>
<td>1.0126</td>
<td>0.8421</td>
<td>0.8095</td>
</tr>
<tr>
<td>4</td>
<td>90</td>
<td>1.0133</td>
<td>0.8496</td>
<td>0.8190</td>
</tr>
</tbody>
</table>

### Formation Reorientation Average Propellant Mass

<table>
<thead>
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<th>Focal Length (km)</th>
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<th>Drone 31 mass-prop (g)</th>
</tr>
</thead>
<tbody>
<tr>
<td>0.5</td>
<td>90</td>
<td>5.6771</td>
<td>0.8584</td>
<td>0.8252</td>
</tr>
<tr>
<td>4</td>
<td>90</td>
<td>5.6811</td>
<td>0.8661</td>
<td>0.8349</td>
</tr>
</tbody>
</table>
## Results – Formation Reorientation

### Formation Reorientation Average ΔVs (without noise)

<table>
<thead>
<tr>
<th>Focal Length (km)</th>
<th>Slew Angle (deg)</th>
<th>Hub (m/s)</th>
<th>Drone 2 (m/s)</th>
<th>Drone 31 (m/s)</th>
</tr>
</thead>
<tbody>
<tr>
<td>0.5</td>
<td>90</td>
<td>0.0408</td>
<td>0.1529</td>
<td>0.1496</td>
</tr>
<tr>
<td>4</td>
<td>90</td>
<td>0.0408</td>
<td>0.1529</td>
<td>0.1495</td>
</tr>
</tbody>
</table>

### Formation Reorientation Average Propellant Mass (without noise)

<table>
<thead>
<tr>
<th>Focal Length (km)</th>
<th>Slew Angle (deg)</th>
<th>Hub mass-prop (g)</th>
<th>Drone 2 mass-prop (g)</th>
<th>Drone 31 mass-prop (g)</th>
</tr>
</thead>
<tbody>
<tr>
<td>0.5</td>
<td>90</td>
<td>0.2287</td>
<td>0.1623</td>
<td>0.1525</td>
</tr>
<tr>
<td>4</td>
<td>90</td>
<td>0.2287</td>
<td>0.1623</td>
<td>0.1524</td>
</tr>
</tbody>
</table>
Summary
(Using example requirements and constraints)

- The control strategy and Kalman filter using higher fidelity dynamics provides satisfactory results.

- The hub satellite tracks to its reference orbit sufficiently for the SI mission requirements. The drone satellites, on the other hand, track to only within a few meters.

- Without noise, though, the drones track to within several micrometers.

- Improvements for first tier control scheme (centimeter control) for SI could be accomplished with better sensors to lessen the effect of the process and measurement noise.
Summary
(using example requirements and constraints)

- Tuning the controller and varying the maneuver intervals should provide additional savings as well. Future studies must integrate the attitude dynamics and control problem.

- The propellant mass and results provide a minimum design boundary for the SI mission. Additional propellant will be needed to perform all attitude maneuvers, tighter control requirement adjustments, and other mission functions.

- Other items that should be considered in the future include:
  - Non-ideal thrusters,
  - Collision avoidance,
  - System reliability and fault detection
  - Nonlinear control and estimation
  - Second and third control tiers and new control strategies and algorithms.
- A Distant Retrograde Formation

- Decentralized control
DRO Mission Metrics

➤ Earth-constellation distance: >50 Re (less interference) and < 100 Re (link margin).
  ▪ Closer than 100 Re would be desirable to improve the link margin requirement
  ▪ A retrograde orbit of <160 Re (10^6 km), for a stable orbit would be ok

➤ The density of "baselines" in the u-v plane should be uniformly distributed. Satellites randomly distributed on a sphere will produce this result.

➤ Formation diameter: ~50 km to achieve desired angular resolution

➤ The plan is to have up to 16 microsats, each with it's own "downlink".

➤ Satellites will be "approximately" 3-axis stabilized.

➤ Lower energy orbit insertion requirements are always appreciated.

➤ Eclipses should be avoided if possible.

➤ Defunct satellites should not "interfere" excessively with operational satellites.
Distant Retrograde Orbit (DRO)

Why DRO?

- Stable Orbit
- No Skp ΔV
- Not as distant as L1
- Mult. Transfers
- No Shadows?
- Good Environment

☑ Really a Lunar Periodic Orbit
☑ Classified as a Symmetric Doubly Asymptotic Orbit in the Restricted Three-Body Problem
DRO Formation Sphere

- Matlab generated sphere based on S03 algorithm
  - Uniform distribution of points on a unit sphere
  - 16 points at vertices represents spacecraft locations

3-D view

SIRA spacecraft

X-Y view

Y-Z view
Formation Control Analysis

How much $\Delta V$ to initialize, maintain, and resize?

<table>
<thead>
<tr>
<th>Phase</th>
<th>Max [m/s]</th>
<th>Mean [m/s]</th>
<th>Min [m/s]</th>
<th>Std [m/s]</th>
</tr>
</thead>
</table>
a     | 0.635     | 0.607      | 0.592     | 0.014     |
b     | 1.323     | 0.792      | 0.392     | 0.293     |
c     | 0.757     | 0.674      | 0.541     | 0.069     |
d     | 0.679     | 0.616      | 0.582     | 0.031     |
e     | 1.201     | 0.721      | 0.367     | 0.263     |
f     | 0.679     | 0.608      | 0.503     | 0.056     |

Examples:

Initialize & maintain 2 yr:  
= 33 m/s

Initialize, Maintain 2yr,  
& four resizes:  
= 36 m/s

Phase Description
a) Init 25km sphere
b) Maintain 25km sphere (strict PD control) one month
c) Maintain 25km sphere (loose control) one month
d) Resize from 25km to 50km
e) Maintain 50km sphere (strict PD control) one month
f) Maintain 50km sphere (loose control) one month
Earth - Moon L4 Libration Orbit
an alternate orbit location
DRO Formation Control Analysis

- Earth/Moon L4 Libration Orbit
- Spacecraft controlled to maintain only relative separations
- Plots show formation position and drift (sphere represent 25km radius)
- Maneuver performed in most optimum direction based on controller output

Impulsive Maneuver of 16th s/c

Radial Distance from Center
Many nodes in a network can cooperate to behave as single virtual platform:

- Requires a fully connected network of nodes.
- Each node processes only its own measurements.
- Non-hierarchical means no leads or masters.
- No single points of failure means detected failures cause system to degrade gracefully.
- Basic problem previously investigated by Speyer.
- Based on LQG paradigm.
- Data transmission requirements are minimized.
References, etc.

1. NASA Web Sites, www.nasa.gov. Use the find it @ nasa search input for SEC, Origin, ESE, etc.
3. Fundamental of Astrodynamics, Bate, Muller, and White, Dover, Publications, 1971
5. An Introduction to the Mathematics and Methods of Astrodynamics, Battin
7. Orbital Mechanics, Chapter 8, Prussing and Conway
15. “Formation Flight near L1 and L2 in the Sun-Earth/Moon Ephemeris System including solar radiation pressure”, Marchand and Howell, paper AAS 03-596
Backup and other slides
DST/Numerical Comparisons

**Numerical Systems**
- Limited Set of Initial Conditions
- Perturbation Theory
- Single Trajectory
- Intuitive DC Process
- Operational

**Dynamical Systems**
- Qualitative Assessments
- Global Solutions
- Time Saver / Trust Results
- Robust
- Helps in choosing numerical methods
  (e.g., Hamiltonian => Symplectic Integration Schemes?)
Libration Point Trajectory Generation Process

Phase and Lissajous Utilities
Generate Lissajous of Interest

Compute Monodromy Matrix
And Eigenvalues/Eigenvectors
For Half Manifold of Interest

Globalize the Stable Manifold

Use Manifold Information for
a Differential Corrector Step
To Achieve Mission Constraints.

Lissajous
Monodromy
Manifold
Transfer

Patch Points
and Lissajous

Fixed Points and Stable
and Unstable Manifold
Approximations

1-D Manifold

Transfer Trajectory to
Earth Access Region
MAP Mission Design: DST Perspective

- MAP Manifold and Earth Access
- Manifold Generated Starting with Lissajous Orbit
- Swingby Numerical Propagation
- Trajectory Generated Starting with Manifold States
• JWST Manifold and Earth Access with Halo Orbit
• Manifold Generated Starting
• Swingby Numerical Propagation
• Trajectory Generated Starting with Manifold States

JWST DST Perspective
Two – Body Motion

- Motion of spacecraft in elliptical orbit
- Counter-clockwise
- \( x \) and \( y \) correspond to \( P \) and \( Q \) axes in \( PQW \) frame

Two angles are defined
- \( E \) is Eccentric Anomaly
- \( \theta \) is True Anomaly
- \( x \) and \( y \) coordinates are
  \[
  x = r \cos \theta \\
  y = r \sin \theta
  \]

In terms of Eccentric anomaly, \( E \)
- \[
  a \cos E = ae + x
  \]
- \[
  x = a (\cos E - e)
  \]
- From eqn. of ellipse: \( r = a(1 - e^2)/(1 + e \cos \theta) \)
- Leads to: \( r = a(1 - e \cos E) \)
- Can solve for \( y \): \( y = a(1 - e^2)^{1/2} \sin E \)

Coordinates of spacecraft in plane of motion are then
- \[
  r = a(1 - \cos E) = a(1 - e^2)/(1 + e \cos \theta)
  \]
- \[
  x = a(\cos E - e) = r \cos \theta
  \]
- \[
  y = a(1 - e^2)^{1/2} \sin E = r \sin \theta
  \]
- \[
  v_x = (\mu/p)^{1/2} [-\sin \theta]
  \]
- \[
  v_y = (\mu/p)^{1/2} [e + \cos \theta] \quad \text{where} \quad p = a(1 - e^2)^{139}
  \]
Two - Body Motion

Differentiate x, y, and r wrt time in terms of E to get
\[ \frac{dx}{dt} = -asinE \frac{de}{dt} \]
\[ \frac{dy}{dt} = a(1-e^2)^{1/2} \cos E \frac{dE}{dt} \]
\[ \frac{dr}{dt} = ae \sin E \frac{dE}{dt} \]

From the definition of angular momentum \( h = r \times \frac{dr}{dt} \) and expand to get
\( h = a^2(1-e^2)^{1/2} \frac{dE}{dt}[1-e \cos E] \) in direction perpendicular to orbit plane

Knowing \( h^2 = ma(1-e^2) \), equate the expressions and cancel common factor to yield
\[ (\mu)^{1/2}a^{3/2} = (1-e \cos E) \frac{dE}{dt} \]

Multiple across by \( dt \) and integrate from the perigee passage time yields
\[ n(t_0 - t_p) = E - e \sin E \]

Where \( n = (\mu)^{1/2}a^{3/2} \) is the mean motion

We can also compute the period: \( P = 2\pi a^{3/2} / (\mu)^{1/2} \) which can be associated with Kepler's 3rd law
Coordinate system transformation
Euler Angle Rotations

Suppose we rotate the x-y plane about the z-axis by an angle $\alpha$ and call the new coordinates $x'$, $y'$, $z'$.
Coordinate system transformation
Orbital to inertial coordinates

Inertial to/from Orbit plane
\[ \mathbf{r} = \mathbf{r}(x,y,z) \quad \mathbf{r} = \mathbf{r}(P,Q,W) \]

In orbit plane system, position \( \mathbf{r} = (rcosf)P + (rsinf)Q + (0)W \) where P, Q, W are unit vectors

\[
\begin{pmatrix}
x \\ y \\ z \\
\end{pmatrix} =
\begin{pmatrix}
P_x & Q_x & W_x \\ P_y & Q_y & W_y \\ P_z & Q_z & W_z \\
\end{pmatrix}
\begin{pmatrix}
rcosf \\ rsinf \\
0 \\
\end{pmatrix}
\]

\[
\begin{pmatrix}
x \\ y \\ z \\
\end{pmatrix} =
\begin{pmatrix}
\cos \omega & \sin \omega & 0 \\ -\sin \omega & \cos \omega & 0 \\ 0 & 0 & 1 \\
\end{pmatrix}
\begin{pmatrix}
1 & 0 & 0 \\ 0 & \cos i & -\sin i \\ 0 & -\sin i & -\cos i \\
\end{pmatrix}
\begin{pmatrix}
\cos \Omega & \sin \Omega & 0 \\ -\sin \Omega & \cos \Omega & 0 \\ 0 & 0 & 1 \\
\end{pmatrix}
\begin{pmatrix}
rcosf \\ rsinf \\ 0 \\
\end{pmatrix}
\]

\[
\begin{pmatrix}
x \\ y \\ z \\
\end{pmatrix} =
\begin{pmatrix}
\cos \omega & \sin \omega \cos \Omega & \sin \omega \sin \Omega \\ -\sin \omega & \cos \Omega & \cos \omega \sin \Omega \\ \sin \omega \sin i & -\sin \omega \cos i & \cos \omega \cos i \\
\end{pmatrix}
\begin{pmatrix}
-r\cos \Omega \sin \omega & -r\cos \Omega \cos \omega & r\sin \omega \\ -r\sin \Omega \sin \omega & -r\sin \Omega \cos \omega & r\cos \omega \\ -r\sin \omega \sin i & r\sin \omega \cos i & 0 \\
\end{pmatrix}
\begin{pmatrix}
rcosf \\ rsinf \\ 0 \\
\end{pmatrix}
\]

For Velocity transformation, \( v_x = (\mu/p)^{1/2} [-\sin \theta] \) and \( v_y = (\mu/p)^{1/2} [e + \cos \theta] \)

\[ \mathbf{r} = xi + yj + zk \] is the position and \( \mathbf{v} = v_x i + v_y j + v_z k \) is the velocity.
Principles Behind Decentralized Control

- **The state vector is decomposed into two partitions:**
  - **Depends only on the control.**
  - **Depends only on the local measurement data node-\(j\).**

- **A locally optimal Kalman filter operates on**

- **A globally optimal control is computed, using**
  - **And globally optimal data that is reconstructed locally using two additional vectors:**
  
  **A "data vector"**
  
  **is maintained locally in addition to the state.**

  **A transmission vector**
  
  **that minimizes the dimensions of the data which must be exchanged between nodes.**

- **Each node-\(j\) computes and transmits**
  - **To and receives**
  - **From all the other nodes in the network.**
The LQG Decentralized Controller Overview

FROM NETWORK

CONTROL VECTORS

TRANSMISSION VECTORS

CONTROL VECTOR

TO NETWORK & LOCAL PLANT

LOCAL SENSOR

TRANSMISSION VECTOR

TO NETWORK