PoET: Polarimeters for Energetic Transients

Mark McConnell (UNH)
Scott Barthelmy (GSFC)
Joe Hill (USRA)
# POET Science Team

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</tr>
</thead>
<tbody>
<tr>
<td>Matthew Baring</td>
<td>Bing Zhang</td>
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<td>Peter Bloser</td>
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<td>Yamazaki Ryo</td>
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<td>Alice Harding</td>
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<td>Wojtek Hajdas</td>
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<td>Don Kniffen</td>
<td>Nicholas Produit</td>
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<td>Toru Tamagawa</td>
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Overview

- GRB Polarimetry Science
- POET mission
  - GRAPE
  - LEP
- POET Performance
- What Now?
Quest for the holy grail

- X-ray polarimetry will be a valuable diagnostic of high magnetic field geometry and strong gravity.....

- One definitive astrophysical measurement (1978) at two energies
  - Weisskopf et al. 
  - $P=19.2\% \pm 1.0\%$
  - $\@ 156^\circ$

Weisskopf et al., 1978
Other Measurements

- Intercosmos (Tindo)
  - Solar Flares
- Rhessi (Coburn & Boggs)
  - GRB 021206
- BATSE Albedo Polarimetry System (Willis)
  - GRB 930131 P>35%
  - GRB 960924 P>50%
- INTEGRAL (2 groups)
  - 2σ result
  - 98 ± 33%
Current Status

- Recent instruments have not been optimised for polarimetry...
  - …or never launched
- Gazillion papers describing the importance
- Need a way to break the cycle
  - new techniques have lowered the technical barriers
Observed Prompt GRB Properties

- High variability: ~ms
- Prompt Spectrum:
  - Band Function: $\alpha \approx -1 \pm 1 \quad \beta \approx 2^{+1}_{-2}$
- Huge release of energy: $\sim 10^{51}$ erg
- Relativistic process to avoid pair-production opacity paradigm
- Achromatic steepening implies GRB jet

\[ E_0 = E_{\text{peak}} / (2 + \alpha) \]
\[ E^\alpha \exp \left( -E/E_0 \right) \]
\[ E_{\text{break}} = (\alpha - \beta) E_0 \]
Standard Fireball Model

- Explains the late afterglow observations well
- Debates for prompt emission on-going
  - Internal shock model solves the rapid variability problem
  - Energy has to be extracted from KE of shells
    - Low efficiency
    - Requires additional mechanisms

Dar et al., 2006

Synchrotron Emission
Cannon-ball model

Cannon balls ejected from central engine
Inverse Compton scattering of ambient photons
Unclear how the cannon balls would survive accretion over large dynamic range and Lorentz factors

Dar et al, 2006
GRB Unknowns

- Unknown Fire Ball content
  - Kinetic energy or magnetically dominated
- Unknown location of ‘where’ the prompt emission is produced
  - Internal Shocks - favored
  - External Shocks
- Unknown dissipation mechanism
  - Shocks
  - Magnetic reconnection
- Unknown radiation mechanism
  - Synchrotron
  - Comptonization
  - Etc
Motivation for POET

- What is the magnetic structure of the jets?
- What is the geometric structure of GRB jets?
- What is the prompt radiation mechanism of GRBs?

The theories on the GRB production mechanism can be constrained by different degrees of linear polarization (P):

- **P>~80%** Generally difficult to achieve within synchrotron emission models. Could be Compton scattering jet viewed from outside the edge of the jet

- **20%<P<60%** is predicted if synchrotron emission in an ordered B-field or as a result of viewing the burst from near the edge of the jet

- **Low degrees of polarization** can be expected from hydrodynamical models in which the random magnetic fields are generated in the shocks with an on-beam viewing geometry
POET - Proposed SMEX Mission

POET - POlarimeters for Energetic Transients

Institutional Responsibilities

University of New Hampshire
- PI: Mark McConnell
  - GRAPE Instrument

Universities Space Research Association
- Deputy PI: Joanne Hill
  - LEP Instrument

Goddard Space Flight Center
- Mission Scientist: Scott Barthelmy
  - Mission Operations Center (MOC)
  - POET Data Center (PDC)
  - Data Archive (HEASARC)

Charles S. Draper Laboratory
- Project Management
  - Mission and Systems Engineering
  - Safety and Mission Assurance

ATK Space, Inc
- Spacecraft Bus
  - Observatory Integration and Test
POET GRB Science

POET will answer questions about GRBs that can only be answered by X-ray and Gamma-ray polarisation measurements.

What is the composition of GRBs?
What is the prompt radiation mechanism?
What is the small-scale geometry of the prompt emission region?
POET Characteristics

Sakamoto, et al.

WXM    FREGATE

normalized counts/sec/keV

channel energy (keV)

X

-2 0 2

10 20

photoelectric effect dominant

pair production dominant

compton effect dominant

27th June 2008

2008 Nanjing GRB Conference

M.S. Longair
# POET Instrument Suite

## LEP Parameters

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Value</th>
</tr>
</thead>
<tbody>
<tr>
<td>Polarimetry</td>
<td>2-15 keV</td>
</tr>
<tr>
<td>Detectors</td>
<td>Ne:CO₂:CH₃NO₂ Gas (8)</td>
</tr>
<tr>
<td>Spectroscopy</td>
<td>2-15 keV</td>
</tr>
<tr>
<td>Field-of-View</td>
<td>±44° (non-imaging)</td>
</tr>
</tbody>
</table>

## GRAPE Parameters

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Value</th>
</tr>
</thead>
<tbody>
<tr>
<td>Polarimetry</td>
<td>60-500 keV</td>
</tr>
<tr>
<td>Detectors</td>
<td>BGO/plastic scintillator (62)</td>
</tr>
<tr>
<td>Spectroscopy</td>
<td>15 keV - 1 MeV</td>
</tr>
<tr>
<td>Detectors</td>
<td>NaI(Tl) scintillator (2)</td>
</tr>
<tr>
<td>Field-of-View</td>
<td>±60° (non-imaging)</td>
</tr>
</tbody>
</table>
X-ray and Gamma-ray Polarimeters

 CAPITALIZE ON: correlation between the incident photon electric field vector and the photoelectron emission direction or scattered photon direction

FIT FUNCTION TO THE ANGULAR DISTRIBUTION

MODULATION FACTOR, \( \mu \):

\[
\mu = \frac{N_{\text{max}} - N_{\text{min}}}{N_{\text{max}} + N_{\text{min}}} = \frac{B}{2A + B}
\]
GRAPE Prototype

Based on use of flat panel PMT.
Grape Performance


μ = 33% @ 69 keV
μ = 44% @ 129 keV
Wide FoV and off-axis uniformity
Balloon flight of an engineering prototype on June 21, 2007.

Measured background with (preliminary) simulated background.
The TPC Polarimeter

- GEM with strip readout
- Track images formed by time-projection by binning arrival times
- Resolution is (largely) independent of the active depth
- Max depth determined only by degree of X-ray beam collimation

[Image: Diagram of TPC Polarimeter with GEM, drift electrode, charge sensitive amplifiers, and differentiated waveforms.]

Black et al, NIM A, 2007
Prototype TPC Polarimeter Results

- Uniform response
- Modulation consistent with gas pixel detectors
- Unit QE possible

<table>
<thead>
<tr>
<th>Polarization Phase</th>
<th>Measured Parameters</th>
<th>$\chi^2$</th>
</tr>
</thead>
<tbody>
<tr>
<td>unpolarized</td>
<td>Modulation (%)</td>
<td>Phase (degrees)</td>
</tr>
<tr>
<td></td>
<td>$0^\circ$</td>
<td>45.0 ± 1.1</td>
</tr>
<tr>
<td></td>
<td>$45^\circ$</td>
<td>45.3 ± 1.1</td>
</tr>
<tr>
<td></td>
<td>$90^\circ$</td>
<td>44.7 ± 1.1</td>
</tr>
</tbody>
</table>

Black et al., 2007, NIM A, 581, 755
Wide FoV Prototype

Parameter | Value
--- | ---
Active Element | Ne:CO₂:CH₃NO₂
Active Volume | 24 x 24 x 24 cm³
Pressure | 780 Torr
Energy Range | 2-15 keV
Energy Resolution | 38% at 6keV
µ @ 6 keV | 45%
Field of View | ±44°
Mass | 28.5 kg
Power (peak/ave) | 33/31 W
Data Volume | 248 MB/day
Temperature Range | 25 ± 1°C / -10 to 50 °C
Peak Sensitivity | ~3.5 keV
Wide FoV Prototype

Spectra: Ne:CO₂

Spectra: Ne:CO₂:CH₃NO₂
# Mission Concept

## Mission Parameters

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Value</th>
</tr>
</thead>
<tbody>
<tr>
<td><strong>Launch Date</strong></td>
<td>May, 2012</td>
</tr>
<tr>
<td><strong>Launch Vehicle</strong></td>
<td>Standard SMEX</td>
</tr>
<tr>
<td><strong>Orbit</strong></td>
<td>600 km, 28.5° incl.</td>
</tr>
<tr>
<td><strong>Mission Lifetime</strong></td>
<td>2+ years</td>
</tr>
<tr>
<td><strong>Pointing Mode</strong></td>
<td>Zenith-pointed</td>
</tr>
<tr>
<td><strong>Spin Rate</strong></td>
<td>15 rpm</td>
</tr>
</tbody>
</table>
POET Spacecraft
POET Mission Operations

Mission Operations Center (MOC)
- Scheduling
- Command loads
- Observatory health & safety
- SC performance
- Level Zero processing
- Orbit prediction

Cmd Requests

Data Analysis Tools

POET Data Center (PDC)

Instrument Teams

Sciene Team

Science Community

HEASARC

TDRS

POET-019

White Sands Complex (WSC)

Two-Line Elements

Post Launch Orbit

GSFC

TDRSS TLM & Cmds

Ranging Data

FITS data
L0, L1, L2, L3
GSFC

3.6 Mbps downlink
2 kbps uplink

S-band CMD/TLM
5 kbps downlink
2 kbps uplink

Wallops
Ground Station
(NASA)

Perth & Hawaii
Ground Station
(USN)
# POET Performance

<table>
<thead>
<tr>
<th># GRBs S/N&gt;5</th>
<th># GRBs Ep</th>
</tr>
</thead>
<tbody>
<tr>
<td>LEP</td>
<td>99%</td>
</tr>
<tr>
<td>GRAPE</td>
<td>80%</td>
</tr>
<tr>
<td>LEP+GRAPE</td>
<td>78%</td>
</tr>
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</table>
Distinguish GRB Models

Physical Model
Ordered B-field

Geometric Model
Optimum viewing factor

Physical Model
Synchrotron Emission
Ordered B-field

Geometric Model
Synchrotron Emission
Random B-field

OR

Compton Drag
What is the GRB radiation mechanism?

GRAPE and LEP will independently measure $\Pi$ above and below $E_{\text{peak}}$

<table>
<thead>
<tr>
<th>LEP</th>
<th>GRAPE</th>
</tr>
</thead>
<tbody>
<tr>
<td>GRBs</td>
<td>MDP</td>
</tr>
<tr>
<td>8</td>
<td>10%</td>
</tr>
<tr>
<td>40</td>
<td>25%</td>
</tr>
<tr>
<td>*72</td>
<td>50%</td>
</tr>
</tbody>
</table>

$h \nu$ (keV)

$E^2 N(E)$

Band function

Ryde function

SO

syn + bb

CD

$E_{\text{peak}}$
POET was not selected for Phase A so now what?.....

- Improve readiness of GRAPE
  - Balloon flight
- Improve readiness of LEP
  - MidSTAR-2 GRBP (~2011)
  - GEMS in Phase-A (Gravity and Extreme Magnetism SMEX)
- Look for new opportunities
  - e.g. Space Station
The GRBP: A payload for MidStar 2

Area: 144 cm²
Depth: 5 cm
FoV: 1 steradian
Gas: Ne:CO₂:CS₂
Pressure: 1 atm

MDP averaged from 2 - 10 keV
**MidSTAR-2**

**USNA Project**

**High risk Low-cost**

**Make a scientific measurement**

**Several GRBs in 2 yr lifetime**

**Low cost proof-of-concept**

**Launch ~2011**
GRBP Prototype
Detector Design
Prototype chamber
HV Power Supplies
(mBAT - Mini BAT 1/8 scale)

<table>
<thead>
<tr>
<th>mBAT Parameters</th>
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<tbody>
<tr>
<td>Energy Range</td>
</tr>
<tr>
<td>FoV</td>
</tr>
<tr>
<td>Spatial Resolution</td>
</tr>
<tr>
<td>Spectral Resolution</td>
</tr>
<tr>
<td>Position Notice</td>
</tr>
</tbody>
</table>
Solar Flare Science

- How does the Sun release such large quantities of energy in a Solar Flare?
- How does the Sun accelerate electrons and ions with such high efficiency?
- POET will determine the angular beaming of electrons
- Polarimetry measures the electron beaming.
- Models predict 20-30% polarization.

GRAPE will measure polarization direction and magnitude of Solar Flares to answer these questions.

<table>
<thead>
<tr>
<th>Energy Band (keV)</th>
<th>23 July 2002 (X4.8) Δt = 60 s</th>
<th>M5 flare Δt = 300 s</th>
</tr>
</thead>
<tbody>
<tr>
<td>50-500</td>
<td>2.3%</td>
<td>27%</td>
</tr>
<tr>
<td>50-100</td>
<td>3.6%</td>
<td>43%</td>
</tr>
<tr>
<td>100-200</td>
<td>3.4%</td>
<td>40%</td>
</tr>
<tr>
<td>200-500</td>
<td>4.9%</td>
<td>62%</td>
</tr>
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</table>
Pulsar Science

X-ray polarimetry is the **only** way to distinguish between the two leading models of accretion flow onto highly magnetized neutron stars.

Intensity (top), polarization position angle (middle) and degree of polarization (bottom) vs. phase predicted by different models for the Crab pulsar. All reproduce the intensity profile. Only polarization measurements can uniquely differentiate between models.

<table>
<thead>
<tr>
<th></th>
<th>LEP (2-10 keV)</th>
<th>MDP in 10 ksec</th>
<th>MDP in 1 ksec</th>
</tr>
</thead>
<tbody>
<tr>
<td>CRAB</td>
<td>4 %</td>
<td>8%</td>
<td></td>
</tr>
<tr>
<td>1/10 CRAB</td>
<td>8%</td>
<td>15%</td>
<td></td>
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