

## GRB 090417B and its Host Galaxy: A Step Towards an Understanding of Optically-Dark Gamma-Ray Bursts

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### ABSTRACT

GRB 090417B was an unusually long burst with a  $T_{90}$  duration of at least 2130 s and a multi-peaked light curve at energies of 15–150 keV. It was optically

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dark and has been convincingly associated with a bright star-forming galaxy at a redshift of 0.345 that is broadly similar to the Milky Way. This is one of the few cases where a host galaxy has been clearly identified for a dark gamma-ray burst and thus an ideal candidate for studying the origin of dark bursts. We find that the dark nature of GRB 090417B can not be explained by high redshift, incomplete observations, or unusual physics in the production of the afterglow. The *Swift*/XRT *X*-ray data are consistent with the afterglow being obscured by a dense, localized sheet of dust approximately 30–80 pc from the burst along the line of sight. Assuming the standard relativistic fireball model for the afterglow we find that the optical flux is at least 2.5 mag fainter than predicted by the *X*-ray flux. We are able to explain the lack of an optical afterglow, and the evolution of the *X*-ray spectrum, by assuming that there is a sheet of dust along the line of sight approximately 30–80 pc from the progenitor. Our results suggest that this dust sheet imparts an extinction of  $A_V \gtrsim 12$  mag, which is sufficient to explain the missing optical flux. GRB 090417B is an example of a gamma-ray burst that is dark due to the localized dust structure in its host galaxy.

*Subject headings:* gamma rays: bursts

## 1. Introduction

The *Swift* observatory (Gehrels et al. 2004) is a multi-instrument satellite mission that was designed to detect and rapidly localize gamma-ray bursts (GRBs). The observatory contains three telescopes: the Burst Alert Telescope (BAT; Barthelmy et al. (2005)), the *X*-Ray Telescope (XRT; Burrows et al. (2005)), and the UltraViolet/Optical Telescope (UVOT; Roming et al. (2005)). The BAT is used to identify GRBs and localize them to  $\sim 3'$  in the energy range 15–150 keV. Once the BAT has localized a burst *Swift* slews to point the XRT and the UVOT at the burst. The XRT obtains rapid *X*-ray localizations to  $\lesssim 5''$  in the energy range 0.2–10 keV while the UVOT obtains localizations to  $\approx 0''.5$  then cycles through a set of optical and ultraviolet filters. Over the past five years *Swift* observations has been detecting GRBs at the rate of approximately 100 per year. Almost all of these *Swift*-detected GRBs have had *X*-ray afterglows, and optical or infrared afterglows have been detected for  $\approx 60\%$  of them (Roming et al. 2009). The remaining  $\approx 40\%$  of the *Swift* bursts have no reported optical or infrared detections. In general between  $\approx 25$ –40% of GRBs are dark GRBs (Fynbo et al. 2009b).

The gamma-ray burst GRB 090417B was detected by the BAT as an image trigger at 13:17:23 UT on 2009 Apr 17 (Sbarufatti et al. 2009). *Swift* immediately slewed to

GRB 090417B and narrow-field observations began observing 387 (XRT) and 378 (UVOT) s after the BAT trigger. The BAT light curve showed a long period of emission starting 200 s before the trigger.

Several ground-based follow-up observations were taken, but no optical or infrared afterglow was detected. Fynbo et al. (2009a) found a source with  $R = 21.3$  inside the XRT error circle (Sbarufatti et al. 2009) and identified it as an SDSS galaxy (Strauss et al. 2002) and the possible host galaxy of GRB 090417B. A redshift of  $z = 0.345$  was measured for this galaxy by Berger & Fox (2009) based on the forbidden oxygen emission lines [O II](3727) and [O III](5006) as well as  $H\alpha$  emission.

Dark GRBs have been an enduring mystery in the lore of GRB studies. Historically only about one quarter of localized gamma-ray bursts have had optical afterglow (OA) detected (e.g., Fynbo et al. 2001). In the pre-*Swift* era it was sometimes assumed that this was due to the response times of ground-based optical telescopes which allowed most afterglows to fade before they could be observed. *Swift* however, with its ability to train the XRT and UVOT on a burst within  $\approx 100$  s of the initial detection, has been able to provide positions to within  $\approx 5''$  for most bursts that have been detected by the BAT. The rapid distribution of these positions has allowed ground-based optical and infrared follow-up observations for many of these bursts, often within minutes of the BAT trigger. Even with such rapid follow-up observations no optical or infrared afterglows have been found for  $\approx 40\%$  of BAT-detected GRBs, thus demonstrating that observational constraints were not the cause to the dark burst problem (Roming et al. 2006).

Several attempts have been made to quantify the dark burst problem by establishing operational definitions of dark bursts (Jakobsson et al. 2004; Rol et al. 2005; van der Horst et al. 2009). Most recently van der Horst et al. (2009) proposed that dark GRBs be defined as those that have  $\beta_{OX} < \beta_X - 0.5$  measured 11 hours after the burst. They define  $\beta_{OX}$  as the spectral index between the optical and  $X$ -ray régimes and  $\beta_X$  as the spectral index in the  $X$ -ray band. The spectral index is defined by  $f_\nu(\nu) \propto \nu^{-\beta}$  where  $f_\nu(\nu)$  is the flux density at a frequency  $\nu$ . This definition assumes that both the  $X$ -ray and optical afterglows are due to synchrotron radiation from a relativistic fireball. This model predicts that the spectral slope will either be the same in both régimes (if there is no cooling break between them) or will differ by  $\Delta\beta = 0.5$  (if there is a cooling break between them). A difference of  $\Delta\beta > 0.5$  can not be explained in this model and thus requires that something be suppressing the flux at optical wavelengths, resulting in a dark burst (Jakobsson et al. 2004).

Several possibilities have been proposed to explain dark bursts. First, several studies have suggested that dark bursts may be due to extinction along the line of sight, either in the Milky Way, or in the host galaxy (e.g., Piro et al. 2002; Levan et al. 2006; Berger et al. 2007;

Jaunsen et al. 2008; Tanvir et al. 2008). Perley et al. (2009) suggest that this is the case for most dark GRBs. Further,  $X$ -ray observations show that some bursts with no optical afterglows have systematically higher H I column densities than those with optical afterglows (Fynbo et al. 2009b). This suggests that dark bursts may be suffering from higher extinction than optically-bright bursts. However, there is significant overlap in the distribution of  $N_{\text{H}}$  values between dark and bright bursts, so it is not clear if extinction is the sole parameter responsible for the darkness of these bursts.

Schady et al. (2007, 2009) did panchromatic studies of several GRBs that were observed with *Swift*'s XRT and UVOT (i.e., optically-bright GRBs) in order to probe the extinction in their host galaxies. They found that the combined  $X$ -ray/ultraviolet/optical data for most of these GRBs are best fit using a Small Magellanic Cloud (SMC) dust model (Pei 1992). The extinctions in these galaxies were small, with  $A_V < 1$  mag. They also find a mean gas-to-dust ratio of  $N_{\text{H}} = (6.7 \times 10^{21}) A_V$ , which is lower than, but similar to, the value found in the SMC ( $N_{\text{H}} = (15.4 \times 10^{21}) A_V$  using equation (4) and Table 2 of Pei (1992)). Kann et al. (2006) also find that the the gas-to-dust ratio for a large sample of host galaxies of optically-bright GRBs is lower than the SMC value. However, the gas-to-dust ratios in the host galaxies of dark GRBs tend to be more like that of the Milky Way (e.g., Piro et al. 2002) than that of the SMC.

A second possible origin for dark bursts is that they are GRBs located at high redshift, so that the Lyman break (or the start of the Lyman-alpha forest) is observed redward of the optical band. The most distant GRB observed to date is GRB 090423 at  $z = 8.2$  (Salvaterra et al. 2009; Tanvir et al. 2009), so some of the dark bursts could simply be at high redshift. However, Fynbo et al. (2009b) have found that  $\lesssim 24\%$  of GRBs have  $z \gtrsim 6$ , whereas between 25% and 42% of GRBs are dark. Therefore, it is unlikely that all of the *Swift* dark burst lie at very high redshifts.

Third, the hypothesis that afterglow radiation is synchrotron radiation may be wrong. This could lead to a different spectrum from what is expected from the relativistic fireball model, in which case the above definition of a dark burst is invalid.

Long-soft GRBs (Kouveliotou et al. 1993) tend to be found in host galaxies that are small, irregular, and have high specific star-formation rates (Christensen et al. 2004). Le Floc'h et al. (2003) find a median infrared luminosity of  $L \approx 0.08L^*$  for the host galaxies of GRBs with optical afterglows and conclude that most hosts are sub-luminous at optical and near-infrared wavelengths. These GRB hosts tend to be small galaxies that have intrinsic luminosities similar to those of the dwarf galaxies in the Local Group. However, some hosts are significantly larger, such as that of GRB 990705 (Le Floc'h et al. 2002), which was an Sc spiral with  $L \approx 2L_*$ . To date no long-soft GRB has been found in an elliptical galaxy.

All of the hosts of long–soft GRBs have morphologies that are consistent with either exponential discs, or irregular structure (Conselice et al. 2005; Wainwright et al. 2007). GRB hosts follow the size–luminosity relation (e.g., Trujillo et al. 2006) and extend it towards lower luminosities (Wainwright et al. 2007). In spite of their generally small size the host galaxies of long–soft GRBs tend to have high specific star-formation rates (SSFRs). Savaglio et al. (2009) found a median SSFR of  $0.8 \text{ yr}^{-1}$ , similar to what is seen in Lyman-break galaxies.

It is important to realize, however, that these properties have been determined using GRBs that have usually been localized to sub-arcsecond precisions. In practice this usually means that the sample of GRB host galaxies is restricted to GRBs that have had optical afterglows. The host galaxies of dark GRBs may have different properties. Perley et al. (2009) identify host galaxies for 14 dark GRBs and find that they have the same redshift distribution as the hosts of optically-bright GRBs do, and that these hosts do not appear to be significantly different in size, structure, or luminosity from the hosts of GRBs with optical afterglows. They do find that a significant fraction of the hosts of dark bursts show evidence for high internal extinction ( $A_V > 2\text{--}5 \text{ mag}$ ). However, many of their host associations are uncertain, with more than half of the galaxies having a probability of a chance association on the sky of  $P_{\text{ch}} > 0.01$ . As we will show in this paper GRB 090417B is located  $1''.08$  from the centre a bright galaxy with a probability of a chance association of only  $P_{\text{ch}} \approx 10^{-3}$ . This makes GRB 090417B one of the strongest associations of a dark GRB with a host galaxy.

In this paper we present space- and ground-based gamma-ray,  $X$ -ray, ultraviolet, and optical observations of GRB 090417B and its host galaxy. We have adopted a standard  $\Lambda$ CDM cosmology with a Hubble parameter of  $H_0 = 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$ , a matter density of  $\Omega_m = 0.27$ , and a cosmological constant of  $\Omega_\Lambda = 1 - \Omega_m = 0.73$ . For this cosmology a redshift of  $z = 0.345$  corresponds to a luminosity distance of  $1.814 \text{ Gpc}$  and a distance modulus of  $41.29$ . One arcsecond corresponds to  $6.54 \text{ comoving kpc}$ , or  $4.84 \text{ proper kpc}$ . The look back time is  $3.80 \text{ Gyr}$ .

## 2. Data

### 2.1. BAT Data

The BAT scaled-map, event and DPH data were reduced using the standard BAT software available through HEASARC<sup>1</sup>.

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<sup>1</sup><http://swift.gsfc.nasa.gov/docs/software/lheasoft/download.html>

The  $T_{90}$  duration of GRB 090417B was  $> 2130 \pm 50$  s in the 15–150 keV band. The light curve evolved gradually, and there were no features in the unimaged count rates that might indicate how long the burst persisted in the BAT energy range past that time. Even at that lower limit on  $T_{90}$ , this was the longest GRB ever seen by BAT or BATSE (Fishman et al. 1988), although GRB 060218 approached this duration with  $T_{90} = 2100 \pm 100$  (Campana et al. 2006).

The BAT light curve for GRB 090417B showed four broad, overlapping peaks in the flux at approximately  $T - 70$ ,  $T + 400$ ,  $T + 500$ , and  $T + 1600$  s, as shown in Figure 1. The 1-s peak flux of  $0.3 \pm 0.1$  photons  $\text{cm}^{-2} \text{s}^{-1}$  in the range 15–350 keV was approximately  $T + 500$  s. The total fluence from  $T - 320$  to  $T + 2105$  s was  $8.20_{-2.1}^{+1.0} \times 10^{-6}$  erg  $\text{cm}^{-2}$  in the 15–150 keV band. The spectrum was best fitted by a simple power-law function with an average photon index of  $1.89 \pm 0.12$ . The index did not change significantly during the burst. All the above uncertainties are at 90% confidence.

The most impressive characteristic in the prompt emission of GRB 090417B was its extreme length. The burst was marginally above background ( $2.3\sigma$ ) on the 64-s interval starting immediately after a pre-planned slew maneuver ended at  $T - 320$  seconds, although not at a high enough level to be detectable independent of the later emission. From that time, the flux increased steadily, to be detected as a point source on-board on the interval  $T - 64$  to  $T_0$  s, and finally as a sufficiently significant source to trigger a burst response on the interval  $T_0$  to  $T + 320$  s. Although it appeared to be declining toward background levels, the burst continued to be detected until the spacecraft slewed away due to the changing Earth constraint at  $T + 2105$  s.

## 2.2. XRT Data

### 2.2.1. Data Reduction

The XRT on board *Swift* began observations on 2009 April 17 at 15:26:33 UT, 387 s after the BAT trigger (Sbarufatti et al. 2009), and ended on 2009 May 5, with a total net exposure of 1456 s in Windowed Timing mode (WT) and 83.36 ks in Photon Counting mode (PC). The XRT observations are listed in Table 1.

The XRT data were processed using the FTOOLS software package distributed inside HEASOFT (v6.6.3). We ran the task XRTPIPELINE applying calibrations and standard filtering criteria. Events with grades 0–2 and 0–12 were selected for WT and PC data respectively. The analysis was performed inside the 0.3–10 keV energy band.

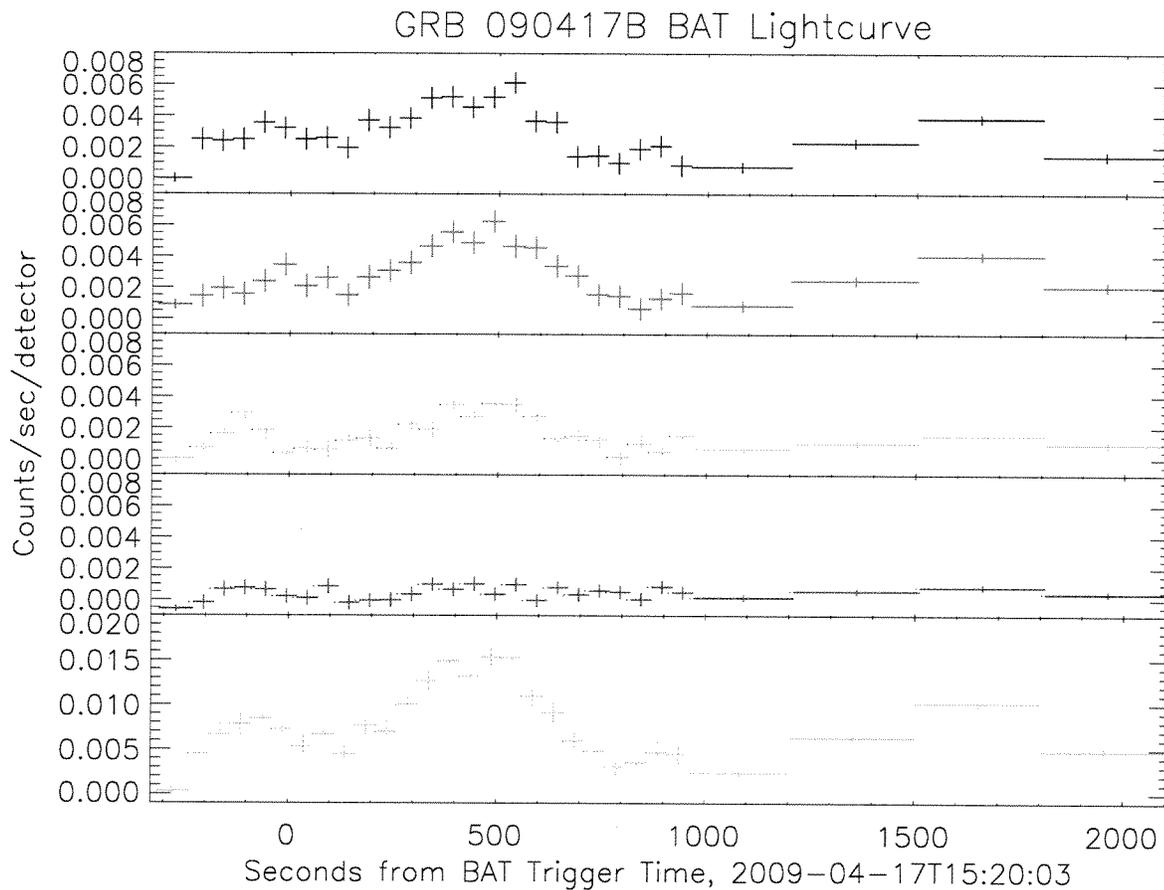


Fig. 1.— The top four panels show the BAT light curves for the prompt emission from GRB 090417B in four energy bands. From the top down the energy bands are: 14–24 keV, 24–51.1 keV, 51.1–101.2 keV, and 101.2–194.9 keV. The bottom panel shows the total light curve in the 14–194.9 keV band.

The best position for the afterglow was obtained using 6972 s of overlapping XRT Photon Counting mode data and 9 UVOT *v*-band images in order to correct the XRT astrometry making use of the XRT–UVOT alignment and matching to the USNO-B1 catalogue, as described by Goad et al. (2007) and Evans et al. (2009). Our best position is RA(J2000.0) = 209°6942 (13:58:46.62), Dec(J2000.0) = +47°0182 (+47:01:05.4) with an error of 1''4 (90% confidence, including boresight uncertainties).

### 2.2.2. Temporal Analysis

Source photons were extracted from a region with a 30 pixel radius (1 pixel = 2''36), with the exception of PC mode data for sequence 00349450000 where we used an annular region with radii 3 pixel and 30 pixel in order to correct for pile-up. The background was estimated from a circular region with a 50 pixel radius located away from any detected source in the field. When the count-rate dropped below the level of  $\approx 10^{-2}$  count s<sup>-1</sup> we used the SOSTA tool of XIMAGE which corrects for vignetting, exposure variations and PSF losses within an optimized box, using the same background region. The 0.3–10 keV light curve thus obtained is shown in Figure 2 (top panel). The light curve has been rebinned in order to achieve a minimum signal-to-noise ratio (S/N) of 3 for each point. The best fit to the light curve is given by a doubly broken power law with indices  $\alpha_1 = 0.86 \pm 0.03$ ,  $\alpha_2 = 1.40 \pm 0.04$ , and  $\alpha_3 = 2.0_{-0.4}^{+0.7}$ , and breaks at  $T + 9100_{-1000}^{+900}$  s and  $T + 4.5_{-1.2}^{+2.8} \times 10^5$  s. The second break could be interpreted as a jet break. However the lack of an optical detection for this burst does not allow us to confirm or deny this hypothesis. Also see § 4.3 for an alternate interpretation of the late-time steepening of the *X*-ray light curve. The first sequence of *X*-ray data is dominated by two strong flares. The first flare began before the start of XRT observations and peaked around  $T + 530$  s with a decay time of 170 s. The second, double-peaked, flare started at  $T + 1260$  s, peaked at  $T + 1410$  s and  $T + 1470$  s, and had a decay time of 240 s. Further flaring activity is apparent up to at least  $T + 4 \times 10^5$  s. The bottom panel of Figure 2 shows the hardness ratio for the afterglow measured using the 0.3–2 and the 2–10 keV bands. The first orbit (up to  $T + 1300$ ) shows a hard hardness ratio with spectral evolution that follows the flaring activity. The same is true for the remaining observations, with the hardness ratio slowly decreasing with time.

### 2.2.3. Spectral Analysis

In order to detect possible spectral variations we extracted spectra separately for the two flares, the early decay part from  $T + 5000$  s to  $T + 9100$  s, the steep decay from  $T + 9100$  s

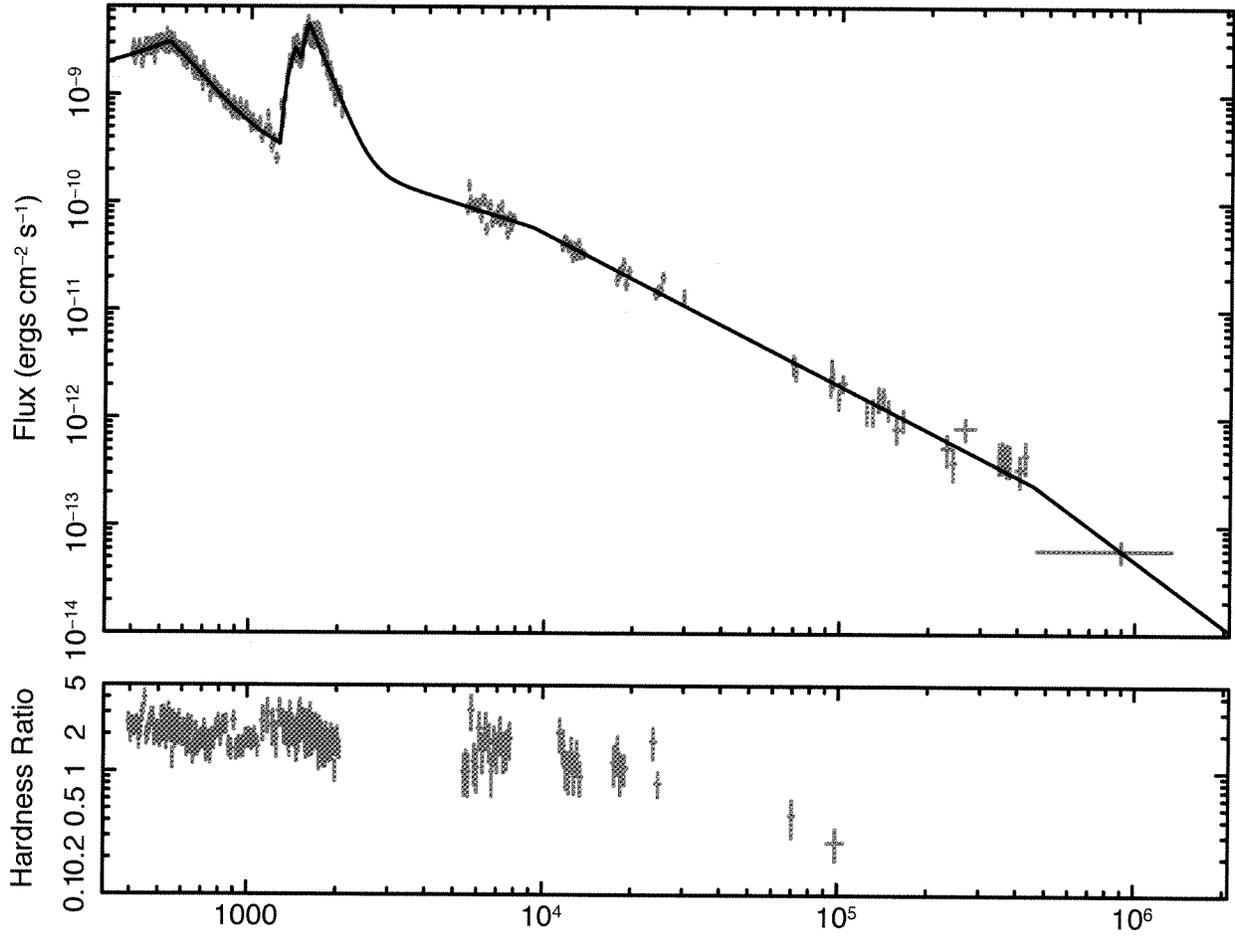


Fig. 2.— Top: The X-ray flux light curve of GRB 090417B. Bottom: The hardness ratio evolution of GRB 090417B in the 0.3–2 and 2–10 keV bands. The blue points are WT data and the red points are PC data.

to  $T + 4.5 \times 10^5$  s, and the late decay from  $T + 4.5 \times 10^5$  s to the end of observations. Source and background spectra were extracted using the same regions used for the light curves, with the exception of the late decay spectrum where the count rate was too low to group the energy channels in order to have a Gaussian distribution of the number of photons per grouped channel. Hence, for this spectrum we replaced the  $\chi^2$  statistic with the  $C$  statistic (Cash 1979), which can be used whenever the Gaussian approximation is not valid, provided that the contamination from background photons is negligible. To ensure this, we extracted photons from a region with a 10 pixel radius.

The ancillary response files were generated using the task XRTMKARF, and, with the exception of the late decay spectrum we grouped channels in order to have at least 20 photons per bin. Spectral fitting was performed using XSPEC (v11.3.2). For each spectrum we modelled the Galactic absorption using a warm absorber with neutral hydrogen column density  $1.60 \times 10^{21} \text{ cm}^{-2}$  (Kalberla et al. 2005) and included an intrinsic warm absorber at the redshift of the host galaxy ( $z = 0.345$ ). The best model for the spectra of the flares was a power law with a high energy cutoff, while for the remaining parts of the afterglow the best fit was obtained using a single power law. The best fit parameters are given in Table 2.

The fact that the first phases of the XRT observations are best fitted by a cutoff power law with the cutoff energy around 3 keV indicates that the peak energy of the flares was passing through the XRT band at the time. After the flares the spectrum is described by a power law with photon index of  $2.0 \pm 0.1$ . After the first break, up to  $T + 29.6$  ks, the spectrum is described by a power law with similar parameters ( $\Gamma = 2.3 \pm 0.2$  and an absorbing column  $N_{\text{H}} = 2.0 \pm 0.4^{22} \text{ cm}^{-2}$ ). At later times the spectrum softens significantly to a photon index of  $\Gamma = 3.8 \pm 0.5$  and shows no significant variations across the break at  $T + 4.5 \times 10^5$  s up to the end of the observations.

The intrinsic absorbing column is in the range  $(1.1 - 2.4) \times 10^{22} \text{ cm}^{-2}$ . From our analysis of the spectral energy distribution (§ 3.2) we find that the best fit to the data is given by a Milky Way extinction law instead of the SMC-like extinction law found in many GRB host galaxies. Therefore, we use the relationship between  $N_{\text{H}}$  and extinction in the Milky Way from Predehl & Schmitt (1995) to find that the observed absorbing column corresponds to an extinction  $A_v \approx 11$  mag. This high value for the extinction in the host along the line of sight to the burst is capable of explaining the dark optical nature of this event.

Table 1. Journal of XRT observations.

Sequence	Start Time UT	End Time UT	Start time <sup>1</sup> (s)	End Time <sup>1</sup> (s)	Obs. Mode	Exposure (s)
00349450000	2009-04-17 15:26:40	2009-04-17 23:29:46	$3.932 \times 10^2$	$2.055 \times 10^3$	WT	$1.456 \times 10^3$
00349450000	2009-04-17 15:38:43	2009-04-17 23:33:26	$1.116 \times 10^3$	$2.960 \times 10^4$	PC	$9.214 \times 10^3$
00349450002	2009-04-18 10:29:48	2009-04-18 11:07:52	$6.898 \times 10^4$	$7.124 \times 10^4$	PC	$2.252 \times 10^3$
00349450003	2009-04-18 16:57:16	2009-04-18 20:20:44	$9.223 \times 10^4$	$1.044 \times 10^5$	PC	$4.871 \times 10^3$
00349450006	2009-04-19 01:17:02	2009-04-19 12:50:57	$1.222 \times 10^5$	$1.638 \times 10^5$	PC	$9.709 \times 10^3$
00349450007	2009-04-20 04:41:58	2009-04-20 22:26:57	$2.209 \times 10^5$	$2.848 \times 10^5$	PC	$6.131 \times 10^3$
00349450008	2009-04-21 14:13:55	2009-04-21 23:59:56	$3.416 \times 10^5$	$3.768 \times 10^5$	PC	$7.910 \times 10^3$
00349450009	2009-04-22 04:29:42	2009-04-22 14:20:47	$3.930 \times 10^5$	$4.284 \times 10^5$	PC	$4.574 \times 10^3$
00349450010	2009-04-23 00:07:52	2009-04-23 22:40:57	$4.637 \times 10^5$	$5.448 \times 10^5$	PC	$2.989 \times 10^3$
00349450011	2009-04-24 00:18:14	2009-04-24 22:57:56	$5.507 \times 10^5$	$6.323 \times 10^5$	PC	$1.344 \times 10^3$
00349450012	2009-04-25 11:39:16	2009-04-25 23:01:57	$6.779 \times 10^5$	$7.189 \times 10^5$	PC	$1.480 \times 10^3$
00349450013	2009-04-26 02:11:18	2009-04-27 23:12:56	$7.303 \times 10^5$	$8.923 \times 10^5$	PC	$2.140 \times 10^3$
00349450014	2009-04-27 00:40:17	2009-04-27 21:39:56	$8.112 \times 10^5$	$8.868 \times 10^5$	PC	$1.224 \times 10^3$
00349450015	2009-04-28 00:37:57	2009-04-28 23:20:56	$8.975 \times 10^5$	$9.792 \times 10^5$	PC	$8.368 \times 10^3$
00349450016	2009-04-29 00:41:42	2009-04-29 23:26:57	$9.841 \times 10^5$	$1.066 \times 10^6$	PC	$4.537 \times 10^3$
00349450017	2009-05-01 00:40:42	2009-05-02 23:37:57	$1.157 \times 10^6$	$1.326 \times 10^6$	PC	$1.604 \times 10^4$

<sup>1</sup>Time since BAT trigger.

Table 2. Spectral parameters.

Start time <sup>1</sup> (s)	End Time <sup>1</sup> (s)	$N_{\text{H}}$ ( $10^{22} \text{ cm}^{-2}$ )	$\Gamma$	Cutoff Energy (keV)	Mean Flux ( $\text{erg cm}^{-2} \text{ s}^{-1}$ )	$\chi^2/\text{d.o.f.}$
$3.932 \times 10^2$	$1.116 \times 10^3$	$1.5 \pm 0.2$	$0.2 \pm 0.2$	$3.0^{+0.7}_{-0.5}$	$1.3 \times 10^{-09}$	446/421
$1.116 \times 10^3$	$2.055 \times 10^3$	$1.3 \pm 0.2$	$0.3 \pm 0.2$	$2.9^{+0.5}_{-0.4}$	$2.8 \times 10^{-09}$	604/523
$5.000 \times 10^3$	$9.100 \times 10^3$	$2.0 \pm 0.4$	$1.9 \pm 0.2$	...	$6.2 \times 10^{-11}$	39/45
$9.100 \times 10^3$	$2.960 \times 10^4$	$1.6 \pm 0.4$	$2.3 \pm 0.2$	...	$1.7 \times 10^{-11}$	15/29
$6.898 \times 10^4$	$4.284 \times 10^5$	$1.9 \pm 0.5$	$3.8 \pm 0.5$	...	$4.4 \times 10^{-13}$	13/22
$4.637 \times 10^5$	$1.326 \times 10^6$	$< 1.9$	$3.1 \pm 0.9$	...	$7.7 \times 10^{-14}$	32/24 <sup>2</sup>

<sup>1</sup>Time since BAT trigger.

<sup>2</sup> $C$ -statistic/bins

### 2.3. UVOT Data

The *Swift*/UVOT began observations 378 s after the BAT trigger (Sbarufatti et al. 2009) with a 9 s settling mode exposure using the *v* filter. This was followed by a sequence that rotated through all seven of UVOT’s lenticular filters. The first of these exposures was a 147 s exposure starting at 395 s taken with the white filter. Observations continued until 1325874 s (= 15.3 days) after the trigger using primarily the *u* and ultraviolet filters. No optical or ultraviolet afterglow was detected in any of the UVOT data. Once it was determined that there was no optical or ultraviolet afterglow for GRB 090417B UVOT observed primarily with its ultraviolet filters in order to obtain ultraviolet magnitudes for the nearby galaxy that can not be obtained from ground-based observatories.

The SDSS galaxy noted by Fynbo et al. (2009a) is detected when we coadd the UVOT images. We determined a centroid for this galaxy of RA, Dec = 13:58:46.66, +47:01:04.4 (J2000.0) with an estimated internal uncertainty of 0<sup>ʹ</sup>.67 and an estimated systematic uncertainty relative to the USNO-B1.0 catalogue of 0<sup>ʹ</sup>.5. These uncertainties are the 90% confidence intervals. This galaxy is 1<sup>ʹ</sup>.08 southeast of the centre of the UVOT-enhanced XRT error circle. The field of GRB 090417B is shown in Figure 3. The SDSS galaxy is well-isolated from other sources in the field, so there is no contamination from neighbouring sources when doing aperture photometry.

We obtained the UVOT data from the *Swift* Data Archive on 12 May 2009. This data has had bad pixels identified, mod-8 noise corrected, and has been transformed into FK5 coordinates. We used the standard UVOT data analysis software distributed with HEASOFT 6.6.2 along with the standard calibration data. Photometry was done on the SDSS galaxy using UVOTSOURCE with circular source aperture of radius 4<sup>ʹ</sup>.0 and a 15<sup>ʹ</sup> radius circular aperture centred at RA, Dec = 13:58:47.57, +47:00:33.3 (J2000.0) for a background region. This region was selected to have similar background properties to those at the location of the galaxy, and to be free of contaminating sources. The UVOT photometry of the SDSS galaxy is presented in Table 3. We found no evidence for any change in the magnitude of the galaxy, in any filter, during the course of the UVOT observations. Our photometry was calibrated to the UVOT photometric system described in Poole et al. (2008).

The UVOT-enhanced XRT position for GRB 090417B is located at RA = 13:58:46.62, Dec = +47:01:05.4 (J2000.0) which corresponds to Galactic coordinates of  $l^{\text{II}}, b^{\text{II}} = 93^{\circ}.7486, +66^{\circ}.1127$ . The line-of-sight Galactic extinction in this direction is  $E_{B-V} = 0.02 \pm 0.01$  mag (Schlegel et al. 1998). This corresponds to extinctions in the UVOT filters of  $A_v = 0.05$ ,  $A_b = 0.07$ ,  $A_u = 0.08$ ,  $A_{uvw1} = 0.11$ ,  $A_{uvw2} = 0.16$ , and  $A_{white} = 0.14$ , and  $A_{white} = 0.08$  mag.

Since there is a candidate host galaxy for GRB 090417B it is not possible to put a

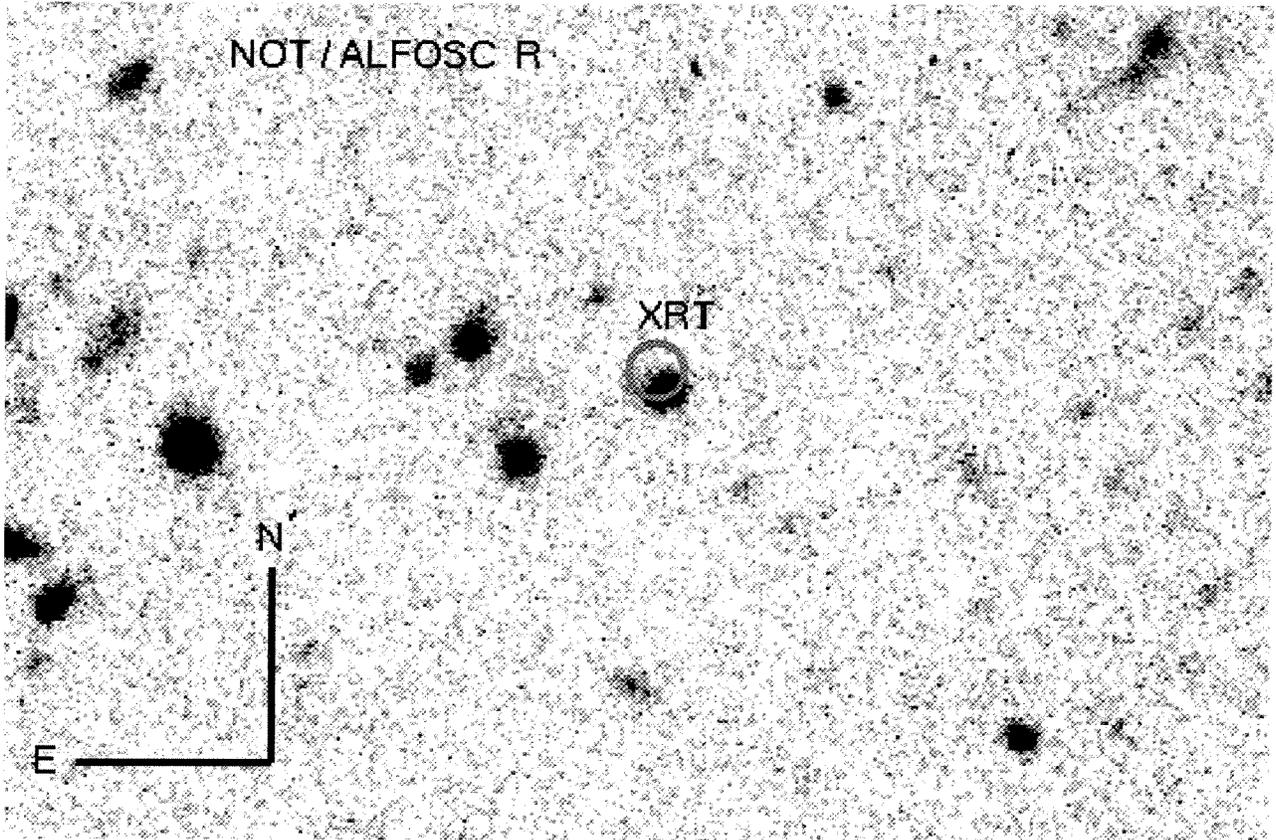


Fig. 3.— This Figure shows the averaged NOT/ALFOSC *R*-band image of the field of GRB 090417B with the XRT error circle. The SDSS galaxy is in the southeast quadrant of the XRT error circle. The arms of the compass have a length of ten arcsec.

direct faint limit on the luminosity of the optical afterglow beyond the observed magnitude of the galaxy itself ( $R = 21.3$  Fynbo et al. (2009a)). In order to constrain the maximum possible luminosity of the afterglow we used the UVOT data to look for variability in the galaxy’s magnitude. To do this we divided the data into early and late epochs. The early epoch consisted of all the UVOT data taken during *Swift*’s initial automated observations of GRB 090417B (OBSID 00349450000), which covers times up to 29 604 s ( $\approx 8$  hr) after the BAT trigger. This is the period when the optical afterglow is expected to be brightest. The late epoch included all observations later than this. The UVOT exposures for each epoch were coadded to produce two deep images for each UVOT filter. We photometered the galaxy in each image and recorded the difference between the two magnitudes and assumed that the late epoch magnitude represented the true magnitude of the galaxy. This provides a limit to any contribution to the magnitude of the host due to an optical afterglow. We found no evidence for any change in the magnitude of the host galaxy larger than  $\pm 0.2$  mag. From this we obtain the following constraints on the magnitude of the optical afterglow at 4 hr (= 14 800 s) after the BAT trigger:  $u_{\text{oa}} > 24.9$ ,  $uvw1_{\text{oa}} > 24.9$ ,  $uvm2_{\text{oa}} > 24.1$ , and  $uvw2_{\text{oa}} > 24.4$ .

## 2.4. Nordic Optical Telescope Data

We obtained optical images of GRB 090417B on 2009 Apr 17–18 with the ALFOSC-FASU mounted on the Nordic Optical Telescope (NOT). ALFOSC-FASU was operated in its high gain state ( $0.726 \text{ e}^-/\text{ADU}$ ). The read-out noise was  $3.2 \text{ e}^-/\text{pixel}$  and the pixel scale was  $0''.19/\text{pixel}$ . Data reduction was done following the standard procedure for optical CCD data and the individual images were aligned by centroiding on several stars in each exposure. Photometry was performed on the combined NOT images using SEXTRACTOR (Bertin & Arnouts 1996). We calibrated the data using the standard ALFOSC zero points<sup>2</sup>. We find  $B = 23.25 \pm 0.13$  (Bessel  $B$ ),  $R = 21.34 \pm 0.03$  (Bessel  $R$ ), and  $i = 20.82 \pm 0.07$  (interference  $i$ ) for the SDSS galaxy. The Galactic extinctions for the NOT data are  $A_B = 0.07$ ,  $A_R = 0.05$ , (Schlegel et al. 1998), and  $A_i = 0.03^3$ .

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<sup>2</sup><http://www.not.iac.es/instruments/alfosc/zpmon/>

<sup>3</sup><http://irsa.ipac.caltech.edu/applications/DUST/>

### 3. The Host Galaxy

#### 3.1. Probability that the SDSS Galaxy is the Host

We used the methodology of Bloom et al. (2002) to estimate the probability that a galaxy would lie within  $1''.08$  of the centre of the XRT error circle by chance,  $P_{\text{ch}}$ . This assumes that galaxies are randomly distributed on the sky and that there is no clustering. This assumption is not formally correct as galaxies do cluster. However, an examination of the field of GRB 090417B suggests that this GRB did not occur in a significant overdensity of background galaxies, so distortions in the computed  $P_{\text{ch}}$  value are likely to be small. Bloom et al. (2002) use the galaxy number counts of Hogg et al. (1997) to generate the surface density of galaxies with  $R$  magnitudes brighter than some limiting magnitude. In this case the  $R$  band is for the Keck 10-m Low Resolution Imaging Spectrometer (LRIS)  $R$  filter. We adopt  $R = 21.34 \pm 0.03$  (§ 2.4) as the magnitude of the SDSS galaxy. Using the Bloom et al. (2002) formula for an error circle containing a galaxy this yields  $P_{\text{ch}} = 0.003$ . This is a fairly small value, which suggests that SDSS J135846.65+470104.5 is likely to be the host galaxy of GRB 090417B.

A chance alignment probability of  $P_{\text{ch}} \approx 10^{-3}$  is large enough that one can not rule out the possibility of a misidentification. A visual examination of the UVOT images shows that the XRT error circle includes a significant fraction of the SDSS galaxy (see Figure 3). The observed separation between the  $X$ -ray afterglow and the centre of the galaxy is  $1''.08$ . At a redshift of  $z = 0.345$  this corresponds to a projected separation of 5.23 proper kpc. The galaxy appears point-like in all of our data. The NOT  $R$ -band data has a seeing FWHM of  $0''.91$ , which represents an upper limit on the effective radius of the galaxy of  $r_e \lesssim 4.4$  proper kpc. Therefore, GRB 090417B likely occurred in the outer regions of the galaxy. Berger et al. (2009) find  $g' = 22.47 \pm 0.16$  for the galaxy, which implies an absolute magnitude of  $M_V \approx -18.8$ . This puts the galaxy inside the observed range of host sizes (see Fruchter et al. (2006), their Fig. 4).

In order to test our hypothesis that SDSS J135846.65+470104.5 is the host galaxy of GRB 090417B we obtained the *Chandra* observations of the  $X$ -ray afterglow that were taken on 2009 May 11, 24 days ( $\approx 2 \times 10^6$  s) after the burst (Proposal Number: 10900117, PI: Andrew Levan). An  $X$ -ray source (consistent with the XRT position) is clearly detected at the position of the SDSS galaxy. The observed flux from this source is consistent with the  $X$ -ray flux that is predicted from the XRT light curve assuming a power-law decay with  $\alpha_3 = 2$ . This suggests that the *Chandra* source is the  $X$ -ray afterglow of GRB 090417B, and that there is minimal (if any) contamination from the galaxy.

Taken together these arguments strengthen the case that SDSS J135846.65+470104.5

is the host galaxy. Therefore, we conclude that we have identified the host of the dark burst GRB 090417B to a high degree of confidence.

### 3.2. Spectral Energy Distribution

We used the combined ultraviolet, optical, and infrared photometry of the SDSS galaxy to determine its spectral energy distribution (SED). Photometry for this galaxy was taken from our observations, the SDSS Web site<sup>4</sup>, and the various GCN Circulars on this burst, and is listed in Table 4. We used XSPEC to fit a power law to the observed SED. The best fit, with no host extinction, has a power-law index of  $\Gamma = 2.67_{-0.03}^{+0.13}$ . However, the goodness of fit is only  $\chi^2/\text{dof} = 81.40/12$ , which is very poor. Therefore, we also tried including extinction due to dust in the host galaxy when fitting the SED. Dust laws for the Small Magellanic Cloud, Large Magellanic Cloud and the Milky Way were tried. These led to a significant improvement in the goodness of fit with the best fit being found for Milky Way extinction in the host galaxy. Our best fit has a spectral index of  $\Gamma \approx -0.80_{-1.19}^{+0.55}$ , and a host extinction of  $A_V = 3.5_{-0.5}^{+1.0}$ .

### 3.3. Luminosity and Star-Formation Rate

SDSS J135846.65+470104.5 has a colour of  $B-R = +1.91 \pm 0.13$ , making it a fairly red galaxy. The typical absolute magnitude for a red galaxy at  $0.2 \leq z < 0.5$  is  $(M_B^*)_{AB} = -20.44$  (Lilly et al. 1995), assuming a cosmology with  $(H_0, \Omega_m, \Omega_\Lambda) = (50, 1, 0)$ . For our adopted cosmology that corresponds to  $(M_B^*)_{AB} = -20$ . At  $z = 0.345$  the rest-frame  $B$  band corresponds roughly to the observed  $r'$  band. Assuming that the galaxy has a power-law spectrum with  $\beta = \Gamma - 1 = -1.8$  (see § 3.2) and  $r' = 21.6$ , then it has a rest-frame luminosity in the  $B$  band of  $L_B \approx 1.3L_B^*$  where  $L_B^*$  is the rest-frame  $B$ -band luminosity of a typical red galaxy at  $z = 0.345$ . This indicates that the host is approximately an  $L^*$  galaxy with a total luminosity that is roughly similar to that of the Milky Way. However, we stress that this result is somewhat uncertain because  $M^*$  is highly correlated with the slope of the faint end of the galaxy luminosity function, and with its normalization (Lilly et al. 1995).

We can estimate the star-formation rate in the host galaxy using Eq. 2 of Madau et al. (1998). At the redshift of the burst  $1500 \text{ \AA}$  approximately corresponds to the UVOT uvw2 filter while  $2800 \text{ \AA}$  approximately corresponds to the SDSS  $u'$  filter. Correcting for Galactic

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<sup>4</sup><http://www.sdss.org/DR7/>

extinction these magnitudes for the host become  $uvw2_0 = 22.55$  and  $u'_0 = 23.15$ . These yield star-formation rates (in  $\mathcal{M}_\odot \text{ yr}^{-1}$ ) of 0.3 (1500 Å) and 0.7 (2800 Å) assuming a Salpeter (1955) initial mass function, and 0.6 (1500 Å) or 1.2 (2800 Å) assuming a Scalo (1986) initial mass function. These values assume that there is no extinction in the host galaxy. This is clearly not correct, so they actually represent a lower limit on the star-formation rate of approximately  $1 \mathcal{M}_\odot \text{ yr}^{-1}$ . Correcting for the mean extinction in the host ( $A_v = 3.5 \text{ mag}$ ) gives a star-formation rate of  $\approx 10^2\text{--}10^4 \mathcal{M}_\odot \text{ yr}^{-1}$ . However, these corrected values should be treated with caution since extinction in the host is likely to be highly variable and may not be affecting all star-forming regions equally.

The host galaxy appears to be roughly similar in luminosity to the Milky Way, but with a significantly higher star-formation rate. This puts it on the massive end of the GRB host galaxy distribution.

## 4. The Gamma-Ray Burst and its Afterglow

### 4.1. Is GRB 090417B a Dark Burst?

No afterglow has been detected for GRB 090417B at infrared, optical, or ultraviolet wavelengths. Changes in the near-ultraviolet flux from the host galaxy between early and late times are consistent with no contribution from an afterglow down to a limiting magnitude of  $u_{\text{oa}} > 24.9$ . We are assuming that the bright limit on the afterglow magnitude is valid for all time after four hours after the BAT trigger. There is no evidence for a brightening of the host galaxy at late times, so this is a reasonable assumption. Therefore we adopt this as an upper limit to the luminosity of the afterglow at 11 hr and compare it to the  $X$ -ray flux at 11 hr. This yields  $\beta_{\text{OX}} \leq -1.9$ , making it a dark burst (Jakobsson et al. 2004). The  $X$ -ray spectrum at this time is evolving from  $\beta_X = 1.3$  to  $\beta_X = 2.3$ , so using the definition of van der Horst et al. (2009) we find that  $\beta_{\text{OX}} < \beta_X - 0.5$ , also making GRB 090417B a dark burst. Our limit on  $\beta_{\text{OX}}$  is somewhat uncertain due to the large uncertainty in our estimate of  $u_{\text{oa}}$ .

Our computation of the spectral index between the optical and  $X$ -ray régimes depends on the limiting magnitude of the afterglow in the UVOT  $u$  band as determined from the statistical error in the observed magnitude of the host galaxy at two epochs. In order to test that GRB 090417B really is a dark burst we compute the expected UVOT  $u$ -band magnitude for the afterglow given the observed  $X$ -ray spectrum and assuming that the cooling frequency is between  $X$ -ray and optical wavelengths. This yields a prediction of  $u \lesssim 21.7$  at 11 hours, which is ruled out by the observations. Therefore, we conclude that GRB 090417B is a dark

GRB.

The difference between the predicted  $u$  magnitude and the upper limit that we derive in § 2.3 is  $\Delta u \gtrsim 3.2$  mag. If we assume a Milky Way extinction law this corresponds to  $A_V \gtrsim 2.5$  mag of extinction in the host galaxy along the line of sight to the burst.

## 4.2. Closure Relations

If we assume the canonical  $X$ -ray light curve for GRB afterglows (Nousek et al. 2006; Zhang et al. 2006) then we can use the temporal decay indices and the spectral slopes to constrain the underlying physics. We stress that the following analysis is only valid if we assume the relativistic fireball model. Alternate models for the  $X$ -ray afterglows of GRBs are available, such as the cannonball model (Dar & De Rújula 2000a,b; Dado et al. 2002a,b), the disc model of Cannizzo & Gehrels (2009), the late internal emission model (Ghisellini et al. 2007; Kumar et al. 2008), the long-lived reverse shock model (Genet et al. 2007; Uhm & Beloborodov 2007), and the prior outflow emission model (Yamazaki 2009; Liang et al. 2009).

The time and post-break decay index of the late-time break in the  $X$ -ray light curve at  $4.5_{-1.2}^{+2.8} \times 10^5$  s after the BAT trigger are consistent with a jet break. However, the dust scattering interpretation of the late-time spectral softening (see § 4.3) suggests that the steepening of the light curve after  $\approx 4.5 \times 10^5$  s is due to scattered emission due to dust dominating over the synchrotron component of the afterglow. Unfortunately there is no data in other wavelength régimes to test if the late-time steepening is achromatic (which would argue for a jet break), therefore we can not assume that this is a jet break.

Sari et al. (1999) give the relationships between the temporal and spectral slopes and the electron power-law distribution index,  $p$ , for a blast expanding into a homogeneous external medium, and Chevalier & Li (1999) give the relationships for a stellar wind ( $\rho(r) \propto r^{-2}$ ) environment. Both of these papers assume that  $p > 2$ . For the case were  $p < 2$  the relationships of Dai & Cheng (2001) are used. Panaitescu & Kumar (2002) found that GRB afterglows exhibit a range of electron distributions with values for the ten GRBs they studied being between  $p = 1.4$  and  $p = 2.8$ .

The  $X$ -ray spectral index during the first part of phase III of the afterglow ( $9100 \leq t \leq 29\,600$  s) is  $\beta_X = \Gamma - 1 = 1.3 \pm 0.2$ . If the cooling frequency is above the  $X$ -ray régime during this time then the predicted electron index is  $p = 2\beta_X + 1 = 2(1.3 \pm 0.2) + 1 = 3.6 \pm 0.4$ , which is well outside the range of values found by Panaitescu & Kumar (2002). If the cooling frequency is below the  $X$ -ray band during this time then  $p = 2\beta_X = 2(1.3 \pm 0.2) = 2.6 \pm 0.4$ ,

which is consistent with the Panaitescu & Kumar (2002) range of electron indices. For  $\nu_c < \nu_X$  the spectral index predicts a temporal decay of  $\alpha_X = 3/2\beta_X - 1/2 = 3/2(1.3 \pm 0.2) - 1/2 = 1.45 \pm 0.30$  for either a homogeneous interstellar medium (ISM) or a stellar wind environment. The observed  $X$ -ray decay during this time is  $\alpha_X = 1.40 \pm 0.04$ , which is consistent with the predicted value. Therefore, we conclude that the cooling frequency is most likely below the  $X$ -ray band at  $9100 \leq t \leq 29600$  s. Unfortunately it is not possible to tell, from the  $X$ -ray data, if the burst is expanding into a uniform external medium or a pre-existing stellar wind. The broad agreement between  $\alpha_2$  and  $\beta_2$  with the predictions of the fireball model suggests that this is a reasonable explanation for the physics of this afterglow during this phase of the afterglow.

Between 68980 s and 428400 s the spectral index softens to  $2.8 \pm 0.5$ . In this interval the fireball model predicts  $\alpha_X = 3.7 \pm 1.0$  for a homogeneous ISM and a wind if  $\nu_c < \nu_X$ . If  $\nu_c > \nu_X$  then the environment must be windy (since  $\nu_c$  is increasing) and  $\alpha_X$  is predicted to be  $4.7 \pm 1.0$ . The observed value is  $\alpha_2 = 1.40 \pm 0.04$ . Neither case is consistent with the data in this time interval.

The data before 68980 s is broadly consistent with a relativistic fireball expanding into either a homogeneous ISM or a wind-stratified external medium, and a cooling frequency below the  $X$ -ray band in the period  $9100 \leq t \leq 29600$  s. However, this model can not explain the data later than 68980 s. Some other physical process is needed to explain the late-time softening of the  $X$ -ray spectrum.

### 4.3. Dust Scattering

For GRBs occurring in dusty star forming regions, the  $X$ -rays emitted during the prompt phase may be scattered by dust grains in small angles, which leads to a delayed  $X$ -ray afterglow component (Mészáros & Gruzinov 2000). Placing circumburst dust at  $\approx 10$ – $100$  pc away from the burst, the calculated dust echo  $X$ -ray lightcurve displays a shallow decay lightcurve followed by late steepening, as is commonly observed in many GRBs (Shao & Dai 2007; Shao et al. 2008). One major feature of this model is the strong spectral softening evolution in the echo emission (Shen et al. 2009). Such a softening feature has not commonly been observed in the majority of GRBs, so the dust scattering model is not favored to interpret those bursts. However, GRB 090417B shows a spectral softening during the shallow decay phase. Also the optically dark nature of the afterglow suggests that there is a substantial amount of dust in the circumburst environment. All these make this burst a very probable candidate for the dust scattering model. In modeling this burst, we assume a dust sheet with distance  $R_d$  from the burst, with other dust grain properties the same as

in Figure 1 of Shen et al. (2009). We find the calculated light curve can reproduce quite well the XRT light curve from  $T - T_0 = 5 \times 10^3$  s to  $10^5$  s (we consider the flares prior to this time as not being due to dust scattering because the  $\gamma$ -ray emission was still on at that time) for  $R_d \approx 30$ – $80$  pc and for a broad range of the source emission spectral index values.

A strong test of the model is whether the spectral index evolution can be reproduced. In Figure 4 we plot the temporal evolution of the observed XRT spectral index. The dust scattered emission spectral index is calculated in the same way as in Shen et al. (2009). The amount of softening predicted by the model is large, i.e.,  $\Delta\beta \approx 3$ – $4$ , and it is independent of the source (prompt) emission spectral index. We tried to use the reported prompt BAT spectral index,  $\beta_{BAT} = 0.89$ , and the even harder spectral indices observed in the two  $X$ -ray flares ( $T - T_0 = 400 - 2 \times 10^3$  s),  $\beta_X = -1$ , as the source  $X$ -ray spectral index, but found that the model predicted emission is significantly softer than what is observed. However, if one assumes an even harder spectral index of  $\beta_X = -2$  for the prompt  $X$ -rays, both the  $X$ -ray light curve and spectral evolution can be properly reproduced (see Figure 4). Such a hard  $X$ -ray spectrum is consistent with the self-absorbed regime of a synchrotron or inverse Compton scattering spectrum. We therefore assume that the GRB prompt emission spectrum has a break above the  $X$ -ray band due to the self-absorption frequency,  $\nu_a$ , below which the spectrum has  $\beta = -2$  or  $-5/2$  (e.g., Granot & Sari 2002). A large  $\nu_a$  is possible if the prompt emission radius is small enough (e.g., Shen & Zhang 2009). The best model fit to the lightcurve and spectral evolution is shown in Figure 4. The agreement is statistically acceptable. We also estimate the amount of dust in terms of  $\tau_0$ —the scattering optical depth to the 1 keV photons, using the same formula as in Shen et al. (2009). The source  $X$ -ray fluence is extrapolated from the observed  $\gamma$ -ray fluence using the assumed spectral form, and the scattered emission  $X$ -ray fluence is estimated from the observed XRT fluence for  $T - T_0 \geq 5 \times 10^3$  s using an intermediate value of observed spectral index  $\beta_X = 1.5$ . We find  $\tau_0 \approx 2.5$ . This means that over half of prompt 1 keV photons would be scattered by the dust. This  $\tau_0$  value also implies  $A_V \approx 15 - 40$  mag, according to some empirical relations between the two dust properties used in Shen et al. (2009). This  $A_V$  is much higher than the average value for the host galaxy that was derived from the spectral energy distribution in Section 3.2, but is consistent with the extinction derived from the  $X$ -ray measurement of the hydrogen column along the line of sight to the burst,  $A_V \approx 11$  mag. This high line-of-sight extinction naturally explains the non-detection of an optical afterglow for GRB 090417B.

In conclusion, the data of GRB 090417B is consistent with the dust scattering model under the condition that the prompt  $X$ -rays have a self-absorbed spectrum in the  $X$ -ray range. The large amount of dust inferred from the  $X$ -ray modeling is consistent with the dark nature of the optical afterglow.

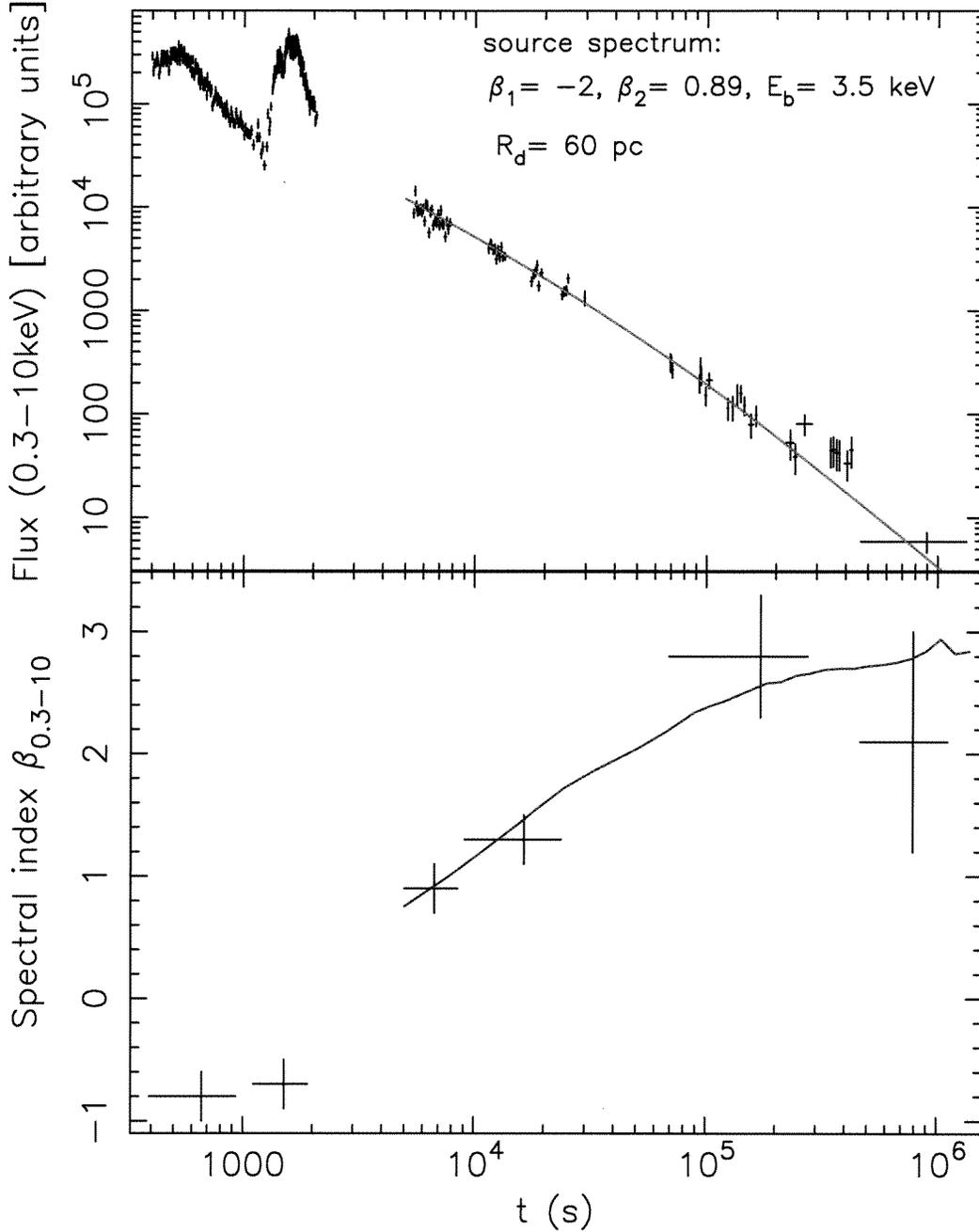


Fig. 4.— Dust scattering model fit to GRB 090417B X-ray data. Top panel: the fit to the light curve. Bottom panel: the fit to the spectral index evolution. The model predictions are shown as solid lines and data as crosses. The model parameter values are same for both panels, and are as labeled. The source spectrum is modeled as a two-piece power law:  $f_\nu \propto \nu^{-\beta_1}$  and  $f_\nu \propto \nu^{-\beta_2}$  with a break  $E_b$ . In the fit, the low energy spectral index  $\beta_1 = -2$  is adopted, suggesting a self-absorbed prompt emission spectrum.  $\beta_2$  is adopted to be equal to the reported prompt  $\gamma$ -ray spectral index.

#### 4.4. Energetics of the Burst

BAT observations of GRB 090417B were cut off at approximately 2105 s after the BAT trigger when *Swift* went into an Earth limb constraint. Therefore, the total duration, and thus the total fluence of the burst is not known. However, the fluence while the burst was observed was  $8.20_{-2.1}^{+1.0} \times 10^{-6}$  erg cm<sup>-2</sup> in the 15–150 keV band. This puts a lower limit on the total fluence. We can use this lower limit to estimate a lower limit on the isotropic equivalent energy of GRB 090417B. The BAT spectrum is best fit by a single power law, so at a redshift of  $z = 0.345$  this corresponds to an isotropic equivalent energy of  $E_{\text{iso}} > 2.4_{-0.6}^{+0.3} \times 10^{51}$  erg. We performed a  $k$  correction following the prescription of Bloom et al. (2001) and find  $k = 2.6$  (with some uncertainty due to the arbitrary choice of cut-off energy for the power-law spectrum). This yields a  $k$ -corrected total isotropic energy of  $E_{\text{iso}} \gtrsim 6.3 \times 10^{51}$  erg in the 20–2000 keV band.

With this estimate on the lower limit of the burst energy, and the constraint that any jet break occurs at no earlier than  $\approx 4.5 \times 10^5$  s, the opening angle of the jet ( $\theta_j$ ) can be constrained (Rhoads 1999; Sari et al. 1999; Frail et al. 2001). For  $t_j \geq 4.5_{-1.2}^{+2.8} \times 10^5$  s ( $\geq 5.2$  days),  $z = 0.345$ , and  $E \approx 6.3 \times 10^{51}$  erg the half-opening angle is

$$\theta_j \gtrsim 0.173(\eta_\gamma/0.2)^{1/8}(n/0.1)^{1/8}\text{rad}, \quad (1)$$

where  $\eta_\gamma$  is the efficiency of converting energy in the ejecta into gamma rays, and  $n$  is the circumburst particle density in cm<sup>-3</sup>. Assuming  $\eta_\gamma = 0.2$  and  $n = 0.1$  cm<sup>-3</sup> (e.g., Frail et al. 2001) we find  $\theta_j \gtrsim 10$  deg. We note that varying these two parameters only has a small effect on the derived jet opening angle. The uncertainty in the isotropic energy has a smaller effect on the half-opening angle than the uncertainty in the jet break time, so it is difficult to make  $\theta_j$  significantly smaller than  $\approx 10$  deg without having more than half of the total energy in the burst being emitted after  $\approx 2000$  s. Therefore we consider 10 deg to be a reasonable estimate of the lower limit of the half-opening angle of the jet. This is consistent with what has been seen for many other afterglows.

The lower limits on the isotropic energy and lower limit on the jet half-opening angle described above correspond to a lower limit on the true gamma-ray energy for GRB 090417B of  $E_\gamma \gtrsim 9.6 \times 10^{49}$  erg. This makes GRB 090417B approximately ten times less energetic than the canonical GRB (Frail et al. 2001), although the discrepancy could be smaller if the jet break occurred significantly later than  $\approx 4.5 \times 10^5$  s. An energy of  $E_\gamma \gtrsim 10^{50}$  erg is on the low end of the observed distribution of  $E_\gamma$  (Kaneko et al. 2007, their Fig. 23), but not unknown. Further, if the jet break occurs significantly later than 5.2 days the energy comes more into agreement with the canonical energy. Therefore, we find no reason to think

that the low  $E_\gamma$  for GRB 090417B is due to it being at a significantly higher redshift than  $z = 0.345$ .

#### 4.5. Constraints on a Supernova Component

UVOT observations were taken up to 15 days after the BAT trigger, so we searched the late-time data for evidence of a supernova component. A type Ib/c supernova like SN1998bw (Patat & Piemonte 1998) at a redshift of  $z = 0.345$  is expected to peak at approximately  $10(1+z) \approx 13$  days after the burst. We find no evidence for a change in the luminosity of the host galaxy down to  $\Delta u \approx 0.2$  mag, which corresponds to an upper limit on the magnitude of the supernova of  $u_{\text{lim}} \gtrsim 24.9$ . At  $z = 0.345$  the central wavelength of the UVOT  $u$  band corresponds to  $\lambda_0 \approx 2600$  Å. No observational data is available for the peak magnitude of SN1998bw at 2600 Å. However *Swift*/UVOT ultraviolet light curves are available for SN2007Y (Brown et al. 2009), a SN Ib/c with a peak absolute magnitude of  $uvw1 \approx -17$  in the rest frame. This corresponds to a rest frame  $uvw1$  of  $\approx 27$  at the redshift of GRB 090417B. This is  $\approx 2$  mag fainter than our limiting magnitude. Therefore, if we assume that SN2007Y is typical of the SNe Ib/c associated with GRBs, we are not able to constrain the existence of a supernova associated with GRB 090417B.

### 5. Discussion

All of the evidence presented in this paper points to GRB 090417B being a dark burst because of localized extinction in the host galaxy. The host galaxy has a redshift of  $z = 0.345$ , and is clearly seen at optical and ultraviolet wavelengths. Therefore Lyman- $\alpha$  absorption can not account for the missing optical afterglow. The  $X$ -ray light curve and spectrum during the first  $\approx 70$  ks obey the conventional closure relationships and are consistent with synchrotron radiation from a relativistic fireball expanding into an external medium (although the structure of that medium can not be determined with the available data). Optical afterglows have been detected for many GRBs with similar  $X$ -ray properties, so it is unlikely that the lack of an optical detection is due to unusual physics during the afterglow phase of the burst. Optical observations of GRB 090417B started 378 s after the BAT trigger, and follow-up observations were made from several observatories for several days after the burst. These observations went deep enough to have detected an optical afterglow if one had been present with a luminosity similar to that of other optical afterglows. Therefore, the dark nature of GRB 090417B can not be explained by a lack of observations.

The dust hypothesis, however, can explain this dark burst. The  $X$ -ray spectrum is consistent with an H I column density of  $N_{\text{H}} \approx 2 \times 10^{22} \text{ cm}^{-2}$ , which corresponds to an extinction of  $A_v \approx 11$  mag along the line of sight in the host galaxy, assuming a Milky Way extinction law as explained in § 2.2. This value is in agreement with the minimum extinction, as explained in § 4.1, that is needed to obscure the expected bright optical afterglow ( $A_V \gtrsim 2.5$  mag). An extremely high extinction ( $A_V \approx 15\text{--}40$  mag) is also predicted by the dust scattering model, which is consistent with the  $X$ -ray data. Such a large amount of dust along the line of sight to the GRB naturally explains the lack of any detection at ultraviolet, optical, or near-infrared wavelengths while allowing a conventional  $X$ -ray afterglow to be observed. It is also consistent with the result of Perley et al. (2009) who find large mean extinctions for the putative host galaxies of several GRBs. Our result, that the extinction along the line of sight in the host GRB 090417B is likely to be at least ten magnitudes in the UVOT  $v$  band, provides strong evidence that local dust concentrations in the host galaxy are responsible for at least some low-redshift dark bursts.

Unlike many GRB host galaxies the dust in the host galaxy of GRB 090417B appears to follow a Milky Way extinction law. In addition the host appears to be an  $L^*$  galaxy with a star-formation rate that exceed  $\approx 1 \mathcal{M}_{\odot} \text{ yr}^{-1}$ . In other words, the host appears to have dust properties, and a luminosity that are similar to those of the Milky Way, and a star-formation rate that is consistent with what is seen in other GRB host galaxies.

GRB 090417B is one of the few dark GRBs where the association with a host galaxy is secure. In most cases associations between a GRB and a galaxy have been made based on the probability of the nearest observed galaxy lying as close to the centre of the  $X$ -ray error circle as it does. There are a few exceptions, such as GRB 000210, GRB 050713A and XRF 050416A. The host galaxy of GRB 000210 was identified from the locations of its radio and  $X$ -ray afterglow. Piro et al. (2002) find a probability of a chance alignment of the galaxy and the burst of  $P_{\text{ch}} = 0.016$  and conclude that if this is the host then the dark nature of GRB 000210 is likely due to dust either at the location of the progenitor or along the line of sight. XRF 050416A occurred at a low redshift ( $z = 0.6535$  Cenko et al. (2005)). Although technically dark an optical afterglow was detected and was consistent with the standard fireball model (e.g., Holland et al. 2007). The presence of an optical afterglow and the low line-of-sight extinction suggests that this burst should not be considered to be a dark GRB. The line-of-sight extinction for XRF 050416A is  $A_V < 1$  mag (Holland et al. 2007; Perley et al. 2009), which is consistent with what is seen for most optically detected GRB afterglows. Perley et al. (2009) note that GRB 050713A also only barely qualifies as a dark burst under the Jakobsson et al. (2004) definition. The internal line-of-sight extinction for this burst is unknown. However, an optical afterglow was seen (Wren et al. 2005), so it is unlikely that the extinction was large. The latter two of these bursts are technically

dark based on the spectral slope between the  $X$ -ray and optical regimes ( $\beta_{OX}$ ) criteria of Jakobsson et al. (2004). However, in these cases optical or infrared afterglows were detected, suggesting that these GRBs were borderline cases of dark bursts.

Further evidence that dust is responsible for GRB 090417B (and perhaps many other dark GRBs) being dark comes from the late-time softening of the  $X$ -ray spectrum. This softening is consistent with what is predicted from the dust scattering model (Shao & Dai 2007; Shao et al. 2008; Shen et al. 2009), and requires approximately 15–40 mag of extinction at optical wavelengths. The effects of dust scattering do not become apparent in the  $X$ -ray spectrum until  $\approx 10^5$  s after the burst. Since  $X$ -ray spectra are usually extracted from early-time data (when there are a large number of photons) it is possible that the softening seen for GRB 090417B has been present in the late-time  $X$ -ray spectra of other dark bursts, but has not been seen due to low count rates. For example, the standard catalogue of *Swift*/XRT spectral fits (Evans et al. 2009) uses data taken within  $\approx 4.3$  ks of the BAT trigger, well before a softening due to dust scattering is expected to appear.

Perley et al. (2009) point out that the host galaxies of dark GRBs are not unusually reddened relative to the hosts of optically bright GRBs. Their analysis assumes that their statistical association of bursts with the nearest detected galaxy on the sky reveals the true hosts of these bursts. It is uncertain if this is valid for all of the GRBs in their sample. However, it is likely that at least some (if not all) of these galaxies actually did host the associated GRBs. Although some of the hosts of dark GRBs have been found to be unusually dusty compared to the hosts of optically-bright GRBs (e.g., GRB 030115 (Levan et al. 2006)), there is no evidence that the hosts of dark GRBs have systematically higher mean extinctions than the hosts of optically-bright GRBs do. However, the mean extinction of the host is probably not the factor that determines if a particular burst is dark or not since any extinction due to material that is not along the line of sight to the burst will not affect our observations of the afterglow. What is important is the extinction along the line of sight. This extinction can be due to dust at the location of the GRB, such as a dusty star-forming region or a giant molecular cloud, or it could be due to unrelated dust in the host that just happens to lie along the line of sight.

For GRB 090417B the late-time evolution of the  $X$ -ray spectrum can be explained if there is a sheet of dust at a distance of  $20 \lesssim R_d \lesssim 80$  pc from the burst. This suggests that GRB 090417B may have been dark because of a chance alignment of the progenitor with a region of high extinction in the host galaxy. The Galactic extinction model of Drimmel et al. (2003) suggest that roughly 25% of the sight lines through the Milky Way have  $A_v \gtrsim 1$  mag. This is within a factor of a few of the fraction of GRBs that are dark ( $\approx 40\%$  (Roming et al. 2006)), which suggests that obscuration by dust along the line of sight may be responsible

for a significant fraction of dark GRBs.

## 6. Conclusions

We have shown that the host of the dark GRB 090417B is very likely to be SDSS J135846.65+470104 an  $L^*$  galaxy at  $z = 0.345$ . This galaxy has an overall star-formation rate that is at least as great as that of the Milky Way, and an overall  $V$ -band extinction of  $A_V = 3.5_{-0.5}^{+1.0}$  mag, which is dustier than typical GRB host galaxies, but not significantly different from that of the Milky Way.

$X$ -ray observations of GRB 090417B show the normal temporal and spectral behaviour seen in the  $X$ -ray afterglows of many *Swift* GRBs. The data up to approximately 70 ks after the BAT trigger are consistent with a relativistic fireball expanding into either a homogeneous ISM or a wind-stratified external medium. The cooling frequency is below the  $X$ -ray band between 9.1 ks and 29.6 ks. The  $k$ -corrected isotropic energy is  $E_{\text{iso}} \gtrsim 6.3 \times 10^{51}$  erg and the jet half-opening angle is  $\theta_j \gtrsim 10^\circ$ . The physics of this burst do not appear to be unusual, and thus are unlikely to explain the dark nature of the burst.

After approximately 70 ks the  $X$ -ray spectrum becomes significantly softer. We find that this can be explained using the dust scattering model of Shao & Dai (2007) and Shen et al. (2009). The observed late-time spectral evolution can be produced by a sheet of dust approximately 30–80 pc from the burst along the line of sight. The model predicts an extinction of  $A_V \approx 15$ –40 mag along the line of sight. This is consistent with the H I column density measured from the  $X$ -ray spectrum. Therefore, we conclude that GRB 090417B is probably dark because of a dense, localized layer of dust along the line of sight between us and the afterglow.

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## REFERENCES

Aoki, K., Tanaka, I., Nakata, F., Ohta, K., Yuma, S., & Kawai, N., 2009, GCNC 9145

- Barthelmy, S. D., Barbier, L. M., Cummings, J. R., Fenimore, E. E., Gehrels, N., Hullinger, D., Krimm, H. A., Markwardt, C. B., et al., 2005, *Sp. Sci. Rev.*, 120, 143
- Berger, E., & Fox, D. B., 2009, GCNC 9156
- Berger, E., Fox, D. B., Kulkarni, S. R., Frail, D. A., & Djorgovski, S. G., 2007, *ApJ*, 660, 504
- Berger, E., Fox, D. B., & Tanvir, N., 2009, GCNC 9158
- Bertin, E., & Arnouts, S., 1996, *A&AS*, 117, 393
- Bloom, J. S., Frail, D. A., & Sari, R., 2001, *AJ*, 121, 2879
- Bloom, J. S., Kulkarni, S. R., & Djorgovski, S. G., 2002, *AJ*, 123, 1111
- Brown, P. J., Holland, S. T., Immler, S., Milne, P., Roming, P. W. A., Gehrels, N., Nousek, J., Panagia, N., et al., 2009, *AJ*, 137, 4517
- Burrows, D. N., Hill, J. E., Nousek, J. A., Kennea, J. A., Wells, A. A., Osborne, J. P., Abbey, A. F., Beardmore, A., et al., 2005, *Sp. Sci. Rev.* 120, 165
- Campana, S., Mangano, V., Blustin, A. J., Brown, P., Burrows, D. N., Chincarini, G., Cummings, J. R., Cusumano, G., et al., 2006, *Nature*, 442, 1008
- Cannizzo, J. K., & Gehrels, N., 2009, *ApJ*, in press, arXiv:0901.3564
- Cash, W., 1979, *ApJ*, 228, 939
- Cenko, S. B., Kulkarni, S. R., Gal-Yam, A., & Berger, E., 2005, GCNC 3542
- Chevalier, R. A., & Li, Z.-Y., 1999, *ApJ*, 520, L29
- Christensen, L., Hjorth, J., & Gorosabel, J., 2004, *A&A*, 425, 913
- Conselice, C. J., Vreeswijk, P. M., Fruchter, A. S., Levan, A., Kouveliotou, C., Fynbo, J. P. U., Gorosabel, J., Tanvir, N. R., et al., 2005, *ApJ*, 633, 29
- Cummings, J. R., Barthelmy, S. D., Baumgartner, W. H., Fenimore, E. E., Gehrels, N., Krimm, H. A., Markwardt, C. B., Palmer, D. M., et al., 2009, GCNC 9139
- Dado, S., Dar, A., & De Rújula, A., 2002a, *A&A*, 388, 1079
- Dado, S., Dar, A., & De Rújula, A., 2002b, astro-ph/0204474

- Dai, Z. G., & Cheng, K. S., 2001, *ApJ*, 558, L109
- Dar, A., & De Rújula, A., 2000a, *astro-ph/0008474*
- Dar, A., & De Rújula, A., 2000b, *astro-ph/0012227*
- de Vaucouleurs, G., 1948, *AnAp*, 11, 247
- de Vaucouleurs, G., 1959, *HDP*, 53, 275 & 311
- Drimmel, R., Cabrera-Lavers, A., & López-Corredoira, M., 2003, *A&A*, 409, 205
- Evans, P. A., Beardmore, A. P., Page, K. L., Osborne, J. P., O’Brien, P. T., Willingdale, R., Starling, R. L. C., Burrows, D. N., et al., 2009, *MNRAS*, 397, 1177
- Fishman, G. J., Meegan, C. A., et al., 1989, *Proc. GRO Science Workshop, GSFC*, 2
- Frail, D. A., Kulkarni, S. R., Sari, R., Djorgovski, S. G., Bloom, J. S., Galama, T. J., Reichart, D. E., Berger, E., et al., 2001 *ApJ*, 562, L55
- Fruchter, A. S., Levan, A. J., Strolger, L., Vreeswijk, P. M., Thorsett, S. E., Bersier, D., Burud, I., Castro Cerón, A. J., et al., 2006, *Nature*, 441, 463
- Fynbo, J. P. U., Jensen, B., Gorosabel, J., Hjorth, J., Pedersen, H., Møller, P., Abbott, T., Castro-Tirado, A. J., et al., 2001, *A&A*, 369, 373
- Fynbo, J. P. U., Malesani, D., Hjorth, J., Leitet, E., Linne, S., Ottosen, T. A., & Jakobsson, J., 2009a, *GCNC* 9140
- Fynbo, J. P. U., Jakobsson, P., Procheska, J. X., Malesani, D., Ledoux, C., de Ugarte Postigo, A., Naradini, N., Vreeswijk, P. M., et al., 2009b, *ApJS*, submitted, *astro-ph/0907.3449*
- Gehrels, N., Chincarini, G., Ciommi, P., Mason, K. O., Nousek, J. A., Wells, A. A., White, N. E., Barthelmy, S. D., et al., 2004, *ApJ*, 611, 1005
- Genet, F., Daigne, F., & Mochkovitch, R., 2007, *MNRAS*, 381, 732
- Ghisellini, G., Ghirlanda, G., Nava, L., & Firmani, C., 2007, *ApJ*, 658, L75
- Goad, M. R., Tyler, L. G., Beardmore, A. P., Evans, P. A., Rosen, S. R., Osborne, J. P., Starling, R. L. C., Marshall, F. E., et al., 2007, *A&A*, 476, 1401
- Graham, A. W. & Driver, S. P., 2005, *PASA*, 22, 118

- Granot, J. & Sari, R., 2002, *ApJ*, 568, 820
- Guidorzi, C., Smith, R., Mundell, C. G., Gomboc, A., O'Brien, P., & Tanvir, N., 2009, GCNC 9144
- Hogg, D. W., Pahre, M. A., McCarthy, J. K., Cohen, J. G., Blandford, R., Smail, I., & Soifer, B. T., 1997, *MNRAS*, 288, 404
- Holland, S. T., Boyd, P. T., Gorosabel, J., Hjorth, J., Schady, P., Thomsen, B., Augusteijn, T., Blustin, A. J., et al., 2007, *AJ*, 133, 122
- Jakobsson, P., Hjorth, J., Fynbo, J. P. U., Watson, D., Pedersen, K., Björnsson, G., & Gorosabel, J., 2004, *ApJ*, 617, L21
- Jaunsen, A. O., Rol, E., Watson, D. J., Malesani, D., Fynbo, J. P. U., Milvang-Jensen, B., Hjorth, J., Vreeswijk, P. M., et al., 2008, *ApJ*, 681, 453
- Kalberla, P. M. W., Berton, W. B., Hartmann, D., Arnal, E. M., Bajaja, E., Morras, R., Pöppel, W. G. L., 2005, *A&A*, 440, 775
- Kaneko, Y., Ramirez-Ruiz, E., Granot, J., Kouveliotou, C., Woosley, S. E., Patel, S. K., Rol, E., in't Zand, J. J. M., et al., 2007, *ApJ*, 654, 385
- Kann, D. A., Klose, S., & Zeh, A., 2006, *ApJ*, 641, 993
- Kouveliotou, C., Meegan, C. A., Fishman, G. J., Bhat, N. P., Briggs, M. S., Koshut, T. M., Paciasas, W. M., & Pendleton, G. N., 1993, *ApJ*, 413, 101
- Kron, R. G., 1980, *ApJS*, 43, 305
- Kuman, P., Narayan, R., & Johnson, J. L., 2008, *MNRAS*, 388, 1729
- Le Floc'h, E., Duc, P.-A., Mirable, I. F., Sanders, D. B., Bosch, G., Rodrigues, I., Covrvoisier, T. J.-L., Mereghetti, S., et al., 2002, *ApJ*, 581, L81
- Le Floc'h, E., Duc, P.-A., Mirable, I. F., Sanders, D. B., Bosch, G., Diaz, R. J., Donzelli, C. J., Rodrigues, I., et al., 2003, *A&A*, 400, 499
- Levan, A., Fruchter, A., Rhoads, J., Mobasher, B., Tanvir, N., Gorosabel, J., Rol, E., Kouveliotou, C., et al., 2006, *ApJ*, 647, 471
- Liang, E.-W., Lv, H.-J., Zhang, B.-B., & Zhang, B., 2009, *ApJ*, submitted, arXiv:0902.3504
- Lilly, S. J., Tresse, L., Hammer, F., Crampton, D., & Le Fèvre, O., 1995, *ApJ*, 455, 108

- Madau, P., Pozzetti, L., & Dickinson, M., 1998, *ApJ*, 498, 106
- Mangano, V., Barthelmy, S. D., Baumgartner, W. H., Burrows, D. N., Evans, P. A., Gehrels, N., Guidorzi, C., Holland, S. T., et al., 2009, GCNC 9133
- Mészáros, P., & Grizinov, A., 2000, *ApJ*, 543, L35
- Nousek, J. A., Kouveliotou, C., Grupe, D., Page, K. L., Granot, J., Ramirez-Ruiz, E., Patel, S. K., Burrows, D. N., et al., 2006, *ApJ*, 642, 389
- Panaiteescu, A., & Kumar, P., 2002, *ApJ*, 571, 779
- Patat, F., & Piemonte, A., 1998, IAUC 6918
- Pei, Y. C., 1992, *ApJ*, 395, 130
- Perley, D. A., Cenko, S. B., Bloom, J. S., Chen, H.-W., Butler, N. R., Kocevski, D., Prochaska, J. X., Brodwin, M., et al., 2009, *AJ*, submitted, arXiv:0905.0001
- Piro, L., Frail, D. A., Gorosabel, J., Garmire, G., Soffitta, P., Amati, L., Andersen, M. I., Antonelli, L. A., et al., 2002, *ApJ*, 577, 680
- Poole, T. S., Breeveld, A. A., Page, M. J., Landsman, W., Holland, S. T., Roming, P. W. A., Kuin, N. P. M., Brown, P. J., et al., 2008, *MNRAS*, 383, 627
- Predehl, P., & Schmitt, J. H. M. M., 1995, *A&A*, 293, 889
- Rhoads, J. E., 1999, *ApJ*, 525, 737
- Rol, E., Wijers, R. A. M. J., Kouveliotou, C., Kaper, L., & Kaneko, Y., 2005, *ApJ*, 624, 868
- Roming, P. W. A., Kennedy, T. E., Mason, K. O., Nousek, J. A., Ahr, L., Bingham, R. E., Broos, P. S., Carter, M. J., et al., 2005, *Sp. Sci. Rev.*, 120, 95
- Roming, P. W. A., Schady, P., Fox, D. B., Zhang, B., Liang, E., Mason, K. O., Rol, E., Burrows, D. N., et al., 2006, *ApJ*, 652, 1416
- Roming, P. W. A., Koch, T. S., Oates, S. R., Porterfield, B. L., Vanden Berk, D. E., Boyd, P. T., Holland, S. T., Hoversten, E. A., et al., 2009, *ApJ*, 690, 163
- Salpeter, E. E., 1995, *ApJ*, 121, 161
- Salvaterra, R., Della Valle, M., Campana, S., Chincarini, G., Covino, S., D’Avanzo, P., Fernandez-Soto, A., Guidorzi, F., et al., 2009, *Nature*, submitted, arXiv:0906.1578

- Sari, R., Piran, T., & Halpern, J. P., 1999, *ApJ*, 519, L17
- Savaglio, S., Glazebrook, K., & Le Borgne, D., 2009, *ApJ*, 691, 182
- Sbarufatti, B., Barthelmy, S. D., Baumgartner, W. H., Burrows, D. N., Curran, P. A., Evans, P. A., Godet, O., Guidorzi, C., et al., 2009, GCNC 9135
- Scalo, J. N., 1986, *Fund. Cosm. Phys.*, 11, 1
- Schady, P., Mason, K. O., Page, M. J., De Pasquale, M., Morris, D. C., Romnanao, P., Roming, P. W. A., Immler, S., et al., 2007, *MNRAS*, 377, 273
- Schady, P., Page, M. J., Oates, S. R., Still, M., De Pasquale, M., Dwelly, T., Kuin, N. P. M., Holland, S. T., et al., 2009, *MNRAS*, submitted, arXiv:0910.2590
- Schlegel, D. J., Finkbeiner, D. P., & Davis, M., 1998, *ApJS*, 500, 525
- Sérsic, J.-L. 1963, *BAAA*, 6, 41
- Sérsic, J.-L. 1968, *Atlas de Galaxias Australes*, Cordoba: Observatorio Astronomico
- Shao, L., & Dai, Z. G., 2007, *ApJ*, 660, 1319
- Shao, L., Dai, Z. G., & Mirabal, N., 2008, *ApJ*, 675, 507
- Shen, R.-F., Willingale, R., Kumar, P., O'Brien, P. T., & Evans, P. A., 2009, *MNRAS*, 393, 598
- Shen, R.-F., & Zhang, B., 2009, *MNRAS*, in press, arXiv:0906.2597
- Strauss, M. A., Weinberg, D. H., Lupton, R. H., Narayanan, V. K., Annis, J., Bernardi, M., Blanton, M., Burles, S., et al., 2002, *AJ*, 124, 1810
- Tanvir, N. R., Levan, A. J., Rol, E., Starling, R. L. C., Gorosabel, J., Priddey, R. S., Malesani, D., Jakobsson, P., et al., 2008, *MNRAS*, 388, 1743
- Tanvir, N. R., Fox, D. B., Levan, A. J., Berger, E., Wiersema, K., Fynbo, J. P. U., Cucchiara, A., Krühler, T., Gehrels, N., et al., 2009, *Nature*, submitted, arXiv:0906.1577
- Trujillo, I., Förster Schreiber, N. M., Rudnick, G., Barden, G., Franx, M., Rix, H.-W., Caldwell, J. A. R., McIntosh, D. H., et al., 2006 *ApJ*, 650, 18
- Uhm, Z. L., & Beloborodov, A. M., 2007, *ApJ*, 665, L93

- van der Horst, A. J., Kouveliotou, C., Gehrels, N., Rol, E., Wijers, R. A. M. J., Cannizzo, J. K., & Racusin, J., 2009, *ApJ*, in press, arXiv:0905.0524
- Wainwright, C., Berger, E., & Penprase, B. E., 2007, *ApJ*, 657, 367
- Wren, J., Vestrand, W. T., Wozniak, P., White, R., & Evans, S., 2005, GCNC 3604
- Yamazaki, R., 2009, *ApJ*, 690, L118
- Yasuda, N., Fukugita, M., & Schneider, D. P., 2007, *AJ*, 134, 698
- Zhang, B., Fan, Y. Z., Dkys, J., Kobayashi, S., Mészáros, P., Burrows, D. N., Nousek, J. A., & Gehrels, N., 2006, *ApJ*, 642, 354

Table 3. This Table lists the UVOT photometry of SDSS J135846.65+470104.5, the host galaxy of GRB 090417B. Exposure is the total exposure time in each filter. All upper limits are  $3\text{-}\sigma$  upper limits.

Filter	Exposure	Mag	Err
<i>v</i>	1237	>20.9	...
<i>b</i>	862	>21.4	...
<i>u</i>	22 554	23.09	0.38
uvw1	10 850	23.12	0.56
uvm2	15 942	22.49	0.29
uvw2	30 827	22.69	0.19
white	841	>22.2	...

Table 4. This Table lists all of the available photometry of SDSS J135846.65+470104.5, the host galaxy of GRB 090417B.

Filter	Mag	Err	Source
uvw2	22.69	0.19	1
uvm2	22.49	0.29	1
uvw1	23.12	0.56	1
$u_{\text{UVOT}}$	23.09	0.38	1
$u'$	23.23	0.53	2
$b_{\text{UVOT}}$	>21.4	...	1
$B$	23.25	0.13	1
$g'$	22.02	0.16	2
$g'$	22.47	0.16	3
$v_{\text{UVOT}}$	>20.9	...	1
$R$	21.34	0.03	1
$r'$	21.9	0.3	4
$r'$	21.62	0.09	2
$r'$	21.62	0.10	3
$i$	20.82	0.07	1
$i'$	21.41	0.11	2
$i'$	21.31	0.12	3
$z'$	20.78	0.25	2
$z'$	21.49	0.30	3
$J$	20.0	...	5
$K_s$	18.5	...	5

References. — (1) This work; (2) SDSS DR7 (3) Berger et al. (2009); (4) Guidorzi et al. (2009); (5) Aoki et al. (2009)