Laboratory spectra of CO₂ vibrational modes in planetary ice analogs

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A B S T R A C T

Laboratory spectra have shown that CO₂ is a powerful diagnostic tool for analyzing infrared data from remote observations, as it has been detected on icy moons in the outer Solar System as well as dust grain surfaces in the interstellar medium (ISM). IR absorption band profiles of CO₂ within ice mixtures containing H₂O and CH₃OH change with respect to temperature and mixture ratios. In this particular study, the ν₃ CO₂ asymmetric stretch mode near 4.3 μm (2350 cm⁻¹), overtone mode near 1.97 μm (5080 cm⁻¹), and the combination bands near 2.7 μm (3700 cm⁻¹), 2.8 μm (3600 cm⁻¹), and 2.02 μm (4960 cm⁻¹) are systematically observed in different mixtures with H₂O and CH₃OH in temperature ranges from 15 K to 150 K. Additionally, some high-temperature deposits (T > 50 K) of H₂O, CH₃OH, and CO₂ ice mixtures were performed and it was discovered that CO₂ may deposit out at higher temperatures than previously recorded. These data may then be used to interpret infrared observational data obtained from icy surfaces in the outer Solar System and beyond.

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1. Introduction

Icy materials in dense clouds in the interstellar medium (ISM) are observed by the absorption of starlight through these clouds (e.g., Chiar et al., 1995). Such ices are typically dominated by H₂O-ice in an amorphous phase, or amorphous solid water (ASW) and often contain other simple species like CO, CO₂, CH₃OH, and NH₃ (e.g., Jenniskens et al., 1995; Lacy et al., 1998; Ehrenfreund et al., 1999; Gibb et al., 2000; Herbst, 2001). Infrared (IR) absorption spectroscopy in the near-IR (λ = ~0.9–2.5 μm) and mid-IR (λ = 2.5–25 μm) are powerful methods by which these icy components can be detected and studied. IR absorption band positions and profiles are used for identification of physical properties such as composition, and temperature Allamandola et al. (1992), Salama et al. (1994), Chiar et al. (1995), Ehrenfreund et al. (1999), Gerakines et al. (1999), Gibb et al. (2000), Bernstein et al. (2005), White et al. (2009), and Öberg et al. (2009). Temperatures of H₂O-dominated ice mixtures can range from 10 to 15 K in quiescent interstellar dust cloud environments to as high as 100–150 K around protostellar disks.

The properties of icy materials in the Solar System and dense clouds in the ISM are typically observed using infrared absorption features from transmitted or reflected light on icy surfaces (e.g., Ehrenfreund, 1999; Dartois et al., 1999; Sandford et al., 2001; Grundy et al., 2006; Cruikshank et al., 2010). Ices on planetary surfaces in the outer Solar System are commonly also dominated by H₂O (Fink et al., 1976; Jenniskens and Blake, 1996; Dartois and Schmitt, 2009; Cruikshank et al., 2010). Observations from instruments such as the Visual and Infrared Mapping Spectrometer (VIMS) instrument aboard the Cassini spacecraft have revealed that most of the H₂O-ice on outer Solar System surfaces, specifically jovian and saturnian satellites, is of crystalline phase (e.g., Cruikshank et al., 2010). As in ISM ices, additional compounds are often present in these H₂O-dominated ices (e.g., Grundy et al., 1999; Cruikshank et al., 2010). For example, spectroscopic data from the Near-infrared mapping spectrometer (NIMS) aboard the Galileo spacecraft that explored satellites such as Ganymede and Callisto indicated the presence of CO₂ within the host ice (Carlson et al., 1996; McCord et al., 1998; Hibbitts et al., 2000, 2002, 2003; Cruikshank et al., 2010). Temperatures of H₂O-dominated ice mixtures typically found on outer Solar System surfaces typically range from 50 to 160 K. These temperatures can vary with time and location, depending on sunlight exposure, local geological activity, etc.

CO₂ is important in studies of icy surfaces on planetary bodies in the Solar System because the IR absorption profiles of CO₂-containing ices created in the laboratory have been shown to be sensitive to conditions of the ice such as temperature, temperature history, ice composition, and radiation exposure (Sandford and Allamandola, 1990a; Ehrenfreund et al., 1998, 1999; Gerakines et al., 1999; Dartois et al., 1999; Moore et al., 2001; Hudson and Moore, 2000, 2001; van Broekhuizen et al., 2006; Hodys et al., 2008; White et al., 2009; Öberg et al., 2007, 2009). These investigations provide a relevant database indispensable to spectral analysis from remote measurements (e.g., Sandford and Allamandola, 1990a;
H₂O–ice mixture around 75 K due to pore collapse as it is heated, H₂O–ice mixtures undergo several phase changes as it is warmed from its amorphous phase a 10 K to its sublimation temperature (e.g., Tielens and Hagen, 1982; Smith, 2000; Baragiola, 2003). These changes can be observed via IR absorption spectroscopy. H₂O-rich ices that contain other molecules can also undergo phase transitions. For example, Öberg et al. (2009) have shown that H₂O + CO₂ ice mixtures undergo both surface and bulk segregation induced by thermal annealing, as observable through reflection–absorption infrared spectroscopy (RAIRS). As another example, thermal processing of H₂O–CH₃OH ices can spontaneously form clathrate hydrates that can trap volatiles such as CO₂ (Blake et al., 1991; Dartois and Schmitt, 2009). The impurity in the H₂O–ice cage may desorb into the ice matrix as the sample is heated (Hornekaer et al., 2005). Also, warming the ice leads to segregation of the H₂O component from CH₃OH in H₂O–CH₃OH mixtures containing H₂O crystals and CH₃OH crystals to form in separate regions (Blake et al., 1991; Tielens, 2005). Trapped volatiles may begin to escape from a H₂O–ice mixture around 75 K due to pore collapse as it is heated, though some of the trapped species may still remain. Above 120 K, a “glass transition” occurs and the remaining pores collapse until ~140 K when the crystalline H₂O structure transitions to a cubic phase, and there is a sudden release of the trapped species near 150 K. The glass transition precedes crystallization of an amorphous material and represents the temperature at which the molecules can begin to migrate more rapidly. Upon heating the amorphous phase ice, the onset of structural relaxation occurs between 120 and 142 K, depending on the heating rate and time since deposition (Jenniskens and Blake, 1996). Reorganization of the surface leads to collapse of the pores and reduction in the number of dangling OH bonds (Cyriac and Pradeep, 2008). Jenniskens and Blake (1996) interpreted the onset of the glass transition as a transformation to a new phase. In this phase, the increasing temperature allows the relaxation of the structure and removes strained bonding angles. Recent calorimetry experiments question whether the glass transition is even detectable (Yue and Angell, 2004). Although, other measurements confirmed the glass transition at 136 K (Kohl et al., 2000). A detailed review of glass transition in pure and aqueous solutions, can be found in Angell (2002).

In the new laboratory study described here, the IR spectra of ices bearing H₂O, CH₃OH, and CO₂ have been measured while systematically varying mixture ratios and temperatures that span the range of values expected on icy surfaces in the outer Solar System as well as the ISM. The experiments in this study also investigate the CO₂ absorptions from some binary and ternary ice mixtures in order to expand upon the existing databases and further facilitate understanding of observations from ground- and space-based instruments.

2. Experiments

2.1. Apparatus

The apparatus used for this study has been described in Hudgins et al. (1993), Bernstein et al. (2006), and Mastrapa et al. (2008). A high vacuum (~1 × 10⁻⁷ mbar) simulation chamber was placed within the sample compartment of a Bio Rad Excalibur FTIR Spectrometer with a spectral range of 0.910–22.2 μm (11,000–450 cm⁻¹) and a spectral resolution of 1 cm⁻¹. The multi-windowed vacuum chamber has IR-transparent KBr windows and houses a ZnSe substrate (window) that is attached to a water-cooled closed-cycle He cryocooler (Air Products Displex IRO2W) that can be cooled to temperatures as low as 15 K. The ZnSe substrate can be rotated within the sample chamber without breaking vacuum. The chamber houses four ports consisting of KBr windows through which the IR beam was sent and a UV transparent window for UV photolysis (not used in this study). The stainless steel vacuum chamber was evacuated using an oil diffusion pump (Edwards Diffstack Series 100) backed by a mechanical pump (Edwards E2MS). The temperature of the ZnSe window was controlled using a small resistive heater mounted on the cold finger just above the substrate and monitored with two Fe–Au/Fe/Au chromel thermocouples attached to the heater and sample holder. A temperature controller (Scientific Instruments 9650) was used to monitor and control sample temperatures. A diaphragm gauge was used on the system to measure sample bulk pressure, while thermocouple and ion gauges allowed constant monitoring of the pressure inside the vacuum chamber. To measure sample thickness, a HeNe laser (λ = 633 μm) was used with a photometer to count interference fringes as the ice was deposited (see Section 2.2).

2.2. Experimental procedure

Gases were mixed in glass bulbs using a separate greaseless glass manifold at room temperature using the same method as, e.g., Hudgins et al. (1993). Ambient pressures in the manifold were typically ~5 × 10⁻⁶ T and bulb pressures after preparation were typically ~19 T. Thus contaminants were limited to about one part in 4 × 10⁶. After the gas mixtures were prepared, the bulb was then mounted onto the cryocooler apparatus. The sample substrate in the cryovacuum system was cooled to the deposition temperature (15 or 50 K for most samples) and a background absorption spectrum was taken of the blank sample substrate. The gas mixture was then vapor-deposited onto the substrate, typically at a rate of about 0.1 μm min⁻¹. Sample thickness is calculated using the equation

\[ d = \frac{m_0}{2n_{ref} \cos \theta_{ref}} \]

where \( d \) is the thickness of the ice in μm, \( m \) is the number of fringes observed in the laser light during deposition, \( \lambda_0 \) is the laser wavelength (0.633 μm), \( n_{ref} \) is the index of refraction of the ice [1.32 for crystalline H₂O–ice (Hale and Querry, 1973), 1.29 for amorphous H₂O–ice (Westley et al., 1998; Dohnálek et al., 2003)], and \( \theta_{ref} \) is the angle of reflection in the ice (taken to be 45°). After the ice was fully deposited, its spectrum was obtained and ratioed to the background spectrum.

In many cases, the ices were subsequently warmed, or warmed and cooled, and additional spectra were taken at higher temperatures. Temperatures where spectra were measured were selected based on where ice segregation is known to occur (e.g., Blake et al., 1991; Jenniskens and Blake, 1996; Ehrenfreund et al., 1997, 1999; Gerakines et al., 1999). The temperature of the sample window inside the chamber was adjusted from 15 K up to 170 K at a rate of ~2 K min⁻¹ and cooled at a rate of ~10 K min⁻¹ in order to observe effects on the ice sample when the temperature is cycled. Samples were held at temperature while absorption spectra were collected. Warming often causes significant changes in the spectra of ices as molecules rearrange and warming rates faster than 2 K min⁻¹ can cause sudden loss of material. However, the cooling rate was not observed to make much difference in the spectral profiles. Thus a faster cooling rate was used. The CO₂ spectra typically reflect the highest temperature at which the ice has been if the
sample was heated (see Section 3). In some experiments involving mostly H$_2$O (where the ice is >60% H$_2$O), samples were deposited at 50 K to reduce the effect of scattering seen in low-temperature amorphous samples deposited at 15 K.

### 2.3. Ice compositions

The ice compositions and temperatures investigated in this study are summarized in Table 1. The mid-IR spectra, $\lambda = 2.5$–22 μm (7500–450 cm$^{-1}$), were measured for several different ice compositions at temperatures ranging from 15 to 150 K. In addition, spectra in the range $\lambda = 0.9$–3.5 μm (11,000–2850 cm$^{-1}$) were studied to identify temperature and mixture ratio effects on the CO$_2$ combination and overtone bands for comparison to previous results of mixtures of H$_2$O + CO$_2$ and CH$_3$OH + CO$_2$ (Bernstein et al., 2005). Thicker samples (~0.2–2 μm) were needed for mid-IR experiments so the asymmetric stretch mode of CO$_2$ was not saturated. Conversely, thinner samples (~2–4 μm) were used in near-IR experiments in order to view the much weaker combination and overtone modes. Some spectra were collected at the University of Alabama at Birmingham (UAB) Astrophysics Laboratory and overtone modes. Some spectra were collected at the University of Alabama at Birmingham (UAB) Astrophysics Laboratory$^1$ and are shown here for comparison. Experimental methods and related experiments from UAB are covered in more depth by White et al. (2009) and White (2010).

### Table 1

<table>
<thead>
<tr>
<th>Mixture and temperature</th>
<th>Assignments$^a$</th>
<th>Position [μm (cm$^{-1}$)]</th>
</tr>
</thead>
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<tr>
<td>H$_2$O + CO$_2$</td>
<td></td>
<td></td>
</tr>
<tr>
<td>24:1$^b$ at 15 K</td>
<td>$v_3$</td>
<td>4.270 (2342)</td>
</tr>
<tr>
<td>20:1 at 50 K</td>
<td>$v_1 + 2v_3 + v_3$</td>
<td>2.017 (4958)</td>
</tr>
<tr>
<td></td>
<td>$v_1 + v_3$</td>
<td>2.703 (3709)</td>
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<td></td>
<td>$2v_1 + v_3$</td>
<td>2.784 (3592)</td>
</tr>
<tr>
<td></td>
<td>$3v_3$</td>
<td>4.270 (2342)</td>
</tr>
<tr>
<td></td>
<td>$v_2$</td>
<td>15.27 (655.0)</td>
</tr>
<tr>
<td>CH$_3$OH + CO$_2$</td>
<td></td>
<td></td>
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<td>$v_1 + 2v_1 + v_3$</td>
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<td>$4v_2 + v_3$</td>
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<td></td>
<td>$v_3$</td>
<td>4.275 (2339)</td>
</tr>
<tr>
<td></td>
<td>$v_2$</td>
<td>15.43 (648.0)</td>
</tr>
<tr>
<td>H$_2$O + CH$_3$OH + CO$_2$</td>
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</tr>
<tr>
<td>10:1:1$^b$ at 50 K</td>
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<td></td>
<td>$v_2$</td>
<td>15.43 (648)</td>
</tr>
<tr>
<td>2:2:1:5:1$^b$ at 15 K (UAB)</td>
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<td>4.270 (2342)</td>
</tr>
<tr>
<td></td>
<td>$v_2$</td>
<td>15.29 (654.0)</td>
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<td>$v_1 + 2v_3 + v_3$</td>
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<tr>
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<td></td>
<td>$v_2$</td>
<td>15.38 (650.0)</td>
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<td>5:3:8:8:1$^b$ at 15 K (UAB)</td>
<td>$v_1$</td>
<td>4.272 (2341)</td>
</tr>
<tr>
<td></td>
<td>$v_2$</td>
<td>2.705 (3697)</td>
</tr>
<tr>
<td>1:10:1 at 50 K</td>
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<td>1.974 (5067)</td>
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<td>$v_1 + v_3$</td>
<td>2.706 (3695)</td>
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<tr>
<td></td>
<td>$2v_2 + v_3$</td>
<td>2.792 (3582)</td>
</tr>
<tr>
<td></td>
<td>$v_3$</td>
<td>4.274 (2340)</td>
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<td>$v_2$</td>
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<tr>
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<td>1.969 (5080)</td>
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<td>$v_1 + 2v_3 + v_3$</td>
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<td>$v_1 + v_3$</td>
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</tr>
<tr>
<td></td>
<td>$v_3$</td>
<td>4.274 (2340)</td>
</tr>
<tr>
<td></td>
<td>$v_2$</td>
<td>15.20 (658.0)</td>
</tr>
</tbody>
</table>

$^a$ Mode assignments from Quirico and Schmitt (1997), Sandford and Allamandola (1990a), and Bernstein et al. (2005). Positions of the modes noted are from when the mixture was deposited. In mixtures where H$_2$O:CO$_2$ was 50:1 or more, the $v_2$ CO$_2$ mode was too weak to be seen.

$^b$ Only mid-IR spectra were recorded in this experiment.

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stretch mode. Fig. 3 contains plots of CO2 stretch mode peak position versus temperature for several mixtures. This absorption split into two peaks in some mixtures above 100 K due to increased mobility of the molecules with temperature and population of CO2 molecules within pores of the ice matrix. Only the position of the stronger low-wavenumber/high-wavelength peaks are noted here for clarity. The implications are discussed in the next section.

3.1.1. H2O-rich mixtures

We first discuss the changes in the band in H2O-dominated ices (mixtures with >50% H2O), since absorption profiles were similar in temperature and appearance in H2O-rich mixtures containing H2O, CH3OH, and CO2. Samples were deposited at a low temperature (15 or 50 K) and heated 2 K min\(^{-1}\) before cooling. For these H2O-rich samples, the CO2 stretch mode would start out near 4.274 \(\mu\text{m}\) (2340 cm\(^{-1}\)) and shift by as much as 0.01 \(\mu\text{m}\) (4 cm\(^{-1}\)) to a higher wavelength as the sample was heated to and above 150 K and cooled (e.g., Fig. 3). In most cases, the changes in band positions were irreversible. However, in mixtures of H2O + CO2 where H2O comprised more than 66% of the mixture, the CO2 stretch mode position could shift back to previous positions during a heating and cooling cycle within the range 70–120 K (Fig. 3—only one heating/cooling cycle is shown here).

An example of CO2 stretch mode profiles in a H2O-dominated mixture [H2O + CH3OH + CO2 (10:1:1)] is shown in Fig. 4. A single peak at 4.272 \(\mu\text{m}\) (2341 cm\(^{-1}\)) is seen in the ice mixture after deposition at 50 K. A second peak appears near 4.263 \(\mu\text{m}\) (2346 cm\(^{-1}\)) when heating between 130 and 140 K, a change only seen in mixtures with >66% H2O. The peak near 4.263 \(\mu\text{m}\) (2346 cm\(^{-1}\)) shifts back to 4.274 \(\mu\text{m}\) (2340 cm\(^{-1}\)) above 160 K as the CO2 sublimes. Between 120 and 140 K, the long wavenumber peak would often intensify as the shorter wavenumber peak would diminish. However, it was generally observed that in H2O-dominated mixtures, the greater the proportion of H2O to CO2, the less likely that more than one peak would appear in the CO2 asymmetric stretch mode profiles as the temperature increased. This is probably because at lower concentrations of CO2 molecules in the ice, individual molecules are more fully isolated from one another.

3.1.2. CH3OH-rich mixtures

Here we discuss CH3OH-dominated mixtures (mixtures with >50% CH3OH), all deposited at 15 K and warmed at a rate of 2 K min\(^{-1}\). Absorption profiles of CO2 were similar to each other in binary mixtures of CH3OH and CO2 so a mixture of CH3OH + CO2 (5:1) is discussed as an example. The position of the \(v_3\) asymmetric stretch mode of the CO2 maximum changed when the sample was cooled as well as when it was heated. Above 100 K, a broad shoulder appeared near 4.248 \(\mu\text{m}\) (2354 cm\(^{-1}\)) (Fig. 5). This shoulder remained throughout the cooling process. Re-heating this mixture revealed that the absorbance peak positions retrace some of the positions observed during the first cooling process, although the primary peak near 4.274 \(\mu\text{m}\) (2340 cm\(^{-1}\)) never returns to its original position at deposition or the first heating process. Absorptions on the \(v_3\) CO2 band that appear on the lower wavelength side at higher temperatures (e.g., 130 K) shift towards shorter wavelengths when cooling down as does the peak near 4.275 \(\mu\text{m}\) (2339 cm\(^{-1}\)). Interestingly, the addition of H2O to the mixture makes a difference and the spectral behavior of CH3OH-rich ices of H2O, CH3OH, and CO2 is not similar to that of the CH3OH + CO2 ices. In CH3OH-dominated H2O + CH3OH + CO2 mixtures (1:10:1), the CO2 stretch mode position only changed by 0.001 \(\mu\text{m}\) (1 cm\(^{-1}\)) or less during the heating process, and little or no change was seen during cooling (Fig. 3).

3.1.3. Equal mixtures of H2O and CH3OH

Though mixtures like H2O + CH3OH + CO2 1:1:1 and 10:10:1 are not astrophysically relevant, experiments were carried out in order to cover the phase space and understand the interaction of CO2 with H2O and CH3OH (e.g., Ehrenfreund et al., 1999). The position of the asymmetric \(v_3\) asymmetric stretch mode in mixtures containing equal or nearly equal parts H2O and CH3OH was shifted to ~0.002 \(\mu\text{m}\) (1 cm\(^{-1}\)) longer wavelength than the same feature in H2O-rich mixtures. In one warm-up experiment of H2O + CH3OH + CO2 (10:10:1), a small peak at 4.30 \(\mu\text{m}\) (2323 cm\(^{-1}\)) appeared and remained throughout the experiment, as did a similar feature with a H2O + CH3OH + CO2 (1:1:1) mixture at 120 K and remained as the sample was cooled (Fig. 6). It should be noted that the CO2 profiles in H2O + CH3OH + CO2 (1:1:1) mixtures at the deposition temperature were at the same position as those of the H2O-rich mixtures. The presence of CH3OH affects the sample such that this particular vibrational mode will not “backtrack” to a previous position from a warming/cooling cycle as it did in the H2O + CO2 (24:1) experiment.

3.1.4. CH3OH in H2O + CH3OH + CO2 mixtures

The asymmetric C–O stretch mode of CH3OH appears as a single absorption at deposition at 50 K in these H2O + CH3OH + CO2...
mixtures, though a shoulder on the short wavelength side appears as the sample is heated (Fig. 7). This shoulder became stronger than the longer wavelength peak above 110 K. The C–O stretch mode profile of CH₃OH changes with temperature, indicative of the tetrahydrofuran (THF) type II clathrate structure forming from H₂O – CH₃OH complexes (Fleyfel and Devlin, 1991 – see also Section 1).

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**Fig. 4.** Selected CO₂ asymmetric stretch mode profiles in a mixture of H₂O + CH₃OH + CO₂ (10:1:1) deposited at 50 K and heated at a rate of 2 K min⁻¹ before cooling down. Spectra were taken while the sample was held at the indicated temperature.

**Fig. 5.** Selected CO₂ asymmetric stretch mode profiles in a mixture of CH₃OH + CO₂ (5:1) deposited at 15 K, warmed at a rate of 2 K min⁻¹, and cooled down. Spectra were taken as the sample was held at the indicated temperature.

**Fig. 6.** CO₂ stretch mode profiles in a mixture of H₂O + CH₃OH + CO₂ (1:1:1) deposited at 15 K and heated at a rate of 2 K min⁻¹ to 120 K before cooling down. Spectra were taken as the sample was held at the indicated temperature. The small feature at 2323 cm⁻¹ at 120 K is shown here.

**Fig. 7.** The asymmetric C–O stretch mode profile of CH₃OH in a mixture of H₂O + CH₃OH + CO₂ (10:1:1) deposited at 50 K and heated at a rate of 2 K min⁻¹. Spectra were taken as the sample was held at the indicated temperature. Spectra became more noisy as the sample was heated. (10:1:1) mixtures, though a shoulder on the short wavelength side appears as the sample is heated (Fig. 7). This shoulder became stronger than the longer wavelength peak above 110 K. The C–O stretch mode profile of CH₃OH changes with temperature, indicative of the tetrahydrofuran (THF) type II clathrate structure forming from H₂O – CH₃OH complexes (Fleyfel and Devlin, 1991 – see also Section 1).
3.2. Near-IR CO₂ modes

Absorption profiles of mixtures of H₂O, CH₃OH, and CO₂ were also examined in the near-IR. Table 1 contains the combination/overton mode that were detected in each mixture. Many of the CO₂ overtone and combination modes appeared as broad, weak absorptions superimposed on the side or on top of H₂O and/or CH₃OH absorptions. As a result, many of these features were difficult to analyze due to the relatively strong features of H₂O near 2 μm and 2.95 μm, as well as CH₃OH features near 2.95 μm (in ternary mixtures – see also Fig. 2).

3.2.1. H₂O-rich mixtures

The CO₂ bands produced by H₂O-rich mixtures (mixtures with >50% H₂O) had similar profiles when deposited at low temperatures (15 or 50 K), so a mixture of H₂O + CH₃OH + CO₂ (10:1:1) is discussed here as an example. Fig. 8a-d contains selected spectra of this mixture deposited at 50 K and heated at a rate of 2 K min⁻¹ to 150 K and cooled back down.

The 2v₁ + v₃ feature near 1.97 μm (5080 cm⁻¹) appeared as a very weak, broad absorption (less than 5% of the height of the strongest absorption near 2.70 μm) at the deposition temperature. When heating up to 125 K, the peak began to grow sharper (Fig. 8a). This profile continued to become stronger by a factor of 2 up to 150 K and remained throughout the cooling process. While cooling, a broad absorption near 1.972 μm (5070 cm⁻¹) also appeared and became more distinct as the sample was cooled. This absorption is most likely due to the interaction of CO₂ and H₂O (Bernstein et al., 2005). This particular mode was very difficult to isolate because it appears directly on top of a large absorption feature due to H₂O.

The 2v₁ + v₃ + v₂ + v₄ mode appeared near 2.017 μm (4958 cm⁻¹) and at least twice the height of the 2v₁ + v₃ absorption at the deposition temperature (50 K). This mode was difficult to detect as it is on the large absorption feature due to H₂O. At 120 K, this absorption changed into a broader peak near 2.018 μm (4956 cm⁻¹) and a smaller “bump” near 2.012 μm⁻¹ (4971 cm⁻¹) (Fig. 8b). The 2.012 μm⁻¹ absorption became sharper (about twice the relative intensity of the 2.018 μm⁻¹ absorption) as the sample was heated to 150 K and remained so even after cooling down to 50 K. The 2.018 μm⁻¹ (4956 cm⁻¹) absorption, however, remained broader (≤ half the intensity of the 2.012 μm⁻¹ feature) when heating to 150 K and shifted to 2.019 μm (4953 cm⁻¹). Interestingly, as the sample was cooled this absorption continued to move to about 2.020 μm (4951 cm⁻¹) and became a sharper peak with about the same height as the 2.012 μm⁻¹ feature and a wing at 2.016 μm (4960 cm⁻¹).

The v₁ + v₃ mode near 2.703 μm (3700 cm⁻¹) remained the strongest CO₂ absorption profile in the spectral range 0.9–3.5 μm (11,000–2850 cm⁻¹) throughout the experiment (Fig. 8c). This absorption shifted to about 2.704 μm (3698 cm⁻¹) as the sample was heated, and at 125 K a second peak appeared at 2.697 μm (3708 cm⁻¹). This new peak became the dominating absorption profile at 140 K. At 150 K, a third peak appeared in between these two at 2.700 μm (3704 cm⁻¹) as the 2.704 μm (3698 cm⁻¹) peak shifted to 2.706 μm (3695 cm⁻¹) and the peak at 2.697 μm (3708 cm⁻¹) once again became the dominating profile. It should be noted that the middle peak that appeared in Fig. 8c after warming to 150 K and cooling back down is most likely due to an interaction of CO₂ and H₂O and may not be solely attributed to CO₂ (Bernstein et al., 2005). The 2.697 μm (3708 cm⁻¹) peak remained the dominating absorption profile while cooling until at 85 K the peak at 2.707 μm (3694 cm⁻¹) became about 10% stronger, though about twice as broad as the 2.697 μm (3708 cm⁻¹) absorption profile.

The 2v₂ + v₁ mode appeared at the deposition temperature around 2.786 μm (3590 cm⁻¹) as a small, broad peak (≤10% the intensity of the v₁ + v₃ feature) on the side of the O–H stretching mode of H₂O and CH₃OH near 2.954 μm (3385 cm⁻¹) making this absorption profile difficult to analyze (Fig. 8d). A second peak appeared around 2.778 μm (3600 cm⁻¹) at 120 K with about the same height as the 2.786 μm (3590 cm⁻¹) feature. This peak became sharper as the sample was heated to 150 K. The two peaks became more distinct after cooling back down to 50 K.

3.2.2. CH₃OH-rich mixtures

Experiments with more than 50% CH₃OH (both binary and ternary mixtures) yielded absorption profiles similar to those of H₂O-dominated mixtures, due to some modes being on the large O–H stretch mode of CH₃OH and/or H₂O and the absorption curve between 1.67 and 2.5 μm (e.g., Fig. 2). As a representation of CH₃OH-dominated ice mixtures, Fig. 9 contains selected spectra of a H₂O + CH₃OH + CO₂ (1:1:1) ice deposited at 15 K, heated at a rate of 2 K min⁻¹ to 120 K and cooled down.

The barely visible 2v₁ + v₃ mode near 1.974 μm (5066 cm⁻¹) became broader and weaker (height about 5% or less than that of
the stronger $v_1 + v_3$ feature), shifting to a higher wavelength when warming up from 15 K (Fig. 9a). This feature continued to weaken and shift by as much as 0.01 μm (8 cm⁻¹), even when cooling down and disappeared completely when the sample cooled to 75 K.

The $v_1 + v_3$ mode appeared near 2.712 μm (3695 cm⁻¹) and remained the strongest CO2 feature in this spectral range throughout the heating/cooling process, as it did in the H2O-dominated mixtures. The position did not change when heating to 120 K. Upon cooling, however, the peak shifted to 2.711 μm (3690 cm⁻¹) at 110 K and gradually to 2.710 μm (3689 cm⁻¹) when cooling to 15 K.

The $2v_2 + v_3$ mode near 2.792 μm (3582 cm⁻¹) appeared as a stronger feature (about half the height of the $v_1 + v_3$ mode) in CH3OH-dominated mixtures compared to the same feature in H2O-dominated mixtures, though it was still difficult to analyze due to its proximity to the large O–H stretch mode. This absorption profile began to split into three peaks as the sample was cooled from 120 K to 15 K (Fig. 9b) with the dominate feature near 2.795 μm (3578 cm⁻¹) and the other two peaks at 2.800 μm (3572 cm⁻¹) and 2.804 μm (3566 cm⁻¹) (about half and one-fourth the height of the dominate peak, respectively). This splitting and increase of intensity of the 2.795 μm feature is likely due to the interaction of H2O–CH3OH complexes trapping CO2 molecules in separate pockets of ice (Section 1).

3.3. High-temperature deposits

Observations of the IR absorption features of H2O in the outer Solar System revealed that most ices exist in a crystalline phase instead of an amorphous structure. Specifically, crystalline phases of H2O–ice such as $l_c$ (cubic phase) and $l_h$ (hexagonal phase) are the dominate forms of H2O in the Solar System (e.g., Cruikshank et al., 1998; Schmitt et al., 1998). However, this does not necessarily mean that the absorption profiles of the ice indicate that the temperatures of the ices on planetary surfaces are currently at or above the phase transition temperature of the H2O–ice component. Rather, it is possible that the ice had been heated to those temperatures at some point in the past and cooled back down (e.g., Ehrenfreund et al., 1999; White et al., 2009). This raises the question as to whether CO2 seen in these ices is a trapped remnant of an earlier, lower temperature interval or whether it can be directly deposited at or above the crystalline phase transition temperatures of solid H2O. Previous experiments have demonstrated that the trapping efficiency and residence time of CO2 in H2O drops rapidly at temperatures above 90 K and little or no CO2 is trapped above 100 K at deposition rates of about 0.09 μm min⁻¹ (Sandford and Allamandola, 1990b; Gálvez et al., 2008). It is possible that CO2 could be trapped in H2O-rich ices at higher temperatures if deposition rates are higher. At higher temperatures, the residence time of a CO2 molecule on an H2O–ice surface are decreased, but the CO2 could still be trapped if the next overlying layer of H2O is deposited quickly enough (Sandford and Allamandola, 1990b). We varied the deposition rates of some H2O, CH3OH, and CO2 mixtures in this paper to explore the extent to which CO2 can be deposited in H2O–ices above 100 K.

Experiments were performed with mixtures of H2O + CO2 (20:1), H2O + CH3OH + CO2 (10:1:1), H2O + CH3OH + CO2 (10:10:1), and H2O + CH3OH + CO2 (1:10:1) with three different deposition rates; a “fast” rate of ~0.5 μm min⁻¹, a “slow” rate of <0.03 μm min⁻¹, and an “intermediate” rate that corresponds to our normal rate of ~0.1 μm min⁻¹. These studies were performed to verify the earlier observations of Sandford and Allamandola (1990a) and to see whether or not CO2 could be trapped in H2O ices at higher temperatures with higher deposition rates. Thicknesses were similar to previous high-temperature deposit experiments (~0.3–1.0 μm). Table 2 contains the highest deposition temperatures at which CO2 was detected in the resulting ice mixtures at the three different deposition rates.

The CO2 asymmetric stretch mode was just visible as a small peak when a sample of H2O + CO2 (20:1) was deposited at 110 K at a rate of about 0.1 μm min⁻¹, indicating a trace amount of residual CO2 (<1% relative to H2O, Fig. 10). All of the CO2 asymmetric stretch mode profiles in successful deposits of CO2 above 100 K resembled this experiment. Note the deformation of the O–H stretch mode in Fig. 10 that suggests a crystalline phase of H2O. CO2 could be trapped at higher temperatures if higher deposition rates are used, although the differences are not dramatic. At the fastest deposition rate of 0.5 μm min⁻¹ CO2 could be partially trapped at temperatures as high as 120 K for H2O:CO2 ratios of 10:1 or 20:1. However, faster deposits of H2O + CH3OH + CO2

![Fig. 9. Spectra from a mixture of H2O + CH3OH + CO2 (1:10:1) deposited at 50 K and heated at a rate of 2 K min⁻¹ to 120 K and cooled down. Spectra were taken as the sample was held at the indicated temperature. The CO2 mode profiles of (a) 2$v_1 + v_3$ and (b) 2$\nu_2 + \nu_3$ are shown. The noisiness in (a) is due to the weakening of the feature.](image-url)

Table 2

Highest temperatures at which CO2 was present when H2O, CH3OH, and CO2 mixtures were deposited at temperatures above 50 K. CO2 was determined to have deposited out if the $\nu_1$ asymmetric stretch mode was visible in the mid-IR spectrum.

<table>
<thead>
<tr>
<th>Mixture</th>
<th>Temperature (K)</th>
<th>Deposition rate (μm min⁻¹)</th>
</tr>
</thead>
<tbody>
<tr>
<td>$H_2O + CO_2$ (20:1)</td>
<td>110</td>
<td>&lt;0.03 and 0.1</td>
</tr>
<tr>
<td></td>
<td>120</td>
<td>~0.5</td>
</tr>
<tr>
<td>$CH_3OH + CO_2$ (5:1)</td>
<td>100</td>
<td>&lt;0.03 and 0.1</td>
</tr>
<tr>
<td></td>
<td>115</td>
<td>~0.5</td>
</tr>
<tr>
<td>$H_2O + CH_3OH + CO_2$ (10:1:1)</td>
<td>115</td>
<td>&lt;0.03 and 0.1</td>
</tr>
<tr>
<td></td>
<td>120</td>
<td>~0.5</td>
</tr>
<tr>
<td></td>
<td>(100:10:1)</td>
<td>~0.5</td>
</tr>
<tr>
<td></td>
<td>(1:10:1)</td>
<td>&lt;0.03 and 0.1</td>
</tr>
<tr>
<td></td>
<td>(10:10:1)</td>
<td>~0.5</td>
</tr>
</tbody>
</table>
at 65 K and above. Evidently, the more CH$_3$OH present, the less likely CO$_2$ will reside in an ice deposited above 100 K. However, the amount of resident CO$_2$ decreases exponentially with temperature (note the differences in the spectral stretch mode in Fig. 10).

The position of the vibrational mode of CO$_2$ in astronomical observations is consistent with laboratory spectra of a H$_2$O-dominated ice sample containing CO$_2$ as seen in this study, and, for example, experiments by Sandford and Allamandola (1990a), Cruikshank et al. (2010), and in theoretical models from Chaban et al. (2007). The addition of CH$_3$OH also induces a slight shift (≤0.01 μm or 1 cm$^{-1}$) but mainly changes the appearance of the absorption profiles (e.g., Ehrenfreund et al., 1998; White et al., 2009; Cruikshank et al., 2010).

Some examples of observed temperatures and spectral positions of the CO$_2$ stretch mode from some jovian and saturnian satellites along with the ranges of temperatures and positions of this feature in laboratory data from this study are shown in Fig. 11. Observations were taken from published data from ground-based telescopes, NIMS aboard the Galileo spacecraft, and VIMS aboard the Cassini probe (Grundy et al., 1999; Hibbitts et al., 2000, 2003; Ostro et al., 2006; Cruikshank et al., 2007, 2010). Temperatures for the surfaces of the jovian satellites, Dione, and Iapetus were derived from an average disk temperature and the IR 1039

![Fig. 10. Spectrum from a mixture of (a) H$_2$O + CO$_2$ (20:1) deposited at 110 K at a rate of −0.1 μm min$^{-1}$ to a thickness of about 0.2 μm compared with (b) a spectrum from H$_2$O + CO$_2$ (24:1) deposited at 15 K at a rate of −0.1 μm min$^{-1}$ to a thickness of about 0.3 μm.](image-url)

(100:10:1 and 100:50:1) yielded CO$_2$ at temperatures no greater than 115 K (Table 2).

Some CO$_2$ was visible in the CH$_3$OH-rich spectra at the same temperatures as observed in the H$_2$O-dominated mixtures (e.g., 100 K, 110 K). Faster deposits (0.5 μm min$^{-1}$) of H$_2$O + CH$_3$OH + CO$_2$ (1:10:1) yielded CO$_2$ at temperatures no greater than 115 K, similar to the H$_2$O-rich experiments above. In H$_2$O + CH$_3$OH + CO$_2$ (10:10:1) experiments, CO$_2$ was not visible when depositing above 100 K. No double peaks on the CO$_2$ stretch mode were visible when depositing at 65 K and above. Evidently, the more CH$_3$OH present, the less likely CO$_2$ will reside in an ice deposited above 100 K.

The ν$_3$ CO$_2$ asymmetric stretch mode was clearly visible at deposition temperatures up to 100 K when samples were deposited at a rate of ≤0.1 μm min$^{-1}$. The position of the CO$_2$ asymmetric stretch mode at mid-IR experiments and the concentration of H$_2$O:CO$_2$ was 2:1 or so (Section 2.3).

Comparison to astronomical data

The position of the vibrational mode of CO$_2$ in astronomical observations is consistent with laboratory spectra of a H$_2$O-dominated ice sample containing CO$_2$ as seen in this study, and, for example, experiments by Sandford and Allamandola (1990a), Cruikshank et al. (2010), and in theoretical models from Chaban et al. (2007). The addition of CH$_3$OH also induces a slight shift (≤0.01 μm or 1 cm$^{-1}$) but mainly changes the appearance of the absorption profiles (e.g., Ehrenfreund et al., 1998; White et al., 2009; Cruikshank et al., 2010).

Some examples of observed temperatures and spectral positions of the CO$_2$ stretch mode from some jovian and saturnian satellites along with the ranges of temperatures and positions of this feature in laboratory data from this study are shown in Fig. 11. Observations were taken from published data from ground-based telescopes, NIMS aboard the Galileo spacecraft, and VIMS aboard the Cassini probe (Grundy et al., 1999; Hibbitts et al., 2000, 2003; Ostro et al., 2006; Cruikshank et al., 2007, 2010). Temperatures for the surfaces of the jovian satellites, Dione, and Iapetus were derived from an average disk temperature and the IR 1039

4. Discussion

The position(s) of the ν$_3$ CO$_2$ asymmetric stretch mode feature shifted to a longer wavelength (smaller wavenumber) as the sample was heated in both H$_2$O- and CH$_3$OH-rich ices (Fig. 3) in response to molecular migration (Section 1). Shifting of the absorption bands observed in these experiments with all ice mixtures, both H$_2$O- and CH$_3$OH-rich, indicate a restructuring of the ice matrix as the samples are heated. Unlike most mixtures, the position of the CO$_2$ ν$_3$ asymmetric stretch mode seems to migrate back and forth when heating and cooling between 70 and 120 K in the H$_2$O + CO$_2$ (24:1) mixture. This is perhaps due to the low to high density amorphous phase changes happening in this range (Blake et al., 1991). Though the mixture H$_2$O + CH$_3$OH + CO$_2$ (10:1:1) was often used as a typical example, mixtures with H$_2$O:CO$_2$ ratios of 100:1 yielded spectra with fewer visible combination and overtone modes of CO$_2$ (1.25–2.04 μm or 8000–4900 cm$^{-1}$), – see also Table 1. It should be noted that saturation of the bands makes it difficult to determine position and other properties such as column density of the species (the asymmetric CO$_2$ stretch mode, for example). This was only an issue if ice thicknesses exceeded 2 μm for the mid-IR experiments and the concentration of H$_2$O:CO$_2$ was 2:1 or so (Section 2.3).

Sometimes this feature split into multiple peaks, particularly in binary mixtures of H$_2$O + CO$_2$ and CH$_3$OH + CO$_2$. This is most likely due to molecular rearrangement, segregation, and phase changes in the ice (Sandford and Allamandola, 1990a; Öberg et al., 2009). Peak splitting is also affected by how much H$_2$O is present and the thickness of the sample. Also, some of the double-peak absorption profiles of the ν$_3$ CO$_2$ asymmetric stretch mode may suggest the formation of clathrate hydrates (type I) (Blake et al., 1991; Dartois and Schmitt, 2009). The splitting and increase of intensity of the 2ν$_2$ + ν$_3$ CO$_2$ mode is likely due to the interaction of H$_2$O – CH$_3$OH complexes trapping CO$_2$ molecules in separate pockets of ice (Section 3.2.2, Blake et al. (1991); Tielens (2005)).

The presence of CH$_3$OH appears to strongly influence many absorption features of CO$_2$, as has been demonstrated in previous studies (e.g., Sandford and Allamandola, 1993; Marchand et al., 2006; Maté et al., 2009; White et al., 2009; White, 2010). The influence of CH$_3$OH on the absorption profiles is strong due to the possibility of stronger bonds between CH$_3$OH and H$_2$O, thus forming more complex structures trapping CO$_2$ molecules (Collings et al., 2004). Also, the carbon atom of CO$_2$ can act as a Lewis acid because part of the electron density is removed due to the existence of two electronegative O atoms (Kazarian et al., 1996; Klotz et al., 2004). The presence of CH$_3$OH may also affect the sticking efficiency of CO$_2$ in ice mixtures deposited at higher temperatures (>50 K) as suggested in Section 5.

5. Comparison to astronomical data

The position of the vibrational mode of CO$_2$ in astronomical observations is consistent with laboratory spectra of a H$_2$O-dominated ice sample containing CO$_2$ as seen in this study, and, for example, experiments by Sandford and Allamandola (1990a), Cruikshank et al. (2010), and in theoretical models from Chaban et al. (2007). The addition of CH$_3$OH also induces a slight shift (≤0.01 μm or 1 cm$^{-1}$) but mainly changes the appearance of the absorption profiles (e.g., Ehrenfreund et al., 1998; White et al., 2009; Cruikshank et al., 2010).
absorption features of H$_2$O-ice from the surfaces (Grundy et al., 1999). Temperatures for Phoebe and Hyperion were derived from observed albedo and average disk temperature (Ostro et al., 2006; Cruikshank et al., 2007). “2nd Peak” in the figure refers to the second feature that appears at temperatures greater than 110 K in some of the H$_2$O- and CH$_3$OH-rich mixtures in addition to the primary feature, or “1st Peak”.

In H$_2$O- and CH$_3$OH-rich ices, a second peak would often appear in the $\nu_3$ asymmetric stretch mode of CO$_2$ at a shorter wavelength (larger wavenumber) and is labeled as e.g., “H$_2$O-ice 2nd Peak” in addition to the primary peak or “1st Peak” (e.g., Section 3.1.1). From Fig. 11, it is obvious that pure CO$_2$ is not likely a component on the surfaces of these satellites. It is also unlikely that CH$_3$OH-rich ices are components and none of the data match up to “Peak 2” of the CH$_3$OH-rich laboratory data. However, the CO$_2$ asymmetric stretch mode appears near ~4.255 μm ($\sim 2350$ cm$^{-1}$) for the Ganymede and Callisto surfaces getting close to the Peak 2 of H$_2$O-rich laboratory ices. It is evident that the surfaces of Phoebe, Dione, and Hyperion likely contain ices dominated by H$_2$O and perhaps CH$_3$OH. It should also be noted, however, that though many of the observational data do not confirm the presence of CH$_3$OH, the laboratory data in this study offer a means of exploring that possibility. Experiments in this study have demonstrated that in CO$_2$ ices with more than H$_2$O present (CH$_3$OH, in this case) the position of the CO$_2$ $\nu_3$ asymmetric stretch mode changes very little when the sample is cooled from 140 K after heating. This may suggest that the ice on many of the outer planetary surfaces are at (or once were at) temperatures near or above the glass transition temperature of H$_2$O (130–140 K).

These data and the data from this paper demonstrate that the spectral characteristics of an ice are dominated by the highest temperature to which the ice has been. However, CO$_2$ $\nu_3$ asymmetric stretch band positions do not match laboratory band positions for Dione, Iapetus, or Ganymede at all. This means there are more species present in the ice affecting the CO$_2$ band positions, there is or have been UV processing of the icy surfaces of these satellites, or a combination of both. Though the laboratory data in this study begin to explain the chemistry of these satellites’ surfaces, many more parameters need to be explored in the laboratory to get a more complete picture.

6. Conclusions

These data explore parameters and permutations of astrophysically relevant ice mixtures in the laboratory. Absorption spectra from solid CO$_2$ are sensitive indicators of temperature and composition. Position of the absorption features of CO$_2$ were different at deposition depending on the ratios of H$_2$O and CH$_3$OH. The $\nu_3$ asymmetric stretch mode of CO$_2$ yielded different results, depending on the ratio of H$_2$O and CH$_3$OH. Mixtures where H$_2$O was the dominate species yielded similar results, as did CH$_3$OH-dominated mixtures. Interestingly, the H$_2$O + CO$_2$ (24:1) mixture seemed to allow the $\nu_3$ asymmetric stretch mode to move back and forth when heating and cooling between ~70 and 120 K. CH$_3$OH-rich mixtures thermally processed yielded similar results in the amount the CO$_2$ asymmetric stretch mode shifted, though the positions were different. When combined with H$_2$O, however, CH$_3$OH seems to have a strong influence on position of the $\nu_3$ asymmetric stretch mode as it seems result in shifting during the heating and cooling process, likely due the strong contribution of CH$_3$OH in the clathrate forming process (e.g., Fig. 3 and Section 1).

The near-IR vibrational modes of CO$_2$ were more difficult to analyze, as expected, due to the much lower relative intensity and widths of the absorption peaks compared to the strong absorption features of H$_2$O and CH$_3$OH. Some of the extra peaks that appear such as the “broad feature” in Fig. 8a, the middle peak in Fig. 8c, and the multi-peak features in Fig. 9b are most likely due to an interaction of H$_2$O and CO$_2$, as discussed earlier (Bernstein et al., 2005). Still, a distinct correlation between temperature and mixture ratios were noted in all binary and ternary H$_2$O, CH$_3$OH, and CO$_2$ mixtures.

In the high-temperature deposition experiments, we concluded that it is possible to deposit CO$_2$ in H$_2$O-dominated ice mixtures at temperatures up to 120 K. This is only possible, however, if the deposition rate is 0.5 μm min$^{-1}$ or at least greater than 0.1 μm min$^{-1}$ (Table 2). This pushes the previously calculated
highest temperature that CO₂ will deposit out in an ice (Sandford and Allamandola, 1990b) by 10 K higher to 120 K. Though only appearing in trace amounts (<1%), the ν₃ asymmetric stretch mode of CO₂ appears at the same position as it would have the mixture been deposited at a low temperature (<50 K) and warmed up at a rate of 2 K min⁻¹ to the deposition temperature. Clearly the IR absorption spectrum reflects the highest temperature to which the ice has been in the past. However, the ice may have deposited at a higher temperature in the first place. These results may provide another variable and perspective on ice formation, particularly in the Solar System. This may help to explain some of the ν₃ asymmetric stretch mode positions from the surfaces of the jovian and saturnian bodies: Theoretical and experimental studies of molecular complexes. Icarus 227, p. 579.


References


2 e.g., http://www.astrochemistry.org/db.php.