THE OXYGEN ISOTOPE COMPOSITION OF DARK INCLUSIONS IN HEDs, ORDINARY AND CARBONACEOUS CHONDRITES. R. C. Greenwood1, M. E. Zolensky2, P.C. Buchanan3, I.A. Franchi1.

1Planetary and Space Sciences, The Open University, Milton Keynes MK7 6AA, UK (r.c.greenwood@open.ac.uk).
2ARES, Johnson Space Center, Houston TX, USA.3Kilgore College, Kilgore, Texas 75662 USA.

Introduction: Dark inclusions (DIs) are lithic fragments that form a volumetrically small, but important, component in carbonaceous chondrites [1,2]. Carbonaceous clasts similar to DIs are also found in some ordinary chondrites and HEDs [3,4]. DIs are of particular interest because they provide a record of nebular and planetary processes distinct from that of their host meteorite [1,5]. DIs may be representative of the material that delivered water and other volatiles to early Earth as a late veneer [6]. Here we focus on the oxygen isotopic composition of DIs in a variety of settings with the aim of understanding their formational history and relationship to the enclosing host meteorite.

Materials and methods: DIs and related materials were obtained from the following meteorites: CV3s (Allende, NWA 2140, NWA 2364): Samples from Allende cover all categories of the four-fold classification scheme [2] (inclusion numbers analyzed in brackets). Type A clasts (1a1, 4b1, 25s1-TW1) contain chondrules, inclusions and matrix, but are somewhat finer grained than normal Allende material. Type A/B clasts (MZB) are transitional between Types A and B (Fig. 1). Type B clasts (12b1) contain opaque matrix and olivine-rich aggregates and may have experienced a hydration-dehydration cycle [2]. Type C clasts (5a1, ekph4b1, MZ15, USNM 3876) consist of fine-grained, opaque material similar to Allende matrix. Full descriptions for of the Allende DI samples are given in [2]. DI material from NWA 2140 analyzed for this study is Type A/B and Type A for NWA 2364. 

HEDs: (Bholghati, PRA 04401, SCO 06040). DIs from the howardite Bholghati have not been analyzed by us for oxygen isotopes. Instead a sequence of 5-10 mg representative whole rock samples have been run to assess its carbonaceous chondrite content. PRA 04401 is an extremely coarse-grained howardite containing a high concentration (~40-50% in places) of nearly cm-sized angular carbonaceous clasts that texturally resemble CM2 material [7]. SCO 06040 is also a coarse-grained howardite breccia, contains a lower amount (~10%) of rounded, up to 2mm diameter, CM2-like clasts [7]. 

OCs: Sharps (H3.4) is a fragmental breccia containing accessory carbonaceous clasts up to 1 cm in diameter [8].

Oxygen isotope analysis was performed by infrared laser-assisted fluorination [9]. All analyses were obtained on untreated whole rock samples (0.5-2 mg). System precision, as determined on an internal obsidian standard is: ±0.05‰ for δ17O; ±0.09‰ for δ18O; ±0.02‰ for Δ17O (2σ).

Results: All of the Allende DIs are displaced slightly to the right of the CCAM line in Fig.2 and define a linear trend with a slope of y = -4.30 + 0.89x R^2 = 0.99. The less altered chondrule-bearing clasts (A and A/B) plot closest to bulk Allende analysis in Fig. 2, with the most altered clast (Type B) (12b1) being approximately 5‰ heavier with respect to δ18O. Matrix-rich Type C clasts show a narrow range of oxygen isotope variation, plotting roughly halfway between the most and least altered DIs in Fig. 2. These results are in agreement with previous studies of Allende DIs [1].

DIs in CV3s NWA 2140 and NWA 2364 (Type A/B and A respectively) plot at the higher δ18O values than their equivalents in Allende (Fig. 2). This is consistent with the results of previous studies [1] which found that Allende DIs are relatively 18O-enriched compared to inclusions in other CV3s (Fig.3).

Carbonaceous chondrite clast material from the howardite PRA 04401 plots at the extension of the CM2 field in Fig. 2 consistent with results of textural and mineralogical studies [7]. The other howardite samples analyzed in this study (SCO 06040 and Bhol-
DIs analyzed in Sharps are clearly distinct from those in either the CV3s or howardites and plot on the CR Mixing Line [11]. The composition of these DIs is not well understood, although they do not appear to be mineralogically related to CR chondrites [8].

**Discussion:** Mineralogical studies of DIs have played a crucial role in highlighting the importance of parent body processes in modifying the composition of CV chondrites [2, 13]. In particular, DI studies have shown that CV3s, previously regarded as pristine nebular condensates, underwent extensive aqueous alteration [2, 13]. Furthermore, textural evidence indicates that, following aqueous alteration, some DIs experienced a phase of thermal metamorphism, resulting in phyllosilicate dehydration and the formation of secondary Fe-rich olivine (Type B inclusions) [13]. This model was later extended to explain the origin of CV3 Fe-rich matrix olivines in general [14].

The results of previous oxygen isotope studies of DIs have also pointed to the role of aqueous alteration in modifying their primary compositions [1]. As shown in this study, DIs define linear arrays with a shallower slope than the CCAM and with less altered material at the $^{18}$O-rich end (Type As) and more altered material (Type Bs) at the $^{16}$O-poor end [1]. However, with respect to phyllosilicate dehydration and secondary Fe-rich olivine formation, oxygen isotope evidence appears less clear-cut [15]. Experimental evidence indicates that dehydration will result in heavy-isotope enrichments, which for Allende matrix olivines are not seen, if the CCAM line is used as a reference [15].