THE GALEX TIME DOMAIN SURVEY. I. SELECTION AND CLASSIFICATION OF OVER A THOUSAND ULTRAVIOLET VARIABLE SOURCES


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ABSTRACT

We present the selection and classification of over a thousand ultraviolet (UV) variable sources discovered in ~40 deg2 of GALEX Time Domain Survey (TDS) NUV images observed with a cadence of 2 days and a baseline of observations of ~3 years. The GALEX TDS fields were designed to be in spatial and temporal coordination with the Pan-STARRS1 Medium Deep Survey, which provides deep optical imaging and simultaneous optical transient detections via image differencing. We characterize the GALEX photometric errors empirically as a function of mean magnitude, and select sources that vary at the 5σ level in at least one epoch. We measure the statistical properties of the UV variability, including the structure function on timescales of days and years. We report classifications for the GALEX TDS sample using a combination of optical host colors and morphology, UV light curve characteristics, and matches to archival X-ray, and spectroscopy catalogs. We classify 62% of the sources as active galaxies (358 quasars and 305 active galactic nuclei), and 10% as variable stars (including 37 RR Lyrae, 53 M dwarf flare stars, and 2 cataclysmic variables). We detect a large-amplitude tail in the UV variability distribution for M-dwarf flare stars and RR Lyrae, reaching up to Δm = 4.6 mag and 2.9 mag, respectively. The mean amplitude of the structure function for quasars on year timescales is five times larger than observed at optical wavelengths. The remaining unclassified sources include UV-bright extragalactic transients, two of which have been spectroscopically confirmed to be a young core-collapse supernova and a flare from the tidal disruption of a star by dormant supermassive black hole. We calculate a surface density for variable sources in the UV with NUV < 23 mag and |Δm| > 0.2 mag of ~8.0, 7.7, and 1.8 deg−2 for quasars, active galactic nuclei, and RR Lyrae stars, respectively. We also calculate a surface density rate in the UV for transient sources, using the effective survey time at the cadence appropriate to each class, of ~15 and 52 deg−2 yr−1 for M dwarfs and extragalactic transients, respectively.

Key words: surveys – ultraviolet: general

Online-only material: color figures, machine-readable table

1. INTRODUCTION

Unlike the optical, X-ray, and γ-ray sky, which have been systematically studied in the time domain in the search for supernovae (SNe) and gamma-ray bursts (GRBs), the wide-field ultraviolet (UV) time domain is a relatively unexplored parameter space. The launch of the GALEX satellite with its 25 degree field of view, and limiting sensitivity per 1.5 ks visit of ~23 mag in the FUV (λ_eff = 1539 Å) and NUV (λ_eff = 2316 Å; Martin et al. 2005; Morrissey et al. 2007), enabled the discovery of UV variable sources in repeated observations over hundreds of square degrees for the first time.

The UV waveband is particularly sensitive to hot (~10^4 K) thermal emission from such transient and variable phenomena as young core-collapse SNe, the inner regions of the accretion flow around accreting supermassive black holes (SMBHs), and the flaring states of variable stars. The characteristic timescales of variables and transients in the UV range from minutes to years. M dwarf flare stars have strong magnetic activity that manifests itself in flares of thermal UV emission on the timescale of minutes (Kowalski et al. 2009). RR Lyrae stars have periodic pulsations which drive temperature fluctuations from ~6000 to 8000 K that cause periodic variability in the UV on a timescale of ~0.5 d (Wheatley et al. 2012). Core-collapse SNe remain bright in the UV for hours up to several days following shock breakout, depending on the radius of the progenitor star (Nakar & Sari 2010; Rabinak & Waxman 2011), and the presence of a dense wind (Ofek et al. 2010; Chevalier & Irwin 2011; Svirski et al. 2012). Active galactic nuclei (AGNs) and quasars demonstrate stronger variability with decreasing wavelength and longer timescales of years (Vanden Berk et al. 2004).

UV variability studies of GALEX data observed as part of the All-Sky, Medium, and Deep Imaging baseline mission surveys (AIS, MIS, DIS) from 2003 to 2007, yielded the detection of M-dwarf flare stars (Welsh et al. 2007), RR Lyrae stars, AGN, and quasars (Welsh et al. 2005, 2011; Wheatley et al. 2008), and flares from the tidal disruption of stars around dormant SMBHs (Gezari et al. 2006, 2008a, 2009). Serendipitous overlap of four GALEX DIS fields with the optical Canada–France–Hawaii Telescope (CFHT) Supernova Legacy Survey enabled the extraction of simultaneous optical light curves from image differencing for two of the tidal disruption
event (TDE) candidates (Gezari et al. 2008a), and enabled the association of transient UV emission with two Type IIP SNe within hours of shock breakout (Schawinski et al. 2008; Gezari et al. 2008b). Chance overlap of GALEX observations with the survey area of the optical Palomar Transient Factory (PTF) detected a Type IIn SN whose rising UV emission over a few days was interpreted as a delayed shock breakout through a dense circumstellar medium (Ofek et al. 2010).

Motivated by the promising results from the analysis of random repeated GALEX observations, and the demonstrated value of overlap with optical time domain surveys, we initiated a dedicated GALEX Time Domain Survey (TDS) to systematically study UV variability on timescales of days to years with multiple epochs of NUV images observed with a regular cadence of 2 days. The GALEX TDS fields were selected to overlap with the Pan-STARRS1 Medium Deep Survey (PS1 MDS; Kaiser et al. 2010). GALEX TDS and PS1 MDS are well matched in field of view, sensitivity, and cadence (shown in Table 1). Here we present the analysis of 42 GALEX TDS fields which intersect with the PS1 MDS footprint, for a total area on the sky of ∼39.91 deg², which were monitored over a baseline of 3.32 years (2008 February–2011 June). In this paper we use PS1 MDS deep stack catalogs to characterize the optical hosts of GALEX TDS sources. Simultaneous UV and optical variability of GALEX TDS sources culled from matches with the PS1 transient alerts (Huber et al. 2011) will be presented in future papers.

The paper is structured as follows. In Section 2 we describe the GALEX TDS survey design, and in Section 3 we describe our statistical methods for selecting UV-variable sources and characterizing their UV variability. In Section 4 we describe the multiwavelength catalog data used to identify the hosts of the UV-variable sources, including archival optical data, a deep PS1 MDS catalog, and archival redshift and X-ray catalogs. In Section 5 we describe our sequence of steps for classifying the GALEX TDS sources, in Section 6 we summarize our classification results, and the UV variability properties of our classified sources. In Section 7 we conclude with implications for future surveys.

### Table 1

**GALEX TDS and PS1 MDS**

<table>
<thead>
<tr>
<th>Survey</th>
<th>Field of View (°)</th>
<th>Plate Scale (arcsec pixel⁻¹)</th>
<th>PSF FWHM (°)</th>
<th>m_{lim} per Epoch (mag)</th>
<th>Cadence (days)</th>
<th>Seasonal Visibility (months)</th>
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<td>5.3</td>
<td>23.3</td>
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<td>1.0</td>
<td>23.0</td>
<td>3</td>
<td>~6</td>
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### Table 2

**GALEX TDS Fields**

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February to 2011 June. PS1 began taking commissioning data of the MDS fields in 2009 May, but did not begin full survey operations until a year later. The GALEX far-UV (FUV) detector became non-operational in 2009 May, and so we only include near-UV (NUV) images in our study.

3. STATISTICAL MEASUREMENTS

3.1. Selection of Variable Sources

Since most galaxies are unresolved by the GALEX NUV 5′′ FWHM point-spread function (PSF), we can use simple aperture photometry instead of image differencing to measure variability. We create a master list of unique source positions from the pipeline-generated catalogs (Morrissey et al. 2007) for all the individual epochs, as well as deep stacks of all the epochs, using a clustering radius of 5′′. This radius is chosen such that for the typical astrometry error of GALEX of 0.5, the Bayesian probability that the match is real is larger than the Bayesian probability that the match is spurious (Budavári & Szalay 2008). The final master list includes 419,152 sources.

In order to select intrinsically variable sources in our survey, we first need to characterize the photometric errors. Although the GALEX images are Poisson-limited, the Poisson error underestimates the total error in the GALEX catalog magnitudes by a factor of ~2 (Trammell et al. 2007). This discrepancy is attributed to systematic errors such as uncertainties in the detector background and flat-field. Thus, we measure the photometric error empirically by calculating the standard deviation of aperture magnitudes in bins of mean magnitude, \(\bar{m}\).

We only include objects in the pipeline-generated catalogs that are detected in all or \(\geq 10\) epochs. In each bin of \(N\) objects with \(\bar{m}_i = (\sum_{k=1}^{n} m_{i,k})/n\) (each bin typically has \(N = 50\) to 1000 sources), we calculate for each epoch \(k\) of a total of \(n\) epochs,

\[
\sigma(\langle m \rangle, k) = \sqrt{\frac{1}{N-1} \sum_{i=1}^{N} (m_{i,k} - \langle m \rangle_i)^2},
\]

where \(m_{i,k}\) is the magnitude (in the AB system) given by \(m_{i,k} = -2.5 \log(f_6) + zp + C_{ap} - 8.2 E(B - V)\), \(f_6\) is the background-subtracted flux in a 6′′ radius aperture, \(zp = 20.08\), the aperture correction is \(C_{ap} = -0.23\) mag (Morrissey et al. 2007), and we correct for Galactic extinction using the values for \(E(B - V)\) listed in Table 2. We use 3\(\sigma\) clipping to remove outliers in the calculation of \(\sigma(\langle m \rangle, k)\) which can arise from artifacts.

The astrometric precision depends on the signal-to-noise of the source, thus we also empirically measure a magnitude-dependent clustering radius. We do so by measuring the cumulative distribution of spatial separations between the position in each epoch and the mean position for sources in bins of \(\langle m \rangle\), and record \(d_{95}(\langle m \rangle, k)\), the value for which 95% of sources have a separation less than or equal to that value. The resulting value for \(d_{95}\) is a strong function magnitude, increasing from \(\sim 1''\) for...
Figure 2. Dates of GALEX TDS NUV observations compared to PS1 MDS observations in the $g_P\,1$, $r_P\,1$, $i_P\,1$, and $y_P\,1$ bands. Dotted lines show yearly intervals. (A color version of this figure is available in the online journal.)

$\langle m \rangle = 18.0$ mag to $\sim 4''$ for $\langle m \rangle = 23.0$ mag. Figures 3 and 4 show $\sigma(\langle m \rangle, k)$ and $d_{95}(\langle m \rangle, k)$ for an example GALEX TDS field PS_COSMOS_MOS23, a quadratic fit to the median function for all epochs in that field, and the median function fit over all fields.

In our master list of source positions, we include all sources detected, including sources detected in only one epoch, and fix the centroid to the epoch for which the source is detected with maximum flux. We measure forced aperture magnitudes at the positions of each source in epochs where the source was not detected by the pipeline or the spatial separation of the matched source is greater than $d_{95}(\langle m \rangle, k)$. When the aperture magnitude is fainter than $m_{\text{lim}}$ in an epoch, it is flagged as an upper limit and replaced with $m_{\text{lim}}$, where $m_{\text{lim}} = -2.5 \log(5 \sqrt{(B_{\text{sky}} N_{\text{pix}}/T_{\text{exp}}) + zp + C_{\text{ap}}})$, where $B_{\text{sky}} = 3 \times 10^{-3}$ counts s$^{-1}$ pixel$^{-1}$, $N_{\text{pix}} = 16\pi$, and $T_{\text{exp}}$ is the exposure time of that epoch in seconds.

We select sources that have at least one epoch for which $|m_k - \langle m \rangle| > 5\sigma(\langle m \rangle, k)$, where $\langle m \rangle$ is calculated only from epochs that have a magnitude above the detection limit of that epoch. We use this selection method to be sensitive to short-term and long-term variability, as well as transients. This $5\sigma$ selection is quite conservative, and requires variability amplitudes increasing from $|\Delta m| > 0.1$ mag for $\langle m \rangle \sim 18$ mag up to $|\Delta m| > 1.0$ mag for $\langle m \rangle \sim 23$ mag.

(A color version of this figure is available in the online journal.)

Figure 3. Empirical determination of $1\sigma$ photometric errors as a function of mean magnitude from the standard deviation of sources detected by the pipeline catalogs in $\geq 10$ epochs for one of the GALEX TDS fields PS_COSMOS_MOS23. Solid green line shows a quadratic fit to the error function for the field PS_COSMOS_MOS23, and dashed blue line shows a quadratic fit to the median error function for all of the GALEX TDS fields. Solid red line shows the expected $1\sigma$ Poisson error, which underestimates the total photometric error by a factor of $\sim 2$. (A color version of this figure is available in the online journal.)
We make the following cuts to the \(5\sigma\) variable source sample to remove artifacts.

1. We remove sources with pipeline artifact flags indicating window bevel reflections or ghosts from the dichroic beam splitter.

2. We remove the brightest objects, with \((m) < 18.0\), due to the large area subtended by the PSF which causes uncertainty in the background subtraction.

3. We select sources within a radius of 0.55 of the center of the field, in order to avoid glints and PSF distortions, which are more prominent on the edges of the image, from un-corrected spatial distortions of photons recorded by the detectors.

4. We do not include objects that are within 1.5 of a \(m < 17\) mag source, to avoid regions affected by the bright source’s PSF and ghost artifacts. Ghost artifacts can appear within 30–60 above and below a bright source in the Y detector direction. Ghosts are point-like, and thus can only be identified from their Y detector position relative to a bright source. While ghosts do not usually appear in the \(GALEX\) pipeline catalogs, we apply this cut since our forced aperture photometry could mistake ghosts for transient sources.

5. We veto objects for which in the epoch of maximum \(\mid m_i - (m) \mid / \sigma((m), k)\) or maximum flux, the aperture flux ratio of the object has \(f_6 / f_{3.8} > R_{00}\), where \(f_6\) is the 6/0 radius aperture flux, \(f_{3.8}\) is the 3/8 radius aperture flux, and \(R_{00}\) is the maximum cumulative aperture flux ratio measured for 90% of the sources in the reference source sample used to calculate \(\sigma((m), k)\) in that epoch. This cut removes fluctuations in the background due to reflections from bright stars just outside the field-of-view, as well as epochs where the PSF is distorted due to a degradation in resolution which sometimes occurs in the Y detector direction. This also vetoes cases when the pipeline sheds a source into multiple sources, and the source is detected as variable because the center of the aperture is off-center from the peak source flux.

6. Finally, we visually inspect all of the remaining variable sources to remove any remaining artifacts that passed through the cuts above.

Figure 5 shows a gallery of good sources and vetoed variable sources from our automated cuts (for a bright reflection in panel (a), an off-center source in panel (b), and a likely ghost in panel (c)) and manual cuts (for a diffuse reflection in panel (d)). Our final \(GALEX\) TDS \(5\sigma\) variable sample after the cuts listed above has a total of 1078 sources.

### 3.2. Variability Statistics

We characterize the variability of each \(\sigma\) variable UV source using several statistical measures. We measure the structure function (following di Clemente et al. (1996)),

\[
V(\Delta t) = \sqrt{\frac{\pi}{2} \left( \langle | \Delta m_i | \rangle \right)^2 + \langle \sigma_i^2 \rangle / \Delta t},
\]

where brackets denote averages for all pairs of points on the light curve of an individual source with \(i < j\) and \(t_i - t_j = \Delta t\). The two-day cadence of the observations combined with the seasonal visibility of the fields results in a distribution of time intervals between observations (shown in Figure 6) that fall into six characteristic timescale bins: \(\Delta t_{2d} = 2 \pm 0.5\) d, \(\Delta t_{4d} = 4 \pm 0.5\) d, \(\Delta t_{8d} = 6 \pm 0.5\) d, \(\Delta t_{16d} = 8 \pm 0.5\) d, \(\Delta t_{32d} = 9.6 \pm 0.4\) yr, and \(\Delta t_{64d} = 1.7 \pm 0.4\) yr. We measure the structure function in these six bins, and define \(S_6\) to be the maximum value of the structure function evaluated for \(\Delta t_{2d}, \Delta t_{4d}, \Delta t_{6d}, \Delta t_{16d}\), and \(\Delta t_{32d}\), and \(S_6\) to be the maximum value of the structure function evaluated for \(\Delta t_{16d}\) and \(\Delta t_{64d}\). We also measure the intrinsic variability as defined by Sesar et al. (2007),

\[
\sigma_{in} = \sqrt{\Sigma^2 - \xi^2},
\]

where \(\Sigma^2 = (1/n - 1) \sum_{i=1}^{n} (m_i - \langle m \rangle)^2\) and
\[ \xi^2 = \frac{1}{n} \sum_{k=1}^{n} \sigma((m_k), k)^2 \], and the maximum amplitude of variability, \( \text{max}(|\Delta m|) \).

### 3.3. UV Light Curve

In order to flag possible transient UV events that may be associated with a SN or TDE, we differentiate between stochastic variability and flaring variability. We identify flaring UV variability as sources that show a constant flux \( \geq 10 \) days before the peak of the light curve, and do not fade more than \( 2\sigma \) below the faintest pre-peak magnitude (measured \( \geq 10 \) days before the peak). This selection criteria is tailored to the NUV rise times observed in SNe (Gezari et al. 2008b, 2010; Brown et al. 2009; Milne et al. 2010; Ofek et al. 2010) and TDE candidates (Gezari et al. 2006, 2008a, 2009). We define constant pre-peak flux as a light curve with a reduced \( \chi^2_r \leq 3 \), or that fade below \( 2\sigma \) of the faintest magnitude measured \( \geq 10 \) days before the peak as stochastically variable. We flag 116 flares, 145 transients, 595 stochastically variable sources, and remain with 222 sources with neither light-curve classification flag. In Figure 7 we show example light curves of sources flagged as stochastically variable (“V”), flares (“F”), and transient flares (“T”). Epochs \( \geq 10 \) days before the epoch of peak flux are circled in red, and the epoch of peak flux is circled in blue. The thick dashed line indicates the mean flux \( \geq 10 \) days before the peak.

(A color version of this figure is available in the online journal.)

### 4. HOST PROPERTIES

#### 4.1. Archival Optical Imaging Catalogs

We first characterize the host properties of the UV variable sources using archival optical \( u, g, r, i, \) and \( z \) photometry and morphology from matches to the Sloan Digital Sky Survey (SDSS) Photometric Catalog, Release 8 (Aihara et al. 2011; \( m_{\text{lim}} \sim 22 \) mag), the CFHTLS Deep Fields D1, D2, and D3 (\( m_{\text{lim}} \sim 26.5 \) mag) and Wide Fields W1, W3, and W4 (\( m_{\text{lim}} \sim 25 \) mag) merged catalogs version T0005,\(^9\) and the SWIRE ELAIS N1 and CDFS Region catalogs (\( m_{\text{lim}} \sim 24 \) mag; Surace et al. 2004). For sources with matches in multiple catalogs,

\[ \xi^2 = \frac{1}{n} \sum_{k=1}^{n} \sigma((m_k), k)^2 \], and the maximum amplitude of variability, \( \text{max}(|\Delta m|) \).

\(^9\) http://terapix.iap.fr/rubrique.php?id_rubrique=252

\(^{10}\) www.ast.cam.ac.uk/~wfcsur/technical/photom/colours

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**Figure 6.** Histogram of the time intervals between all pairs of observations for the 42 fields in GALEX TDS. Hatched regions show the time intervals over which the structure function is measured for all of the sources.

**Figure 7.** Example light curves of sources flagged as stochastically variable (“V”), flares (“F”), and transient flares (“T”). Epochs \( \geq 10 \) days before the epoch of peak flux are circled in red, and the epoch of peak flux is circled in blue. The thick dashed line indicates the mean flux \( \geq 10 \) days before the peak. (A color version of this figure is available in the online journal.)

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The GALEX TDS fields overlap with the PS1 MDS fields MD01 (PS_XMMLSS), MD02 (PS_CDFS), MD04 (PS_COSMOS), MD07 (PS_GROTH), MD08 (PS_ELAISN1), and MD09 (PS_VVDS22H). The Pan-STARRS1 observations are obtained through a set of five broadband filters, \((g_{\text{PS1}}, r_{\text{PS1}}, i_{\text{PS1}}, z_{\text{PS1}}, y_{\text{PS1}})\). Further information on the passband shapes is described in Stubbs et al. (2010). The PS1 MD fields are observed with a typical cadence in a given filter of three days, with an observation in the \( g_{\text{PS1}} \) and \( r_{\text{PS1}} \) bands on night one, in the \( i_{\text{PS1}} \) band on night two, and the \( z_{\text{PS1}} \) band on night three, with \( y_{\text{PS1}} \)-band observations during each of three nights on either side of the Full Moon. Image differencing is performed on the nightly stacked images, reaching a typical \( 5\sigma \) detection limit of \( \sim 23.3 \) mag per epoch in the \( g_{\text{PS1}}, r_{\text{PS1}}, i_{\text{PS1}}, \) and \( z_{\text{PS1}} \) bands and \( \sim 21.7 \) mag in the \( y_{\text{PS1}} \) band. Image difference detections from the PS1 Image Processing Pipeline (IPP; Magnier 2006) and an independent pipeline hosted by Harvard/CfA (Rest et al. 2005) are internally distributed to the PS1 Science Consortium as transient alerts for visual inspection and classification.
Deep stacks of the multi-epoch observations were generated to provide deep imaging with a 5σ point-source limiting magnitude of \( \sim 24.9, 24.7, 24.7, 24.3, \) and 23.2 mag in the \( g_p1, r_p1, i_p1, z_p1, \) and \( y_p1 \) bands, respectively. The magnitudes are in the “natural” Pan-STARRS1 system, \( m = -2.5 \log(\text{flux}) + m' \), with a relative zero-point adjustment \( m' \) made in each band for each individual epoch (Schlafly et al. 2012) before stacking to conform to the absolute flux calibration in the AB magnitude system (Tonry et al. 2012). We convert the PS1 magnitudes to the SDSS absolute flux calibration in the AB magnitude system (Tonry et al. 2012) before stacking to conform to the epoch (Schlafly et al. 2012) before stacking to conform to the absolute flux calibration in the AB magnitude system (Tonry et al. 2012). We convert the PS1 magnitudes to the SDSS system using the bandpass transformations measured for stellar extinction using the Schlegel et al. (1998) dust map for Galactic extinction using the Schlegel et al. (1998) dust map. We also take advantage of the many archival X-ray and spectroscopic catalogs available from the overlap of the GALEX TDS survey with legacy survey fields. In the PS_CDFS field we use X-ray catalogs from the 0.3 deg\(^2\) Chandra Extended CDFS survey (Giacconi et al. 2002; Lehmer et al. 2005; Virani et al. 2006), and redshift catalogs from the VIMOS VLT Deep Survey (VVDS; Le Fèvre et al. 2004), and a compilation of redshift catalogs from GOODS and SWIRE.\(^{11}\) In the PS_XMMLSS field we use X-ray catalogs from the 5.5 deg\(^2\) XMM-LSS survey (Chiappetti et al. 2005; Pierre et al. 2007), and redshift catalogs from VVDS (Le Fèvre et al. 2005). In the PS_COSMOS field, we use X-ray catalogs from the 1.9 deg\(^2\) XMM-Newton Wide-Field Survey (Hasinger et al. 2007) and the 0.9 deg\(^2\) Chandra COSMOS survey (Elvis et al. 2006), and redshifts from the Magellan COSMOS AGN survey (Trump et al. 2007, 2009), the VLT zCOSMOS bright catalog (Lilly et al. 2007, 2009), and the Chandra COSMOS Survey catalog (Civano et al. 2012). In the PS_GROTH field we use X-ray catalogs from the 0.67 deg\(^2\) Chandra Extended Groth Strip (Nandra et al. 2005; Laird et al. 2009) and redshift catalogs from the DEEP2 Galaxy Redshift Survey (Newman et al. 2012). For PS_ELAISN1 we use the X-ray catalog from the 0.08 deg\(^2\) Chandra ELAIS-N1 deep X-ray survey (Manners et al. 2003). Figure 1 shows the overlap of the GALEX TDS fields with the archival X-ray surveys. Finally, we also match the sources with the ROSAT All-Sky Bright Source and All-Sky Survey Faint Source catalogs (Voges et al. 1999, 2000).

5. CLASSIFICATION

We classify the GALEX TDS sources using a combination of optical host photometry and morphology, UV variability statistics, and matches with archival X-ray and redshift catalogs. Table 3 summarizes the sequence of steps we use to classify the sources, which we describe in detail below.

5.1. Cross-match with Optical Catalogs

We first cross-matched our 1078 GALEX TDS sources with the archival \( u, g, r, i, z \) catalogs described in Section 4.1 with a

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\(^{11}\) http://www.eso.org/sci/activities/garching/projects/goods/MASTERCAT_v3.0.dat
matching radius of 3′′. This radius is recommended for matches between GALEX and ground-based optical catalogs (Budavári 
& Szalay 2008), and corresponds to a spurious match rate of only 1%–2% at the high Galactic latitudes of the GALEX TDS fields (Bianchi et al. 2011). However, we found that there was a population of “orphans” (no optical match within 3′′) that were detected in their NUV low-state, and had a match between 3′′–4′′ with an optically identified quasar. Given the strong likelihood that these are real matches, we increased our matching radius to 4′′. We used the star/galaxy classifications from the catalogs to label sources as point sources (pt) or extended sources (ext). This results in 878/1078 optical matches (81%), with the majority of sources without matches in PS_CDFS, which is only partially covered by the archival catalogs.

We then match the sources that do not have archival optical matches to the PS1 MDS catalog described in Section 4.2, this increases the number of sources with optical photometry and morphology (albeit without the r band) to 1057/1078 (98%). Figure 9 shows a histogram of the r-band magnitude of the optical hosts, and the NUV - r colors of the GALEX TDS sources in their low-state. The optical hosts have a distribution that peaks at r ~ 21 mag, over 3 mag brighter than the detection limit of PS1 MDS, and NUV - r ~ 1 mag. Sources not detected in their NUV low-state are shown as upper limits in the NUV - r color histogram, and peak at NUV - r > 2 mag.

5.2. Orphans

We visually inspected the PS1 stack images at the locations of the 21 sources with no optical matches, and confirm that they are true orphan events. Furthermore, all of the orphans are undetected in their low-state in the NUV, with upper limits of NUV > (22.3–23.1) mag. Thus the orphan hosts are likely distant stars or faint galaxies (i.e., dwarf galaxies or high-redshift galaxies) that are undetected during their low-state in the optical and NUV.

5.3. Color and Morphology Cuts

We first use the color and morphology of the optical hosts to classify the GALEX TDS 5σ UV variable sources. We define quasars as sources with optical point-source hosts with

$$u - g < 0.7 
-0.1 < g - r < 1.0$$

in order to avoid the stellar locus and white dwarfs (Richards et al. 2002). Note that this color selection can be contaminated by cataclysmic variable stars (CVs), which overlap in color–color space with the quasar sample. Indeed, two of the sources classified by color as quasars are in fact spectroscopically confirmed CVs (VVDS22H_MOS05-05 and ELAISN1_MOS15-02). VVDS22H_MOS05-05 is ROTSE3 J221519.8-003257.2, a confirmed cataclysmic variable star with a dwarf-nova type spectrum. We observed ELAISN1_MOS15-02 with the APO 3.5 m telescope Dual Imaging Spectrograph (DIS) on 2011 May 3 and detected broad Balmer emission lines from a Galactic source, characteristic of a CV/dwarf nova spectrum. While these sources stood out easily because of their extreme magnitude of variability of |Δm| > 4 mag, shown in Figure 10, there

Figure 9. Left: histogram of r magnitudes of optical matches. Right: histogram of NUV - r colors of optical matches for those sources detected in the NUV during their low-state (solid lines), and those sources with upper limits during their low-state (dashed lines).

Table 3

<table>
<thead>
<tr>
<th>Step</th>
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<th>Archive Pt</th>
<th>Archive Ext</th>
<th>PS1 Pt</th>
<th>PS1 Ext</th>
<th>Orphan</th>
<th>QSO</th>
<th>RRL</th>
<th>Mdw</th>
<th>Star</th>
<th>AGN</th>
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<td>$\sigma_u/\sigma_g \geq 3$</td>
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<td>19</td>
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<tr>
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<tr>
<td>X-ray/spec match</td>
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</table>

302 21 358 37 53 22 305 93 189
may be lower amplitude CV events still hiding in our quasar sample. However, given the low surface density of CVs relative to quasars in the sky (Szkody et al. 2011), the expected contamination rate is consistent with the two CVs identified.

We define RR Lyrae stars as sources with optical point source hosts with

\[
\begin{align*}
0.75 < u - g < 1.45 \\
-0.25 < g - r < 0.4 \\
-0.2 < r - i < 0.2 \\
-0.3 < i - z < 0.3
\end{align*}
\] (Sesar et al. 2010). Note that the color cuts for quasars and RR Lyrae require the \( u \)-band, which is not available for sources with PS1-only matches. However, we define M dwarf stars as point sources with \( r - i > 0.42 \) (West et al. 2011), which does not require \( u \)-band data. We classify stars on the main stellar locus as those with

\[
\begin{align*}
1.0 < u - g < 2.25 \\
0.4 < g - r < 1.0 \\
-0.2 < r - i < 2.0 \\
-0.3 < i - z < 1.0
\end{align*}
\] (modified from Yanny et al. (2009)). This color and morphology selection results in 37 RR Lyrae, 53 M-dwarf flare stars, 17 stars, and 325 quasars. Figure 11 shows the optical color–color diagram of the sources with optical point-source hosts and \( u \)-band data, and their classifications as quasars, RR Lyrae, M dwarfs, and stars.

5.4. UV Variability Cuts

Figure 12 shows the structure function on timescales of days and years described in Section 3 for the sources classified as RR Lyrae and quasars in Section 5.3. In Figure 13 we show the ratio of the NUV structure function on timescales of years to days \( (S_{yr}/S_d) \). While quasars demonstrate a wide range of \( S_{yr}/S_d \), all RR Lyrae have \( S_{yr}/S_d < 3 \). We use this UV variability property to relax our color constraints, and increase our photometric sample of quasars to all sources with optical point-source hosts.
Figure 13. Histogram of the ratio of the NUV structure function on timescales of years and days for sources classified as RR Lyrae (red) and quasars (blue). The dotted line shows the selection criteria of $S_y/S_d \geq 3$ used to classify sources with optical point source hosts as quasars. (A color version of this figure is available in the online journal.)

Figure 14. Colors of archival optical matches to UV variables with extended optical hosts. Sources classified as AGN (by either stochastic UV variability, archival spectra, and/or an X-ray match) are color coded in green. Sources with matches with archival X-ray matches are circled in purple. Note that all galaxy sources with an archival X-ray match are defined as AGNs. (A color version of this figure is available in the online journal.)

Figure 15. Maximum NUV variability amplitude for sources without a classification. The dashed blue line shows the median 5σ variability selection function as a function of mean magnitude. The nature of the optical host is indicated (point source, galaxy, or orphan), and flaring sources (including transients) are marked in yellow. (A color version of this figure is available in the online journal.)

5.5. X-Ray Sources

The archival X-ray catalogs overlap with $\sim 8.45$ deg$^2$ of the GALEX TDS survey area. Within this area, 81/89 quasars, 92/105 AGNs, and 8/9 M-dwarf stars are detected in the X-rays. In addition, there are nine optical point sources with X-ray matches that are likely quasars and M dwarfs just outside the quasar and M-dwarf color–color selection regions. UV variability selection appears to be selecting a similar population of active galaxies and M dwarfs as X-ray detection, since $\sim 90\%$ of the UV variability-selected active galaxy and M dwarf sample is also detected in the X-rays. However, only 2% of all the X-ray sources (the majority of which are active galaxies) are detected as UV variable at the selection threshold of the GALEX TDS catalog.

5.6. GUVV Catalog

We also cross-match our GALEX TDS sample with the first and second GALEX Ultraviolet Variability Catalogs (GUVV-1 and GUVV-2) from Welsh et al. (2005) and Wheatley et al. (2008). These catalogs include 894 UV variable sources ($\Delta m > 0.6$ mag in the NUV) selected from an analysis of archival GALEX AIS, MIS, DIS, and Guest Investigator (GI) fields with repeated observations. With a cross-matching radius of 4″, we find a match with 36 GUVV sources. For the 15 matches that have GUVV classifications, all are classified by GUVV as active galaxies (AGN or quasar), which are in agreement with our GALEX TDS classifications. Of the 21 matches without GUVV classifications, we find 5 sources classified by GALEX TDS as RR Lyrae, 9 classified as active galaxies (AGN or quasar), 6 with optical point-source hosts, and 1 with a galaxy host.

5.7. Unclassified Sources

The remaining 302 unclassified sources include 91 with optical point-source hosts, which are likely stars, quasars with non-standard colors (high-redshift or reddened), or unresolved galaxies, 190 with galaxy hosts, and 21 orphans. The 190 galaxy hosts may either be faint AGNs with poorly constrained UV light curves, or hosts of UV-bright extragalactic transients. In Figure 15 we show the maximum $|\Delta m|$ in the NUV as a function of low-state NUV magnitude of the remaining unclassified sources. The unclassified UV source with the most extreme amplitude, ELAISN1_MOS15-09 with $|\Delta m| > 4.2$ mag, was
spectroscopically confirmed to be from the nucleus of an inactive galaxy at $z = 0.1696$, and its UV/optical flare detected by GALEX TDS and PS1 MDS (PS1-10jh) was attributed to the tidal disruption of a star around an SMBH (Gezari et al. 2012). Also in this sample is a UV transient spectroscopically confirmed to be a Type IIP SN 2010aq at $z = 0.086$ (COSMOS_MOS26-29), whose UV/optical light curve from GALEX TDS and PS1 MDS was fitted with early emission following SN shock breakout in a red supergiant star (Gezari et al. 2010). Both of these spectroscopically classified extragalactic transients are labeled in Figure 15.

Our $5\sigma$ selection criteria translates to a limiting sensitivity to transients in a host galaxy with a magnitude $m_{\text{host}}$ of a magnitude of

$$m_{\text{trans}} = -2.5 \log \left(10^{\frac{m_{\text{host}} - 5\sigma_{\text{host}}}{25}} - 10^{\frac{m_{\text{host}}}{25}}\right),$$

which ranges from $m_{\text{trans}} \sim 20.0$ mag for $m_{\text{host}} = 18$ mag to $m_{\text{trans}} \sim 22.7$ mag for $m_{\text{host}} = 23$ mag. Thus, our variability selection threshold is less sensitive to transients in host galaxies with bright NUV fluxes. On the red sequence of galaxies, where $M_{\text{NUV}} \approx -14.5$ (Wyder et al. 2007), this selection effect is not as much of an issue, since already for $z > 0.05$ one gets $m_{\text{host}} > 22$ mag. However, star-forming galaxies on the blue sequence are 2.5 mag brighter in the NUV, and thus the host galaxy brightness can be a factor in reducing the sensitivity to faint transients. For example, our GALEX TDS $5\sigma$ sample does not include SN 2009kl, a luminous Type IIP SN in a star-forming galaxy at $z = 0.182$ which we reported our GALEX TDS detection of in Botticella et al. (2010). This source varied at only the $4.25\sigma$ level in the NUV during its peak. However, because this transient was selected from a spatial and temporal coincidence with a PS1 transient alert, we could lower our threshold for variability selection in the UV. The systematic selection of SN and TDE candidates from the joint GALEX TDS and PS1 MDS transient detections will be presented in future papers.

6. DISCUSSION

6.1. Classification Demographics

Figure 16 shows a pie diagram of the source classifications. Out of the total of 1078 GALEX TDS sources, 62% are classified as actively accreting SMBHs (quasars or AGN), and 10% as variable and flaring stars (including RR Lyrae, M dwarfs, and CVs). Note that the relative fraction of the different classes of sources is sensitive to both their intrinsic magnitude distribution, and the magnitude-dependent variability selection function of the sample. Table 4 gives the GALEX TDS catalog, sorted by decreasing NUV amplitude, with the GALEX ID, R.A., decl., low-state NUV magnitude, maximum amplitude of NUV variability ($\Delta m_{\text{max}}$), intrinsic variability ($\sigma_{\text{int}}$), the structure function on day ($S_d$) and year ($S_y$) timescales; the characteristics of the NUV light curve: flaring (F) or stochastically variable (V); the morphology of the matching optical host: point-source (pt) or extended (ext); the color classification of the matching optical host: RR Lyrae (RRL), M dwarf star (Mdw), star (Star), quasar (QSO), and active galactic nuclei (AGN), and sources with no classification with galaxy hosts (Galaxy), no host (Orphan), and point-source hosts (Point).

(A color version of this figure is available in the online journal.)
AGNs, and RR Lyrae, respectively. For the transient sources, we can calculate a total surface density rate, #/(area × t_{eff}), where t_{eff} is the effective survey time at the cadence that matches the characteristic timescale of the transient. For extragalactic transients such as young SNe and TDEs, which vary on a timescale of days, we use the time intervals for which the fields were observed with a cadence of 2.0±0.5 days. If we include all flaring or transient GALEX TDS sources with a galaxy host, this yields a surface density rate of 52±38 deg^{-2} yr^{-1} for extragalactic transients. For M dwarfs which vary on timescales shorter than an individual observation, we use the total exposure time for each epoch. If we assume a survey with a cadence of two days and t_{exp} = 1.5 ks, this translates to a surface density rate for M dwarfs of 15±10 deg^{-2} yr^{-1}.

6.2. UV Variability Properties of Classified Sources

Various optical studies of rest-frame UV variability in high redshift quasars have demonstrated that the variability of AGN increases with decreasing rest wavelength (di Clemente et al. 1996; Vanden Berk et al. 2004; Wilhite et al. 2005). In Figure 18, we show histograms of σ_{int} for the UV variable sources with classifications. Quasars show a distribution of σ_{int} with a mean that is ~2 times larger than measured at optical wavelengths from the SDSS Stripe 82 sample from Sesar et al. (2007). This effect is even more pronounced in the structure function on years timescales (S_{r}), which has a mean that is five times larger than the mean measured in the r-band (λ_{eff} = 6231) from Schmidt et al. (2010). This trend is consistent with the wavelength-dependent rise in variability amplitude observed in the structure function for quasars in the rest-frame UV (Vanden Berk et al. 2004) and observed UV (Welsh et al. 2011).

The fact that AGN become bluer during high states of flux (Giveon et al. 1999; Geha et al. 2003; Gezari et al. 2008a) has been attributed to increases in the characteristic temperature of the accretion disk in response to increases in the mass accretion rate (Pereyra et al. 2006; Li & Cao 2008). However, Schmidt et al. (2012) argue that the color variability observed in individual quasars in their SDSS Stripe 82 sample is stronger than expected from just varying the accretion rate (M) in accretion disk models. In a future study, we will use simultaneous UV and optical light curves from GALEX TDS and PS1 MDS for our 358 individual quasars to test this result with a larger dynamic range in wavelength.

For the subsample of 95 quasars with archival redshifts (z_{mean} = 1.26, σ_{z} = 0.39), in Figure 19 we plot σ_{int} versus the low-state NUV absolute magnitude, and find a steep negative correlation fitted by log(σ_{int}) = (1.6 ± 0.1) + (β/2.5)M_{NUV}, where β = 0.24 ± 0.04, in excellent agreement with the trend for increased variability in lower luminosity quasars seen from optical observations with β = 0.246 ± 0.005 (Vanden Berk et al. 2004), and shallower than expected for a Poissonian process which has β = 0.5 (Cid Fernandes et al. 2000).

We also show the subset of 68 AGN with archival redshifts (z_{mean} = 0.64, σ_{z} = 0.55), which clearly do not show a relation between σ_{int} and low-state NUV absolute magnitude. This is most likely a result of dilution of the variability amplitude from the contribution of the host galaxy in the NUV.

The largest values of |Δm| (plotted in Figure 10) are found in RR Lyrae and M dwarfs, with a tail of large amplitude variations reaching up to |Δm| = 2.9 mag in RR Lyrae and up to |Δm| = 4.6 mag in M dwarfs. In the optical, the RR Lyrae structure function is very weakly dependent on timescale, with an amplitude of 0.1–0.2 mag (Schmidt et al. 2010). The NUV structure function also shows a weak dependence of amplitude on timescale when comparing the structure function on days to years timescales, however, with an amplitude that is...
the expectation for higher temperatures during higher states
factors of UV variability amplitudes in RR Lyrae and quasars that are
Lyraes, 358 quasars, and 305 AGNs. We find median intrinsic
redshift catalogs. This yields a sample of 53 M dwarfs, 37 RR
variability behavior, and cross-matches with archival X-ray and
and their classifications based on optical host properties, UV
variability is driven by variations in surface temperature from
dependence on variability amplitude can be explained if the
∼M ∼1

Figure 19. Intrinsic NUV variability as a function of low-state NUV absolute magnitude. M_{NUV} = \text{low(NUV)} - DM where DM is the distance modulus, for the subsample of quasars (blue dots) and AGNs (green circles) with catalog redshifts. Dashed blue line shows the fit to the quasars to log(\sigma int) \propto (\beta/2.5)M. The spectroscopically classified extragalactic transients (TDE PSI-10jh and SN 2010aq) are labeled.

(A color version of this figure is available in the online journal.)

~3 times larger in the NUV than in the optical. This wavelength
dependence on variability amplitude can be explained if the
variability is driven by variations in surface temperature from
pulsations of the stellar envelope, where higher states of flux are
associated with higher temperatures (Sesar et al. 2007).

7. CONCLUSIONS

We provide a catalog of over a thousand UV variable sources
and their classifications based on optical host properties, UV
variability behavior, and cross-matches with archival X-ray and
redshift catalogs. This yields a sample of 53 M dwarfs, 37 RR
Lyraes, 358 quasars, and 305 AGNs. We find median intrinsic
UV variability amplitudes in RR Lyrae and quasars that are
factors of >3 larger than at optical wavelengths, consistent with
the expectation for higher temperatures during higher states
of flux. The regular cadence and wide area of GALEX TDS
enables us to systematically discover persistent and transient
(i.e., tidal disruption of a star) accreting SMBHs over wide fields
of view, study the contemporaneous UV and optical variability
of variable stars, and catch young core-collapse SNe within
the first days after explosion. The overlap of the GALEX TDS
with PS1 MDS and multiwavelength legacy survey fields will
continue to be helpful for classifying transient sources in these
heavily observed fields. We also measure the surface densities
of variable sources and the surface density rates for transients
as a function of class in the UV for the first time.

With GALEX TDS we are only scratching the surface of UV
variability. Our 3σ sample of 1078 sources is less than 0.3% of
the 419,152 UV sources in the field (with an average density of
1.3 \times 10^4 UV sources per square degree down to m_{int} = 23 mag).
Looking to the future, the discovery rate for UV variable sources
and UV transients could increase by several orders of magnitude
with the launch of a space-based UV mission with a wide
field of view (several deg^2), a survey strategy of daily cadence
observations over ~100 deg^2, and detectors with an order of
magnitude improved photometric precision relative to GALEX.
In the optical sky, 90% of quasars vary with \sigma int > 0.03 mag on the timescales of years (Sesar et al. 2007). Given the factor of ~2 larger \sigma int observed for quasars in the NUV, one could achieve a nearly complete sample of low-redshift quasars with

photometric errors of \sigma(m) \sim 0.01 mag. In coordination with
a ground-based optical survey, such as Pan-STARRS2 (Burgett
2012) or LSST,12 this could yield the simultaneous UV and
optical detection of \sim 10^4 variable quasars and \sim 10^3 RR Lyrae
and M dwarfs, as well as increase the discovery rate of UV-
bright extragalactic transients (young SNe and TDEs) by a factor
of ~100.

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