INVESTIGATING CO₂ RESERVOIRS AT GALE CRATER AND EVIDENCE FOR A DENSE EARLY ATMOSPHERE. P. B. Niles1, P. D. Archer2, E. Heil3, J. Eigenbrode4, A. McAdams5, B. Sutter6, H. Franz7, R. Navarro-Gonzalez8, D. Ming1, P. Mahaffy4, F. J. Martin-Torres5,7, and M. Zorzano6; 1Astromaterials Research and Exploration Science, NASA Johnson Space Center, Houston, TX 77058; (paul.b.niles@nasa.gov); 2Jacobs, NASA Johnson Space Center, Houston, TX 77058; 3HX5-Jacobs JETS Contract, NASA Johnson Space Center, Houston, TX 77058; 4NASA Goddard Space Flight Center, Greenbelt, MD; 5Instituto Andaluz de Ciencias de la Tierra (CSIC-UGR), Grenada, Spain; 6Centro de Astrobiología (INTA-CSIC), Madrid, Spain; 7Division of Space Technology Department of Computer Science, Electrical and Space Engineering, Lulea University of Technology, Kiruna, Sweden.

Introduction: One of the most compelling features of the Gale landing site is its age. Based on crater counts, the formation of Gale crater is dated to be near the beginning of the Hesperian near the pivotal Hesperian/Noachian transition [1, 2]. This is a time period on Mars that is linked to increased fluvial activity through valley network formation and also marks a transition from higher erosion rates/clay mineral formation to lower erosion rates with mineralogies dominated by sulfate minerals [3].

Results from the Curiosity mission have shown extensive evidence for fluvial activity within the crater suggesting that sediments on the floor of the crater and even sediments making up Mt. Sharp itself were the result of longstanding activity of liquid water [4].

Warm/wet conditions on early Mars are likely due to a thicker atmospheric and increased abundance of greenhouse gases including the main component of the atmosphere, CO₂ [5]. Carbon dioxide is minor component of the Earth’s atmosphere but plays a major role in surface water chemistry, weathering, and formation of secondary minerals. An ancient martian atmosphere was likely dominated by CO₂ and any waters in equilibrium with this atmosphere would have different chemical characteristics.

Studies have noted that high partial pressures of CO₂ would result in increased carbonic acid formation and lowering of the pH so that carbonate minerals are not stable [6]. However, if there were a dense CO₂ atmosphere present at the Hesperian/Noachian transition, it would have to be stored in a carbon reservoir on the surface or lost to space. The Mt. Sharp sediments are potentially one of the best places on Mars to investigate these CO₂ reservoirs as they are proposed to have formed in the early Hesperian, from an alkaline lake, and record the transition to an aeolian dominated regime near the top of the sequence [1].

This study seeks to better understand the CO₂ content of the soils and sediments investigated by the MSL rover at Gale crater with the goal of trying to piece together the nature of the atmosphere and climate during the early Hesperian.

Methods: The SAM instrument on the MSL rover provides the capability of analyzing drilled sample powders via pyrolysis and evolved gas analysis (EGA) via mass spectrometry [7]. This capability provides an excellent insight into the nature of CO₂ in the samples as the temperature at which CO₂ evolves is highly indicative of the nature of the phase in which it is stored. In general carbonate minerals evolve between 500 and 800 °C while reduced carbon phases evolve at temperatures lower than 400 °C. In both cases there can be exceptions to this relationship, so it is important to interpret EGA results with care.

Results: The SAM instrument has examined 5 different samples from the Gale crater region and CO₂ has been a major volatile component of each sample, however the carbon contents have not indicated that a carbon containing phase such as carbonate minerals make up a significant portion of the samples. All samples have CO₂ contents below 1 wt% but not below 0.1 wt% (Table 1).

The Rocknest soils are the only samples which show a substantial CO₂ evolution above 500 °C, indicating the possible presence of carbonate minerals. All of the other samples analyzed to date have shown lower temperature CO₂ releases indicating that the CO₂ is not evolving from a carbonate mineral (see discussion below).

Discussion: The presence of significant CO₂ (> 0.1 wt%) within all of the samples analyzed to date suggests that these samples do show some interaction with the atmosphere. However, the amount of CO₂ present and the temperatures at which it evolve are substantially different from what was expected.
The total amount of CO₂ in these sediments does not indicate the presence of a substantial reservoir of CO₂ (Table 1). Assuming that these samples may have similar carbon content to the average martian soil, then a 50 m global equivalent layer (GEL) yields only 50 mbar of CO₂ storage. This is well short of the ~1 bar estimates of the ancient martian atmosphere [5].

The CO₂ releases in all of the samples analyzed to date are also well below 500° C. While it might be possible that this is from a fine grained carbonate mineral [8], this has not yet been demonstrated and most carbonates decompose above 500° C, even Fe/Mg rich varieties.

CO₂ releases are often associated with oxygen releases in the martian data (Fig. 2) suggesting possible combustion of organic matter [9]. However, if this is true it suggests substantial quantities of organic carbon (up to 2000 ppm) (Table 1), and GC-MS analyses do not detect substantial organic fragments that might be associated with kerogen or other complex organic carbon species [10]. Instead it is evolved entirely as CO₂ and CO. This suggests some other form of oxidized carbon may be the carbon source.

One intriguing possibility is Fe(II) oxalate (Fig 1-2) which releases CO₂ at a very similar temperature [10, 11]. Oxalate minerals could form as a result of the breakdown of more complex organic species [11], or could possibly be the result of radiolysis of CO₂ hydrates [12].

Conclusions: The total amount of CO₂ in the Gale crater soils and sediments is significant but lower than expected if a thick atmosphere was present at the Hesperian/Noachian boundary. Likewise, the absence of carbonates suggests that CO₂ weathering processes similar to those present on Earth were not dominant. Instead it is possible that more exotic CO₂ deposition has occurred driven by atmospheric photochemistry and/or degradation of organic carbon.

Table 1. Average CO₂ evolved from samples analyzed by the SAM instrument.

<table>
<thead>
<tr>
<th>Sample</th>
<th>Evolved CO₂ (µmol)*</th>
<th>C (ppm)</th>
<th>Wt% CO₂ GEL 10m</th>
<th>GEL 50 m</th>
</tr>
</thead>
<tbody>
<tr>
<td>Rocknest</td>
<td>9.9 ± 1.9</td>
<td>2640</td>
<td>1.0% 11 mbar</td>
<td>54 mbar</td>
</tr>
<tr>
<td>John Klein</td>
<td>6.9 ± 1.8</td>
<td>1840</td>
<td>0.7%  8 mbar</td>
<td>38 mbar</td>
</tr>
<tr>
<td>Cumberland</td>
<td>2.7 ± 1.0</td>
<td>720</td>
<td>0.3%  3 mbar</td>
<td>16 mbar</td>
</tr>
<tr>
<td>Windjana</td>
<td>9.2 ± 2.3</td>
<td>2453</td>
<td>0.9% 10 mbar</td>
<td>48 mbar</td>
</tr>
<tr>
<td>Confidence Hills</td>
<td>4.3 ± 1.5</td>
<td>1147</td>
<td>0.4%  4 mbar</td>
<td>21 mbar</td>
</tr>
</tbody>
</table>

References: