MODELING THE GEOLOGIC HISTORY OF MT. SHARP. A. Pascuzzo1 and C. Allen2, 1Department of Geosciences, Smith College, Northampton, MA 01063 (apascuzz@smith.edu), 2Astroaterials Research and Exploration Science, NASA-JSC (carlton.c.allen@nasa.gov).

Introduction: Gale is a ~155 km diameter crater located on the martian dichotomy boundary (5°S 138°E). Gale is estimated to have formed 3.8 – 3.5 Gya, in the late Noachian or early Hesperian [1,2,3]. Mt. Sharp, at the center of Gale Crater, is a crescent shaped sedimentary mound that rises 5.2 km above the crater floor (Fig. 1). Gale is one of the few craters that has a peak reaching higher than the rim of the crater wall [1,2,4]. The Curiosity rover is currently fighting to find its way across a dune field at the northwest base of the mound searching for evidence of habitability.

This study used orbital images and topographic data to refine models for the geologic history of Mt. Sharp by analyzing its morphological features. In addition, it assessed the possibility of a peak ring in Gale. The presence of a peak ring can offer important information about how Mt. Sharp was formed and eroded early in Gale’s history.

Methodology:

Craters with sedimentary mounds. Craters located between 60°N - 60°S and 100 km-230 km in diameter with interior deposits were searched for using the Mars Crater Database [5]. The latitude limits were chosen to avoid ice processes that may affect mound formation, erosion and exposure. This search yielded 255 craters, of which 6 were selected (Fig 2, Table 1) to observe closely using orbital images. Context Camera (CTX), High Resolution Imaging Science Experiment (HiRISE), and Thermal Emission Imaging System (THEMIS) images were used to identify similarities and differences in mound morphology with Mt. Sharp.

The six craters were chosen based on their unique mound morphology as well as their similarity to Mt. Sharp. Topographic profiles for each crater were constructed using ArcMap10.1 software and Mars Orbital Laser Altimeter (MOLA) derived digital elevation models.

Craters with peak rings. Craters between 60°N - 60°S and 100 km-230 km in diameter with peak rings were cataloged using the same Mars Crater Database [9] as well as examining MOLA data and GoogleMars. Four craters were selected for close analyses of morphologic features and textural characteristics (Table 2).

Results and Discussion: Craters with mounds. Craters with strong evidence of fluvial processes, for example Firsoff (Fig 2d), have dried up river channels and layered sediment along the walls. This suggests that the crater was once partially or completely filled followed by erosion to form a circular mound in the center of the crater. Prominent yardangs show recent prevailing wind directions.

Cerulli (Fig 2a) and Becquerel (Fig 2c) have sedimentary mound deposits showing little correlation to the other four craters or to Gale. There isolated mounds show evidence for glaciation and periodic deposition in addition [6].

Nicholson as Gale analog. Out of the six craters analyzed, Nicholson (Fig 2f) is the closest analog to Gale in terms of mound morphology. Both Gale and Nicholson essentially lie on the dichotomy boundary [3]. Both craters are younger than the Hesperian Noachian transition unit and are located at the equator. Therefore they may have been exposed to similar climate evolution when Mars transitioned from a wet to arid climate [3]. Nicholson has a large central peak and a surrounding sedimentary mound similar to Gale. The cross section (Fig 3) of the Nicholson mound shows the rim and crater floor slope down towards the north-northwest [4]. The peak of the Nicholson mound rises higher than the surrounding rim, similar to Gale.

Both craters show evidence of deposition of wind-transported sediments on the mound and central peak.
On Gale, there is a thin young layer of wind transported possibly an outer layer of the Medusae Fossae formation (MFF) [4, 7]. On the east and west lobes of the Nicholson mound the surfaces exhibit crescentic scour features, similar to features found in the MFF, indicating a young layer of MFF may also present on the Nicholson mound [8].

Another distinguishing feature on Nicholson’s mound is the large deep channel, ~18 km long, running from the center to the near southeast base of the mound. A similar large canyon incises the western lobe of the LM of Mt. Sharp. Such canyons suggest large quantities of water were once present.

The major morphological feature that makes Gale unique is the crescent shape of Mt. Sharp. The comparison mounds are relatively concentric with their crater rims, but Gale is a half-circle mound with the central peak exposed along the southern face. This unique erosion of Mt. Sharp may suggest the presence of a partial peak ring that allowed protection of the sediment [9].

Peak ring morphology. A peak ring typically has a diameter that is half of its craters diameter. Peak rings are composed of deep-seated material the look fundamentally different than bedded sedimentary mounds. The maximum depth of excavation for the peak ring in Gale is ~17 km [10].

The peak rings do have in common is their style of erosion into massive unlayered conical hills, of which some areas erode into boulders only seen in peak rings. Lowell (Fig 4d/Fig 5b) is the ideal representation of a well formed peak ring crater. A thin mass of sediment is corralled inside the northeast quadrant of the peak ring. This may be analogous to the much thicker mass of sediments that forms the LM of Mt. Sharp.

An arc of low hills southeast of Mt. Sharp may be the eroded remains of a peak ring in Gale crater (Fig 5c). Holden Crater has conical hills (Fig 5d) analogous to Gale’s peak ring hills.

**Conclusions:** The present findings support the conclusions of previous studies that fluvial deposition and erosion, followed by aeolian deposition characterize the history of Mt. Sharp [1-4].

Nicholson is morphologically the closest analog to Gale. Studies focusing on morphology and mineral composition should be done to further compare Nicholson with Gale.

The presence of a peak ring within Gale is plausible based on the analyses of peak rings in other craters. Peak rings erode to very distinct, unlayered conical hills, especially in craters smaller than 200 km in diameter. Peak rings may allow protection for sediments within the ring, resulting in the formation of a mound high in topography in relation to the eroded crater floor [4].

**References:**