RARE EARTH ELEMENT PARTITIONING IN LUNAR MINERALS: AN EXPERIMENTAL STUDY
E. C. McIntosh1, J.F. Rapp2 and D. S. Draper3, 1University of South Carolina, Columbia, SC, 29208, 2Jacobs, NASA Johnson Space Center, Houston, TX 77058, 3NASA Johnson Space Center, Houston, TX 77058.

Introduction: The partitioning behavior of rare earth elements (REE) between minerals and melts is widely used to interpret the petrogenesis and geological context of terrestrial and extra-terrestrial samples. REE are important tools for modelling the evolution of the lunar interior. The ubiquitous negative Eu anomaly in lunar basalts is one of the main lines of evidence to support the lunar magma ocean (LMO) hypothesis, by which the plagioclase-rich lunar highlands were formed as a flotation crust during differentiation of a global-scale magma ocean [1]. The separation of plagioclase from the mafic cumulates is thought to be the source of the Eu depletion, as Eu is very compatible in plagioclase. Lunar basalts and volcanic glasses are commonly depleted in light REEs (LREE), and more enriched in heavy REEs (HREE) [1]. However, there is very little experimental data available on REE partitioning between lunar minerals and melts. In order to interpret the source of these distinctive REE patterns, and to model lunar petrogenetic processes, REE partition coefficients (D) between lunar minerals and melts are needed at conditions relevant to lunar processes. New data on DREE for plagioclase [2], and pyroxenes [3,4] are now available, but there is limited available data for olivine/melt DREE [5-12], particularly at pressures higher than 1 bar, and in Fe-rich and reduced compositions – all conditions relevant to the lunar mantle.

Based on terrestrial data, REE are highly incompatible in olivine (i.e. D << 1), however olivine is the predominant mineral in the lunar interior, so it is important to understand whether it is capable of storing even small amounts of REE, and how the REEs might be fractionated, in order to understand the trace element budget of the lunar interior.

This abstract presents results from high-pressure and temperature experiments investigating REE partitioning between olivine and melt in a composition relevant to lunar magmatism.

Methods: The starting composition for this study was based on A14b, a composition used by [13] based on lunar picritic glass sample 14307. Currently no lunar samples have been identified as mantle rocks, but A14B is based on the most ultramafic green glass sample, and has therefore likely undergone the least modification from its initial primary mantle melt. The starting composition was synthesized from mechanical mixtures of anhydrous oxides and carbonates, with 200ppm of each trace element added using ICP-MS standard solutions. The mixture was then melted at 1600°C and re-ground, before conditioning at 1 log unit below the iron-wüstite (IW-1) buffer. This additional melting step ensures complete homogenization of the trace elements, and removes all water from the composition.

Experiments from 0.5 GPa to 1.5 GPa were performed in a 13mm QuickPress piston cylinder using BaCO3 as a pressure medium, graphite sample capsules and heaters. Experiments at 3.5 GPa to 4.0 GPa were performed in a multi-anvil press using graphite sample capsules and rhenium heaters. All experiments were quenched isobarically.

Experimental run products were analyzed for major and minor elements using the JEOL 8530F electron microprobe at NASA Johnson Space Center, and for trace elements, using Laser Ablation Inductively Coupled Mass Spectrometry (LA-ICP-MS) at the University of Houston using USGS glasses BHVO and BIR as standards.

Results: 16 successful experiments are discussed here, although not all had a grain size large enough for LA-ICP-MS analysis. At 0.5 – 1.5 GPa, olivine is the liquidus phase, with orthopyroxene (opx) co-crystallizing with olivine at temperatures below 1550°C. The liquidus at 0.5 GPa is approximately 1600°C, rising to 1650 °C by 1.8 GPa. At 3.5 GPa, however, it appears to be depressed, although there is some uncertainty in the temperatures. Thermal modeling of the experimental assembly shows that the hotspot in the capsule could be more than 100°C hotter than the temperature measured by the thermocouple at the top of the capsule.

The forsterite content of experimental olivines ranged from Fo81 in near-liquidus experiments down to Fo76 in equilibrium with opx. Orthopyroxene compositions show a similar trend, ranging from enstatite-rich (En83Fs17Wo1) in near liquidus experiments at high pressure to more pigeonitic (En68Fs19Wo12) at low pressure co-crystallizing with olivine.

Runs CM-8, CM-10, CM-12, and CM-16 contained crystals that were suitable to analyze using LA-ICP-MS to determine partition coefficients.

Olivine/melt partition coefficients increase from the LREE (La – Gd) to HREE (Y – Lu) with Lu being the most compatible (Fig. 1), although all REE can still be considered very incompatible in olivine. The concentration of LREE such as La and Ce is so low in olivine that it is close to the detection limit of the instrument. Therefore the uncertainty on these measurements is much larger, and the resultant partition coeffi-
which could more readily take part in substitutions low, and at the low flotation crust. However, this that is no modelling suggests during LMO crystallization (e.g. anomalies in lunar basals have a negative Eu anomaly (Fig. 2). Negative Eu that olivine crystallizing from a chondritic source cumulates. en-riched i site to that observed in lunar basalts, which are e relative to HREE in the melt. This distribution is opp as Yb are slightly more compatible in olivine from HREE. Sm/Yb increases from chondritic 0.905 to gest that olivine has the potential to fractionate LREE composition [15], our new partition coefficients su (after [7]) reported tallization using the Rayl lunar interior at increasing degrees of fractional cry scheming, as REE abundances in olivine are very major of Eu should be in the reduced Eu$^{2+}$ form, with Fe$^{2+}$ and Mg$^{2+}$ in the olivine structure over other REE$^{3+}$ cations. The cause of this apparent small negative Eu anomaly in our experimental data is not yet clear, but if it is a real effect it could have wider implications for the composition of the LMO cumulates and lunar basalt sources.

Conclusions: Using the new experimentally determined olivine/melt $D_{\text{REE}}$ we can begin to understand the trace element composition of the lunar mantle, and the petrogenesis of lunar magmas. Our new $D_{\text{REE}}$ allow LREE to be enriched in the residual melt during olivine crystallization, and crystallizing olivine to have inherent negative Eu anomalies, adding further to the debate on LMO crystallization processes.

Figure 1: REE partition coefficients for experimentally grown olivine in lunar composition 14307 compared to literature data [8-13]

Figure 2: Chondrite normalized REE abundances in modeled olivine (blue) and opx (orange). Experimental $D_{\text{REE}}$ produce olivine with a marked negative Eu anomaly.