

# Shock Radiation Tests for Saturn and Uranus Entry Probes

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This paper describes a test series in the Electric Arc Shock Tube at NASA Ames Research Center with the objective of quantifying shock-layer radiative heating magnitudes for future probe entries into Saturn and Uranus atmospheres. Normal shock waves are measured in Hydrogen/Helium mixtures (89:11 by mole) at freestream pressures between 13-66 Pa (0.1-0.5 Torr) and velocities from 20-30 km/s. No shock layer radiation is detected below 25 km/s, a finding consistent with predictions for Uranus entries. Between 25-30 km/s, radiance is quantified from the Vacuum Ultraviolet through Near Infrared, with focus on the Lyman- $\alpha$  and Balmer series lines of Hydrogen. Shock profiles are analyzed for electron number density and electronic state distribution. The shocks do not equilibrate over several cm, and distributions are demonstrated to be non-Boltzmann. Radiation data are compared to simulations of Decadal survey entries for Saturn and shown to be significantly lower than predicted with the Boltzmann radiation model.

## Nomenclature

$A_{ij}$	= Einstein Coefficient for transition $i \rightarrow j$ ( $s^{-1}$ )
$c$	= Speed of Light ( $2.998 \times 10^8$ m/s)
$B(\lambda, T)$	= Blackbody (Planck) Function ( $W/cm^2\text{-sr-}\mu\text{m}$ )
$D$	= Shock Tube Diameter (m)
$g_i$	= degeneracy of state $i$
$h$	= Planck's Constant ( $6.626 \times 10^{-34}$ J-s)
$k$	= Boltzmann Constant ( $1.381 \times 10^{-23}$ J/K)
$I$	= Spectral Radiance ( $W/cm^2\text{-sr-}\mu\text{m}$ )
$I_{ij}$	= Spectral Radiance from transition $i \rightarrow j$ ( $W/cm^2\text{-sr-}\mu\text{m}$ )
$L(\lambda)$	= Lorentzian Function
$n_i$	= number density of state $i$ ( $cm^{-3}$ )
$T$	= Temperature (K)
$\Delta\lambda_{1/2}$	= Full Width at Half Maximum (FWHM, nm)
$\lambda$	= Wavelength (nm)
$\lambda_0$	= Center Wavelength (nm)
$\lambda_{ij}$	= Wavelength from transition $i \rightarrow j$ (nm)
$\lambda_{\text{shift}}$	= Stark Shift (nm)
$\tau$	= Optical Depth

## I. Introduction

THE 2013-2022 Decadal Survey for planetary exploration has identified probe missions to Uranus and Saturn as high priorities.[1] Recently, using the state-of-the-art simulation tools at NASA Ames Research Center, uncertainties in predictive models for these entries were evaluated through Monte Carlo analysis.[2] For both cases, the uncertainty in convective heating was shown to be relatively low. Margin would likely be driven by uncertainty in trajectory and atmospheric composition rather than the aerothermal tools. For Uranus, the entry velocities and shock-layer temperatures are low enough that radiative heating is negligible. For Saturn entries however, depending on trajectory, the radiative heating could make up a significant fraction of total heating. Furthermore, the uncertainty in radiative heating from the CFD model alone (i.e. not including radiative modeling uncertainty) was substantial, with the range of Monte Carlo solutions encompassing from 0.5-4 $\times$  of the nominal value.

A shock tube test in the Electric Arc Shock Tube (EAST) facility was therefore constructed to examine this radiative heating uncertainty, with the results reported in this work. H<sub>2</sub>/He mixtures have been studied in the

Electric Arc Shock Tube previously in the 1970s for Galileo (Jupiter) entry,[3] in 1998-99 in support of the Giant Planet Facility, and in 2004-05 for Neptune Aerocapture.[4] In these latter tests, electron densities within the shock were found to be much larger than expected from equilibrium considerations. The conjecture was that the radiation from the arc driver can heat the shock to a temperature well above that predicted by equilibrium shock wave equations.[5] As a result of this, much of the data previously collected in H<sub>2</sub>/He mixtures in the EAST facility has been discarded. Approaches to mitigate the driver heating phenomenon were successfully employed and are described in this paper. The paper also describes the radiative measurements at flight relevant conditions, analysis of spectral data, and recommendations for improving radiative heating models for Giant Planet Entries.

## II. Experimental

The EAST facility has been described in detail in previous work,[6] and is summarized here. The facility consists of a driver, a driven tube (with optional buffer section), and a dump tank. The EAST electric arc driver is conical with a 10.16 cm exit diameter and 1292 cm<sup>3</sup> volume. The cross-wire grounding configuration discussed in [7] was used for conditions reported here. The energy for the arc discharge is supplied by a 1.2 MJ, 40 kV capacitor bank with a total capacitance of 1530  $\mu$ F. The buffer is a 3 m section of 10.16 cm tube that is located between, and isolated from, the driven and driver gas. Helium gas was chosen for both the driver and the buffer so as to avoid a resonance excitation of the Hydrogen Lyman- $\alpha$  line. The driver pressure was fixed at 0.68 MPa (100 psi), while several tests were performed to optimize the buffer section and diaphragms. Ultimately, a buffer pressure of 0.5 Torr was selected. The buffer/driven diaphragm was a 3  $\mu$ m polyester sheet with thin aluminum coating. The coating was intended to reflect driver radiation, thus preventing radiative heating of the test gas, while the thinness of the diaphragm was necessary to achieve high velocity. The driver/buffer diaphragm was selected as 0.25 mm (0.010") Mylar, which gave a slight velocity increase over the more typical 0.30 mm (0.012") Aluminum diaphragm. With these settings, the discharge voltages employed were between 21-35 kV, depending on driven gas pressure and velocity obtained.

The driven tube diameter is 10.16 cm and the total length of the driven tube, from the main diaphragm to the dump tank, (including the buffer) is 12.5 m. Two turbomolecular pumping stations at either end of the driven tube evacuate the tube to base pressures less than 10<sup>-5</sup> Torr prior to charging with the test gas. The test gas was an 89:11 mixture of H<sub>2</sub>:He (by mole), which is the approximate composition of Saturn's atmosphere (generally accepted value is from 9.9-13.8% He.[8] Uranus is believed to have between 11.9-18.5% He by mole.[9]) The test gas and vacuum quality is monitored using a residual gas analyzer from Stanford Research Systems prior to each shot. The optical test section is located approximately 6.75 m downstream of the main diaphragm. A third diaphragm at the end of the driven tube separates the driven tube from the dump tube and tank.

Along the tube are several round diagnostic ports for time of arrival instrumentation. The ports accommodate flat windows for optical instrumentation or pressure sensors. The shock arrival times are monitored by up to 18 high speed pressure transducers. At a limited number of locations, additional transducers are installed for redundancy and/or planarity measurement. The majority of pressure sensors are model 132A35 piezoelectric shock sensors from PCB electronics. Three of the stations employ the model 113A21 transducer which has a lower temporal sensitivity but provides more accurate measurements of shock pressure. The time of arrivals in the test section is fit with a second order polynomial, the derivative of which gives shock velocity. For these high velocities, it is found that small errors in the time of arrival detection may have significant impact on the velocity measurement. Therefore the reported velocities may have uncertainty up to  $\pm 0.4$  km/s. To monitor radiation from the driver, a 90 degree mirror mount was inserted in the center of the shock tube, downstream from the test section. The mirror mount was used to redirect driver (and shock) radiation through the wall window mount, where a PMT sensitive to 190-320 nm is mounted. This configuration duplicates that employed in Ref. [5] and its values are compared directly to those obtained in that test series.

Radiation measurement in the EAST is made via 4 spectrometers in two vacuum optical boxes coupled to opposite sides of the test section. The spectrometers image the same axial location in the shock tube through a slot window in the test section. For about half of the tests reported here, fused silica windows were employed, which are transparent between 165-4200 nm. In the other tests, magnesium fluoride (MgF<sub>2</sub>) windows were employed in the vacuum ultraviolet, which are transparent to below 120 nm and could therefore measure the resonant atomic Lyman- $\alpha$  line and Lyman band radiation from H<sub>2</sub>. Spectrally and spatially resolved emission from the shocked gas is monitored through these ports. The vacuum box is a rectangular aluminum box which contains all the imaging optics for the EAST spectroscopy implementation. The vacuum box couples to the spectrometers through window ports for ultraviolet (UV), visible (VIS) and near infrared (NIR) spectrometers and through a sealed port for vacuum ultraviolet (VUV) spectrometers. Each vacuum box contains two optical paths, which are each created by a series of

six mirrors. The two optical paths enter the test section through the slot windows at slightly different angles, and pass through a focal point at the center axis of the tube. The first two mirrors rotate the image of the slot to a vertical orientation. The remaining mirrors direct the image and focus it onto the spectrometer slits. Each set of mirrors is chosen to optimize particular wavelength ranges depending on the spectrometer they are associated with.

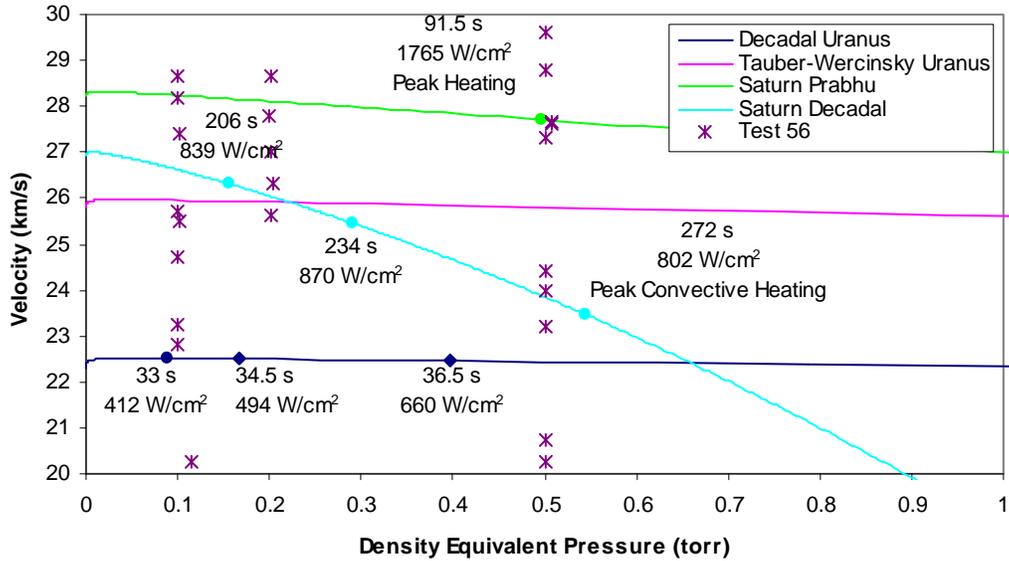
The four spectrographs each consist of a CCD array coupled to the focal plane of an imaging spectrometer. They are denoted VUV, UV (or Blue), visible (or Red) and near infrared (NIR). The VUV and Blue spectrometer are coupled to the east side of the shock tube, where the optics box is maintained under high vacuum. The Red and NIR spectrometers are coupled to the west side, but were not maintained under vacuum for this test series. The four spectrometers employed are from PI-Acton and include three 0.5 m spectrometers (SP2500i) and one 0.39 m vacuum spectrometer (VM504). Each spectrometer has a dedicated computer that operates both the camera and spectrometer simultaneously. Intensity calibration is accomplished using either an integrating sphere ( $> 300$  nm), producing highly uniform output over a 10.2 cm (4-inch) diameter, or a D<sub>2</sub> lamp source ( $< 300$  nm). The D<sub>2</sub> lamp source is purged with Argon to remove oxygen absorption and mounted to a translating stage to obtain the spectra at different spatial locations. The D<sub>2</sub> lamp source is always referenced against the integrating sphere to convert relative intensities to absolute ones. The measurement ranges and features observed for each spectrometer is given in Table I.

**Table I. Spectrometer Settings and Features Observed**

Spectrometer (Window)	Spectral Range (nm)	Species	Spectral Feature	Transition	Wavelength (nm)
VUV (MgF <sub>2</sub> )	117-173	H	Lyman- $\alpha$	$n=2 \rightarrow 1$	121.57
		H <sub>2</sub>	Lyman Band	$B^1\Sigma-X^1\Sigma$	100-160
VUV (SiO <sub>2</sub> )	164-218	None			
UV	323-497	H	Balmer- $\beta$	$n=4 \rightarrow 2$	486.14
			Balmer- $\gamma$	$n=5 \rightarrow 2$	434.05
			Balmer- $\delta$	$n=6 \rightarrow 2$	410.17
			Balmer- $\epsilon$	$n=7 \rightarrow 2$	397.00
			Balmer- $\zeta$	$n=8 \rightarrow 2$	388.91
			Balmer Continuum	$n=\infty \rightarrow 2$	$< 364.6$
Visible	653-659	H	Balmer- $\alpha$	$n=3 \rightarrow 2$	656.28
NIR	1022-1652	H	Paschen- $\beta$	$n=5 \rightarrow 3$	1281.81
			Paschen- $\gamma$	$n=6 \rightarrow 3$	1093.82

### III. Results

The test conditions obtained in this work are shown in terms of velocity and freestream (laboratory) pressure, and compared to different proposed Uranus/Saturn trajectories in Figure 1. The shock tube is assumed to have similarity to stagnation line profiles for matched freestream velocity and density. For the trajectories, the atmospheric density has been converted to an equivalent pressure at laboratory conditions. A total of 25 tests were performed at pressures of 0.1, 0.2 and 0.5 Torr, three of which produced unusable data. Of the remaining 22 tests, 6 had test time but no observable signal. These tests were all less than 25 km/s, confirming predictions that radiative heating would be negligible at these velocities (which encompass most Uranus trajectories.) From the 16 tests with signal, test times were short or unclear in 4 of them, leaving 12 tests with good data quality. Of these 12 tests, 8 included measurement of the Lyman- $\alpha$  line in the VUV.



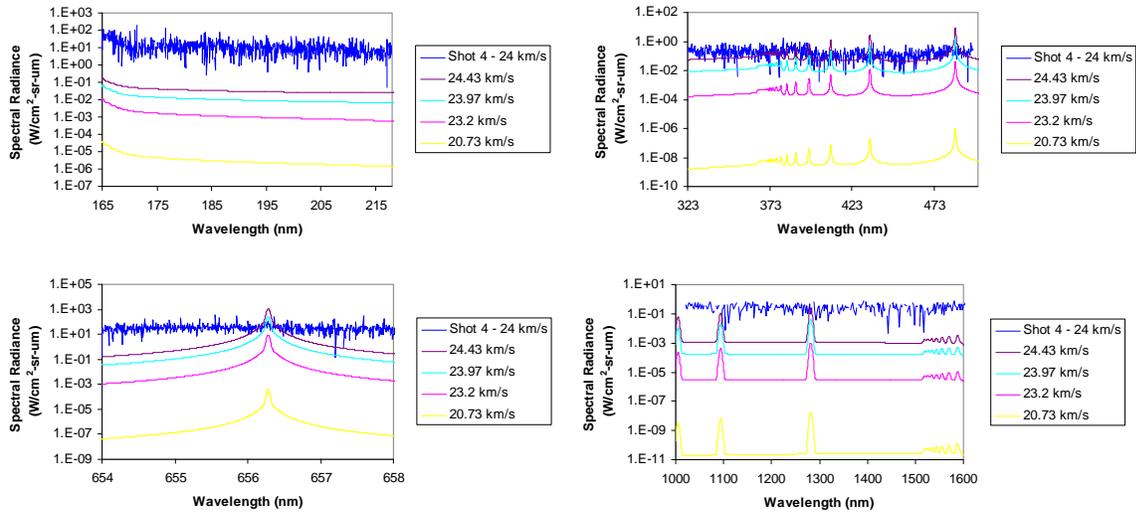
**Figure 1. Plots of candidate trajectories for Uranus and Saturn entry overlaid with test conditions performed in this work.**

Somewhat surprisingly, no evidence of driver gas preheating was seen during any of the optimization tests. Radiance was monitored using the 90 degree mirror and compared to those obtained in the test from Ref. [5]. For all tests in this work, the radiance was 15-80× lower than in the previous work. This is attributed to the blocking affect of the aluminized polyester buffer diaphragm. The shock pressures were always within 20% of the equilibrium pressure, while in the previous test campaigns they were up to a factor of 3 higher, and the widths of Hydrogen Balmer series lines agreed with equilibrium values to within 20%, whereas before they were an order of magnitude larger. As will be shown below, almost all of the tests were below equilibrium predictions, suggesting this phenomenon to have been effectively eliminated.

The data in the following is presented in order of ascending velocity. Only selected data sets are shown in this work, but the full set of data can be obtained by contacting the lead author.

#### A. 20-25 km/s

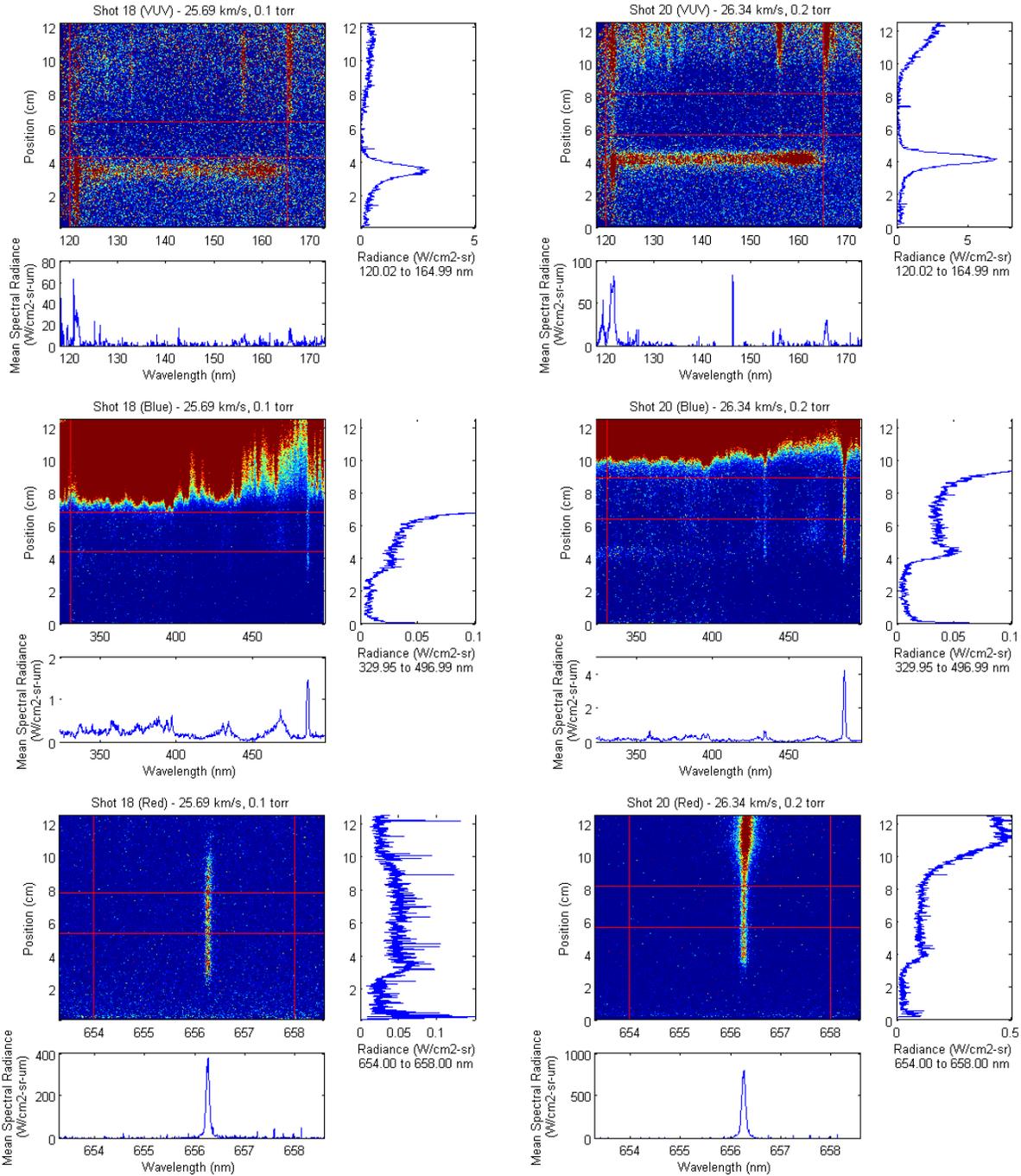
Six shots with valid test time, one at 0.1 Torr and five at 0.5 Torr, were obtained in the range of 20-25 km/s. None of these displayed any measurable radiation, however the cameras did pick up radiation from the contamination/driver gas. Because of this characteristic, it was initially thought there was no test time at all. By interpolating the camera trigger pulse onto the shock time of arrival curve, it is possible to estimate the location of the shock at the time the camera triggered. The shock would then continue moving during the camera exposure time (generally 0.25 μs for VUV/Blue/Red cameras and 0.98 μs for IR.) Since the camera acquisition signals are monitored on the low speed acquisition card at a rate of 5 MHz, the accuracy of this determination is no better than 0.2 μs. To verify whether these results were legitimate, the equilibrium radiance level has been computed for several of the shock speeds between 20-25 km/s and pressure of 0.5 Torr, and compared to the spectrometer noise floor in Figure 2. For most of these cases, the expected equilibrium level is orders of magnitude less than the noise limit of the spectrometer. Minor exceptions are apparent above 24 km/s, where the equilibrium prediction of the Balmer series lines at 656, 486, 434, and 410 nm suggests they may be within the detection limits. There may be several valid reasons why they may not be seen, as discussed below. This includes the slow equilibration rate/induction time in the shock and NEQAIR's Stark broadening calculation.



**Figure 2. Predicted equilibrium radiances from 20-25 km/s compared to the spectrometer noise floor. Each simulation corresponds to an actual shot condition in the test series. The experimental data from only one of these tests are shown, as the other tests are essentially identical.**

### B. 25-27 km/s

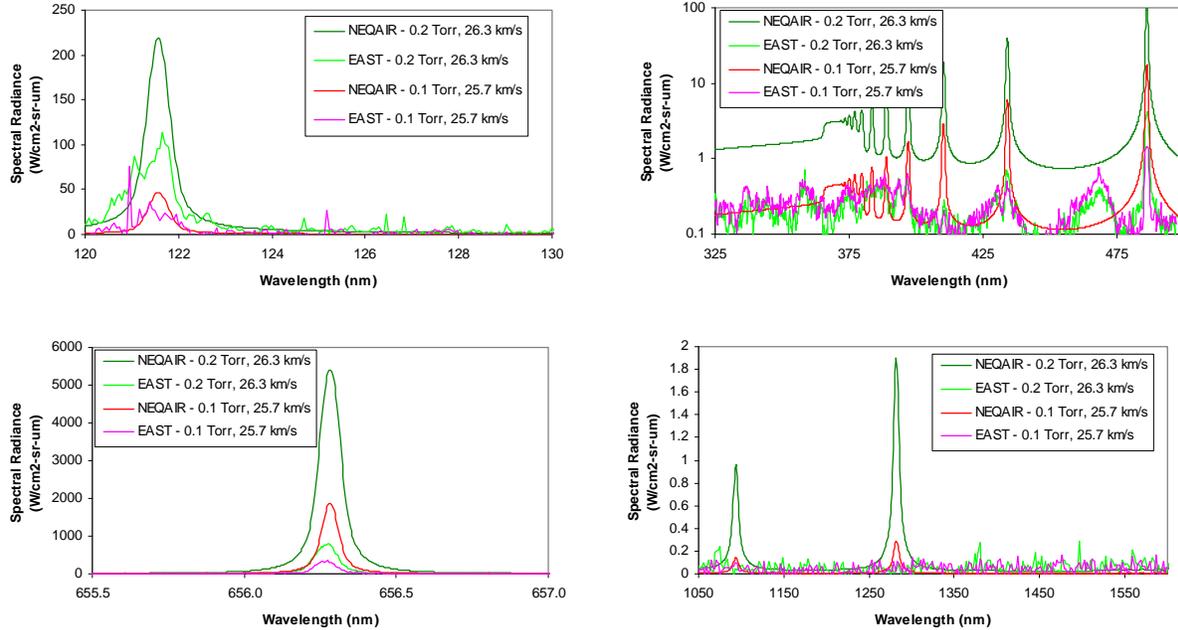
There were four shots obtained between 25-27 km/s, including two at 0.1 Torr and two at 0.2 Torr. A representative data set is given from each of these pressures in Figure 3. In all these cases, radiation in the near infrared was below the detection limit, so is not included. The presence of the shock front is seen with a broad non-equilibrium radiance in the VUV at all conditions, near  $y = 4$  cm. Analysis of the camera trigger and time of arrival shows that this feature actually precedes the shock front. The broad emission is attributed to the Lyman band (B-X) of molecular hydrogen, which is presumably quenched when the shock density increases. The Lyman- $\alpha$  line at 121.6 nm is also seen and precedes the shock front by as much as 1-2 cm, with the distance being greater at lower pressure. A pre-shock excitation would be consistent with absorbed radiation in the freestream, as suggested by Park[10], and will be examined in more detail in the full paper. Evidence of molecular emission is also faintly seen at the shock front on the Blue camera. More clearly evidence of the shock front is ascertained by the appearance of the Balmer- $\alpha$ ,  $\beta$ , and  $\gamma$  lines on the Red and Blue cameras. Following the non-equilibrium peak, the radiance settles into a steady level for several cm. The contamination is observed at about  $y = 8$  and 10 cm at 0.1 and 0.2 Torr, respectively. This corresponds to test times of approximately 4-6 cm.



**Figure 3. Test data at (left) 25.7 km/s, 0.1 Torr and (right) 26.3 km/s at 0.2 Torr. The rows represent the spectral ranges of Deep VUV, UV and Visible.**

The radiation intensities from these tests (during the valid test time) are compared against the equilibrium radiance predictions for these cases in Figure 4. In the VUV, the only interesting feature expected in equilibrium is the 121.6 nm Lyman- $\alpha$  line, therefore only the region from 120-130 nm is shown. Both lines are somewhat below equilibrium levels, though the 25.7 km/s condition may be considered to be within the noise of the data. The Blue region from 325-500 nm is shown on a logarithmic scale, as otherwise most features, other than the Balmer- $\beta$  line (486 nm), would be difficult to observe. These features are all significantly less than equilibrium radiation levels, again suggesting that the shock has not equilibrated over the few cm of test distance. Weak molecular band features near 470, 430 and 390 nm are noted. These may originate from water impurities in the tube (OH visible B-X band)

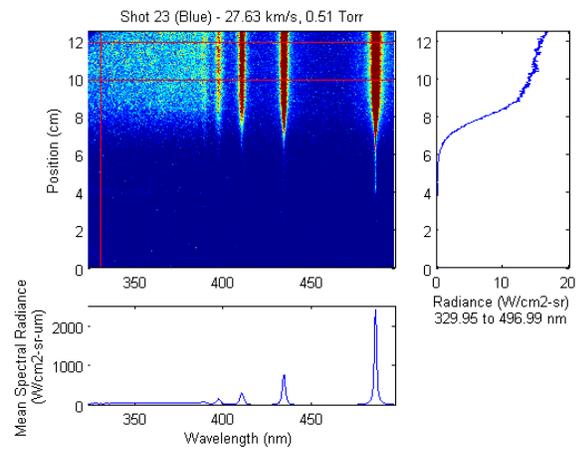
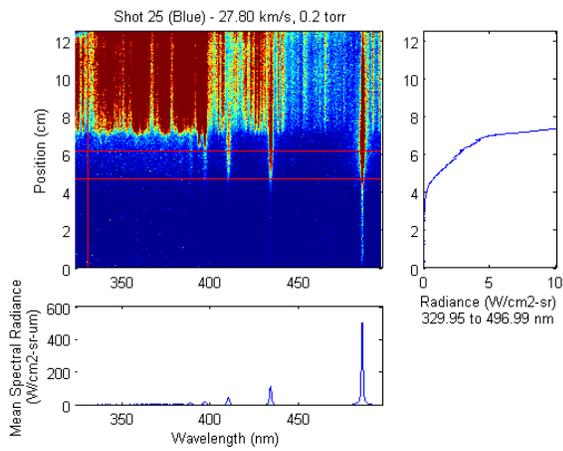
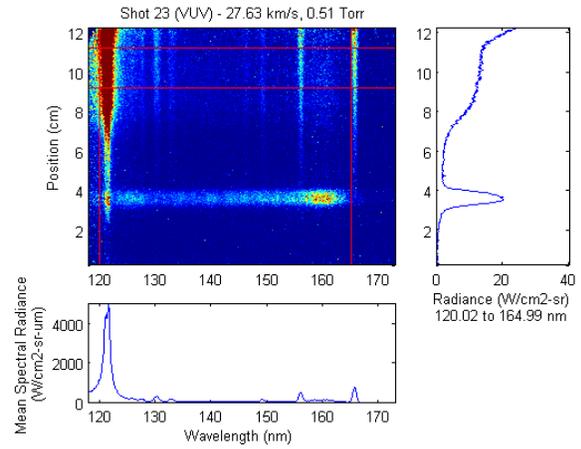
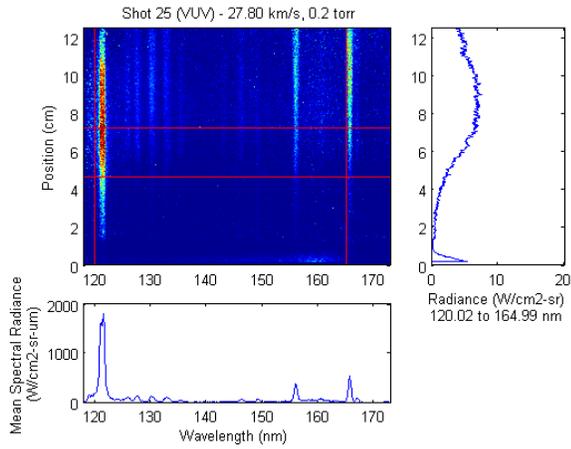
which is not included in the prediction. The Red camera measurement of the Balmer- $\alpha$  line similarly shows significantly lower radiance than expected at equilibrium, with the peaks being 6-7 $\times$  lower than equilibrium radiance values. There is no measurable infrared signal in these tests, even though the equilibrium levels are above the camera noise floor. The summary of all this data indicates that the steady region behind the shock front is not an equilibrium condition. It also confirms that significant driver gas radiative heating is not present in these tests.

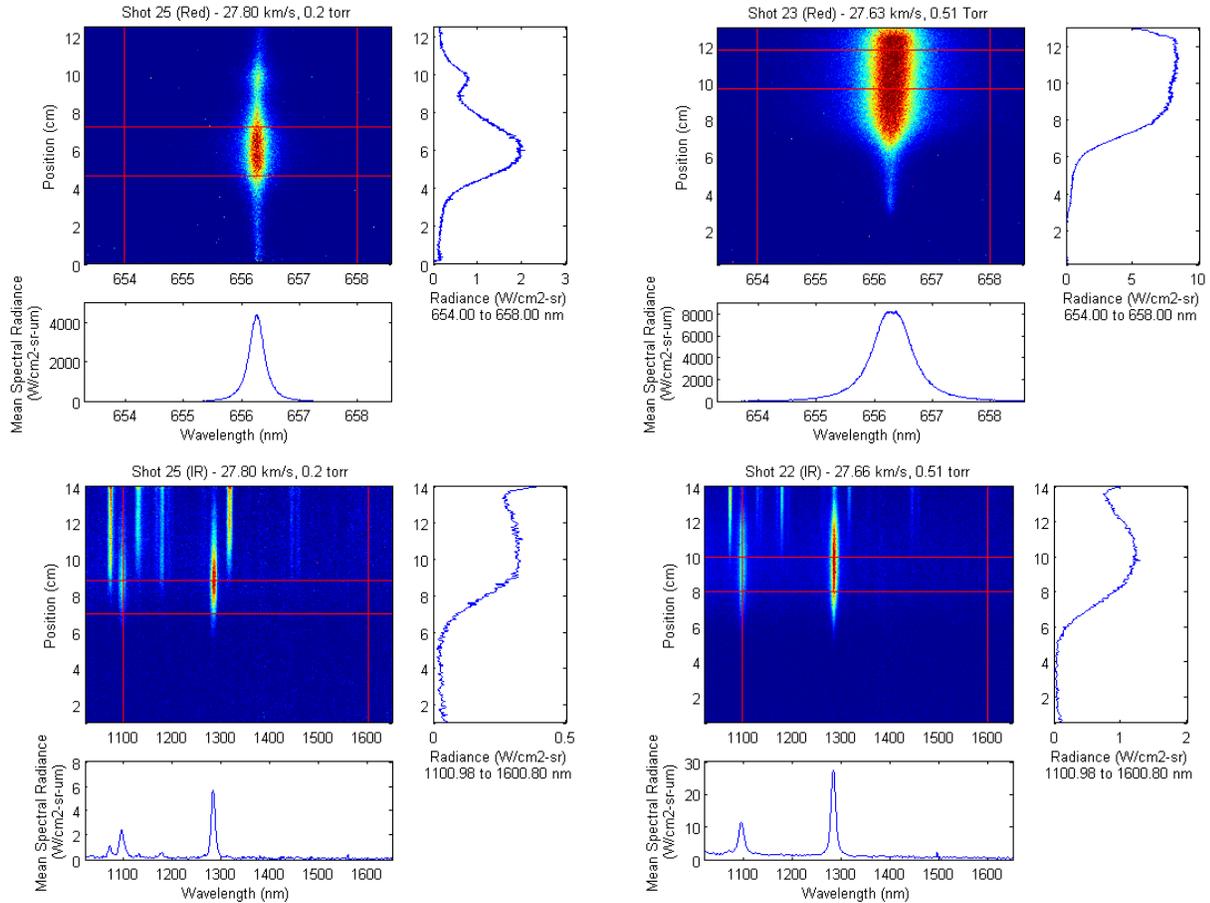


**Figure 4. Comparison of EAST spectra to equilibrium radiance levels for velocities of 25.7-26.3 km/s.**

### C. 27-28 km/s

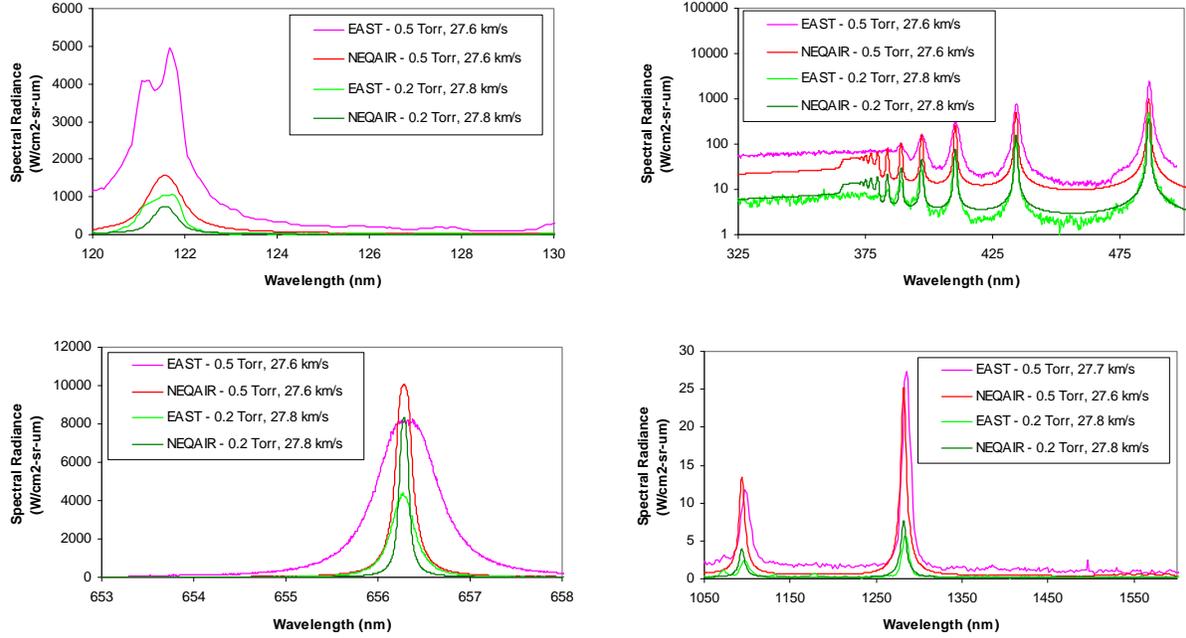
Six shots were obtained between 27-28 km/s, including three at 0.5 Torr, two at 0.2 Torr and one at 0.1 Torr. One shot from each of 0.2 Torr and 0.5 Torr conditions are shown in Figure 5. In the 0.2 Torr data (Shot 25), on all 4 cameras, the trigger was such that the shock front is near, or off of, the bottom edge of the CCD array. At 0.5 Torr (Shot 23), the shock front is near  $y=4$  cm. The Lyman- $\alpha$  line again precedes the molecular emission but not by as much as at the lower pressure conditions. The contact front for Shot 25 occurs near  $y=8$  cm and no contact front is observed in Shot 23. Due to a triggering error in the IR camera during shot 23, no IR data was obtained and the near infrared data from a similar condition (Shot 22) is shown instead. A common feature to all these tests is a delayed radiative transition. As at lower velocity, the shock front is observed with molecular emission features in the VUV which then decay to a low value. At around 3-4 cm behind the shock front, the radiation suddenly increases. This trend is observed on all three of the VUV, Blue and Red cameras. The radiation transition is accompanied by an increase in the atomic linewidth. The Balmer- $\alpha$  line at 0.5 Torr displays a flattened top which suggests it is optically thick. In the infrared, no radiation is seen at the shock front. Instead the Paschen- $\beta$  and  $\gamma$  lines appear behind the shock front at the same distance from where the radiation spikes on the other cameras. This induction time is attributed to a relatively slow ionization rate in the shock layer, and will be discussed in the next section. This transition was consistently observed in this testing regime, though the 0.1 Torr test did not have sufficient test time for this to occur.





**Figure 5. Data collected at (left) between 27.8 km/s and 0.2 Torr and (right) 27.6 km/s and 0.5 Torr. The rows are ordered by wavelength.**

The radiance following this transition is compared against equilibrium NEQAIR calculations in Figure 6. Here, the Lyman- $\alpha$  line exceeds equilibrium for both cases. The Blue camera data is of the same order as the equilibrium computation. At 0.2 Torr, the Balmer- $\beta$  line peak is somewhat above equilibrium and the  $\gamma$  line is somewhat below. As lines move to higher energy (lower wavelength), they become increasingly less than equilibrium. At 0.5 Torr, the peak widths are clearly wider in the experiment. Also notable is the absence of any intensity change at about 366 nm, which suggests that the Balmer continuum is not predicted correctly. The Balmer- $\alpha$  line, as measured by the Red camera, is significantly wider in experiment than as given by NEQAIR. The disagreement in lineshapes is not unexpected as the current version of NEQAIR does not rigorously model the line broadening of hydrogen. In terms of the total radiative contribution (i.e. area under the curve), the experimental Balmer- $\alpha$  line data is in excess of equilibrium by about 35% at 0.2 Torr, but 8 $\times$  larger at 0.5 Torr. The comparison in the infrared at 0.5 Torr is fairly close to equilibrium, while at 0.2 Torr, the Paschen series radiance in the infrared is slightly less than equilibrium.



**Figure 6. Comparison of 27.6-27.8 km/s to equilibrium radiance predictions.**

The 0.2 Torr condition has shown certain lines to exceed equilibrium and others to be less than equilibrium. This trend correlates with the principle quantum number ( $n$ ) of the state responsible for radiation. In excess of equilibrium are the Lyman- $\alpha$  ( $n=2$ ), Balmer- $\alpha$  ( $n=3$ ) and Balmer- $\beta$  ( $n=4$ ). Below equilibrium are the Balmer- $\gamma$  and higher ( $n \geq 5$ ) and the Paschen- $\beta$  and  $\gamma$  ( $n=5,6$ ). The simplest interpretation of this trend is that the electronic temperature is lower than the equilibrium shock temperature and that the shock is still equilibrating. An alternative, but more complex, interpretation could be that the upper states are being depleted by preferential processes, such as electron impact ionization. In this case, a state-to-state or QSS model of the hydrogen atom would be required to correctly predict the radiation. This will be examined in more detail in the discussion.

#### D. 28-30 km/s

Five shots were obtained between 28-30 km/s, including two at 0.5 Torr, one at 0.2 Torr, and two at 0.1 Torr. Except for one shot at 29.6 km/s and 0.5 Torr, these all had short test times. Unfortunately, the deep VUV setting was not measured at 29.6 km/s, and the Blue camera was saturated. Therefore, this data is not shown. The induction time becomes increasingly shorter at higher velocity and the Balmer- $\alpha$  line more broadened with an extended flat top.

## IV. Discussion and Analysis

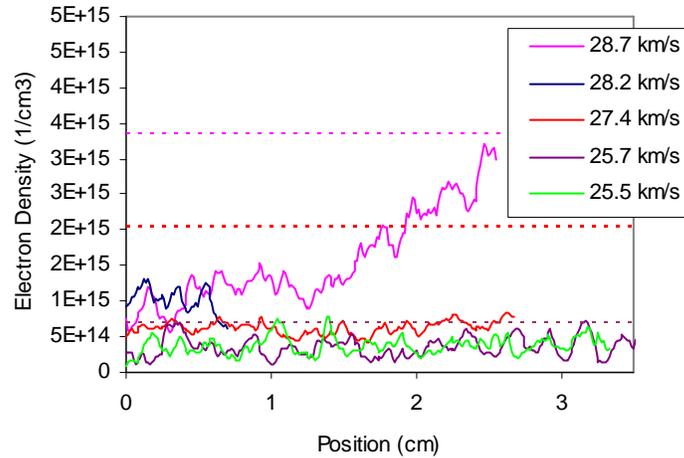
### A. Electron Density and Temperature Trends

The spectra may be analyzed for electron density values, as discussed in [11]. In this case, the Balmer- $\alpha$  line has been measured at high resolution, so is preferred for Stark analysis. However, since many of the spectra are optically thick, this may yield unrealistically high electron densities. In this case, the emission is given by:

$$I = B(\lambda, T) \left[ 1 - \exp \left( -\tau L \left( \frac{\lambda - \lambda_o - \lambda_{shift}}{\Delta\lambda_{1/2}} \right) \right) \right] \quad (1)$$

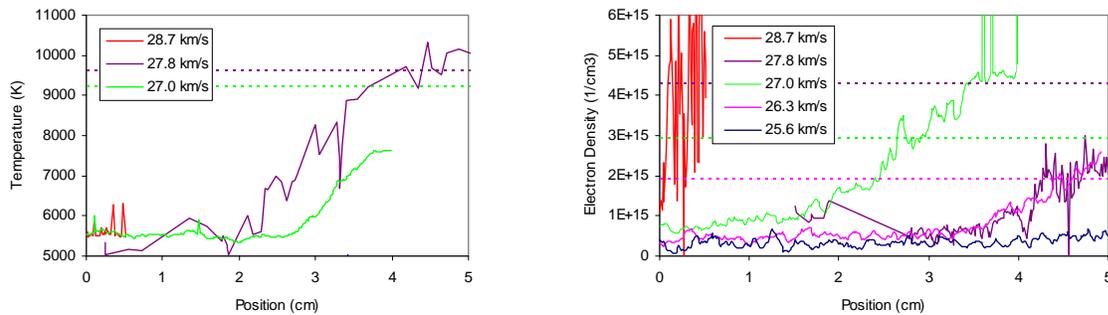
where  $B(\lambda, T)$  is the blackbody (Planck function) and  $L$  is the Lorentzian function. Fit parameters then include  $T$ ,  $\tau$ ,  $\lambda_{shift}$ , and  $\Delta\lambda_{1/2}$ . When optically thick, the fit is fairly insensitive to electron density, while the optically thin case is insensitive to temperature. Therefore only temperature or electron density can be reliably reported for a single line. The fit must include a convolution with the instrument lineshape, and the sensitivity limit of the measurement will depend on how close the linewidth is to the instrument lineshape. The high resolution lineshape used for the  $\alpha$  line

is equivalent to an electron density between  $2\text{-}4 \times 10^{14} \text{ cm}^{-3}$ , depending on the slit settings used. The  $\beta$  and  $\gamma$  lines, being measured at low resolution, had widths corresponding to of  $5.7\text{-}9.2 \times 10^{15}$  and  $4.0\text{-}6.5 \times 10^{15} \text{ cm}^{-3}$ , respectively. Depending on the quality of the signal, it may be possible to detect densities lower than these values, as the line broadening is, to first order, additive.



**Figure 7. Electron Density deduced from Balmer- $\alpha$  linewidth at pressure of 0.1 Torr. Dotted lines display equilibrium electron densities at 28.7, 27.5 and 25.7 km/s. (Densities at 28.2 and 25.5 are omitted for clarity)**

The 0.1 Torr tests are all optically thin, so electron density is obtained from the  $\alpha$  line without temperature data. The  $\beta$  and  $\gamma$  line measurements are limited by instrument sensitivity and thus not shown. The plots in Figure 7 have been shifted so that the shock front (which is known to within  $\sim 0.5$  cm) is at  $y=0$ , and the traces have been truncated at the end of the test time. The dotted lines display equilibrium electron densities at 28.7, 27.4 and 25.7 km/s. Of the five traces shown, all but the 28.7 km/s case are flat over the course of the test and are near the instrument linewidth value. Therefore, it is possible the actual electron density is lower than these curves suggest. The electron density at 28.7 km/s begins to increase noticeably approximately 1.3 cm after the shock front. From here it continues increasing until the end of the test approximately 2.5 cm behind the shock front. At this point, it is nearing the equilibrium electron density but is still lower than equilibrium. This is consistent with emission observations indicating the shock is not yet in equilibrium.

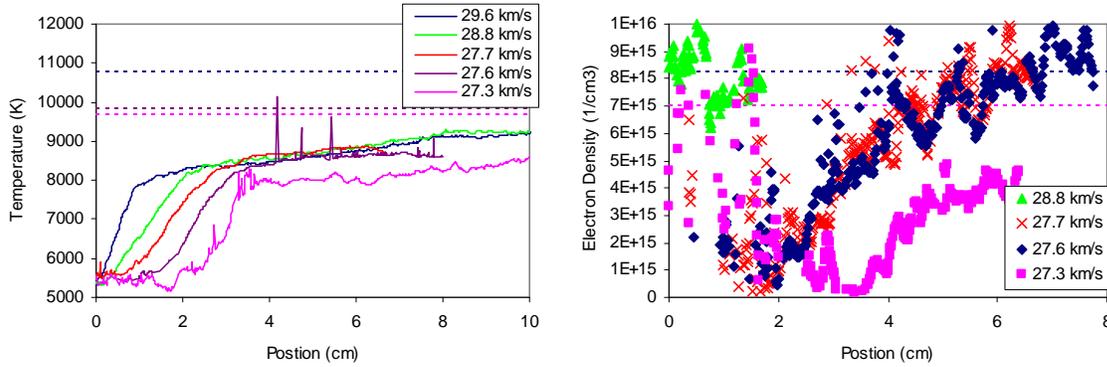


**Figure 8. Electron Density and Temperature deduced from H- $\alpha$  and/or  $\gamma$  linewidth and intensity at pressure of 0.2 Torr. Dotted lines display equilibrium electron densities and temperatures at 27.8, 27.0 and 26.3 km/s.**

At 0.2 Torr, optical thickness is only observed above 27 km/s, therefore temperature plots are given for only three of the velocities. The 28.7 km/s has very short test time, so that only a few points are given. The 27.8 km/s case does not begin increasing in electronic temperature until about 1.4 cm behind the shock front. From here it increases until it has reached equilibrium near the end of the test time. At 27.0 km/s, the shock becomes blackbody limited at approximately 3 cm behind the shock front. Before this point, it is not entirely clear from the data whether it is blackbody limited or not, though the trend is carried out to the shock front. The temperature increases

monotonically from 3-4 cm behind the shock front where the test time ends, still several thousand degrees below equilibrium. The electron densities are shown in Figure 8(b). Here, the  $\alpha$  line cannot be used above 27 km/s due to the optical thickness. At 27.8 km/s, the density from the  $\gamma$  line is shown. The electron density stays low until about 3 cm behind the shock where it begins increasing. The density is still below equilibrium at the end of the test time. At 27 km/s, the  $\alpha$  line density is shown until 3 cm from the shock front where it switches to the  $\beta$  line. The 27.0 km/s data appears inconsistent with the rest, starting to increase near 2 cm and rising above equilibrium at 3 cm. It is possible in this case that the electron density fit was invalid for the various reasons discussed above. At the two lowest velocities, the density obtained from the  $\alpha$  line is given. The 26.3 km/s data follows the same trend as at 27.8 km/s, beginning to increase at 3 cm. By the end of the test, it is at the equilibrium level. At 25.6 km/s, the electron density is near the detection limit the whole test and does not increase over the entire 5 cm.

At 0.5 Torr, all cases are blackbody limited. Temperatures are obtained from the  $\alpha$  line while electron densities are found from the  $\gamma$  lines. The  $\beta$  line density is generally 1.6-1.9x lower, which is attributed to the electron density being too near the instrument linewidth. The curves are shown in Figure 9 and have been shifted so as to place  $y=0$  at the estimated shock front location. The temperature trends appear as a series of S curves, with the delay time increasing at lower shock velocity. The temperature is found to ramp from approximately 5500 K to 8000 K fairly quickly, and continues rising slowly after that. In none of these cases does the temperature reach the equilibrium level. While most of these cases have less than 10 cm of test time, the temperature curve is well behaved, so is carried out that far for most of the plots. The electron density curves are shown only for valid test time in Figure 9(b). Electron density was not obtained at 29.6 km/s due to camera saturation. At 28.8 km/s, the electron density is measured near  $9 \times 10^{15} \text{ cm}^{-3}$  during its short test time. At 27.6-27.7 km/s, it takes about 2 cm before electron density becomes reliably detected. The density continues to rise following this, and is near equilibrium levels at 8 cm beyond the shock front. At 27.3 km/s, the electron density is first detected near 4 cm, and continues rising, but is still far from equilibrium at the end of the test almost 6 cm behind the shock front. It is noteworthy that the electron density is first observed approximately where the temperature begins to plateau. This makes sense because a high electron temperature will be required to drive ionization. As the electron density rises, however, the temperature of the shock should fall, which is contrary to the measured temperature trend. This inconsistency indicates that the electronic temperature has not equilibrated to the translational temperature. It may also be the case that the Hydrogen state population has not yet attained a Boltzmann distribution.



**Figure 9. (a) Electronic temperature and (b) electron densities estimated from Hydrogen Balmer series lines at 0.5 Torr. Equilibrium temperature/densities are shown as dotted lines at 27.3, 27.6 and 29.6 km/s.**

## B. Electronic State Distributions

To check whether the shock may be Boltzmann distributed, line intensities are analyzed for distribution information in Shots 20 and 23. There are as many as 9 lines observed in this work. The intensity under each one may be integrated and, as long as it is not optically thick, used to derive a state density:

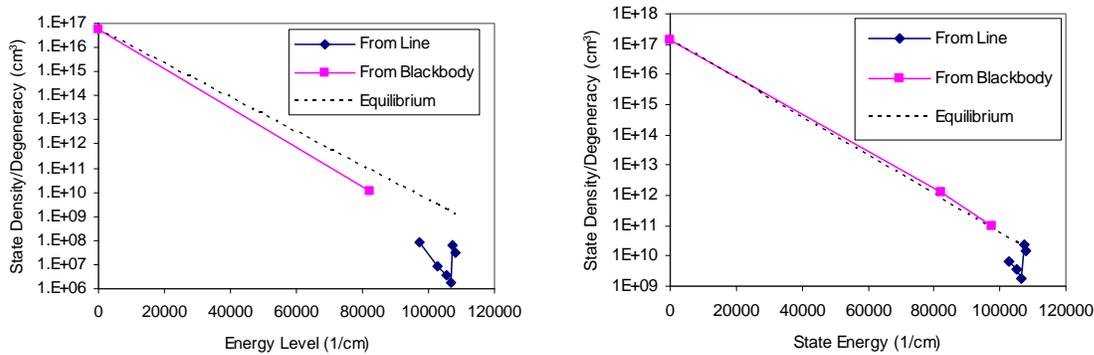
$$n_i = \frac{\int I_{ij} d\lambda}{A_{ij} D} \left( \frac{hc}{4\pi\lambda_{ij}} \right)^{-1} \quad (2)$$

When the line is optically thick, its peak intensity may derive an effective temperature, which in turn may be used to find a ratio of two lines. This is only observed for the  $n=3 \rightarrow 2$  and  $n=2 \rightarrow 1$  transitions:

$$I_{ij} = \frac{2hc^2}{\lambda_{ij}^5} \left[ \exp\left(\frac{hc}{\lambda_{ij}kT_{ij}}\right) - 1 \right]^{-1} \quad (3)$$

$$n_i = n_j \frac{g_i}{g_j} \exp\left(-\frac{hc}{\lambda_{ij}kT_{ij}}\right) \quad (4)$$

For Shot 23, the density for  $n = 4-8$  is found from equation (2) and the relative densities between  $n = 1-3$  is found from (4). If  $n_1$  is assumed to be given by the equilibrium H concentration, absolute values for  $n = 2-3$  can be obtained. For Shot 20, equation (2) is used for finding the density from  $n=3-8$ . The density of  $n=2$  is found relative to the equilibrium H atom density at  $n=1$ . The results of these evaluations, taken several cm behind the shock front, are shown in Figure 10. In both cases, a suppression of the density of high energy states is noted. In Shot 23, the density of  $n=2,3,7$  and  $8$  are all near equilibrium, while in Shot 20 all the states are well below equilibrium. However, states  $2, 7$  and  $8$  are higher relative to states  $3-6$ . The higher densities estimated for states  $7$  and  $8$  may in part be due to the accuracy in integrating over the  $\epsilon$  and  $\zeta$  lines, which have significant underlying continuum radiation. Regardless, the results are non-Boltzmann in nature, requiring a QSS model to accurately predict. This indicates that the effective temperature trends in Section A are not true electronic temperatures and only represent the effective temperature between the  $n=2$  and  $n=3$  states.

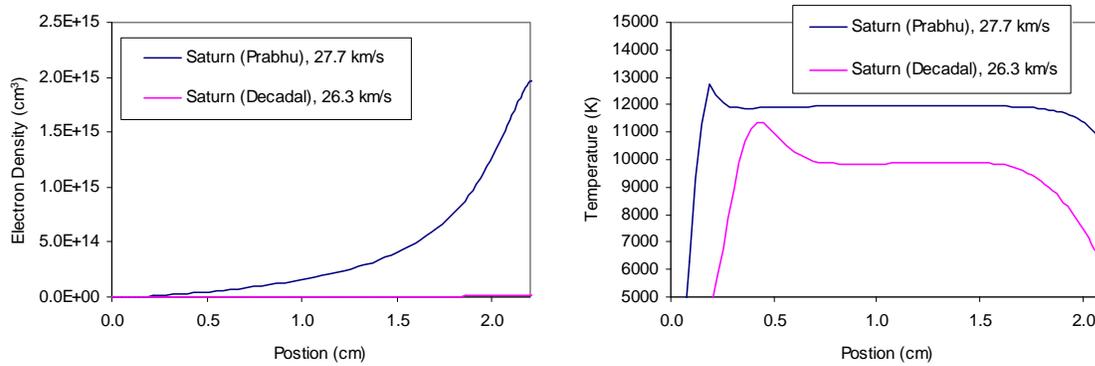


**Figure 10. Boltzmann diagrams from Shots (a) 20 (26.3 km/s, 0.2 Torr) and (b) 23 (27.6 km/s, 0.5 Torr).**

The QSS approach performs a mole balance on each electronic level, accounting for rates of impact excitation/de-excitation, radiative transitions and ionization/recombination. All three of these processes would tend to bring the electronic states toward a Boltzmann distribution, however, depending on the exact properties of the gas phase, the effective temperature they trend toward may differ. Impact excitation with heavy particles and electrons will bring the temperature toward the translational or electron temperature, respectively. Radiative transitions will de-populate the higher states until the transition becomes optically thick, at which point it will tend toward the electronic temperature. Ionization/recombination processes will drive the distribution toward a Saha distribution. Since ionization rates scale as  $n_e$ , while recombination scales as  $n_e^2$ , the two rates are not balanced at low electron density, which leads to a de-population of states. The distribution will tend toward a Saha distribution based on the lower ionization fraction. These processes all compete with each other and have different relative effects on different states, so can conceivably produce a distribution like the one above.

### C. Comparison to flight CFD

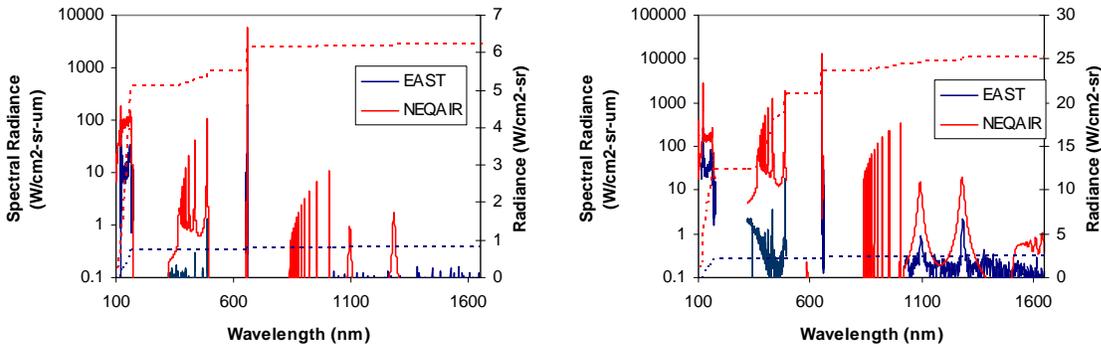
Since the predictions for Saturn probe entry generally have shock stand-off distances between 2-2.5 cm[2], the EAST experiment may be considered to replicate the entire stagnation line. Two relevant points from the uncertainty study were obtained for a velocity of 27.7 km/s and density equivalent pressure of 0.49 Torr and 26.3 km/s and 0.15 Torr. Shot 23 will be compared to the first case, while Shot 20 will be compared to the second.



**Figure 11. Stagnation line predictions corresponding to the 91.5 s point on the Prabhu trajectory and 206 s on the Decadal trajectory. (a) Electron density, (b) Vibrational temperature.**

Stagnation line predictions of electron density and vibrational temperature (assumed equal to electron/electronic temperature) are shown in Figure 11. The shock front, defined as when pressure reaches 1% of its equilibrium value, is placed at the origin. For the 27.7 km/s condition, the CFD simulation predicts the electron density exceeding  $10^{15}$  cm<sup>-3</sup> at about 2 cm beyond the shock front, which is in general agreement with the shock tube data. The 26.3 km/s obtains a maximum electron density near  $10^{13}$  cm<sup>-3</sup> after 2 cm, which would not be detected in the shock tube. The CFD temperature is predicted to rise to 12,000K within a few millimeters of the shock front at the higher pressure condition. For the lower velocity/pressure case, a steady state temperature is obtained within about 7 mm. In both cases, later in the shock, the temperature begins to decrease as ionization becomes significant. This effect is confounded, however, by the presence of the boundary layer in the flight simulation. This temperature observation conflicts with the observations of Figure 9 for reasons discussed in Section B.

As a final test of the predictive capability, the shock tube data is normalized by the shock tube width and integrated in space to obtain an estimated radiative heating profile. The integration is performed over 2.5 cm for shot 23 and 2.0 cm for shot 20, to be consistent with predicted shock stand-off distances. This integrated quantity can then be compared to the stagnation line of sight produced by NEQAIR. This comparison is shown in Figure 12. While the major features are present in the ranges measured, the magnitude is substantially overpredicted. Primarily, this results from having the non-Boltzmann distribution as discussed above. Considering the radiance from the measured regions only, which contributes ~90% of the predicted radiance, the radiance is only about 2.5 W/cm<sup>2</sup>-sr and 0.9 W/cm<sup>2</sup>-sr for shot 23/22 and 20, respectively. In terms of radiative heating, this would contribute at most 16 W/cm<sup>2</sup> and 5.4 W/cm<sup>2</sup>, compared to the nominal predicted heating of 125 W/cm<sup>2</sup> and 35 W/cm<sup>2</sup>. However, the approximations made thus far are only valid for an optically thin flow. Features from the VUV and Red cameras are at least partially optically thick, which could increase the radiative heating by several times over this rough calculation. However, the NEQAIR result will still be an overestimate of the actual heating. Implementation of a QSS model should reduce the radiative prediction significantly and bring it closer in-line with observations.



**Figure 12. Comparison of EAST estimates of stagnation line radiance and DPLR/NEQAIR predicted radiances for (a) Shallow (-8°) Saturn Entry Trajectory and (b) Steep (-15°) Saturn Entry Trajectory. The solid lines display the spectral radiance (primary axis) and the dashed lines are the accumulated radiance (secondary axis).**

## V. Conclusions and Recommendations

These tests have established conditions for future testing of Hydrogen/Helium mixtures in the EAST facility. Previous issues with gas preheating by driver radiation have been eliminated, within detection limits, through the use of light buffer diaphragms and He driver gas composition. Tests below 25.0 km/s display no radiance, as expected, due to the low temperature obtained in hydrogen. The majority of Uranus entries would fall into this category. Between 25-30 km/s, radiance is observed, but equilibrium levels are not obtained over several centimeters. For most shocks, an induction time is observed, whereby radiance levels rapidly increase from low levels several cm behind the shock front. This sudden change in radiance is shown to coincide with an increase in electron density and effective electronic temperature. The electron density trend is predicted by existing CFD tools, where the relatively slow ionization rate causes the electron density to gradually build up over the width of the shock. While the agreement is qualitative, simulating the wider range of EAST tests using larger shock stand-offs would provide further verification of this. The effective temperature trend is not captured by CFD as it is attributed to quasi-steady state phenomena that are typically implemented within the radiation code.

The radiance obtained through NEQAIR with Boltzmann state modeling is shown to overpredict flight-relevant conditions by as much as a factor of 10. The disagreement is attributed to the absence of a quasi-steady state model. Additional gaps in the NEQAIR modeling capability also identified include the H atom Stark broadening mechanism and Balmer continuum prediction. It is possible that the QSS approach would improve the continuum underprediction. The shock tube data indicates that, for a 1 m diameter probe with shock stand-off distances around 2.0-2.5 cm, radiative heating is unlikely to be an issue for most Uranus or Saturn entry trajectories. This is because the ionization rate of Hydrogen is slow enough that the electron density does not obtain the critical density necessary over this distance to excite the emitting states. The agency's current tools are expected to overpredict this heating, though the extent to which they will overpredict is not fully determined. Some consideration should also be given to the possibility that ablation products/impurities may "seed" the shock layer with easily ionized species. In this scenario, the radiative heating may be increased substantially. Entries to Neptune or Jupiter, which occur at higher velocity, are expected to have substantial radiative heating[10, 12-14] and would require revisions to the radiative model for accurate predictions.

## References

1. National Research Council, *Vision and Voyages for Planetary Science in the Decade 2013-2022*, National Academies Press, 2012.
2. Palmer, G., Prabhu, D., and Cruden, B. A., "Aeroheating Uncertainties in Uranus and Saturn Entries by the Monte Carlo Method," *Journal of Spacecraft and Rockets*, Vol. 51, No. 3, 2014, pp. 801-814.
3. Dannenberg, R., "A new look at performance capabilities of arc-driven shock tubes," *Shock tube and shock wave research; Eleventh International Symposium*. Seattle, WA, 1977.
4. Hollis, B., Wright, M., Olejniczak, J., Takashima, N., Sutton, K., and Prabhu, D., "Preliminary Convective-Radiative Heating Environments for a Neptune Aerocapture Mission," AIAA 2004-5177, 16 - 19 August 2004.

5. Bogdanoff, D. W., and Park, C., "Radiative interaction between driver and driven gases in an arc-driven shock tube," *Shock Waves*, Vol. 12, No. 3, 2002, pp. 205-214.
6. Cruden, B. A., "Absolute Radiation Measurements in Earth and Mars Entry Conditions ", Technical Report RTO-EN-AVT-218, NATO.
7. Dannenberg, R. E., "A conical arc driver for high-energy test facilities," *AIAA Journal*, Vol. 10, 1972, pp. 1692-1694.
8. Conrath, B. J., and Gautier, D., "Saturn Helium Abundance: A Reanalysis of Voyager Measurements," *Icarus*, Vol. 144, No. 1, 2000, pp. 124-134.
9. Conrath, B., Gautier, D., Hanel, R., Lindal, G., and Marten, A., "The helium abundance of Uranus from Voyager measurements," *Journal of Geophysical Research: Space Physics*, Vol. 92, No. A13, 1987, pp. 15003–15010.
10. Park, C., "Nonequilibrium Ionization and Radiation in Hydrogen–Helium Mixtures," *Journal of Thermophysics and Heat Transfer*, Vol. 26, No. 2, 2012, pp. 231-243.
11. Cruden, B. A., "Electron Density Measurement in Reentry Shocks for Lunar Return," *Journal of Thermophysics and Heat Transfer*, Vol. 26, No. 2, 2012, pp. 222-230.
12. Park, C., "Nonequilibrium Chemistry and Radiation for Neptune Entry," *Journal of Spacecraft and Rockets*, Vol. 48, No. 6, 2011, pp. 897-903.
13. Matsuyama, S., Ohnishi, N., Sasoh, A., and Sawada, K., "Numerical simulation of Galileo probe entry flowfield with radiation and ablation," *Journal of Thermophysics and Heat Transfer*, Vol. 19, No. 1, 2005, pp. 28-35.
14. Park, C., "Stagnation-region heating environment of the Galileo probe," *Journal of Thermophysics and Heat Transfer*, Vol. 23, No. 3, 2009, pp. 417-424.