INVITED PLENARY TALK

Spring Conference of the Korean Space Science Society (KSSS)
Gangneung-si, Gangwon-do, Republic of Korea

SPACE RADIATION RESEARCH AT NASA

John Norbury

NASA Langley Research Center, Hampton, Virginia, USA

Friday April 29, 2016
OUTLINE

1. 3 SOURCES OF SPACE RADIATION

2. RADIATION & DOSE

3. NUCLEAR & PARTICLE PHYSICS & TRANSPORT

4. MOON, MARS, JUPITER, SATURN

5. CONCLUSIONS
SPE = Solar Particle Event  
GCR = Galactic Cosmic Rays  
Geo = Geomagnetically trapped particles
GCR Composition, Spectrum, Origin

3 regions
- High Energy $< \text{PeV}$
- Very High Energy (knee) $\text{PeV} - \text{EeV}$
- Ultra High Energy (ankle) $> \text{EeV}$

keV $= 10^3 \text{eV}$  MeV $= 10^6 \text{eV}$
GeV $= 10^9 \text{eV}$  TeV $= 10^{12} \text{eV}$
PeV $= 10^{15} \text{eV}$  EeV $= 10^{18} \text{eV}$
ZeV $= 10^{21} \text{eV}$

Large Hadron Collider
14 TeV cm $\Rightarrow$ 400 PeV lab
Space radiation problem

GCR (primary) composition
- 98% nuclei, 2% $e^+ e^-$
- Nuclear component: 
  87% Hydrogen
  12% Helium
  1% heavy nuclei

GCR origin
- Emitted in stellar wind & flares & accelerated by supernova shock waves (within our Galaxy)

GCR High Energy < PeV

Solar System Abundances

GCR Abundances

Relative contribution in fluence, dose and dose equivalent of different elements in the GCR spectrum. Calculation is an average over 1 year in solar minimum behind 5 g/cm² Al shielding.

Solar particles e, p & some heavy nuclei < 1 GeV/N \((v \sim 0.9c)\)

**A. Galactic Cosmic Rays**

The free space GCR environment is made up of heavy and light charged ions originating outside the solar system. This ever-present environment is modulated by the solar wind and, therefore, varies with distance from the sun and to a larger extent, the solar cycle. Maximum GCR intensity is at solar minimum, when the sun is least active while minimum GCR intensity occurs at solar maximum, when the sun is most active. Short duration exposure to GCR provides little health risk, but longer duration exposure may result in late term effects such as cataracts and cancers. The sample calculations described in this document utilize the 1992 Badhwar-O'Neill model which defines a solar maximum GCR environment and a solar minimum GCR environment at 1 AU as shown in Fig. 1. Here the ions are grouped by charge, \(Z\). The GCR environment for a given day is calculated by interpolating between solar maximum and solar minimum. One method for doing this interpolation utilizes the neutron count measured by the Deep River Neutron Monitor (DRNM). The charged ions making up the free-space GCR environment interact with the atoms making up the Earth's atmosphere in two ways. When an atomic interaction occurs, the charged ion strips an electron from an atom and loses energy in the process. When a nuclear interaction occurs, the charged ion collides with or comes very close to an atom's nucleus. Nuclear collisions often result in the destruction of the original ion and the production of a number of smaller ions and neutrons. The neutron count measured on the Earth's surface, in this case at the Deep River station in Canada, is therefore a good predictor of free space GCR intensity, because the number of neutrons produced in the Earth's atmosphere increases when the number of charged ions impinging on this atmosphere increases. Predicted DRNM numbers have been used since the monitor was turned off in 1995. Figure 2 demonstrates the inverse relationship between solar activity and GCR intensity by showing measured and predicted DRNM neutron count numbers on the same plot with measured and predicted sun spot numbers.

**Figure 1.** Freespace GCR environment.

**Figure 2.** Sun spot number (blue) and DRNM neutron count (red), measured (before 1995) and projected (after 1995).

**B. Large Solar Particle Events**

Unlike the GCR environment, solar particle events are isolated events with durations usually measured in hours. Solar particle events occur when a large number of particles, mostly protons, move through the solar system. These events happen during periods of increased solar activity and appear to correspond to large coronal mass ejections. Large SPE have occurred only rarely, one or two per eleven year solar cycle in the past sixty years, but exposure to a large SPE could be lethal if...
Solar particles 11 year cycle - 2016 near maximum!

NASA, 2013

John Norbury (NASA Langley)

Space Radiation Research at NASA

Friday April 29, 2016
**Geomagnetically Trapped Particles**

**Inner belt:** 0 - 3 $R_E = 18,000$ km - mainly p
- Starts about 3,000 km
- But SAA dips down to 400 km

**Outer belt:** 3 - 12 $R_E = 36,000$ km (e,p) - mainly e
(LEO: 200 - 500 km  GEO: 22,000 miles = 35,000 km)

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Wikipedia, 2014  
Unit of absorbed dose:
1 Gray == 1 J/kg

Radiation quality factor $Q$
Sievert = Gray $\times Q$

ICRP estimate: 5% per Gy
1 in 20,000 risk of fatal cancer per 1 mSv dose (lifetime)
| Dose Equivalent (mSv) |  
|-----------------------|-------------------------|--------------------------|--------------------------|
| Chest x-ray           | 0.1                     | USA annual background    | 4                        |
| Public annual limit   | 1                       | International Airline crews | 4                        |
| Radiation worker annual limit | 50                   | No observed effects (Abomb, instant) | 200                    |
| Death (instantaneous dose) | 3,000                 | ISS (with shield) annual | 150                     |
| Astronaut career limit effective dose* | 470                   | Mars (3 year, incl. surface) annual | 1,000                   |
| Large solar flare (free space) | 10,000                | ICRP cancer risk estimate: 5% per Gy ~ 5% per Sv (for Q=1) | 1 in 20,000 risk of fatal cancer per 1mSv dose (lifetime) |
|  |  | * 30 year old female, 1 year mission (50 yr m/f ~ 1,000 mSv) |
An investigation into the nature of high altitude cosmic radiation in the stratosphere.

Figure 7. The Pfotzer curve. Adapted from [3].

\[
0.003388579 \times 17 = 0.05760585 = \text{expected number of lost counts per 17 s (package)}
\]

\[
227 + 0.05760585 = 227.05760585 = \text{expected number of counts if lost counts are included.}
\]

\[
0.05760585 \times 227.05760585 \times 100 = 0.025370588 = \text{percentage of counts lost.}
\]

As only 0.25% of the counts are lost at the maximum count rate, we can safely assume that there was no significant effect on our readings due to an influence by dead time.

Conclusion

We concluded that the peaks in the graphs are to be attributed to the maximum flux of the components of cosmic rays that are of secondary origin. In other words, at an approximate altitude of 20 km, there is a peak amount of ionizing material produced from air shower cascades that are propagating downwards to the surface. This peak can be understood by identifying two competing effects. Firstly, cosmic ray intensity will decrease as altitude decreases, because there have already been collisions higher up. Secondly, atmospheric density will increase as altitude decreases, increasing the likelihood of collisions. The product of these two competing effects produces a maximum in cosmic ray flux at \( \sim 20 \) km. As we did not use multiple coincidence arrays to isolate our cosmic radiation measurements solely to the vertical direction, we are unable to ascertain the impact that ionizing particles interacting with the Geiger–Muller tube from the side had. We can only assume, based on previous research, that the majority of the radiation detected was in the vertical plane. Looking at previous investigations, we discovered an article produced by cosmic ray physicist Pfotzer in the 1930s [3], based on his work with Regener. Their published graph (see figure 7) has remarkable similarities with our own results. It should be noted that the count rate on the Pfotzer curve (see figure 7) was measured in counts per 4 min, and our count rates were measured per every approximate 17 s. The two graphs seem to be supporting each other. Our Geiger–Muller tube recorded omnidirectionally, whereas their apparatus consisted of threefold coincidences which isolated their recordings to the vertical plane. This does not seem to have caused a major difference between our graphed results. This suggests that the majority of cosmic radiation is propagating downwards in a vertical direction.

Figure 8. Payload arrangement. The payload equipment was contained in an insulated polystyrene box, separated with dividers (as pictured) and fastened down with a secondary polystyrene 'roof'.

Figure 9. Balloon burst at 31,685 m.
Figure 7. The Pfotzer curve. Adapted from [3].

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http://www.scienceinschool.org/2010/issue14/cloud/maltese
• Domestic crews 1 - 2 mSv /yr

• International crews < 4 mSv / yr

• Pregnant woman < 5 mSv (to fetus per pregnancy)
AIRCRAFT - WHY ALL THE CONCERN NOW?

- NCRP & ICRP have lowered radiation worker exposure
  - 50 mSv / yr to 20 mSv / yr
- Air crews most highly exposed of any occupation group
- FAA criticized for not paying enough attention
- Many more polar flights
- Future High Speed Civil Transport (HSCT) radiation levels
  - 3 times higher than for crews of subsonic transport
- Only solution available now:
  - reduce flight hours
- NAIRAS - Mertens (Langley)
Computers
- Junction density increasing
- Switching energy decreasing

Need for predicting Single Event Upsets (SEU)
- satellite electronics
- aircraft electronics (civilian & military)

Shuttle - several hundred SEU / mission

http://holbert.faculty.asu.edu/eee560/see.html
Electronics on Spirit, Opportunity, Curiosity etc. are radiation hardened

Shielding very important for Jupiter, Saturn

http://mars.jpl.nasa.gov/msl/multimedia/images/?imageid=3504

Credit: NASA/JPL/Space Science Institute

http://photojournal.jpl.nasa.gov/catalog/PIA04866
These are the doses received

- How were results obtained?
- How to design spacecraft & aircraft shields so dose is minimal?

Need

- Accurate atomic, nuclear, particle physics theory
- Accurate transport theory
- Biological models
Solve Boltzmann transport eqn (HZETRN)

Deterministic, not Monte Carlo

- Want quick answers
- Real time dose as function of position & time
- Both transport & nuclear physics must run fast
- → Applied nuclear physics

Wilson et al., NASA-RP 1257, 1991
The effects on the radiation environment of localized magnetic fields on the Martian surface environment due to GCR have been examined in this paper. The proton spectra for three of the largest historical events are shown in Fig. 3. The spectral shape of the "design basis" SPE depend on mission characteristics, such as duration and time in the solar system. This ever-present environment is modulated by the solar wind and, therefore, varies with distance from the Sun. The free space GCR environment at 1AU is made up of heavy nuclei and light charged ions originating outside the solar atmosphere. Since the lunar atmosphere is negligible, the free space environment is only affected by the lunar regolith. The charged ions lose energy due to ionization, and nuclear collisions occur, producing secondary ions as well as neutrons. These charged ions and neutrons also interact with atoms making up the Martian surface material, including H2O ice, CO2 ice, or some combination. Calculated Mars surface environments for solar maximum and solar minimum, when the sun is most active. Short duration coronal mass ejections can produce SPE within a few days. Long duration events such as large SPEs require months or years, but exposure to a large SPE could be lethal if enough shielding is not provided. Development of a probabilistic model, in which the size and possibly the charge, Z, of the GCR environment for a given day is calculated by interpolating between solar maximum and solar minimum GCR environments at 1 AU as shown in Fig. 1. Here the ions are grouped by charge, Z. The GCR environment for a given day is made up of neutrons produced during nuclear collisions between free space ions and the atoms making up the surface point, and on the atomic make-up of the material below the surface point (regolith, atmosphere above the surface point), and on the atomic make-up of the material below the surface point (regolith, atmosphere above the surface point), and on the atomic make-up of the material below the surface point (regolith, atmosphere above the surface point). The charged ion environment on the surface is, therefore, approximately half that of the free space environment due to GCR. The charged ions making up the free space environment and the atoms making up the atmosphere, primarily CO2, provide the primary particles that produced them, but some lower energy, lighter particles (primarily neutrons) produced in the Earth's atmosphere increases when the GCR intensity, because the number of neutrons collides with or comes very close to an atom's nucleus. Nuclear collisions often result in the destruction of the original ion and the production of a number of smaller, charged ions of various charges. The GC R environment for a given day is approximately half that of the free space environment due to GCR. The proton spectra for three of the largest historical events are shown in Fig. 3. This plot demonstrates the inverse relationship between solar activity and increased solar activity and appear to correspond to large SPE. The proton spectra for three of the largest historical events are shown in Fig. 3. This plot demonstrates the inverse relationship between solar activity and increased solar activity and appear to correspond to large SPE. The proton spectra for three of the largest historical events are shown in Fig. 3. This plot demonstrates the inverse relationship between solar activity and increased solar activity and appear to correspond to large SPE.
Figure 1. Primary galactic cosmic ray spectra for 1977 solar minimum.

Left: Primary GCR spectra at Mars for 1977 solar minimum Kim et al., NASA TP 208724, 1998

Right: Martian surface environment due to GCR Clowdsley et al., AIAA, 2006
Dose Equivalent as a function of depth for various materials

From radiation point of view, safest place is inside liquid hydrogen fuel tank!

Major result
- Low Z materials required for weight reduction necessary for future High Speed Civil Transport (HSCT) and for future spacecraft are also the best radiation protection materials
- Thank goodness!
Short duration

- Solar particle events

Long duration

- Solar particle events
- Galactic cosmic rays
## Lunar regolith composition

<table>
<thead>
<tr>
<th>Material</th>
<th>Mass percentage</th>
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<tbody>
<tr>
<td>SiO$_2$</td>
<td>52.6%</td>
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<tr>
<td>FeO</td>
<td>19.8%</td>
</tr>
<tr>
<td>Al$_2$O$_3$</td>
<td>17.6%</td>
</tr>
<tr>
<td>MgO</td>
<td>10.0%</td>
</tr>
</tbody>
</table>
The Surface of the Moon is Slightly Richer in Fe, Ca, and Mg Compared to the Earth's Crust

http://fti.neep.wisc.edu/neep602/lecture12.html
Figure 12. Calculated fluence of projectile fragments after traversal of 18 g/cm² thick polyetherimide shield irradiated with 33.88 GeV 56 Fe ions.

Figure 13. Attenuation of dose equivalent due to 1977 solar minimum GCR fluence behind regolith and regolith-epoxy shield as a function of areal density [27].

Simonsen et al., NASA Conference Publication 3360, 1997
### Mars

<table>
<thead>
<tr>
<th></th>
<th>Earth</th>
<th>Mars</th>
</tr>
</thead>
<tbody>
<tr>
<td>Atmospheric thickness (g/cm²)</td>
<td>1000</td>
<td>20</td>
</tr>
<tr>
<td>Magnetic field (Gauss)</td>
<td>1</td>
<td>0</td>
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</tbody>
</table>

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http://www-k12.atmos.washington.edu/k12/resources/
Mars

### Chemical composition of Martian atmosphere

<table>
<thead>
<tr>
<th>Component</th>
<th>Percentage (%)</th>
</tr>
</thead>
<tbody>
<tr>
<td>CO₂</td>
<td>95.32</td>
</tr>
<tr>
<td>N₂</td>
<td>02.70</td>
</tr>
<tr>
<td>Ar</td>
<td>01.60</td>
</tr>
<tr>
<td>O₂</td>
<td>00.13</td>
</tr>
<tr>
<td>CO</td>
<td>00.08</td>
</tr>
</tbody>
</table>

De Angelis et al., Rad. Meas. vol.41, p.1097, 2006
### Mars

#### Chemical composition of Martian surface

<table>
<thead>
<tr>
<th>Component</th>
<th>Percentage (%)</th>
</tr>
</thead>
<tbody>
<tr>
<td>SiO₂</td>
<td>44.2</td>
</tr>
<tr>
<td>Fe₂O₃</td>
<td>16.8</td>
</tr>
<tr>
<td>Al₂O₃</td>
<td>08.8</td>
</tr>
<tr>
<td>CaO</td>
<td>06.6</td>
</tr>
<tr>
<td>MgO</td>
<td>06.2</td>
</tr>
<tr>
<td>SO₃</td>
<td>05.5</td>
</tr>
<tr>
<td>Na₂O</td>
<td>02.5</td>
</tr>
<tr>
<td>TiO₂</td>
<td>01.0</td>
</tr>
</tbody>
</table>

De Angelis et al., Rad. Meas. vol.41, p.1097, 2006
GCR Environment

Model prediction of dose equivalent from GCR. Calculations are shown at average skin depth near solar maximum. Cucinotta, Rad. Res. vol.43, p.S35, 2002
Mars transit inside vehicle: $1.84 \pm 0.33 \text{ mSv/day}$

⇒ MSL (one way) 253 days gives 466 mSv

⇒ 331 mSv for 180 day cruise DRM

⇒ 662 mSv return trip

• Plus surface exposure 200 mSv?

• Approaching and exceeding limits
Most robotic missions to Mars have failed, but those few that have succeeded, such as NASA's Mars Exploration Rover Spirit, which landed on the red planet in January 2004, have encountered a menace from space: radiation. The measurements—from the Radiation Assessment Detector (RAD) piggybacking the Mars Science Laboratory (MSL) spacecraft that carried Curiosity to Mars provided much the clearest look yet at the nature of the radiation that near-Earth spacecraft will encounter as they voyage beyond Earth's protective magnetic field.

Astronauts beyond low Earth orbit will encounter energetic charged particles including those from solar energetic particles (SEP) events. These can occur in cycles, with the most recent cycle peaking in late 2012. The protons' energies tend to range up to a few hundred megaelectron volts. Astronauts could well spend a year or more on the surface of Mars, where the radiation risk is expected to be an order of magnitude higher than on the International Space Station (ISS). Solar cycle-driven SEP events can cause a spike in radiation, which can range from a factor of 2 to 5 times from December 2011 into July 2012, and again in mid-2013.

Comparative Radiation Exposures

<table>
<thead>
<tr>
<th>Exposure Source</th>
<th>Radiation Dose (in Millisieverts)</th>
</tr>
</thead>
<tbody>
<tr>
<td>U.S. Annual Average All Sources</td>
<td>3.6</td>
</tr>
<tr>
<td>Abdominal CT Scan</td>
<td>8</td>
</tr>
<tr>
<td>DOE Radiation Worker Annual Limit</td>
<td>20</td>
</tr>
<tr>
<td>6 Months on ISS (Average)</td>
<td>75</td>
</tr>
<tr>
<td>MSL-RAD Transit to Mars</td>
<td>466</td>
</tr>
</tbody>
</table>

MSL-RAD= Mars Science Laboratory Radiation Assessment Detector


Frankly, it's a modest portion of the risks on a Mars mission.
Intense radiation

- On Jupiter moon Io humans could not survive for more than a few hours
- Callisto
In addition to applications to Mars science, the models will allow the use of these models to assess the impact on human exposure and allow design considerations for all components in either surface operations, in lava tubes, or in Mars orbit.

3.3. Jupiter neighborhood

The solar wind generated convective currents produce less effect on the GCR near Jupiter relative to Earth. In addition, the jovian magnetic field traps particles with electron intensities to large distances from Jupiter (70 R_J). The electron flux spectrum near Callisto is shown in Fig. 9. Near a jovian moon, the intensities of all components are reduced except the induced neutron fields.

We are developing an Anytime/Anywhere software package for use in mission analysis and an example of a mission from the Earth/Lunar L1 to Callisto and return is shown in Fig. 10 for a fixed mass of four shield materials. The mission starts near an assumed solar maximum (2045 AD) using the same projection model as in Fig. 1. The large exposure rates mid-mission are due to the jovian electron belts until arrival on the Callisto surface where there is shielding below the horizon. The continued increase of the environment on the return to L1 is due to decreasing solar activity. This software will soon include the Mars models presented in the previous section.

4. Conclusions

The deep space environments developed for mission analysis use physical models to extrapolate the limited environmental data in both space and time. Induced fields are evaluated using high-speed transport models. Interesting dependence of local induced fields on atmosphere and ground composition are found which may be validated using Mars orbital data. The broader software being prepared will allow radiation exposure evaluation for arbitrary spectra in the future but limited to near Earth, near Jupiter, and interplanetary space for now. The near Mars environment will be added in the near future.

References


Clowdsley, M.S., Heinbockel, J.H., Kaneko, H., Wilson, J.W., Singleterry, R.C., Shinn, J.L. A comparison of the multigroup and...
Saturn

Saturn’s Intense Main Radiation Belt

http://www.universetoday.com/15381/radiation-on-saturn/

**Health Effects on Human Body**

- **Carcinogenesis**
  - Earlier appearance and more aggressive tumors not seen with controls, gamma-rays or proton induced tumors
  - Persistent oxidative damage and inflammatory pathway responses
  - New genomics data showing distinct gene expression profiles in HZE versus $\gamma$-ray or x-ray irradiated cell models

- **Acute Radiation Syndrome due to Solar Particle Events**
  - Research addresses dose threshold, dose-rate effects with countermeasure evaluation
  - Future work to understand impact of possible high skin dose and microgravity on immune system and blood forming organs
Health effects on Human Body

- Acute or Late Central Nervous System (CNS) Effects
  - Concern for CNS risks originated with the prediction of the light flash phenomenon from single high-Z high-energy nuclei traversals of the retina; this phenomenon was confirmed by the Apollo astronauts
  - Major uncertainty how to extrapolate results from animals to humans

- Degenerative Tissue or other health effects
  - Occupational radiation exposure from the space environment may result in degenerative tissue diseases (non-cancer or non-CNS) such as cardiac, circulatory, or digestive diseases, & cataracts
  - Mechanisms & magnitude of influence of radiation leading to these diseases are not well characterized
Brookhaven National Lab on Long Island
- protons: 4 GeV
- protons - Fe: 50 MeV/n - 1.5 GeV/n
- up to Au: 165 MeV/n

Pions can contribute almost 50% to dose.
NEW DEVELOPMENTS - NEUTRONS & LIGHT IONS

Percent contribution to BFO dose equivalent by charge group


Neutrons and light ions (H, He) can dominate Dose Equivalent
Two approaches to Galactic Cosmic Ray (GCR) simulation:

**External field approach**

Beams selected to represent external, free space field before shielding.

**Local tissue field approach**

Beams selected to directly represent shielded tissue field.
**NEW DEVELOPMENTS - MINIMUM DOSE EQUIVALENT VS. DEPTH**

- **Geometry:** Slab with equal thickness of aluminum shielding in front and behind a thin (0.03 mm) water target
- **Boundary conditions:** 1977 solar minimum GCR environment
  - Considered each ion individually with greater emphasis on more abundant ions
  - Connects with ions considered in measurement piece
- **Computed flux, dose, dose equivalent, and LET spectra in all codes**
  - Carefully went through Monte Carlo options to ensure output quantities are defined consistently

**External GCR boundary condition spectrum**

- Front shield
- Back shield
- Infinite lateral dimensions

\[ X \text{ g/cm}^2 \text{ Al} \equiv X \text{ g/cm}^2 \text{ Al front} + 0.03 \text{ mm water} + X \text{ g/cm}^2 \text{ Al back} \]

Recent High Impact Discoveries

- **Minimum in dose equivalent vs. depth**
  - Near 40 g/cm$^2$
  - Increased shielding increases exposure
  - Material optimization more important than previously thought

- **If forward/backward (FB) neutron transport and pions turned ON**
  - Local minimum in dose equivalent response near 40 g/cm$^2$
  - Increased shielding (mass) changes/amplifies exposure
  - Design implication: material optimization may be more important than previously thought

**Contribution to dose rate for particle groups at different locations on the ISS**

<table>
<thead>
<tr>
<th>MDU</th>
<th>Dose Rate, µGy/day</th>
</tr>
</thead>
<tbody>
<tr>
<td></td>
<td>Pions + Muons + EM</td>
</tr>
<tr>
<td></td>
<td>Ions + Neutrons</td>
</tr>
<tr>
<td>MDU 1</td>
<td>41%</td>
</tr>
<tr>
<td>MDU 2</td>
<td>49%</td>
</tr>
<tr>
<td>MDU 3</td>
<td>47%</td>
</tr>
<tr>
<td>MDU 4</td>
<td>26%</td>
</tr>
<tr>
<td>TEPC</td>
<td>48%</td>
</tr>
<tr>
<td></td>
<td>53%</td>
</tr>
<tr>
<td></td>
<td>74%</td>
</tr>
<tr>
<td></td>
<td>52%</td>
</tr>
<tr>
<td></td>
<td>51%</td>
</tr>
</tbody>
</table>

**MDU=Mobile Dosimetry Unit (ISS)**

**Important discovery in space radiation**

Blattning, Slaba, Bahadori, Norman, Clowdsley, Space Radiation Investigators’ Workshop, Galveston, TX, 2014

Walker, Townsend, Norbury, Advances in Space Research 51, 1792, 2013

Pions make large contributions to dose rate for particle groups at different locations on the ISS.
KOREA CONTRIBUTIONS

- Human presence throughout solar system in 21 century
  - International endeavour
  - New frontier - many economic rewards - reaped by participants

- Korea Institute of Radiological and Medical Sciences (KIRAMS)
  - Expertise in proton and heavy ion therapy and radiological sciences
  - Highly relevant to space radiation

- Space Radiation
  - Major uncertainties associated with low dose rate
  - KIRAMS contribution?
CONCLUSIONS

- Human exploration of solar system
- Radiation protection is a major issue
- Fundamental studies in physics and radiobiology still needed
- Republic of Korea could make major contributions
THE END

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