NIRCam Intro (from STScI Jdox)

Introduction

The JWST Near Infrared Camera (NIRCam) observes from 0.6 to 5.0 \( \mu \text{m} \) and offers imaging, coronagraphy, and grism slitless spectroscopy. NIRCam has 2 modules pointing to adjacent fields of view. Each module uses a dichroic to observe simultaneously in a short-wavelength channel (0.6–2.3 \( \mu \text{m} \)) and a long-wavelength channel (2.4–5.0 \( \mu \text{m} \)).

NIRCam has 5 observing modes for science:

- Imaging of two 2.2' \( \times \) 2.2' fields separated by 44'' covering 9.7 arcmin² in total
- Coronagraphic imaging at multiple wavelengths
- Wide field slitless spectroscopy (2.4–5.0 \( \mu \text{m} \)) using grisms with resolving power \( R = \lambda/\Delta\lambda \sim 1500 \)
- Time series imaging (photometric monitoring)
- Grism time series (spectroscopic monitoring)

NIRCam will also obtain wavefront sensing measurements used to align and phase JWST's primary mirror.

Focus of this talk

From https://jwst-docs.stsci.edu/display/JTI/
NIRCam Fields of View (from STScI Jdox)

Module A
- coronagraph masks
- overlapping FOVs obtained simultaneously
- short wavelength detectors
- long wavelength detectors

Module B
- Time series
- imaging
- when projected on detectors

2.2'
5.1'
44''
4.5''
42''
20''
129''
64''
NIRCam modes: selectable with wheels

No Short Wavelength Spectroscopic Capabilities in Cycle 1

From OTE

0.6 – 2.4 μm
Short Wavelength Channel

2.4 – 5 μm
Long Wavelength Channel

Simultaneous SW imaging is possible!

2 LW grisms in each module provide R~1500 slitless spectroscopy: Chose dispersion orientation and filters to suit your science

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NIRCam LW Grism Spectra

Left: NIRCam spectral image of the OSIM super-continuum lamp point source taken with the LWA R grism and F444W filter during JWST instrument testing.

Right: Extracted spectrum. The continuum decreases toward longer wavelengths due to low fiber transmittance, and the broad feature near 4.27 μm is due to CO$_2$ absorption. These are artifacts of the test equipment and not NIRCam itself.

* NIRCam FOV is 2.’2 x 2.’2 with dispersion of 10 Å per 0.”065 x 0.”065 pixel
NOTE: Total spectroscopic throughput is the **product** of Grism curve and selected filter!

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Figure 3. Left: Total system throughput including all OTE and NIRCam optics and the detector quantum efficiency for several NIRCam filters. The theoretical LW grism efficiency curve (shown for the A module) must be multiplied by the filter curves to produce the system throughput at each wavelength. The Module B LW grisms are anti-reflection coated on only 1 side and therefore have throughputs approximately 25% lower than the LWA grisms. Right: Grism FWHM spectral resolving power vs. wavelength for point sources, limited by pixel sampling of the PSF at shorter wavelengths ($\lambda \lesssim 4 \mu m$) and limited by the circular beam factor and diffraction at longer wavelengths ($\lambda \gtrsim 4 \mu m$).
Module A (TSO) Spectral Saturation Values

<table>
<thead>
<tr>
<th>$\lambda$ ((\mu\text{m}))</th>
<th>$K_{\text{sat}}$ (A0V)$^c$</th>
<th>$K_{\text{sat}}$ (M2V)$^c$</th>
<th>Filter$^d$</th>
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<tr>
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<td>4.3</td>
<td>4.2</td>
<td>F322W2</td>
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</table>

$^c$: K-band Vega magnitudes for saturation (80% full well or 65,000 electrons) for 0.68 s integrations (2 reads) of 2048 x 64 pixel NIRCam can observe bright stars!

See Greene+ (2017) JATIS article for more Module A & B saturations and sensitivity values
Time-series imaging is also possible

- $\lambda < 2.4 \ \mu m$ TSO imaging can be done simultaneously with either $\lambda > 2.4 \ \mu m$ imaging or spectroscopy
- SW observations can be done with weak lenses for better bright limits and potentially higher precision photometry
- Show HAT-P-18 b APT example???
Setting TSO parameters

- Determine how much dwell time for each object
- Set subarrays and exposure parameters
- Set SW filter: simultaneous $\lambda < 2.4 \, \mu m$ imaging
- Consider target acquisition
  - Offset acquisition required for bright targets in Cycle 1
- Visibility, position angles, and spectral overlaps
- Enter values into APT
NIRCam grism time series options (APT)

- Can choose from 64, 128, 256, & 2048 x 2048 subarrays
- 1 or 4 outputs (4 for very bright stars)
- Simultaneous short wavelength imaging with weak lens to spread the light over many pixels is possible
- No dithering
- Flexible detector MULTIACCUM exposure & readout parameters
Select Subarray Size

mag > bright limit + 0.75?

Y → Want > 128 subarray?

N → 2048 x 64

N Outputs

Y → Is mag > bright limit + 1.5?

N → 2048 x 256

N Outputs

Y → 2048 x 128

1 or 4 Outputs

Y → 2048 x 256

1 or 4 Outputs
Select Detector Readout Parameters

- RAPID exceeds data limit?
  - Y: BRIGHT1 > limits?
    - Y: BRIGHT 2 > limits?
      - Y: Set # groups from:
        - host star brightness
        - mode saturation limit
        - subarray size
        - # of outputs
      - N: BRIGHT2
    - N: BRIGHT1
  - N: RAPID

- RAPID Ngroups > limit?
  - Y: RAPID
  - N: RAPID

- Set # Ints to fill dwell time
Set SW Filter: Simultaneous $\lambda < 2.4 \mu m$ Imaging

Currently Available SW Filters:
- CLEAR + WLP4
- WLP8 + 182M
- WLP8 + 210M
- WLP8 + 187N
- WLP8 + 212N
Target Acquisition Note

• In Cycle 1, grism time series target acquisition is done with F335M filter, 32 x 32 subarray, and Ngroups ≥ 3
  – Saturation limit is K = 7.0 mag
• Stars with K < 7.0 may require offset target acquisition
  – Offset from nearby fainter star with known coordinates
• Using a narrow-band acquisition filter would allow acquiring on K < ~4.5 mag stars (likely Cycle 2 and later)
Check spectral overlap of nearby objects

We are working on an automated tool for this (NIRCam + MIRI LRS)

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NIRCam TSOs
NIRCam uses the JWST time-series data pipeline

- Users can download & re-run the pipeline with different options, additions, or removals
Future Possible Simultaneous 1 – 2 μm Spectra

- Dispersed Hartmann Sensor (DHS) elements in the SW channel of NIRCam provide 1 – 2 μm spectra using 10 sub-apertures of the JWST pupil, potentially allowing simultaneous spectra of bright stars during LW grism observations.

- This is not an approved science mode for Cycle 1; it may be approved for later cycles. There may be limitations on spectra.

See Schlawin+ (2017) PASP
The End